



Two Quasi-periodic Fast-propagating Magnetosonic Wave Events Observed in Active Region NOAA 11167

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Abstract

We report a detailed observational study of two quasi-periodic fast-propagating (QFP) magnetosonic wave events that occurred on 2011 March 9 and 10, respectively. Interestingly, both the events have two wave trains (WTs): a strong main one (WT-1) and a small and weak secondary one (WT-2). Peculiar and common characteristics of the two events are observed, namely, (1) the two QFP waves are accompanied with brightenings during the whole stage of the eruptions; (2) both the two main WTs are nearly propagating along the same direction; (3) extreme-ultraviolet (EUV) waves are found to be associated with the two events. Investigating various aspects of the target events, we argue that (1) the second event is accompanied with a flux rope eruption during the whole stage; (2) the second event eruption produces a new filament-like dark feature; (3) the ripples of the two WT-2 QFP waves seem to result from different triggering mechanisms. Based on the obtained observational results, we propose that the funnel-like coronal loop system is indeed playing an important role in the two WT-1 QFP waves. The development of the second WT-2 QFP wave can be explained as due to the dispersion of the main EUV front. The coexistence of the two events offers thereby a significant opportunity to reveal what driving mechanisms and structures are tightly related to the waves.

Unified Astronomy Thesaurus concepts: [Solar corona \(1483\)](#); [Magnetohydrodynamics \(1964\)](#); [Solar flares \(1496\)](#); [Solar extreme ultraviolet emission \(1493\)](#); [Solar active regions \(1974\)](#); [Solar coronal waves \(1995\)](#); [Space plasmas \(1544\)](#)

Supporting material: animations

1. Introduction

Magnetohydrodynamic (MHD) waves carry vital information of the source region and they propagate across structured waveguides; therefore, the analysis of MHD waves could be used to infer the key parameters of both the source and waveguide on the Sun, which are not usually measurable in practice. Various types have been detected and studied during the last decades, such as coronal extreme-ultraviolet (EUV) waves (e.g., Thompson et al. 1998, 1999; Liu et al. 2010; Warmuth 2010, 2015; Warmuth & Mann 2011; Yuan & Nakariakov 2012; Yang et al. 2013; Liu & Ofman 2014; Muhr et al. 2014; Goddard et al. 2016, 2019; Kumar et al. 2017; Pascoe et al. 2017, 2019; Cheng et al. 2018; Shen et al. 2018b, 2019), chromospheric Moreton waves (Moreton 1960; Chen & Wu 2011; Krause et al. 2018), fast mode (Ofman & Thompson 2002; Liu et al. 2011, 2012; Yuan et al. 2013; Zhang et al. 2015; Ofman & Liu 2018), and slow mode (e.g., Nakariakov & Zimovets 2011; Yuan et al. 2015b) magnetosonic waves. There are various types of waves that can lead to coronal loop and filament oscillations (Nakariakov & Ofman 2001; Nakariakov & Verwichte 2005; Chen et al. 2008; Liu et al. 2012; Li & Zhang 2012; Shen et al. 2014a, 2014b; Zhou et al. 2018). Many solar physicists believe that waves and oscillatory phenomena are very important and crucial for the coronal heating (Nakariakov et al. 1999a, 1999b; Yuan & Van Doorsselaere 2016a, 2016b). Waves and oscillatory phenomena are also used to detect the magnetic field in which they are propagating (Shen et al. 2012, 2013a; Ofman & Liu 2018).

The quasiperiodic fast-propagating (QFP) magnetosonic waves are usually along the funnel-like loops (Liu et al. 2011, 2012; Yuan et al. 2016a, 2016b; Qu et al. 2017; Shen et al. 2018a; Miao et al. 2019). In addition, the QFP wave trains (WTs) may be triggered by impulsive energy releases in solar flares (Liu et al. 2010, 2012; Shen & Liu 2012a; Yuan et al. 2013; Shen et al. 2014b; Kumar et al. 2017; Yu & Chen 2019). The first unambiguous observation of QFP WTs were reported by Liu et al. (2012) using the high resolution observations taken by the *Solar Dynamics Observatory* (SDO; Pesnell et al. 2012)/Atmospheric Imaging Assembly (AIA; Lemen et al. 2012). The authors found two components of the multiple arc-shaped WTs that propagate, simultaneously, ahead and behind a coronal mass ejection (CME). The two components of the WTs appear to have different periods and speeds. This phenomenon has attracted significant increasing interests since its early observations (Liu et al. 2010). Nakariakov et al. (2004) presented the characteristic time evolution of these short-period QFP WTs. Due to the low temporal and spatial resolution in early observations, reported detections of QFP waves are still very scarce so far. Williams et al. (2001, 2002) studied a QFP magnetosonic wave with a period of 6 s and a phase speed of 2100 km s⁻¹. Liu et al. (2011, 2012) found that multiple arc-shaped QFP wavefronts sequentially emanate from the kernel of the accompanying flare. The QFP wave is found to possess some common periods with the accompanying flare. Therefore, the authors considered that the QFP wave has a tight relationship with the accompanying flare. Shen & Liu (2012a) and Shen et al. (2013b) confirmed that not

only common periods simultaneously are detected in both the QFP waves and the accompanying flares, but interestingly some extra periods in QFP waves are also detected without being associated with the accompanying flares. Liu & Ofman (2014) summarized the characteristics of the QFP waves. The authors indicated that the speed, period, and deceleration of QFP waves are in ranges 500–2200 km s⁻¹, 25–400 s, and 1–4 m s⁻², respectively. A QFP wave event was reported by Liu et al. (2011) that was successfully reproduced through a three-dimensional MHD model by Ofman et al. (2011). They presented the three-dimensional MHD modeling to interpret the nature and evolution of the QFP wave. In addition, Liu et al. (2010, 2011, 2012) and Ofman & Liu (2018) considered that the periods of the QFP waves have a tight relationship with flares. Using the two-dimensional MHD model, Pascoe et al. (2013) and Pascoe et al. (2014) reported that fast-mode waves propagating in funnel-like waveguides can also dispersively evolve into QFP WTs. A numerical simulation study by Yuan et al. (2015a) pointed out that the ripples of the QFP waves can be generated by the dispersion of the EUV waves. The authors provided a new approach to detect the relationship between the filamentations and the waves, which can also be used to diagnose the presences of a true QFP wave.

Many observations of QFP waves are usually accompanied with EUV waves (e.g., Liu et al. 2012; Shen & Liu 2012a; Miao et al. 2019). Large-scale EUV wave was first observed by the *Solar and Heliospheric Observatory*/Extreme-ultraviolet Imaging Telescope (EIT; Delaboudinière et al. 1995), and was initially dubbed “EIT wave” (Thompson et al. 1998, 1999). The debate on the EUV-waves physical nature is still open (Chen 2017). Early observations of the EUV waves indicated that they are probably the coronal counterpart of the chromospheric Moreton waves (Thompson et al. 1998, 1999; Wang 2000; Wu et al. 2001; Ofman & Thompson 2002; Schmidt & Ofman 2010; Shen & Liu 2012b; Shen et al. 2017b, 2018a, 2018b). In particular, Chen et al. (2002, 2005) considered that there should be two kinds of EUV waves associated with a CME event, namely, a slowly moving obvious wave and a fast-mode wave. The authors considered that the faster one corresponds to the coronal counterpart of Moreton wave, while the slower one is triggered by the erupting flux rope. The triggering mechanism of EUV waves is also an open question. Some solar scientists believe that EUV waves are driven by the pressure pulse inside the flare (e.g., Cliver et al. 1999; Li et al. 2012; Shen & Liu 2012b; Shen et al. 2017a), while others propose that they are indeed excited by CMEs (e.g., Cliver et al. 1999; Chen et al. 2002, 2008; Chen 2006, 2009, 2016; Li et al. 2012; Shen & Liu 2012b; Shen et al. 2017a; Miao et al. 2018). It is now widely accepted that the large-scale EUV waves, both the faster and the slower are driven by CMEs (see Warmuth 2015; Chen 2016; Liu et al. 2018, 2019, for reviews). The EUV wave and the QFP wave may show some close relationship. Miao et al. (2019) presented a QFP wave that was most probably triggered by a CME as the piston-driven shock wave interacts with funnel-like coronal loops, reminiscent to what was previously reported by Shen et al. (2018b). The authors indicated that the original broadband pulse could dispersively develop into multiple QFP wavefronts. Pascoe et al. (2013) performed an interesting work to highlight that process, which was observationally confirmed later-on by Nisticò et al. (2014). It is worth noticing here that Shen et al. (2018b) reported that EUV waves can be driven by sudden loop expansions with the lifetimes of the waves shorter than those driven by CMEs.

