

The Type Ia Supernova Hubble Diagram

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The original Hubble Diagram (Galaxies)

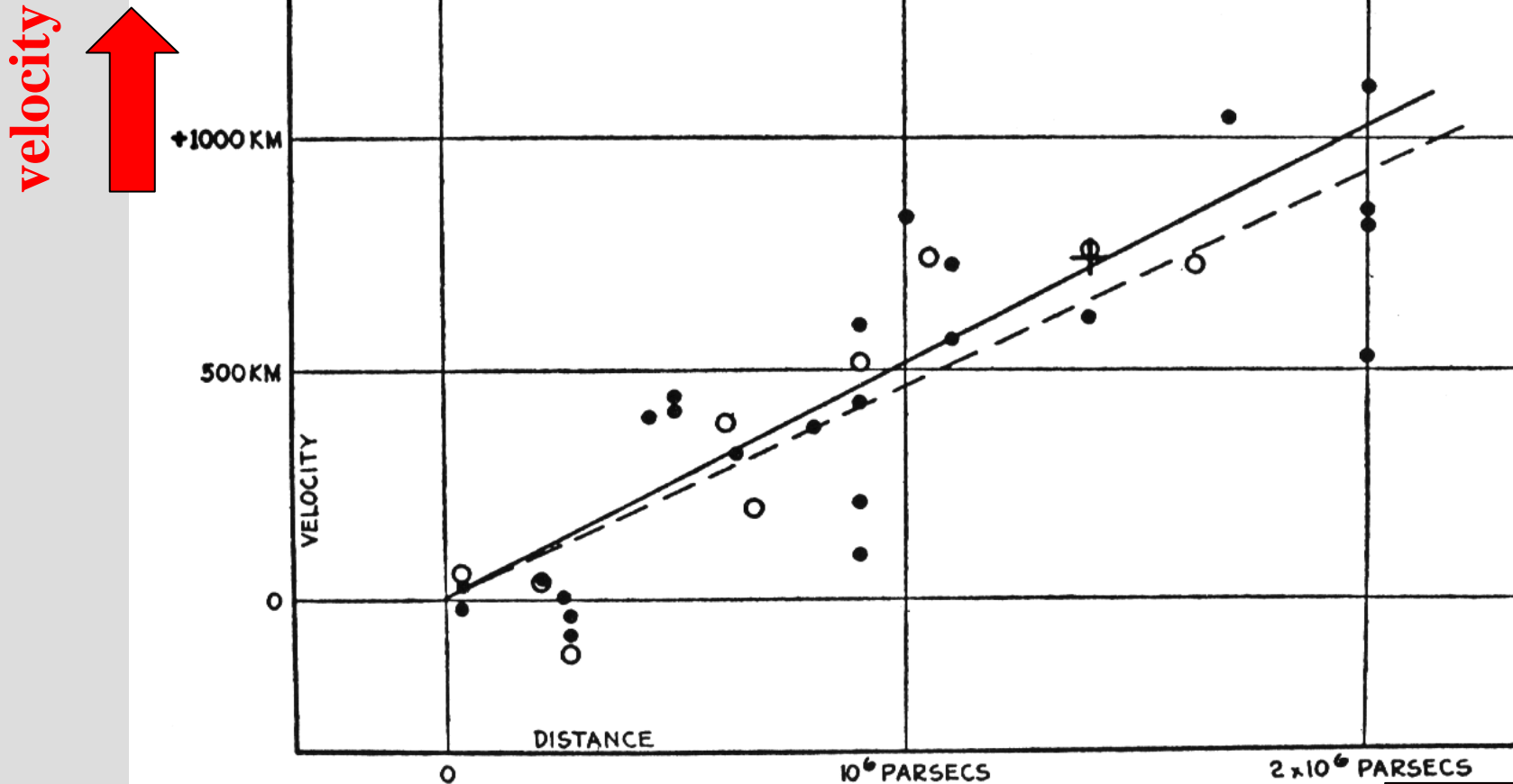
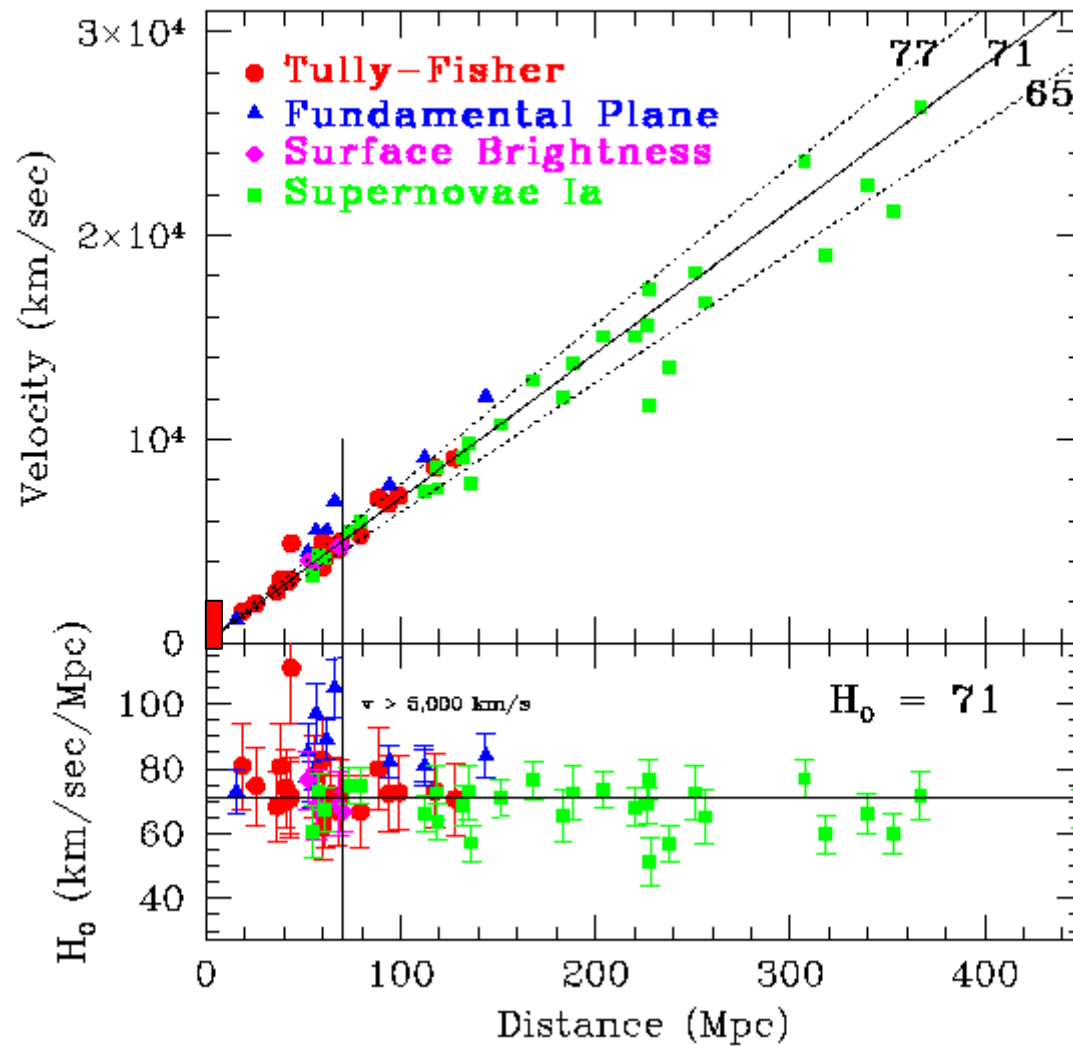


