



High-resolution Spectroscopic Study of Dwarf Stars in the Northern Sky: Lithium, Carbon, and Oxygen Abundances

E. Stonkutė¹, Y. Chorniy¹, G. Tautvaišienė¹, A. Drazdauskas¹, R. Minkevičiūtė¹, Š. Mikolaitis¹, H. Kjeldsen^{1,2},
 C. von Essen^{1,2}, E. Pakštienė¹, and V. Bagdonas¹

¹ Astronomical Observatory, Institute of Theoretical Physics and Astronomy, Vilnius University, Sauletekio av. 3, 10257 Vilnius, Lithuania; edita.stonkute@tfai.vu.lt

² Stellar Astrophysics Centre, Department of Physics and Astronomy, Aarhus University, Ny Munkegade 120, DK-8000 Aarhus C, Denmark

Received 2019 December 10; revised 2020 January 3; accepted 2020 January 4; published 2020 February 6

Abstract

Abundances of lithium, carbon, and oxygen have been derived using spectral synthesis for a sample of 249 bright F, G, and K Northern Hemisphere dwarf stars from the high-resolution spectra acquired with the VUES spectrograph at the Molėtai Astronomical Observatory of Vilnius University. The sample stars have metallicities, effective temperatures, and ages between $(-0.7 \div 0.4)$ dex, $(5000 \div 6900)$ K, $(1 \div 12)$ Gyr, accordingly. We confirm a so far unexplained lithium abundance decrease at supersolar metallicities— $A(\text{Li})$ in our sample stars, which drop by 0.7 dex in the $[\text{Fe}/\text{H}]$ range from $+0.10$ to $+0.55$ dex. Furthermore, we identified stars with similar ages, atmospheric parameters, and rotational velocities, but with significantly different lithium abundances, which suggests that additional specific evolutionary factors should be taken into account while interpreting the stellar lithium content. Nine stars with predominantly supersolar metallicities, i.e., about 12% among 78 stars with C and O abundances determined, have the C/O number ratios larger than 0.65, thus may form carbon-rich rocky planets. Ten planet-hosting stars, available in our sample, do not show a discernible difference from the stars with no planets detected regarding their lithium content.

Unified Astronomy Thesaurus concepts: Chemical abundances (224); Dwarf stars (2053); Solar neighborhood (1509); High resolution spectroscopy (2096)

Supporting material: machine-readable table

1. Introduction

In this work, we focus on a homogeneous abundance determination of the three important chemical elements—lithium, carbon, and oxygen—in a sample of dwarf stars and analyze the results in the light of extrasolar planet search surveys (e.g., NASA’s Transiting Exoplanet Survey Satellite (*TESS*; Ricker et al. 2015) and the upcoming ESA’s PLANetary Transits and Oscillations of stars (*PLATO*; Rauer et al. 2014) which will focus on the bright targets, of which only a third has any spectroscopic studies.

The complex pattern of lithium abundance— $A(\text{Li})$, observed in different types of stars in the Milky Way is a subject of ongoing discussions. ${}^7\text{Li}$ has three different nucleosynthesis sites—Galactic cosmic rays, stars, and primordial nucleosynthesis. Lithium is one of four elements synthesized in the primordial nucleosynthesis with the initial value of $A(\text{Li}) = 2.72$ dex. This value is predicted from the Big Bang nucleosynthesis models based on the results from the *Wilkinson Microwave Anisotropy Probe* (Cyburt et al. 2008) and is approximately three to four times higher than the observationally determined value from halo dwarfs (Spite & Spite 1982). Stars are producing the majority of ${}^7\text{Li}$ in the Galaxy. It is still unclear whether the sources are red giants, asymptotic giant branch stars, novae, or core-collapse supernova. Carbon and oxygen, on the other hand, were synthesized in the post Big Bang epoch. Both elements, however, were formed by different processes. Carbon acts as a catalyst in the CNO-cycle converting H to He. Carbon production occurs in stellar interiors, then it is dredged up from cores and mainly released into the interstellar medium by massive stellar winds, driven by radiation pressure from high-mass stars (e.g., Gustafsson et al. 1999). The low- and intermediate-mass stars contribute as well,

but their relative importance and yields are not well understood. Carbon is one of the required elements for life, as we know it, and plays an important role in searching for habitable exoplanets.

Oxygen is the third most common element in the universe and is produced by hydrostatic burning in massive stars and then essentially mostly dispersed by the Type II supernovae. The oxygen abundance is one of the key tracers of the formation and evolution of planets, stars, and galaxies. The abundances of oxygen in atmospheres of stars, differently from carbon, remain almost constant during the lifetime of stars and thus can be studied to trace back the Galactic chemical evolution.