Until now, the triggering mechanism, evolution processes and the physical nature of the QFP wave events are still unclear, essentially, due to the rarity of such detected events in the literature. Certainly the richer the sample of the studied events the better our understanding of the associated triggering physical mechanisms. In this paper, we present two QFP wave events that occurred in the same active region, with one event interestingly accompanied with the eruption of a magnetic flux rope. Each of the two EUV waves were in front of the main WT (WT-1) and produced a weak halo CME. Observations and instruments, used in our investigation, are introduced in Section 2. The observational results of the two QFP wave events are presented in Section 3. Discussions and conclusions are highlighted in the last section.

2. Observations and Instruments

The present two QFP wave events were observed by *SDO*/AIA from 2011 March 9 to 10. The seven EUV channels and three UV-visible channels full-disk images are taken by the AIA instrument, whose temporal cadences are 12 s and 24 s, respectively. The field of view and the spatial resolution of the AIA instrument are, respectively, 1.3 R_{\odot} and 0''6 pixel⁻¹. The line-of-sight (LOS) magnetograms and continuum intensity images are taken by the Helioseismic and Magnetic Imager (HMI; Scherrer et al. 2012) on board *SDO*. The spatial and temporal resolutions of HMI LOS images are 45 s and 0''5, respectively. The measurement precision of the HMI LOS magnetograms is 10 Gauss. The EUV waves were also observed by the Extreme Ultraviolet Imager (EUVI) of the Sun Earth Connection Coronal and Heliospheric Investigation (SECCHI; Howard et al. 2008) on board the *Solar Terrestrial Relations Observatory* (*STEREO*; Kaiser et al. 2008) which captures full-disk 195 and 304 Å images, with 5 and 10 minute cadence and a pixel width of 1''6.

3. Observational Results

In this paper, we present two QFP wave events hosted by AR NOAA 11167. This section is divided into two parts. The first part introduces the first QFP wave event, while the second part reports the second QFP wave event. Both events were interestingly accompanied by a process of brightening in the active region.

3.1. The First QFP Wave Event

The first QFP wave event, hosted by AR NOAA 11167, started at about 23:48 UT on 2011 March 9 and ended at about 00:10 UT on 2011 March 10. This event involved two multiple arc-shaped WTs, namely, one main WT (WT-1) and a weak WT (WT-2), within a broad EUV wave, as shown in Figure 1. At about 23:47:45 UT, the active region began to erupt and gradually brightened. The brightening almost accompanied the whole stage of the QFP wave event. During the brightening lifetime, multiple arc-shaped wavefronts are observed to be continuously emanating from it. These wavefronts propagated along a cluster of funnel-like coronal loops rooted in the center of the active region. At about 23:52 UT, a weak and small scale WT (WT-2) was detected through 171 and 193 Å running difference images (see Figures 1(a3) and (b3)), respectively. It should be noted that the WT-2 was identified only partially on the 193 Å channel. Additionally, it was also visible later on in 171 Å images. It might imply that the QFP wave was probably inclined to propagate along the lower corona (Miao et al. 2019).

At the beginning of the eruption, an EUV wavefront can be identified as a bright semicircle encompassing these WTs (see

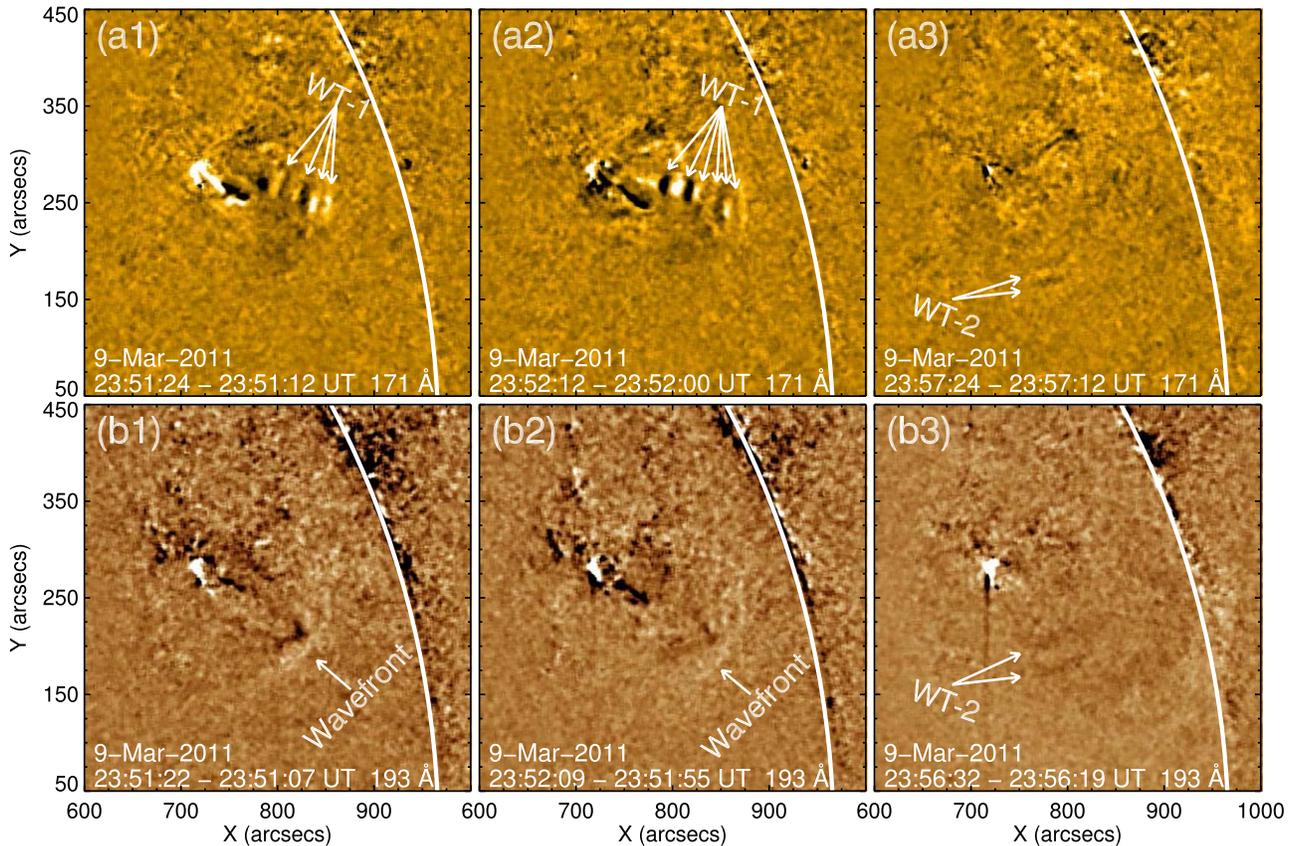


Figure 1. AIA 171 and 193 Å running-difference images. Panels (a1)–(b3) show the WT-1, WT-2, and EUV wave in the first event. The arrows in panels (a1) and (a2) point to the multiple arc-shaped main WT (WT-1). Panels (a3) and (b3) report the weak WT (WT-2) that is indicated by two white arrows in the two channels, respectively. The EUV wavefront is displayed in panels (b1) and (b2).

Figures 1(b1), (b2)). Fortunately, the EUV wave was also observed by the *STEREO*/EUVI-A 195 Å channel (see Figure 2). The evolution of the EUV wavefront is reported in Figures 2(b1)–(b6). From Figure 2, it seems that the configuration of the EUV bright wavefront is symmetrical. The green arrows indicate the position of a filament-like (FL) dark feature, which refers to a preexisting structure. The profile of the structure is indeed clearly discernible and is labeled as FL1 in Figure 3. The evolution of the eruption is highlighted in Figure 3 using the 195 and 304 Å raw images. Starting at about 23:50:30 UT (see Figure 3(a2)), near the footpoint of FL1, a brightening is reported in the active region (see also the animation of Figure 2).

In order to quantify the kinematics of the QFP wave, we use a semiautomatic method to construct two stack plots from 6° wide sectors (“A” and “B”; see Figure 4(a)), made from AIA 171 Å running-difference images on the solar surface. The resulting stack plots are reported in Figures 4(a1) and (a2). Adopting the same method, six stack plots from 15° wide sectors (“A1” to “A6”; see Figure 4(b)) made from AIA 193 Å running-difference images are displayed in Figures 4(a3)–(a8). The average speed of the WT-1 and the EUV wave is about 718 km s⁻¹ in the 171 Å channel (see Figure 4(a2)). The speed of the WT-2 is in the range of 300–404 km s⁻¹ (see Figures 4(a1) and (a3)). The speed of the EUV wave is found in the range of 378–802 km s⁻¹ in the 193 Å channel as shown in Figures 4(a4)–(a8). For a better emphasis of the evolution of the waves, a corresponding constructed animation is available (Figure 4). The red dashed lines in Figure 4 denote the positions where we analyze the periodicities of the two WTs. The detailed results are shown in Figure 9.