FIG. 9. *The Formulation of the Velocity-Distance Relation.*

Distance

A modern Hubble Diagram



The expansion of the universe

Luminosity distance in an isotropic, homogeneous universe as a Taylor expansion

$$D_L = \frac{cz}{H_0} \left\{ 1 + \frac{1}{2}(1 - q_0)z - \frac{1}{6} \left[1 - q_0 - 3q_0^2 + j_0 \pm \frac{c^2}{H_0^2 R^2} \right] z^2 + O(z^3) \right\}$$

$$H_0 = \frac{\dot{a}}{a} \quad q_0 = -\frac{\ddot{a}}{a} H_0^{-2} \quad j_0 = \frac{\ddot{\dot{a}}}{a} H_0^{-3}$$

Supernovae

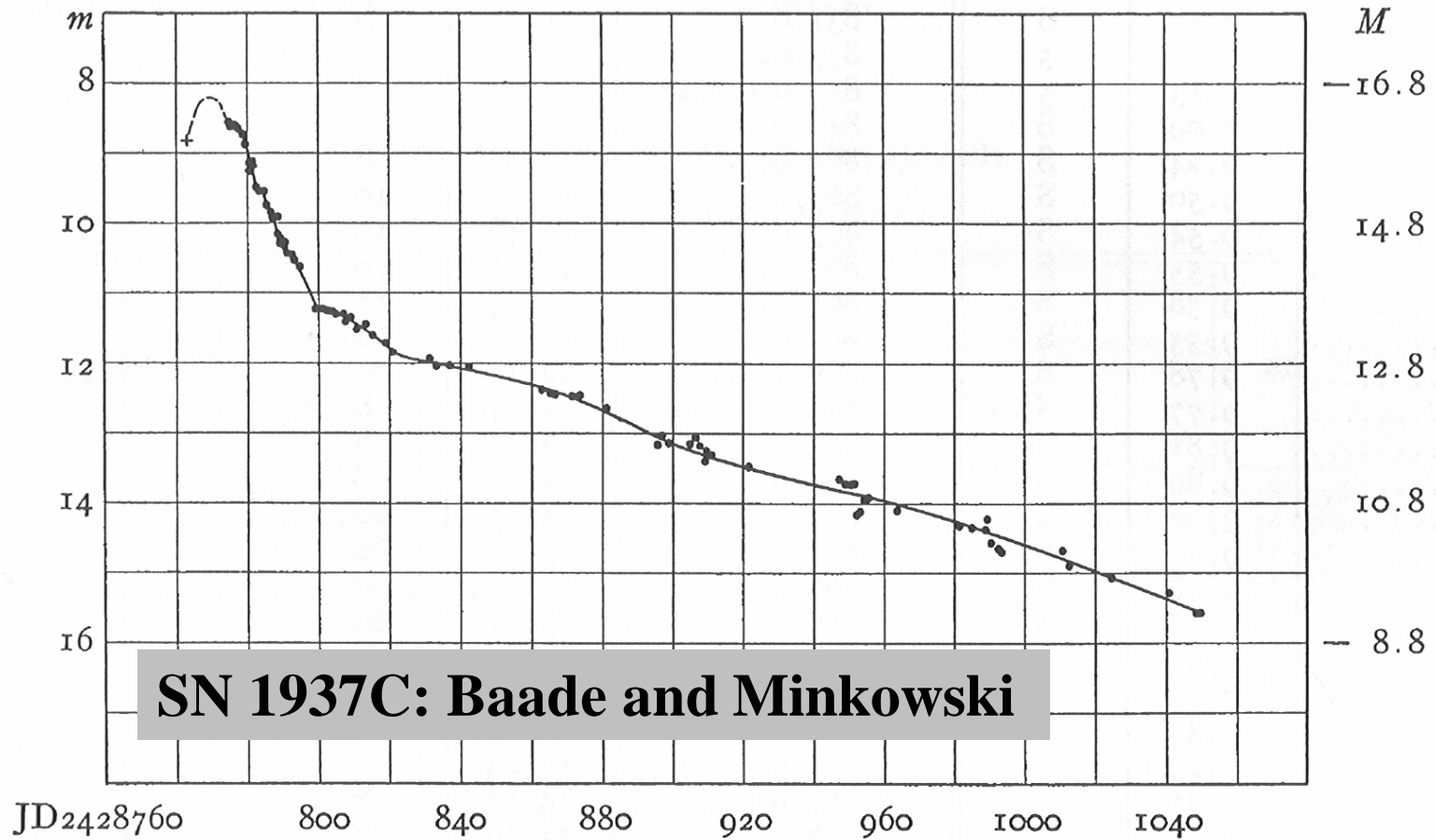


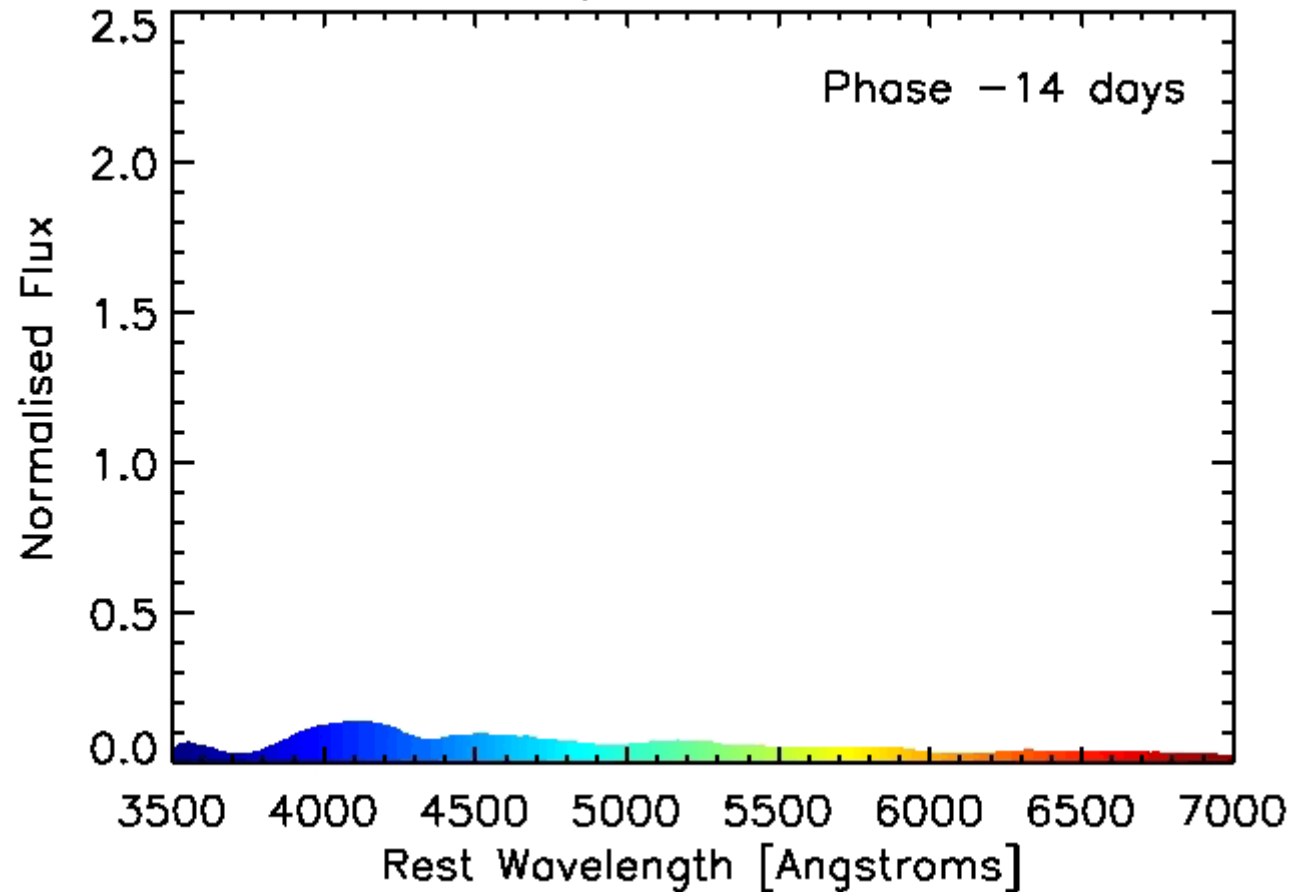
FIG. 1.—Photographic light-curve of supernova in IC 4182