Every new analysis of these chemical elements in stellar atmospheres contributes to the ongoing discussion of their importance in understanding the Galaxy evolution and whether they constitute a star–planet connection (see, e.g., Delgado Mena et al. 2010; González Hernández et al. 2010; Brugarner et al. 2011; Petigura & Marcy 2011; Schuler et al. 2011; Ramírez et al. 2012; Brewer et al. 2016; Spina et al. 2016; Luck 2017; Adibekyan et al. 2018; Bensby & Lind 2018; Fu et al. 2018; Luck 2018a; Nissen & Gustafsson 2018; Guiglion et al. 2019; Pavlenko et al. 2019).

For example, the sites of lithium production are identified, but their input for the lithium enrichment of the ISM is still debatable (Bensby & Lind 2018; Fu et al. 2018; Guiglion et al. 2019). The poorly understood mechanisms of ${}^7\text{Li}$ depletion and/or synthesis observed in F-, G-, and K-type stars make it difficult to constrain the Galactic chemical evolution models (Cescutti & Molaro 2019). Classical evolutionary models predict that the Li abundances in the main-sequence (MS) stars should depend uniquely on the stellar main parameters such as

the effective temperature, metallicity, age, or chromospheric activity. According to the work by Israelian et al. (2004), the T_{eff} (or mass) is the main parameter responsible for the Li depletion in solar-type stars. The second depletion driving parameter is the stellar age, while the third parameter could be metallicity and/or rotation. Classical chemical evolution models show a smooth increase of lithium in relation to metallicity (Romano et al. 2001; Prantzos 2012). Those models have some limitations and neglect several key physical processes such as gravitational settling, thermal diffusion, and radiative acceleration, rotational mixing, mass loss—any of which could be critical for the interpretation of ^7Li abundances in solar-type stars (Fu et al. 2015).

The same stars that produce ^7Li , destroy it as well. Lithium nuclei are burned through the proton capture when they are exposed to temperatures 2.5×10^6 K. ^7Li depletion observed in the Sun is inconsistent with the classical models. In order to explain the observations, lithium must be transported from the convection zone to the hot layers where the temperature is more than 2.5×10^6 K (Israelian et al. 2004).

From observational data we also see that ^7Li abundances in stars above the solar metallicity show lower values than those at the solar metallicity, suggesting a puzzling decrease of lithium. Recent high-resolution spectroscopic surveys of the Milky Way disk have reported that lithium abundances in the solar neighborhood decrease at supersolar metallicities. Delgado Mena et al. (2015) were some of the first to show that the ^7Li abundance decreases with $[\text{Fe}/\text{H}] > 0$ and later works by Guiglion et al. (2016) in the AMBRE project, Fu et al. (2018) in the *Gaia*-ESO Survey, and Bensby & Lind (2018) from combined data of various spectrographs (e.g., FEROS and MIKE) confirmed this result. The studies by Guiglion et al. (2016) and Prantzos et al. (2017) tried to explain the ^7Li decreases at supersolar metallicities with stellar migration. The behavior of ^7Li observed in high-resolution spectroscopic surveys can be explained by the interplay of mixed populations that originate in the inner regions of the Galaxy disk. The recent work by Guiglion et al. (2019) claims that stars have lower ^7Li content as a consequence of inside-out disk formation and radial migration.

Other mechanisms such as the presence of planets have been proposed to be responsible for additional ^7Li depletion. The work of King et al. (1997) was the first to suggest a connection between ^7Li depletion and planet hosts after finding a difference in ^7Li abundance for the very similar stars of the double system 16 Cyg, in which one of the stars is known to be a host to a Jupiter sized planet.

Israelian et al. (2004) and Chen & Zhao (2006) have suggested that Li depletion could be a result of planet migration, creating instability that produces effective mixing. The protoplanetary disk retains a large amount of angular momentum and therefore creates some rotational breaking in the host stars during the pre-MS phase inducing an increased mixing. Théado & Vauclair (2012) suggested that the accretion of metal-rich planetary material onto a star in its early phases could induce the thermohaline convection below the convective zone and lead to the extra ^7Li depletion in stars.

The discussion on ^7Li depletion on exoplanets hosts was revived by Figueira et al. (2014) and Delgado Mena et al. (2015), who concluded that exoplanet hosts show a significant ^7Li depletion when compared to hosts without planets. Later work by Bensby & Lind (2018) found that there is no

difference in ^7Li abundances in stars with detected planets compared to those with no (at the moment) detected planets.