3.2. The Second QFP Wave Event Associated with the Eruption of a Flux Rope

The second QFP wave event was associated with an eruption of a magnetic flux rope at about 04:08 UT on 2011 March 10. The *SDO*/AIA 1700, 171, 193, 304, 131, 94 Å images are displayed in the two top rows of Figure 5 to illustrate the structures in the active region of AR11167. Clearly, the small active region (see Figure 5(a)), the funnel-like coronal loops (see Figure 5(b)), and the flux rope, are highly related to the event. In addition, the flux rope is unequivocally identified by multiple channels imaging (see Figures 5(b)–(f)).

The second QFP wave event is also characterized by two WTs. It should be noted that we also use the same names WT-1 and WT-2 to represent the two WTs. The WT-1 and the flux rope almost simultaneously erupted with associated observed brightening. The flux rope is not only observed in the hot 131 and 94 Å channels (see Figures 5(e) and (f)) but also in the cool 304 Å channel (see Figure 5(d)), indicative of a clear flux rope topology (Zhang et al. 2012; Cheng et al. 2013; Filippov et al. 2015; Awasthi et al. 2018; Joshi et al. 2018; Shen et al. 2018a). The flux rope eruption did not cause any CME. The contours of the positive (red) and negative (blue) magnetic fields are overlaid on the 131 Å image as shown in Figure 5(e), where the contour adopted levels are ±300 and ±100 G. The profile of the flux rope is also highlighted by a green curve line in panel (e) of Figure 5. In panel (h) of Figure 5, the profile of the flux rope is also overlaid on the HMI LOS image. The two footpoints of the flux rope roots in the positive magnetic field and negative magnetic field, respectively. Using the 171 and

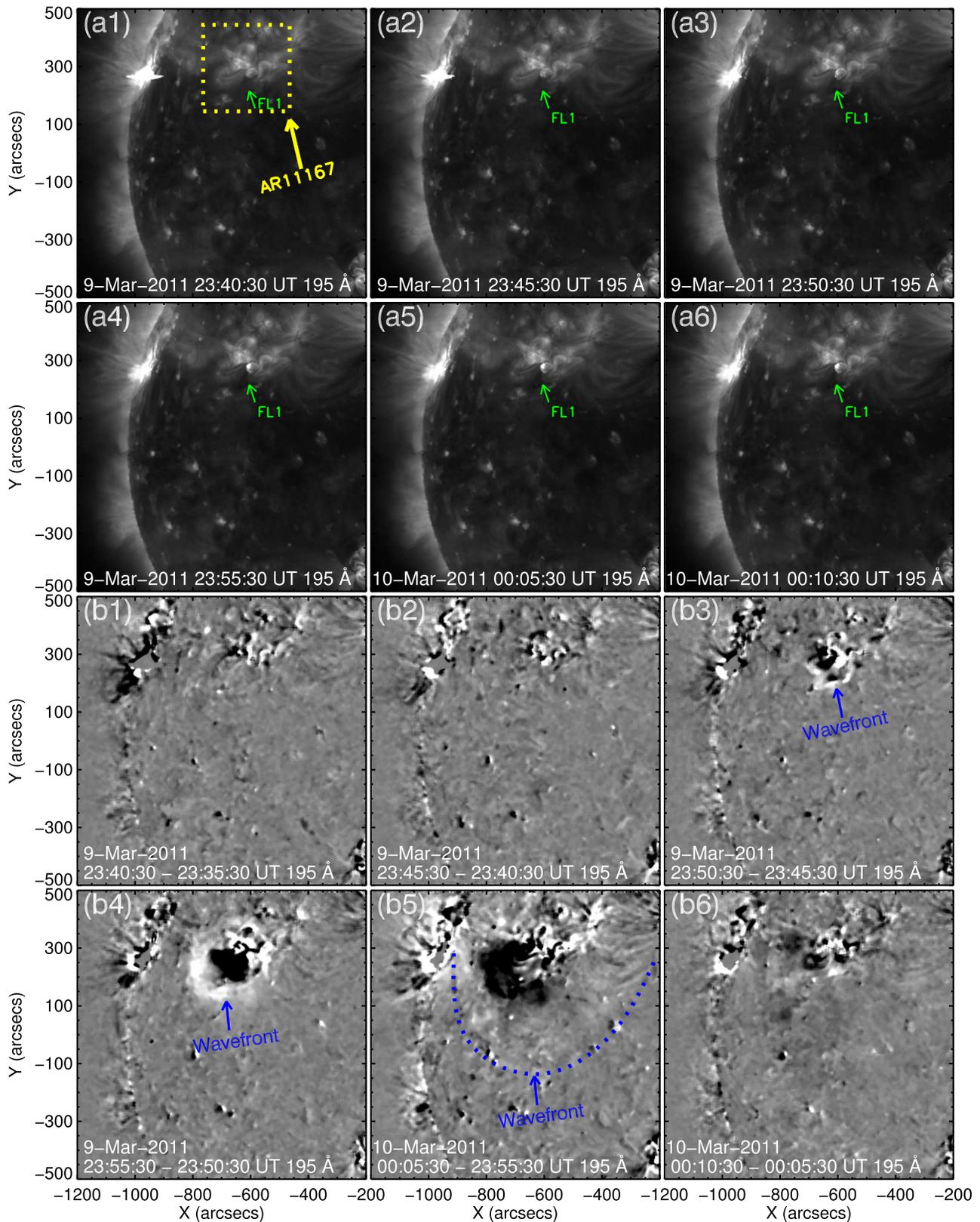


Figure 2. STEREO/EUVI-A 195 Å images. Panels (a1)–(a6) show the evolution of the EUV wave in the first event. The green arrows in panels (a1)–(a6) display the FL dark feature (FL1). The yellow dotted box, in panel (a1), marks the location of AR11167. Panels (b1)–(b6) display the evolution of the running-difference images of the EUV wave in EUVI-A 195 Å channel. The blue arrows indicate the wavefront in panels (b3)–(b5), while the blue dotted arc represents the configuration of the EUV wave. An animation of this figure is available. The video begins at 2011 March 9, 23:25:30 UT, and ends at 2011 March 10, 02:00:30 UT. The realtime duration of the video is 2 s.

(An animation of this figure is available.)