The near

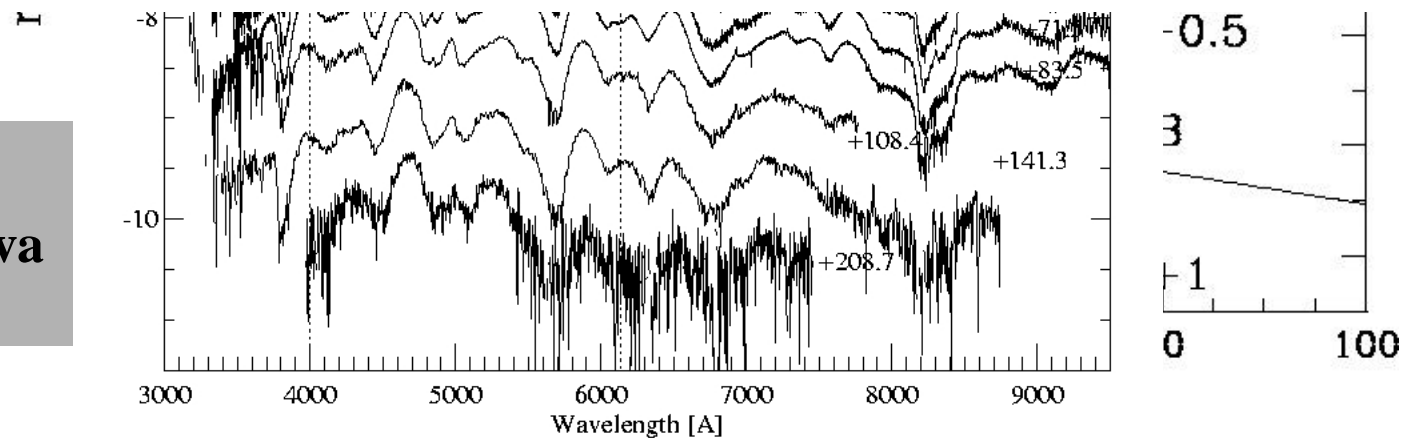
excellent cov
few objects

- fairly comp
- allows deta
comparison

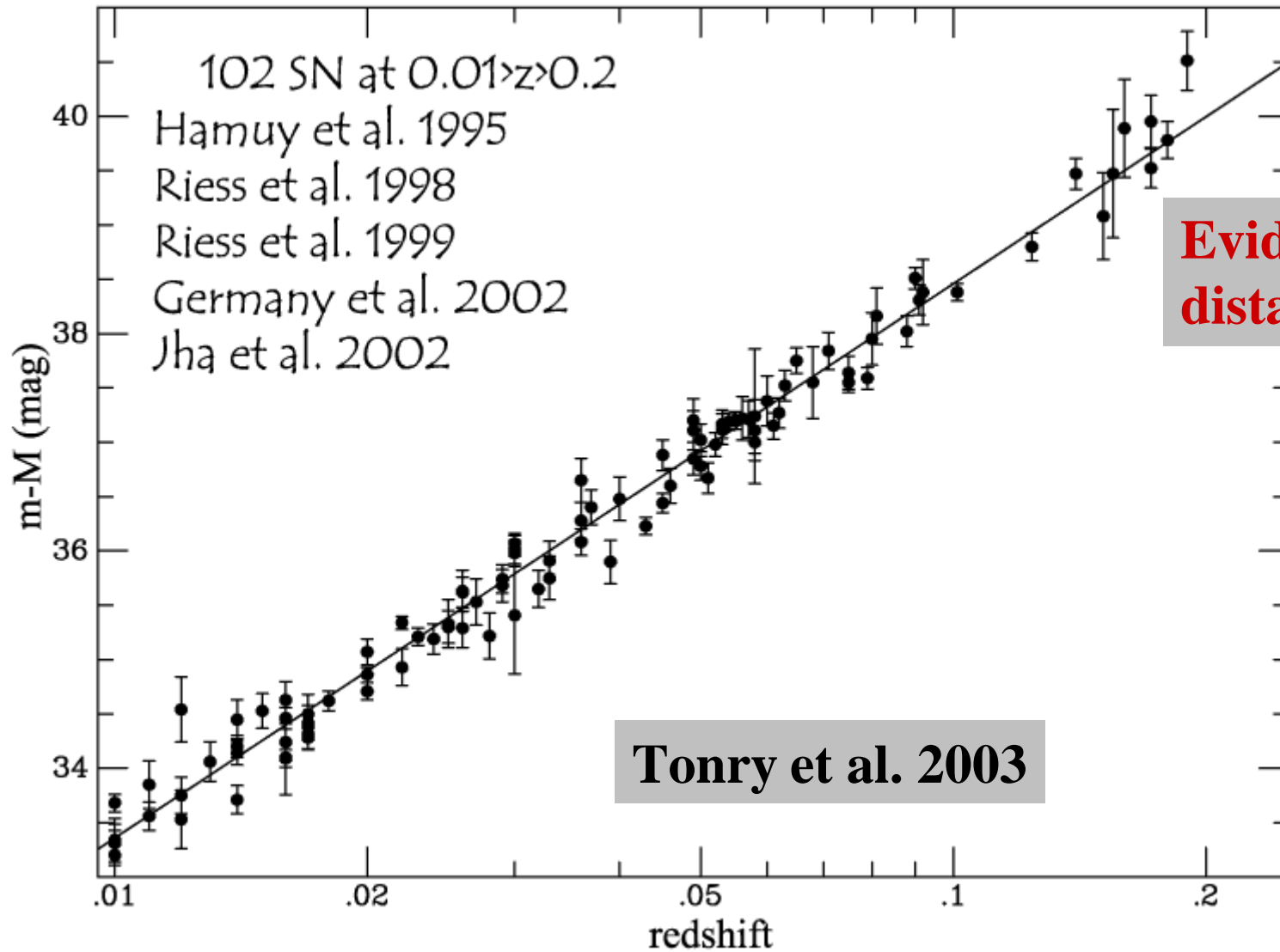
SN Ia spectral evolution



SN 2003du
European Supernova
Collaboration

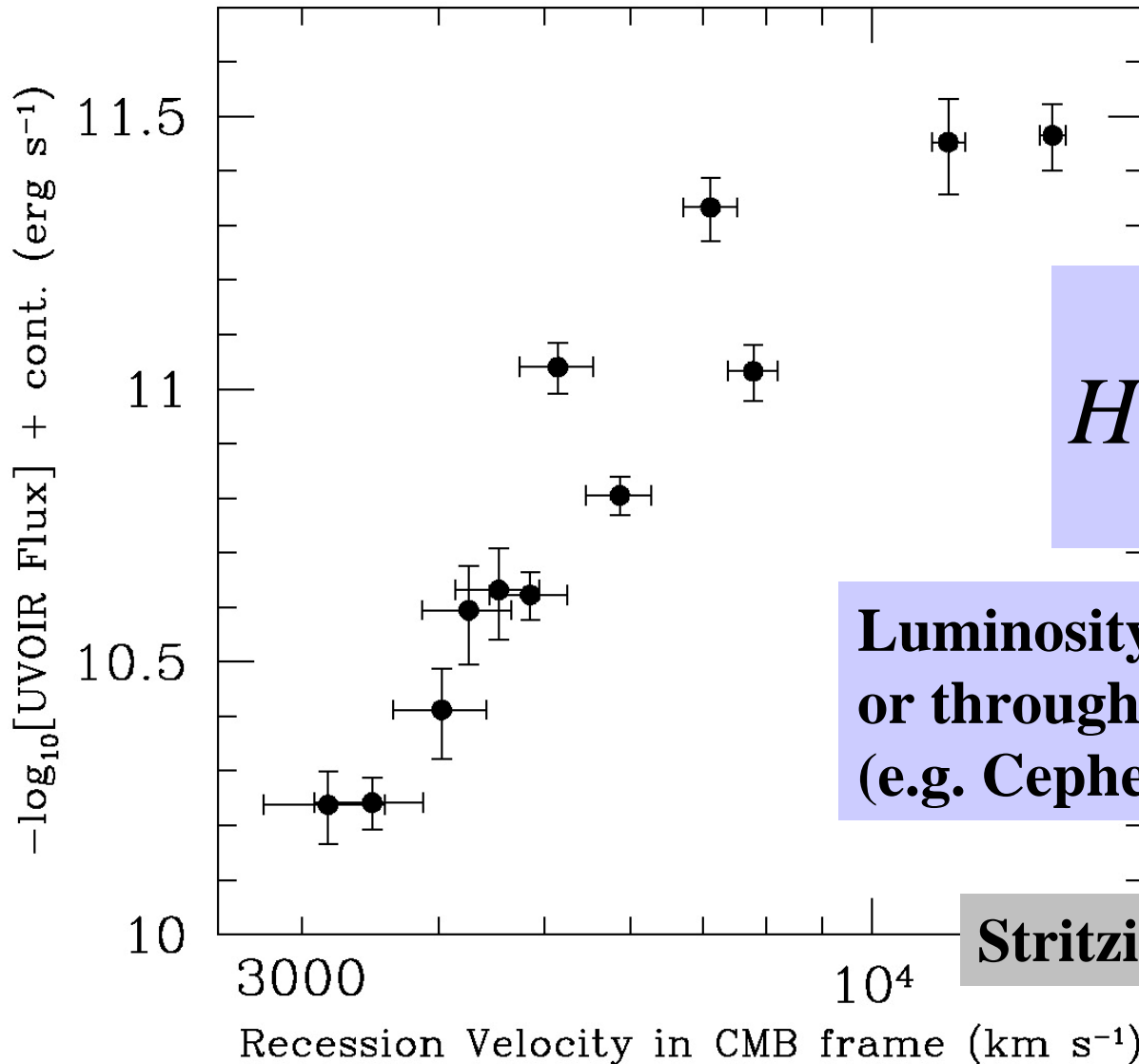


The nearby SN Ia sample



Evidence for good distances

The nearby SN Ia Hubble diagram