In the case of the evolution of carbon and oxygen abundances in the Galaxy, they can be used to set constraints on stellar nucleosynthesis and to help understand the formation and evolution of the Milky Way. The information about the origins and evolution of carbon and oxygen may be obtained from differences in the elemental ratios $[\text{C}/\text{Fe}]$, $[\text{O}/\text{Fe}]$, and $[\text{C}/\text{O}]$ when looking at two stellar populations—thin and thick disk. The study by Reddy et al. (2006) has found evidence of a systematic difference in $[\text{C}/\text{Fe}]$ between thin- and thick-disk stars. This difference was not confirmed by Bensby & Feltzing (2006). Furthermore, in the work by Cescutti et al. (2009) it was shown that there is a clear distinction between the trends of $[\text{C}/\text{O}]$ in the thin and thick Galactic disk comparing the results from observational data with the predictions of Galactic chemical evolution models. Cescutti et al. (2009) suggest that the differences in $[\text{C}/\text{O}]$ versus $[\text{O}/\text{H}]$ for the two disks show that the thick disk is not made from the thin-disk material by dynamical heating of the thin disk, or that it was not possible to make the lower metallicity thin disk from the thick-disk material.

The past decade induced the theoretical and observational studies of carbon and oxygen abundances in stars in the context of determining a composition of extrasolar terrestrial planets, assuming that the composition of the host star and its protoplanetary disk is interrelated (see, e.g., Bond et al. 2010; Petigura & Marcy 2011; Nissen et al. 2014; Delgado Mena et al. 2015; Madhusudhan et al. 2016; Bedell et al. 2018, and references therein). C and O and their elemental number ratios C/O^3 could also be used to derive information related to the star–planet connection as in the case of lithium depletion. The proportions of carbon, oxygen, and rock-forming elements like magnesium, silicon, or iron determines the structure and composition of planets. The C/O elemental number ratio controls the amount of carbides and silicates formed in planets (Larimer 1975).

In the theoretical work by Bond et al. (2010) the authors concluded that C/O and Mg/Si elemental ratios are important in determining the mineralogy of extrasolar terrestrial planets. The work shows that if the C/O ratio is greater than 0.8 (under the assumption of equilibrium), Si exists in a solid form primarily as carbide, and is a factor of two higher than the solar ratio, $\text{C}/\text{O}_\odot \simeq 0.55$ (Caffau et al. 2010). On the other hand, if this ratio of C/O is below the 0.8 value, Si will form silicates (SiO etc.) and be present in rock-forming minerals. This has led to suggestions that there should exist exoplanets consisting of carbides and graphite instead of Earth-like silicates (Bond et al. 2010).

However, as a result of the chemical properties of C the gaseous C/O number ratio in planets can vary from the stellar value depending on different parameters (temperature, pressure, etc.) and processes during planet formation, including the initial location of formation of the planetary embryos, the migration path of the planet, and the evolution of the gas phase of a protoplanetary disk (Thiabaud et al. 2015a, 2015b; Madhusudhan et al. 2016).

The recent works by Suárez-Andrés et al. (2018) and Bedell et al. (2018) aimed to determine the mineralogy of planetary

³ C/O is defined as $N_{\text{C}}/N_{\text{O}}$, where N_{C} and N_{O} are the number densities of carbon and oxygen nuclei, respectively. $[\text{C}/\text{O}]$ is the solar-normalized logarithmic ratio.

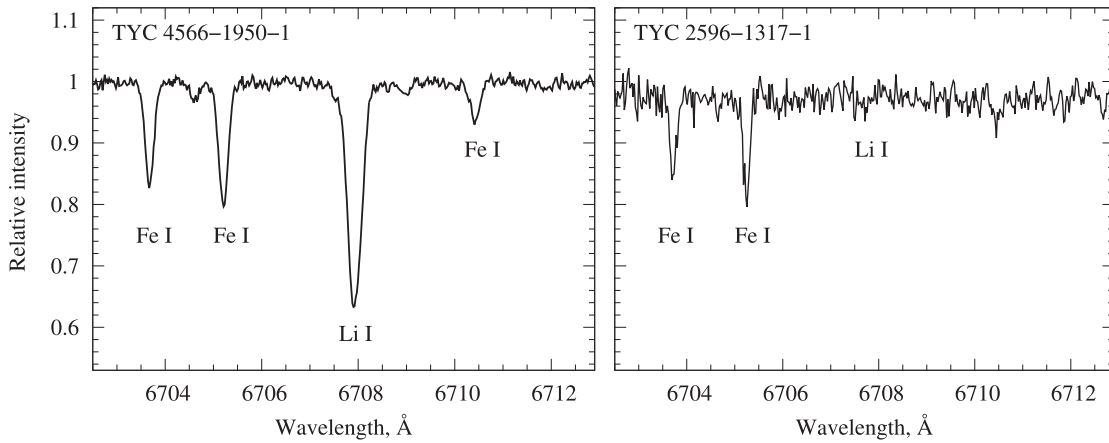


Figure 1. Examples of observed spectra for stars TYC 4566-1950-1 with the notable lithium line at 6708 Å and TYC 2596-1317-1 with the weak lithium line. In the latter and other similar cases only $A(\text{Li})$ upper limits had been determined.