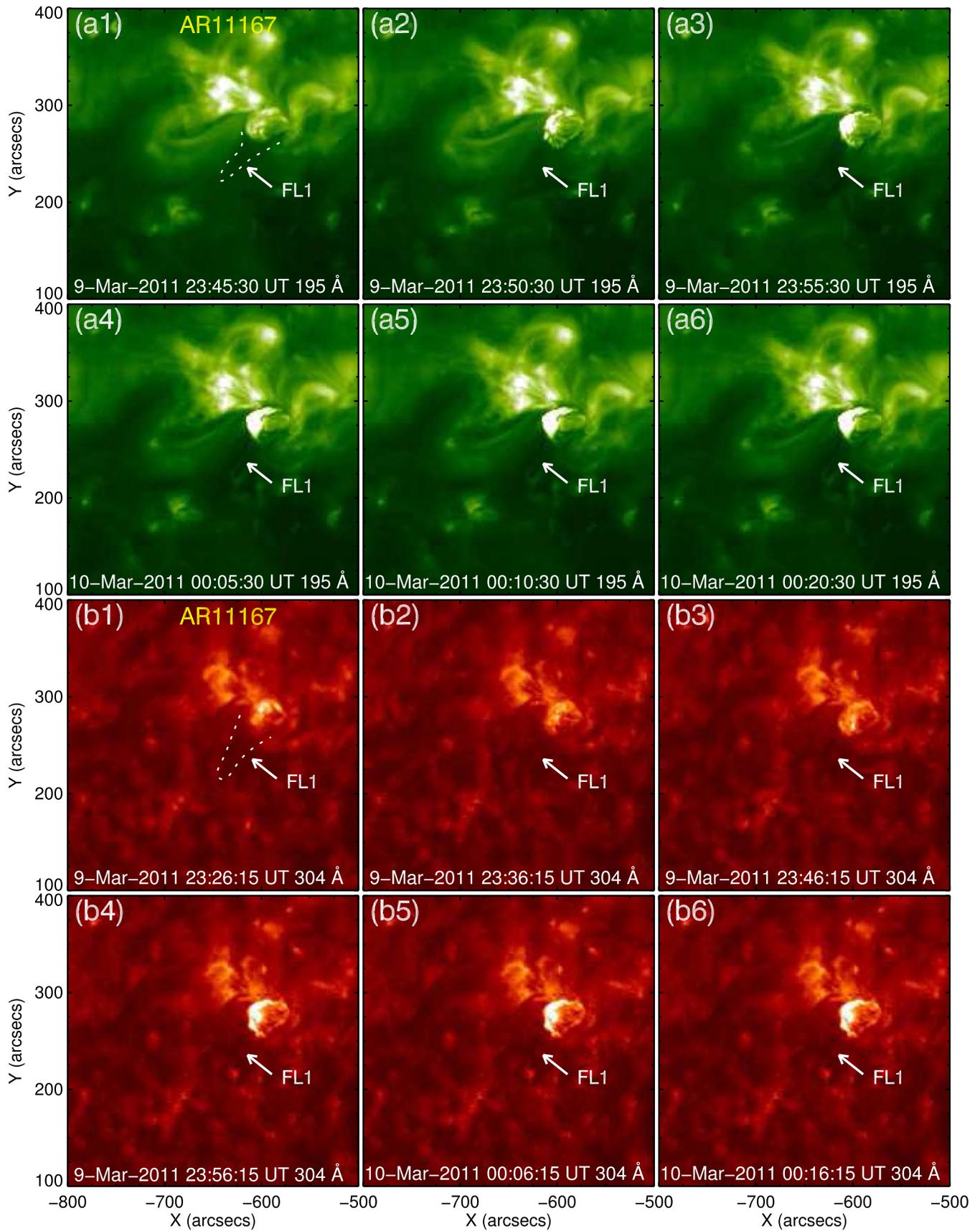


Figure 3. STEREO/EUVI-A 195 and 304 Å images, highlighting zoomed view of the active region AR11167. Panels (a1)–(a6) display the evolution of the active region in the EUVI-A 195 Å channel. The white arrows indicate the FL dark feature (FL1) in panels (a1)–(b6). The dotted curves of the profiles mark the FL dark feature (FL1) in panels (a1) and (b1). The evolution of the brightening is clearly recognizable from the figure.

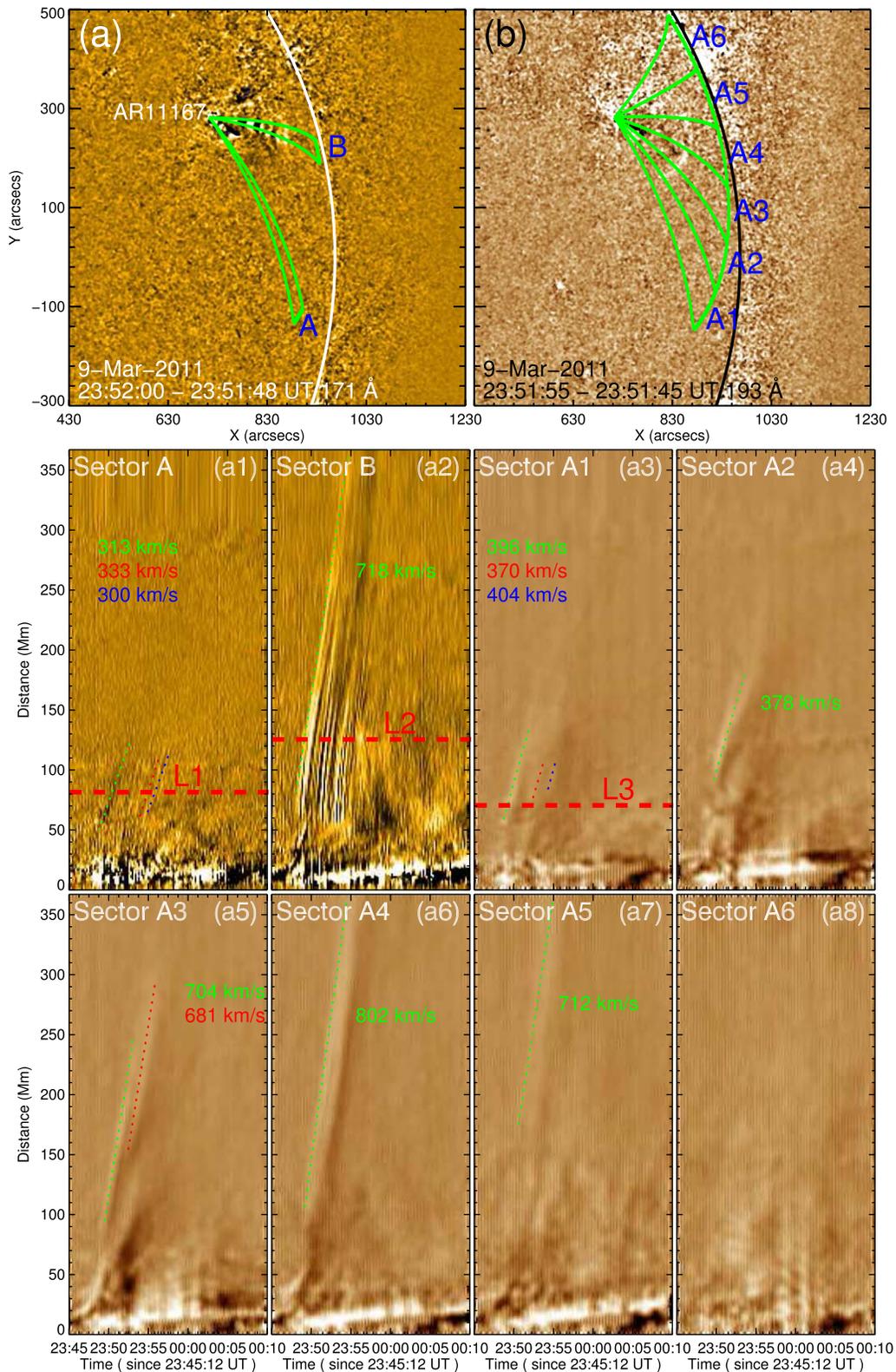


Figure 4. *SDO/AIA* 171 and 193 Å running-difference images showing sectors of the two event. The top row displays the two 6° wide sectors (“A” and “B”) and six 15° wide sectors (“A1”–“A6”). Two time-distance diagrams showing the kinematics of the two WT (WT-1, WT-2) and EUV wave in AIA 171 Å running-difference images (see panels (a1) and (a2)). Panel (a1) displays the speeds of the WT-2 in the range of 300–333 km s⁻¹. The average speed of the WT-1 and the EUV wave is about 718 km s⁻¹ as shown in panel (a2). Six time-distance diagrams illustrating the kinematics of the WT-2 and EUV wave in AIA 193 Å running-difference images. The speeds of the EUV wave are in the range of 378–802 km s⁻¹ in different directions. Panel (a3) shows the speeds of the WT-2 in the range of 370–404 km s⁻¹. The wave signal along the sixth sector is not detected in panel (a8). The three red dashed lines in panels (a1)–(a3) mark the positions where the periodicities of the two WT. An animation of the top two panels (a) and (b) is available. The video begins at 2011 March 9, 23:45:22 UT and ends at 2011 March 10, 00:09:40 UT. The realtime duration of the video is 3 s.

(An animation of this figure is available.)