$$H_0 = cz \sqrt{\frac{4pF}{L}}$$

Luminosity L either from models
or through the distance ladder
(e.g. Cepheids)

Stritzinger & Leibundgut 2005

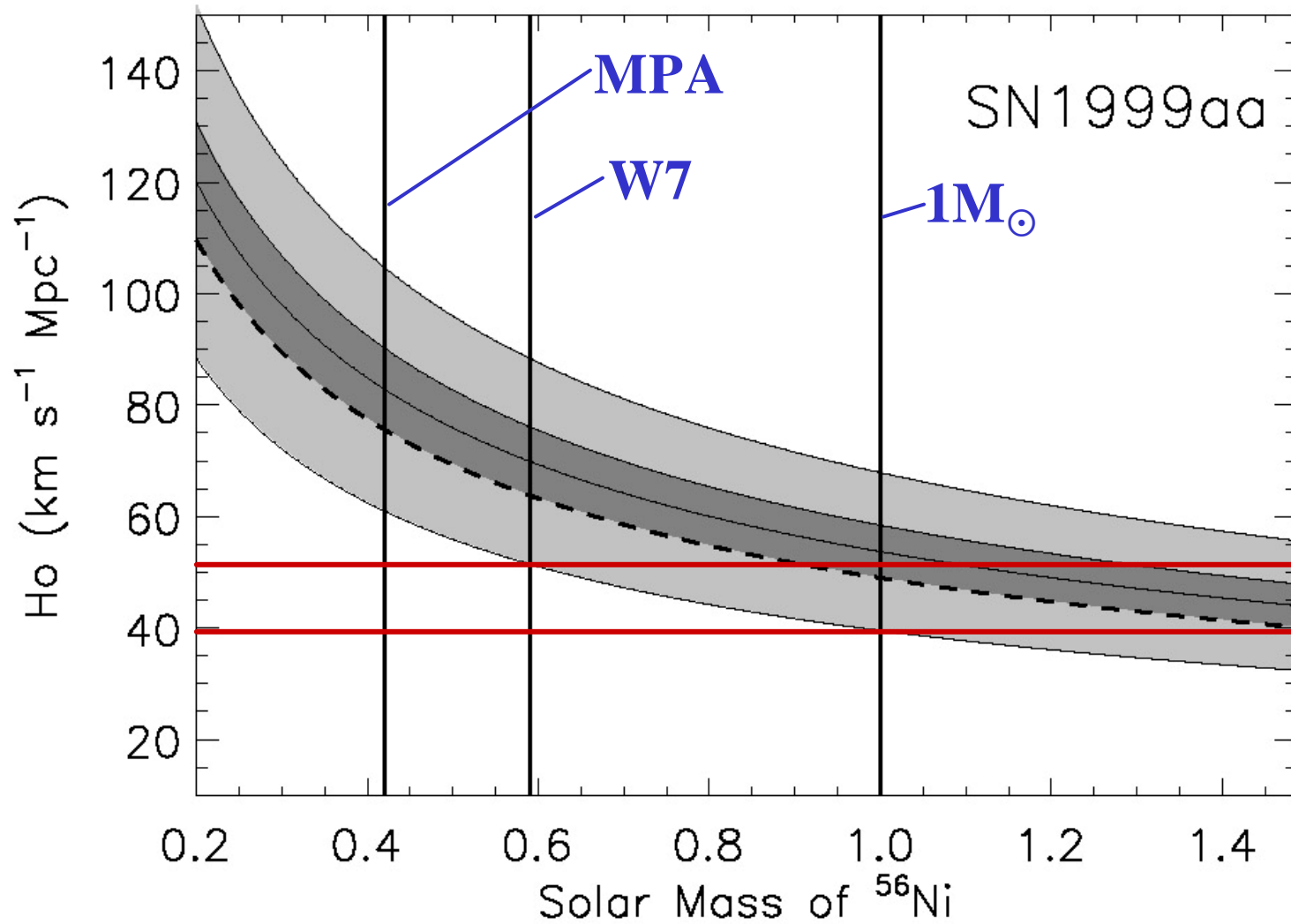
H_0 from the nickel mass

$$H_0 = \frac{cz}{D} = cz \sqrt{\frac{4pF}{L}} = cz \sqrt{\frac{4pF}{E_{Ni}}} \propto cz \sqrt{\frac{4pF}{M_{Ni}}}$$