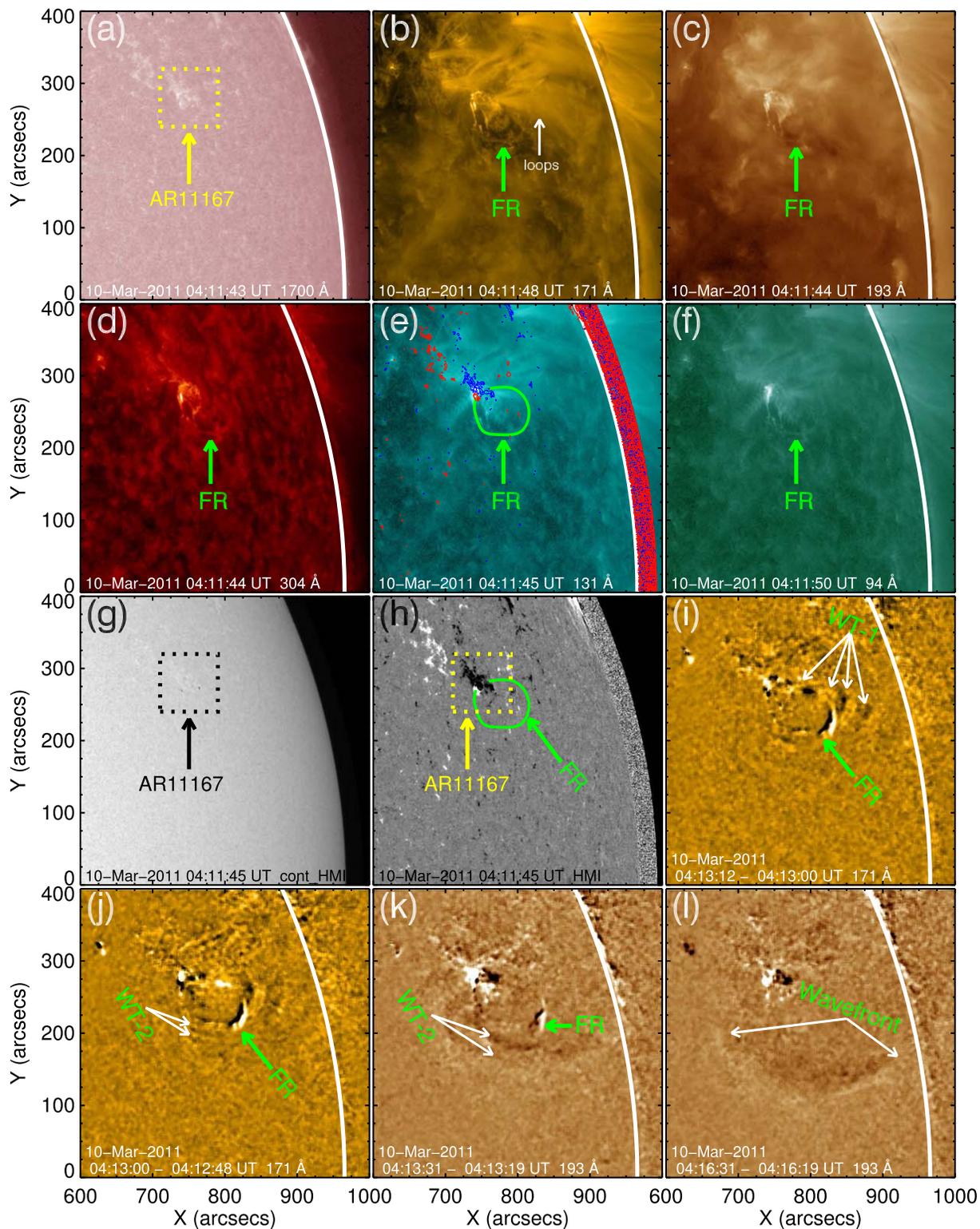


Figure 5. *SDO/AIA* 1700, 171, 193, 304, 131, and 94 Å images, illustrating the flux rope (FR) in panels (a)–(f). The panels (g) and (h) display the active region AR11167 in HMI images. The cluster of funnel-like coronal loops is also indicated in panel (b). The profile of the flux rope is also highlighted by a green curve line in panel (e). The red contours and the blue contours represent the positive magnetic field and the negative magnetic field in panel (e), respectively. The contour levels are ± 300 G and ± 100 G. In panel (h), the profile of the flux rope is also overlaid on the HMI LOS image. The main WT is represented by WT-1 in the second event in 171 running-difference images. The weak WT is represented by WT-2 in panels (j) and (k) in 171 and 193 Å running-difference images, respectively. The EUV wavefront is also indicated by the white arrows in panel (l). The boxes in panels (a), (g), and (h) highlight the location of AR11167.

193 Å running-difference images, the WT-1 and WT-2 are highlighted in Figures 5(i)–(k). During the eruption of the flux rope, multiple arc-shaped wavefronts of the WT-1 propagation

along the funnel-like coronal loops were observed. It is noted that the wavefronts of the WT-2 signals were faint and quick and can be seen in the AIA 171 and 193 Å running-difference

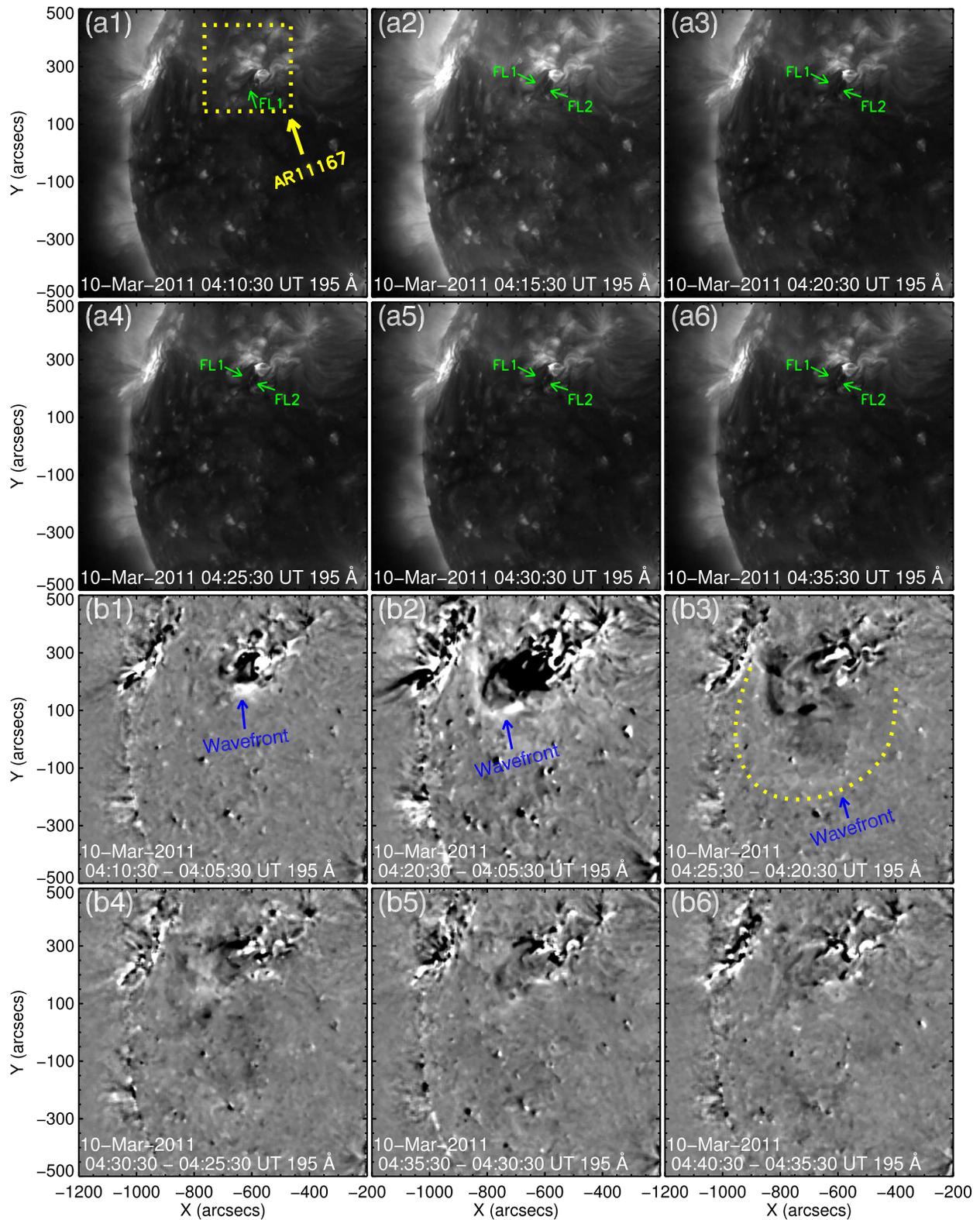


Figure 6. STEREO/EUVI-A 195 Å images, where panels (a1)–(a6) show the evolution of the EUV wave in the second event. The yellow dotted box in panel (a1) marks the location of AR11167. The green arrows in panels (a1)–(a6) display the preexisting FL dark feature (FL1) and a new FL dark feature (FL2) that can be seen from panels (a1) to (a6). Panels (b1)–(b6) display the evolution of the EUV wave of the second event in EUVI-A 195 Å running-difference images. The blue arrows indicate the wavefront in panels (b1), (b2), and (b3). The EUV wave occurring around 04:25:30 UT is highlighted by yellow dotted arc. An animation of this figure is available. The video begins at 2011 March 10, 03:35:30 UT and ends at 2011 March 10, 06:10:30 UT. The realtime duration of the video is 2 s.

(An animation of this figure is available.)

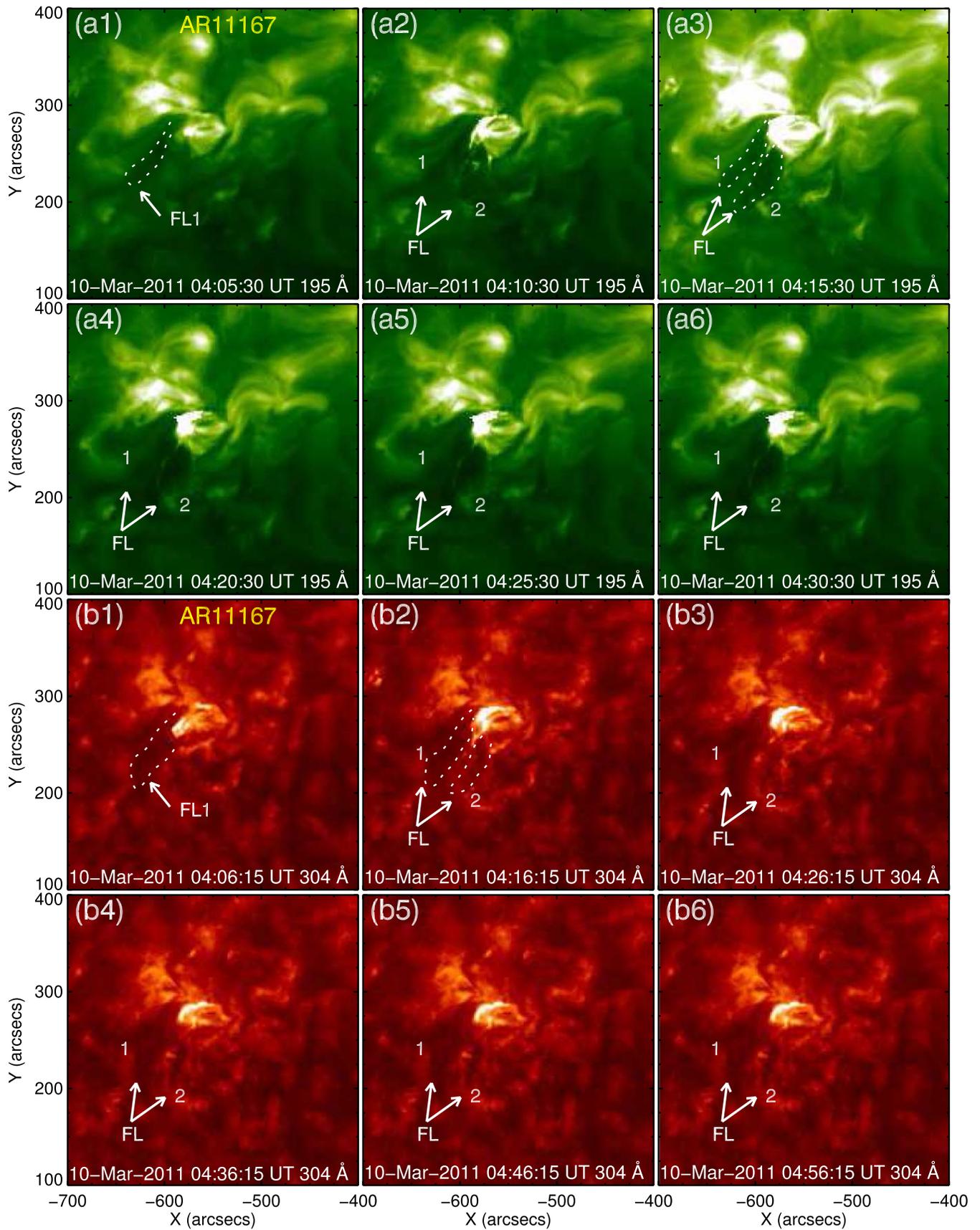


Figure 7. Zoomed view of the active region AR11167 in the *STEREO*/EUVI-A 195 and 304 Å raw images. Panels (a1)–(a6) and (b1)–(b6) display the evolution of the FL dark features in EUVI-A 195 and 304 Å images, respectively. The white dotted curves indicate the profile of the preexisting FL dark feature (FL1) in panels (a1) and (b1). The profiles of the preexisting FL dark feature (FL1) and the new FL dark feature (FL2) are sketched by dotted curves shown in panels (a3) and (b2).

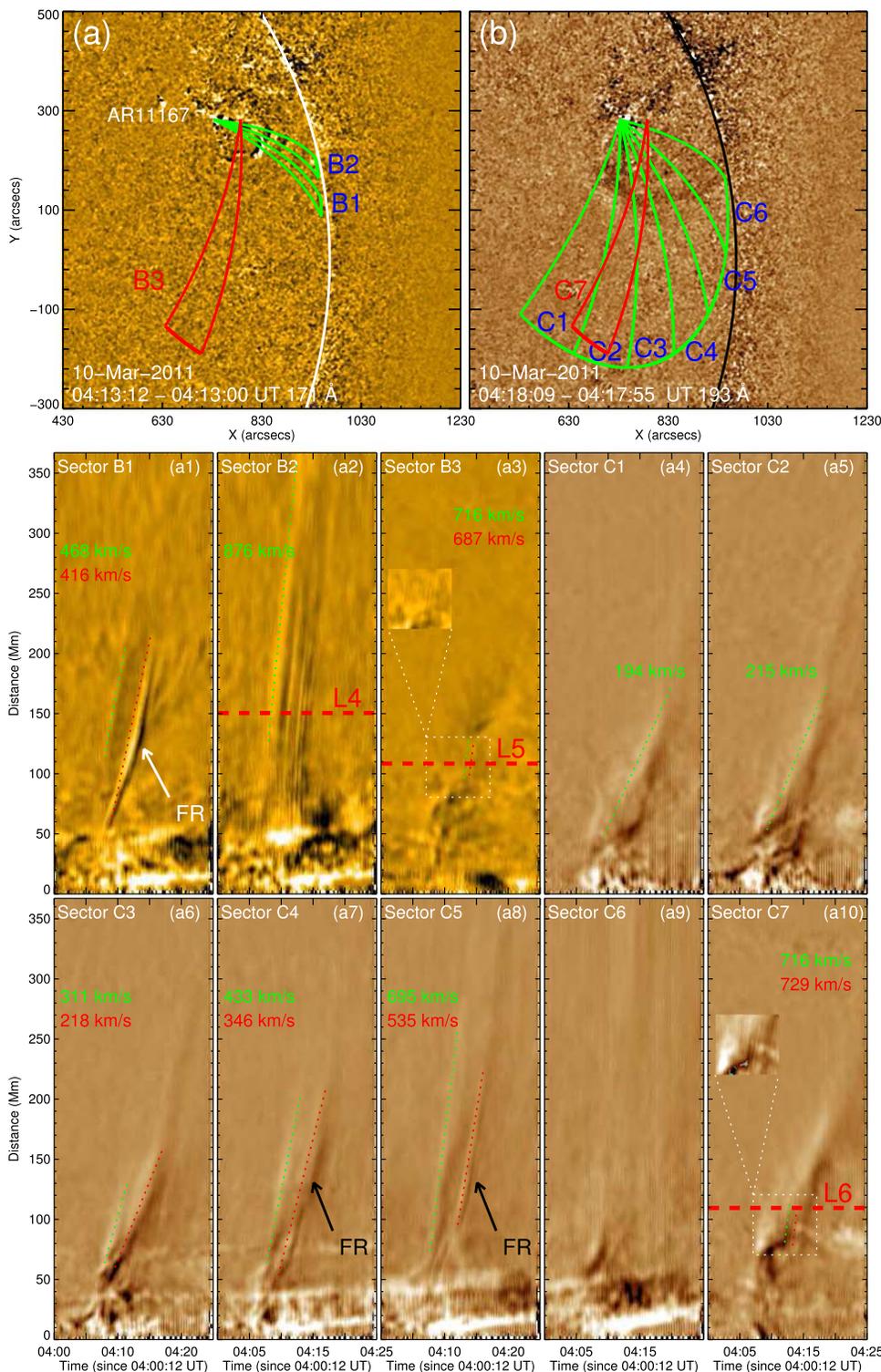


Figure 8. The top row shows the two 5° wide sectors (“B1” and “B2”), six 20° wide sectors (“C1”–“C6”), and two 15° wide sectors (“B3” and “C7”) (see panels (a) and (b)). Panels (a1)–(a3) are time–distance diagrams obtained from AIA 171 Å running-difference images along sectors “B1,” “B2,” and “B3” as shown in panel (a), respectively. Panels (a4)–(a10) are time–distance diagrams obtained from AIA 193 Å running-difference images along sectors “C1”–“C7” as shown in panel (b), respectively. The speeds of the EUV wave and flux rope are about 468 and 416 km s^{-1} as shown in panel (a1). The average speed of the WT-1 and the EUV wave is about 876 km s^{-1} as reported in panel (a2). Panels (a4)–(a9) show the kinematics of the EUV wave and the flux rope from AIA 193 Å running-difference images as shown in panel (b). The speeds of the EUV wave are in the range of 194–695 km s^{-1} in different directions (see panels (a4)–(a9)). The speeds of the flux rope are in the range of 346–535 km s^{-1} in panels (a7) and (a8). The speeds of the WT-2 are in the range of 687–716 and 716–729 km s^{-1} in panels (a3) and (a10), respectively. The red dashed lines in panels (a2) (“L4”), (a3) (“L5”), and (a10) (“L6”) mark the positions where the periodicity of the two WTs have been estimated. An animation of the top two panels (a) and (b) is available. The video begins at 2011 March 10, 04:00:20 UT and ends at 2011 March 10, 04:24:34 UT. The realtime duration of the video is 3 s.