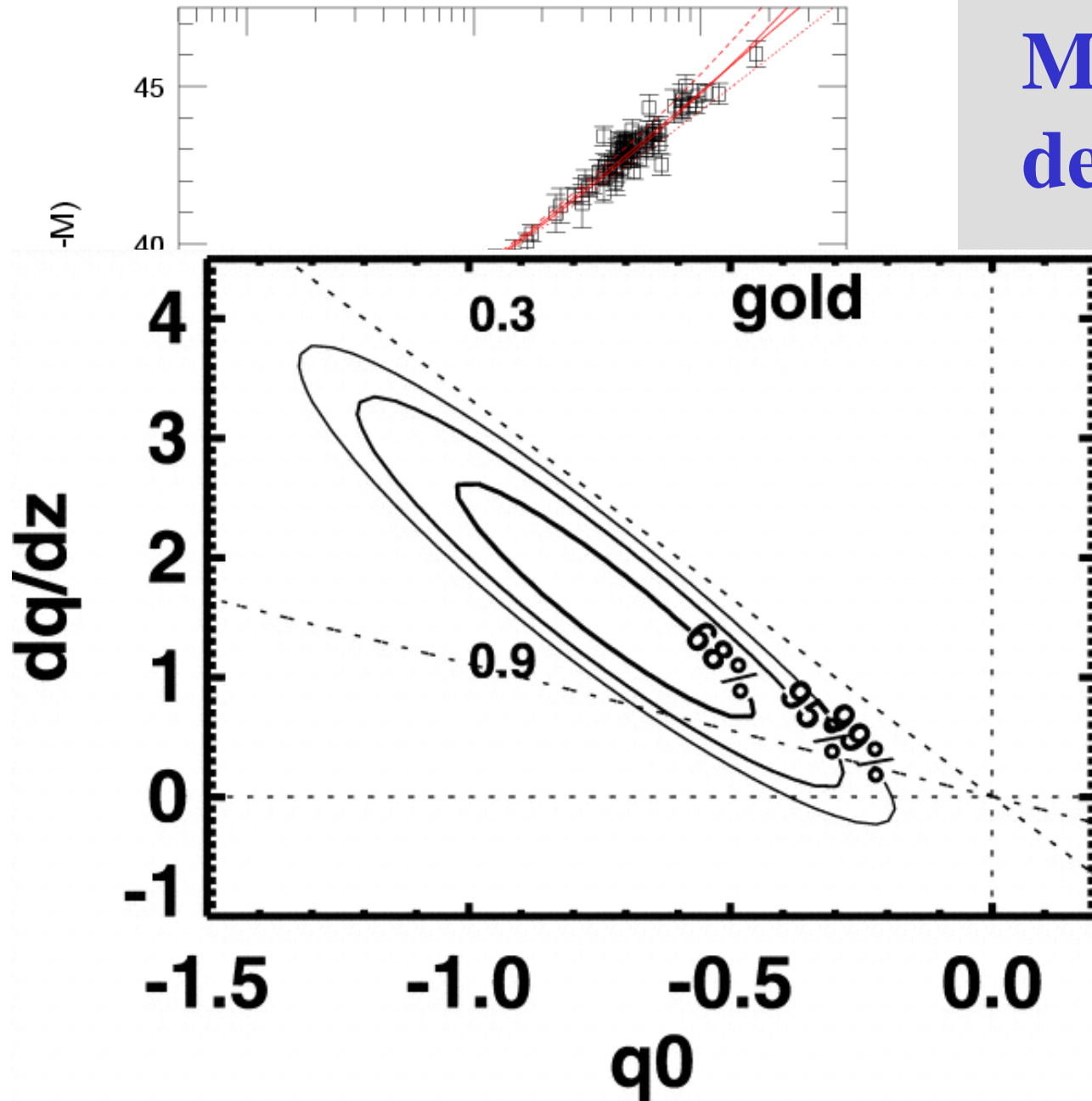
Hubble luminosity distance **Arnett's rule** **Ni-Co decay and rise time**

Need bolometric flux at maximum F and the redshift z as observables

Comparison with models

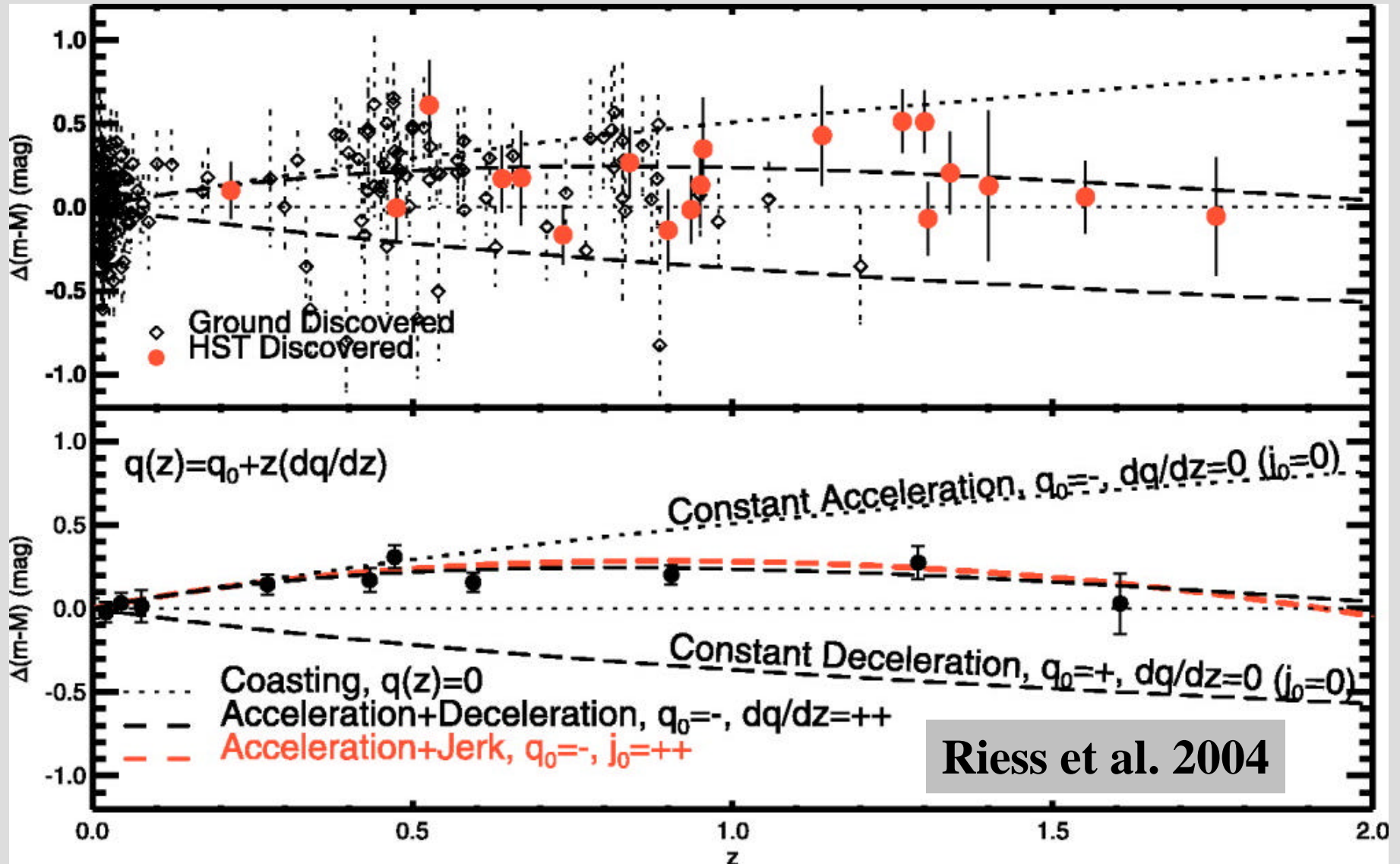


Measure deceleration



Riess et al. 2004

Adding jerk ...