(An animation of this figure is available.)

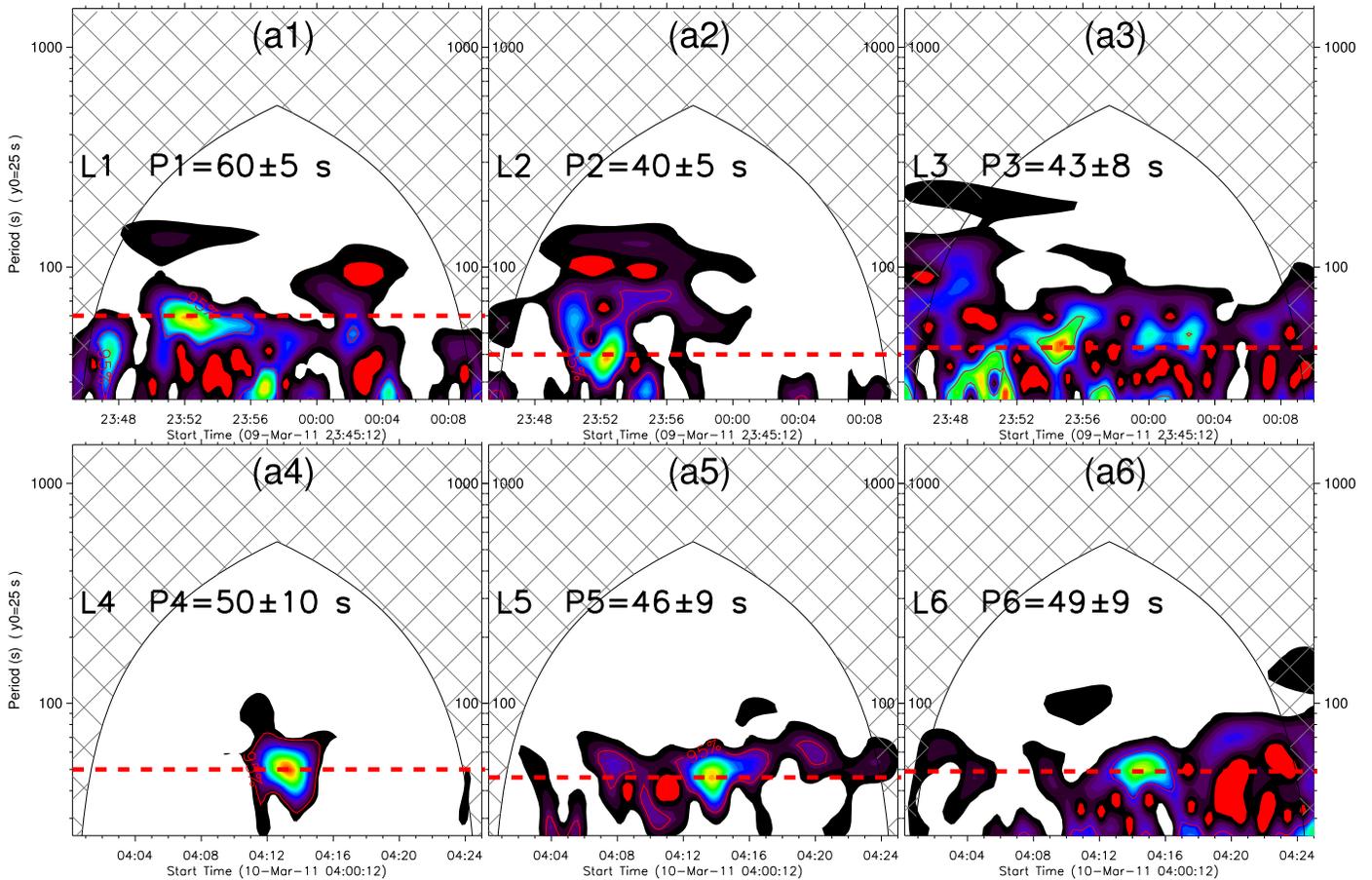


Figure 9. Panels (a1)–(a6) are the wavelet power diagrams of the detrend intensity profiles of “L1”–“L6,” respectively. Panels (a1)–(a3) show the periods of the first event, while panels (a4)–(a6) show the periods of the second event. The period of the first WT-1 is about 40 ± 5 s, while the periods of the WT-2 are 60 ± 5 and 43 ± 8 s as indicated in panels (a2)–(a3). The period of the second WT-1 is about 50 ± 10 s, shown in panel (a4). The periods of the second WT-2 are about 46 ± 9 and 49 ± 9 s as displayed in panels (a5) and (a6), respectively.

images. Actually, the active region NOAA 11167 consists of a very small bipolar magnetic field structure. According to panels (a), (g), and (h) of Figure 5, one can distinguish the active region having a small scale magnetic field and few small-scale sunspots.

We also exploit the *STEREO*/EUVI-A data to study the second EUV wave. Due to the low cadence of the data, only the EUV wave was detected from the *STEREO*. Panels (a1)–(a6) of Figure 6 are 195 \AA images and panels (b1)–(b6) of Figure 6 are in EUVI-A 195 \AA running-difference images, respectively. The green arrows indicate the FL dark features in Figures 6(a1)–(a6). More details of those features are displayed in Figure 7. From the panels (b1) to (b3) of Figure 6, we recognize a bright wavefront appearing on the southeastern side of the edge of AR11167. The animation of Figure 6 highlights the second EUV wave, made using 195 \AA images. In Figure 7 we display the evolution of the EUV wave from the viewpoint of the *STEREO*/EUVI-A in 195 and 304 \AA channels. At about 04:05:30 UT, we identify an FL dark feature rooted in the periphery of the accompanying brightening (see panel (a1) of Figure 7. The FL dark feature in panel (a1) and (b1) of Figure 7, is indeed nothing but the FL1 structure reported in the first event. As the flux rope erupts, a new FL dark feature (FL2) appeared at about 04:10:30 UT (see panel (a3) of Figure 7).

To inspect the kinematics of the QFP waves and the flux rope, we again utilize a semiautomatic method to obtain two stack plots from 5° wide sectors (see “B1” and “B2” in Figure 8(a)) and one stack plot from 15° wide sector (see “B3” in Figure 8(a)) in AIA 171 \AA running-difference images. Similarly, we get six stack plots from 20° wide sectors (see “C1”–“C6” in Figure 8(b)) and one stack plot from a 15° wide sector (see “C7” in Figure 8(b)) in AIA 193 \AA running-difference images. The results are reported in Figure 8(a1)–(a10). An animation is also made to show the evolution and the process of the second event (see Figure 8).

The stack plot of sector “B1” is shown in Figure 8(a1). The green and red dotted lines show the speeds of the EUV wavefront and flux rope to be 468 and 416 km s^{-1} , respectively. The stack plot of sector “B2” is then used to measure the speed of the WT-1 as shown in Figure 8(a). The average speed of the WT-1 and the EUV wave is estimated to be 876 km s^{-1} as reported in Figure 8(a2). The speed of the WT-2 is found in the range of 687 – 729 km s^{-1} (see Figures 8(a3) and (a10)). We exploit the AIA 193 \AA running-difference images to inspect the kinematics of the EUV wave as well as of the flux rope. The resulted stack plots are displayed in Figures 8(a4)–(a9). According to Figure 8, the speeds of the EUV wave and of the flux rope are hence evaluated to be in the

Table 1
Parameters of the Two Events

Event	Wavelength	EUV Wave Speed (km s ⁻¹)	QFP Wave Speed (km s ⁻¹)	Position	Period (s)
The First WT-2	171 Å	...	300–333	L1	60 ± 5
The First WT-1	171 Å	718 ^a	718 ^a	L2	40 ± 5
The First WT-2	193 Å	378–802	370–404	L3	43 ± 8
The Second WT-1	171 Å	468–876	876 ^b	L4	50 ± 10
The Second WT-2	171 Å	...	687–716	L5	46 ± 9
The Second WT-2	193 Å	194–695	716–729	L6	49 ± 9

Notes.

^a The average speed of the EUV and QFP waves of the WT-1 in the first event.

^b The average speed of the EUV and QFP waves of the WT-1 in the second event.

range of 194–876 km s⁻¹ and 218–535 km s⁻¹, respectively. The speed range of the EUV wave has a large span, probably because the eruption of the EUV wave is too close to the edge of the solar limb from the view of *SDO*/AIA, making the kinematics related measurements difficult to assess. Moreover, the wavelet-analysis approach is applied along the red dashed line L4, L5, and L6. The detailed results are shown in Figure 9.

Indeed, we apply the wavelet software (Torrence & Compo 1998) to analyze the periodicities of the intensity variations of the two QFP wave events along the six red dashed lines displayed in Figure 4 (L1, L2, L3) and Figure 8 (L4, L5, L6). The first event related results are shown in panels (a1)–(a3) of Figure 9. In the 171 Å channel, the period of WT-1 (WT-2) is about 40 ± 5 (60 ± 5) s. In 193 Å instead, the period of the WT-2 is estimated to be approximately 43 ± 8 s. It should be noted that the wavelet spectrum in panel (a1) of Figure 9 has a tadpole shape. According to Nakariakov et al. (2004) and Nakariakov & Verwichte (2005), the tadpole wavelet spectrum consists of a thin tail and a thick head. The authors indicated that the thin tail may be formed by the rapidly decreasing spectral dependence of the group speed. The thick head of the tadpole may be related to the dispersionless part of the group speed. The similar features of the tadpole wavelet were also detected in some radio sources (Mészárosová et al. 2009a, 2009b, 2011, 2013; Karlický et al. 2011). The dispersion evolution of the QFP WTs probably leads to the appearance of characteristic tadpole wavelet signatures (Nakariakov et al. 2004; Pascoe et al. 2013, 2014). However, the tadpole shape in panel (a1) of Figure 9 can be irregular probably because of the presence of a very strong noisy component of the signal.

The results of the second event, together with the corresponding estimated periods, are shown in Figure 9 (see panels (a4)–(a6)). The period of WT-1 is computed to be about 50 ± 10 s. The wavelet power spectra of the WT-2 detrended intensity profiles along L5 and L6 are displayed in Figures 9(a5) and (a6) from AIA 171 and 193 Å running-difference images, respectively. At the positions L5 and L6, strong powers with periods of 46 ± 9 and 49 ± 9 s are identified. In order to shed light on and emphasize the parameters of the two events, Table 1 reported the speeds of the EUV and QFP waves, respectively. The positions and the periods of the WTs are shown in the fourth and fifth columns, respectively. According to Table 1, the periods are close to or below about one minute.

4. Discussion and Conclusions

Exploring the high spatial and temporal resolution and multiangle observations taken by *SDO* and *STEREO*, we present two QFP events associated with two brightenings and

two EUV waves from 2011 March 9 to 10. Interestingly, one of the two studied events is found to be associated with an eruption of a flux rope. The active region NOAA 11167 appears to be a small region, hosting no large-scale magnetic fields nor intense flares; however, two QFP wave events are found to occur in the same location within the active region.

According to one of the previously mentioned animations, namely the animation of Figure 8, one can clearly recognize that the WT-2 wavefronts seemingly did not emanate from the kernel of the accompanying brightening. The WT-2 appeared to emanate from the flux rope and the FL dark features. The direction of the WT-2 is apparently subject of deflection from the animation. This deflection of the wavefronts is probably due to the refraction effect owing to the changes in the magnetic strength. In fact, the flux rope or FL structure eruption can alter the magnetic structure. The magnetic field strength of the flux rope is stronger than that of the quiet-Sun region. Shen et al. (2013a) and Miao et al. (2019) reported the refraction effect about the EUV wave that was similar to the WT-2 of the second QFP wave event. Yuan et al. (2015a) indicated that some ripples are formed by dispersion of the main EUV wavefront. The authors, through numerical simulation modeling, provided a new way to detect the filamentations of the solar atmosphere. Accordingly, the FL dark features and the flux rope can be influenced by the EUV wave. It is important to notice here that when the main EUV front interacted with the FL dark features and the flux rope, the ripples of the WT-2 appeared in the bottom left of the flux rope. Hence, the WT-2 is probably not a real “WT” as that originating from the FL dark and the flux rope.

In concluding, scrutinizing these atypical events reveals indeed several interesting characteristics and findings that can be summarized as follows:

(1) The two QFP events were observed to be related to two WTs. We also report the presence of two brightenings and two EUV waves. Additionally, the second QFP event was associated with a flux rope eruption, which consequently influenced the second WT (WT-2). A subsequent refraction effect can be clearly recognized from the accompanying animation (Figures 4 and 8, based of 171 and 193 Å running-difference images), reminiscent of what was reported in some previous investigations (Shen et al. 2013a; Miao et al. 2019).

(2) We report an interesting phenomenon namely that the eruption of the flux rope in the second event produced a new FL dark feature (see FL2 in Figure 7). The flux rope and the FL features can change the propagation path and speed of the waves (one can see the eruption of the EUV wave in the second event in Figure 6). Due to the low cadence of the *STEREO* observations, we only detected the EUV waves during the two

events. Hence, the flux rope or filament eruption probably altered the magnetic and thermal structure (local Alfvén speed).

(3) In accordance with some previous studies, the QFP WTs are found to be easily detected in the 171 Å channel running-difference images. We argue that the propagation of the QFP WTs possibly have a tight relationship with the height of the funnel-like coronal loops. This finding reinforces the recent results by Miao et al. (2019). Indeed, the authors have indicated that their studied QFP WTs phenomena were inclined to propagate in the lower corona and that they may be associated with the height of the funnel-like coronal loops. The authors also discussed some exciting emerging ideas of the possible mechanisms of the EUV and QFP waves, although a clear picture is still not fully drawn.

(4) The first weak WT (WT-2 in the first event) is found to be slower than the second weak WT (WT-2 in the second event). The speeds of the WT-2 in the first (second) event are in the range of 300–404 (687–729) km s⁻¹. The difference in the speeds between the first WT-2 and the second WT-2 is probably caused by the different driving mechanisms. The WT-2 of the second event possibly was not a real “WT.” According to Yuan et al. (2015a), the ripples can be formed by dispersion of the main EUV wavefront. In addition, the “WTs” of the WT-2 in the second event did not emanate from the center of the brightening within the active region. In the second event, the eruption of the flux rope and the FL features most likely changed the strength and the configuration of the magnetic field, considering the strengths of the magnetic fields and magnetic configurations can change the direction of propagation (Liu et al. 2012; Shen et al. 2013a, 2018a; Ofman & Liu 2018; Miao et al. 2019). The magnetic field strength of the second event is presumably stronger compared to that of the first event.

(5) The periods of the two weak WTs were detected with some differences in 171 and 193 Å running-difference images, respectively. The period of the first WT-2 in 171 (193) Å is 60 ± 5 (43 ± 8) s, whereas the period of the second WT-2 in 171 (193) Å is 46 ± 9 (49 ± 9) s, respectively. These curious differences in the measured periods between the 171 and 193 Å channels are not completely clear, nevertheless we propose that is seemingly due to the strengths of the corresponding magnetic fields, magnetic configurations and even some temperature effects.

(6) It is worth noting that details about the event-associated flares are absent, probably because the brightenings were too weak to be detected. From the observation by *STEREO*, due to the low temporal resolution, only one wavefront of the initial two EUV waves were detected, similar to the scenario presented in Miao et al. (2019) in which the authors have proposed that a QFP wave can be excited by CMEs. They considered that the QFP wave was produced by CMEs as the piston-driven shock wave interacts with funnel-like coronal loops. Pascoe et al. (2013) and Nisticò et al. (2014), through simulation and observational analyses, support these results. The nature of the QFP wave events is still unrevealed. We need more comprehensive and precise data to discern any relationship between the QFP waves and CMEs (flares) in the future.

In summary, our present investigation reported interesting phenomena from the two QFP events associated with two brightenings and two EUV waves. One event is found to be associated with the eruption of the flux rope. Noteworthy, the eruption of the flux rope led to the strength of the magnetic

field and changed the propagation direction of the WT-2. However, the real configurations of the QFP and EUV waves are clearly still not well understood. To better probe the nature of the QFP and EUV waves more observational inspections are certainly required. These events also probably provide a new example to study the QFP and EUV wave events. In addition, our results suggest that the funnel-like coronal loops may play a relevant role in giving rise to the QFP waves. However, the “WTs” of the so-called “QFP wave” may be caused by different mechanisms. The formation of the ripples of the waves are apparently associated with various factors. Hence, the coexistence of the two events offers a significant opportunity not only to reveal what possible factors and mechanisms are tightly related to the QFP waves, but also to provide us with a novel way to study the relationship between the QFP and the EUV waves. Undoubtedly, with more observationally detected and analyzed similar events, we believe that the nature of this peculiar class of waves with their related physical mechanisms will attract more attention within the solar scientific community and hopefully will be better understood in the near future.

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