WALL TO WALL
ALBERTS
FOR WORLD
CUP
LHO MAM 2003

$R_{\mu\nu} - \frac{1}{2} g_{\mu\nu} R = -\frac{8\pi G}{c^4} T_{\mu\nu}$
A. EINSTEIN

EP 115 HLBA

INDIAN
E 1
02
VOLP

Friedmann cosmology

Assumption:
homogeneous and isotropic universe

Null geodesic in a Friedmann-Robertson-Walker metric:

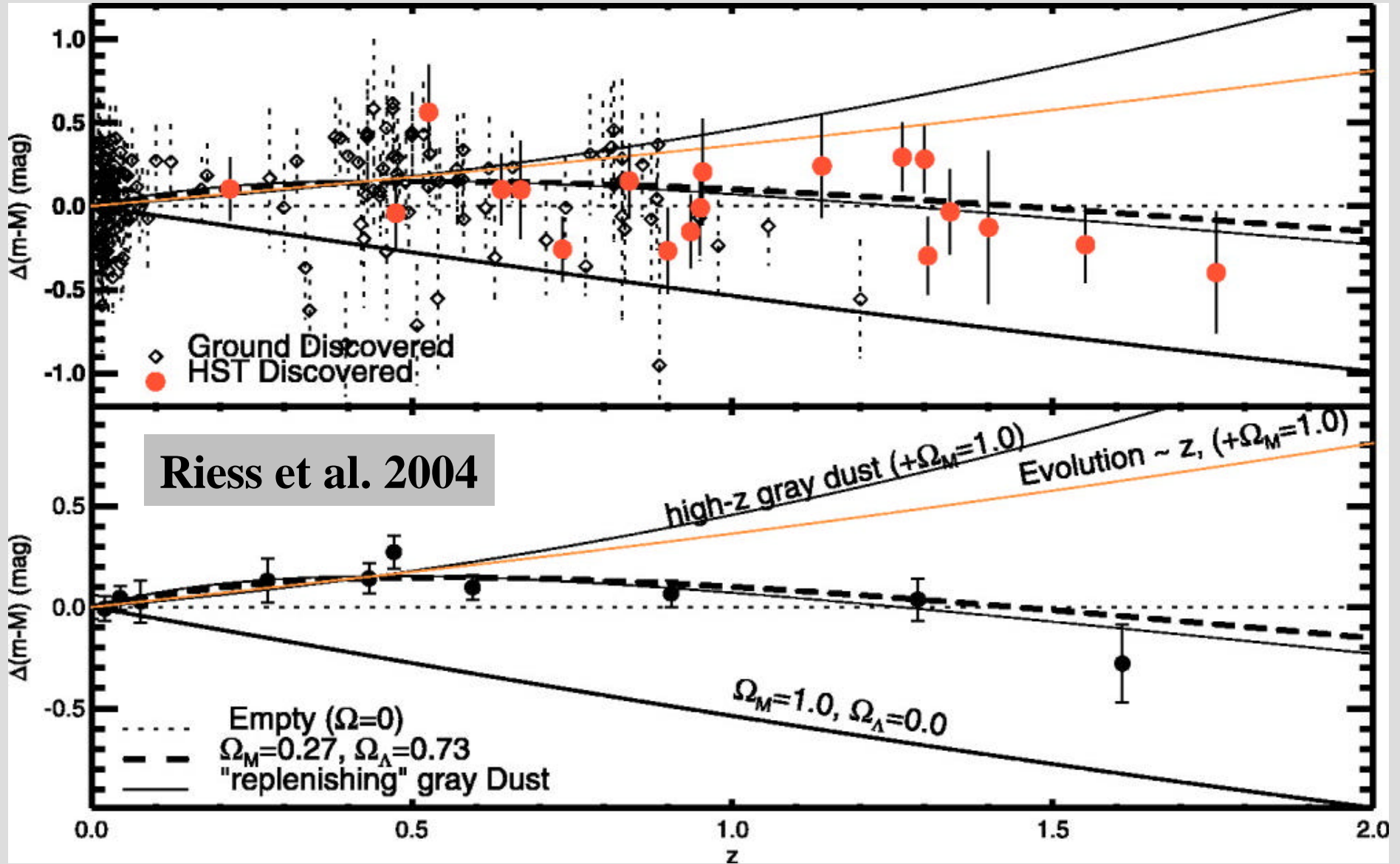
$$D_L = \frac{(1+z)c}{H_0 \sqrt{|\Omega_k|}} S \left\{ \sqrt{|\Omega_k|} \int_0^z \left[\Omega_k (1+z')^2 + \Omega_M (1+z')^3 + \Omega_\Lambda \right]^{-1/2} dz' \right\}$$

$$\Omega_M = \frac{8pG}{3H_0^2} r_M$$

$$\Omega_k = -\frac{kc^2}{R^2 H_0^2}$$

$$\Omega_\Lambda = \frac{\Lambda c^2}{3H_0^2}$$

Evidence for the O_2



The equation of state parameter w

General luminosity distance

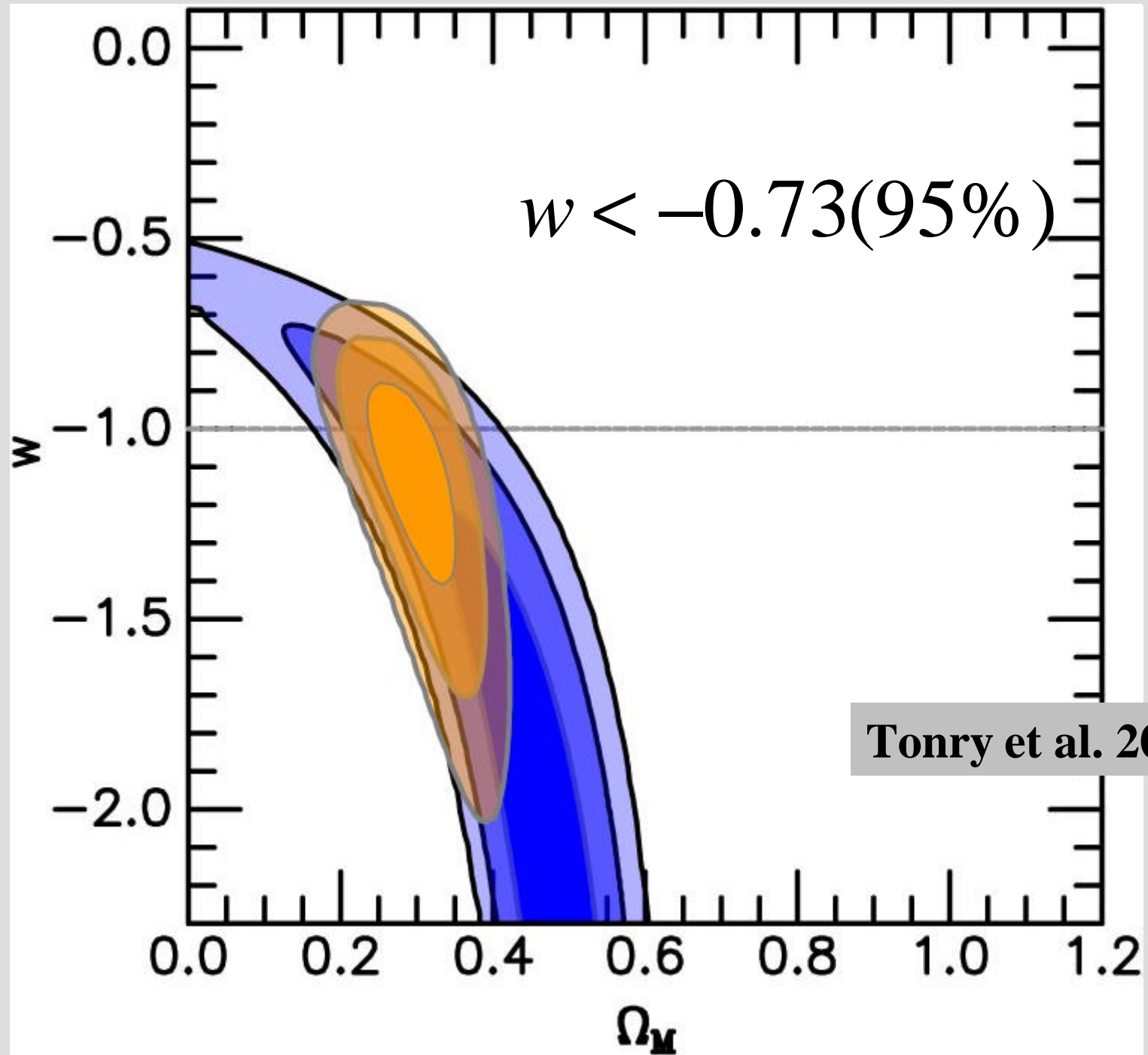
$$D_L = \frac{(1+z)c}{H_0 \sqrt{|\Omega_k|}} S \left\{ \sqrt{|\Omega_k|} \int_0^z \left[\Omega_k (1+z')^2 + \sum_i \Omega_i (1+z')^{3(1+w_i)} \right]^{-1/2} dz' \right\}$$

- **with** $k = 1 - \sum_i \Omega_i$ **and** $w_i = \frac{p_i}{r_i c^2}$

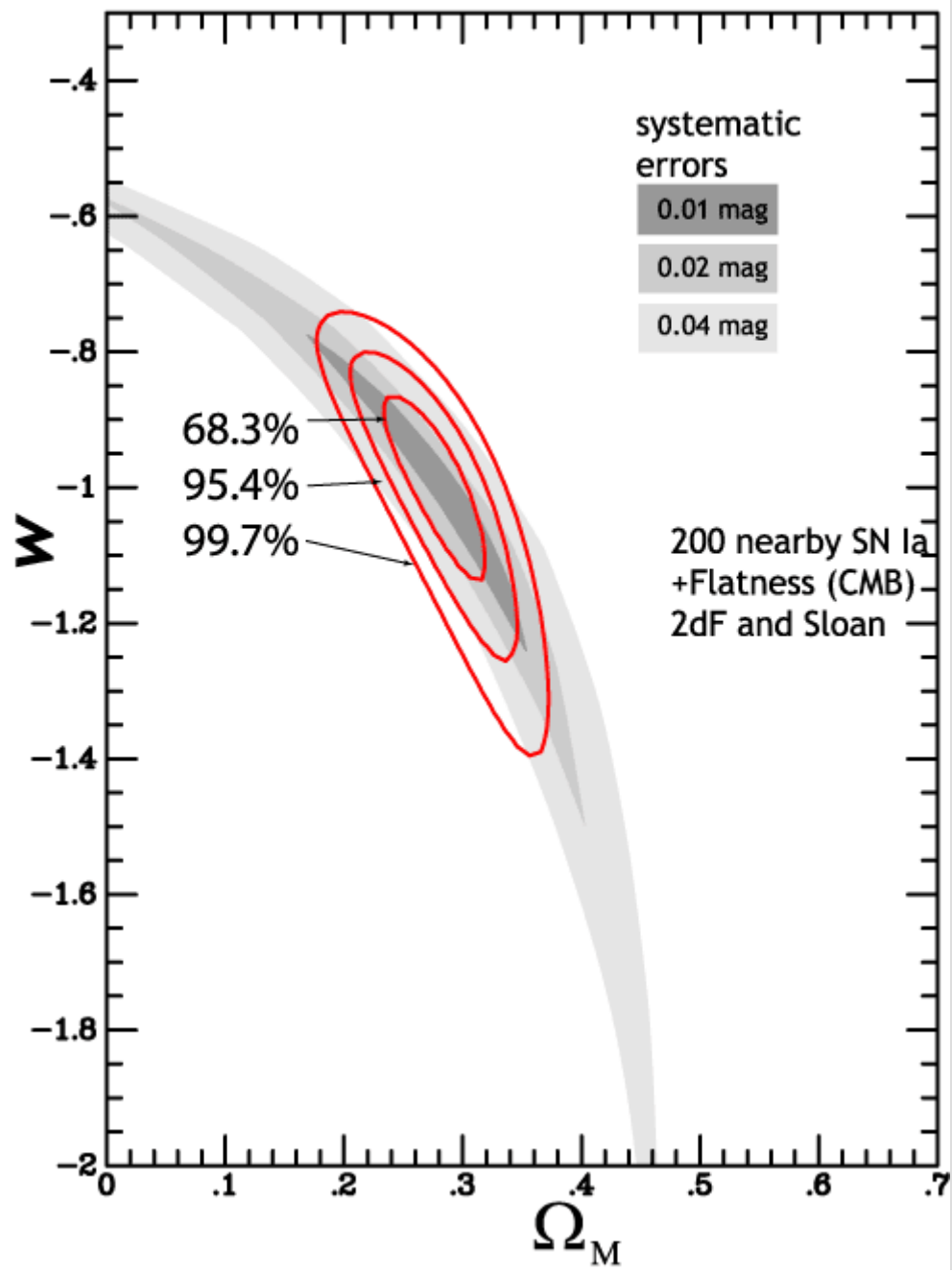
$w_M = 0$ (matter)

$w_R = ?$ (radiation)

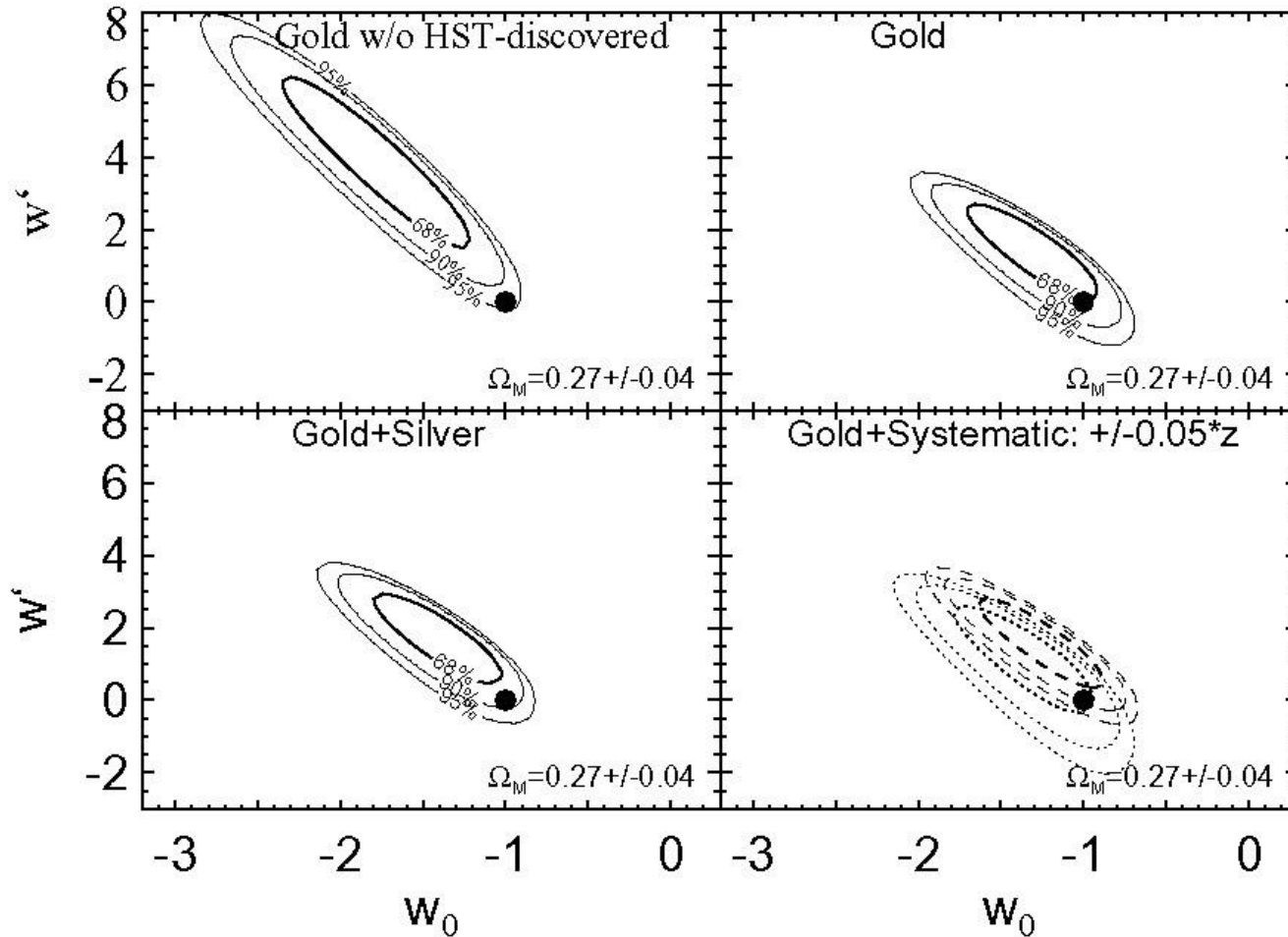
$w_L = -1$ (cosmological constant)



Tonry et al. 2003



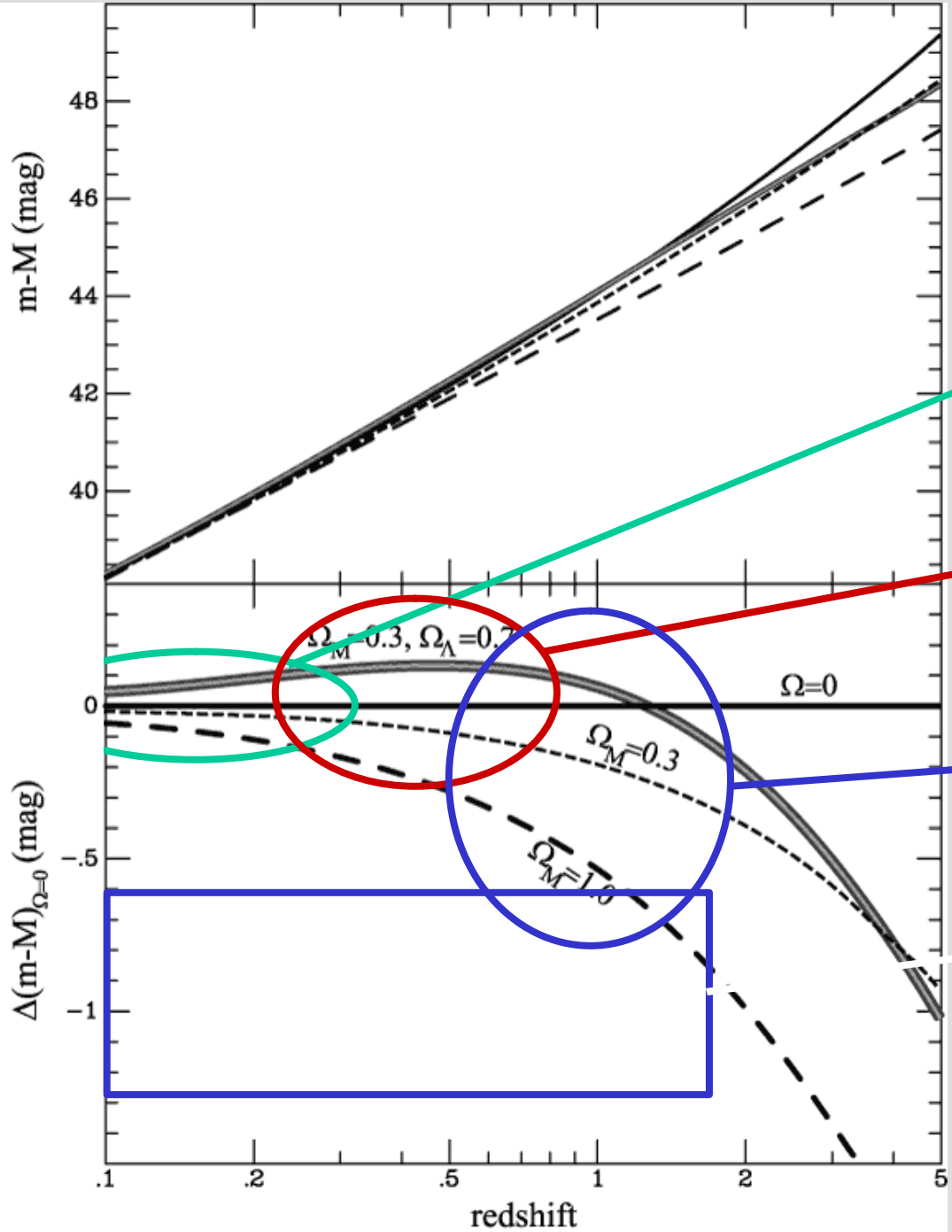
And on to a variable ?



Ansatz:

$$? (z) = ?_0 + ?'z$$

SN Projects



SN Factory
Carnegie SN Project

ESSENCE
CFHT Legacy Survey

Higher-z SN Search
(GOODS)

SNAP

Four redshift regimes

$z < 0.05$

- **Define the characteristics of Type Ia Supernova**
- **Understand the explosion and radiation physics**
- **Determination of H_0**

$z < 0.3$

- **Explore the systematics of SNe Ia**
- **Establish distance indicator**

Four redshift regimes (cont.)

$0.2 < z < 0.8$

- **Measure the strength of the cosmic acceleration (dark energy)**

$z > 0.8$

- **break the degeneracy**
- **measure matter density**

All redshifts

- **Measure details of dark energy**

The SN Ia Hubble diagram

- **powerful tool to**
 - **measure the absolute scale of the universe**
 H_0
 - **measure the expansion history (q_0)**
 - **determine the amount of dark energy**
 - **measure the equation of state parameter of dark energy**

Caveats

Warning to the theorists:

Claims for a measurement of a change of the equation of state parameter w are exaggerated. Current accuracy is inadequate for too many free parameters in the analysis.

Summary

Type Ia supernovae appear currently the most promising route to provide a possible answer to what the Dark Energy is.

All redshifts need to be covered

- distant SNe Ia alone are useless**
- nearby SNe Ia are the source of our understanding of the distance indicator**