companions looking at C/O number ratios of planet hosts. In the work by Suárez-Andrés et al. (2018), the determined C/O ratios revealed different kinds of planetary systems that can be formed, most of them unlike the solar system. They found that 100% of their sample of stars with detected planets has $\text{C/O} < 0.8$. Meanwhile, in a spectroscopic study of Sun-like stars, Bedell et al. (2018) found that the ratios of C/O in solar metallicity stars are homogeneous to within 10% throughout the solar neighborhood, implying that exoplanets may exhibit much less diversity in their composition than previously thought. The recent work by Pavlenko et al. (2019) determined carbon and oxygen abundances in atmospheres of the super-solar metallicity stars with and without detected planets and confirmed that metal-rich dwarf stars with planets are more carbon-rich in comparison with nonplanet host stars, with a statistical significance of 96%.

In this work, we pay special attention to determine the lithium, carbon, and oxygen abundances in a sample of dwarf stars, mainly of spectral classes F, G, and K, in the solar neighborhood with the aim to provide new key insights in the era of exoplanet surveys. Here we address some of the problems mentioned above by homogeneously determining precise abundances of Li, C, and O and analyzing links between the lithium abundance and stellar parameters (e.g., age, T_{eff} , metallicity, and rotation) in dwarf stars. The choice of using only dwarf stars is due to the fact that their atmospheres should present the original chemical composition of their birth places and carbon abundances in the atmospheres of dwarf stars are not affected by internal mixing of material comparing to giants.

2. Stellar Sample and Analysis

2.1. Observational Data Set

This work is based on spectroscopic data from the Spectroscopic and Photometric Survey of the Northern Sky (SPFOT) project Mikolaitis et al. (2018, 2019). The SPFOT survey aims to provide the detailed chemical composition and the main spectroscopic parameters (e.g., T_{eff} , $\log g$, and $[\text{Fe}/\text{H}]$) for bright stars ($V < 8$ mag) in the northern sky. Spectral observations for this project were carried out with the 1.65 m telescope at the Molėtai astronomical observatory of Vilnius University in Lithuania and the high-resolution Vilnius

University Echelle Spectrograph (Jurgenson et al. 2016) covering a wavelength range from 4000 to 9000 Å. From this survey, we took a data set consisting of 249 F, G, and K spectral type dwarf and subgiant stars observed in 2017–2018. The reader is directed to the papers by Mikolaitis et al. (2018, 2019) for a more detailed description of the observations, data reduction, determination of the main atmospheric parameters, stellar ages, and detailed abundances for 17 elements (Na, Mg, Al, Si, S, K, Ca, Sc, Ti, V, Cr, Mn, Fe, Co, Ni, Cu, and Zn). Like in the previous papers, we used a spectral synthesis method and the code TURBOSPECTRUM (Alvarez & Plez 1998) with the MARCS stellar atmospheric models (Gustafsson et al. 2008). Unlike in the previous papers where abundances of chemical elements were determined using an automated software, the lithium, carbon, and oxygen abundances we determined visually inspecting the fits for each star. The lithium abundance $A(\text{Li})$ was determined from profiles of the Li I 6708 Å line (if this line was too weak or too noisy, only the upper limit $A(\text{Li})$ was determined), see Figure 1. The fits for some Li-rich stars are provided in Figure 2 where the best fit was determined by eye and represented by the (red) middle solid line whereas the other two solid lines indicate ± 0.1 dex.

The nonlocal thermodynamic equilibrium (NLTE) effects are negligible for the majority of stars in our sample and have not been taken into account, as for stars with larger lithium abundances they would be lower than 0.1 dex, thus the results are valid within the assumptions of 1D model atmospheres in LTE and hydrostatic equilibrium (see, e.g., Lind et al. 2009).

For the carbon abundance determination we used two regions: C_2 Swan (1, 0) band head at 5135 Å and C_2 Swan (0, 1) band head at 5635 Å. The oxygen abundance was determined from the forbidden [O I] line at 6300 Å (Figure 3). The oscillator strength values for ^{58}Ni and ^{60}Ni , which blend the oxygen line, were taken from Johansson et al. (2003). The NLTE effects for C_2 Swan bands have not been investigated in detail; however, it is argued that NLTE effects should not affect the abundances significantly since the forbidden [C I] line and C_2 lines give the same carbon abundances (Gustafsson et al. 1999). On the other hand, the forbidden oxygen line at 6300.3 Å has been studied extensively, and it is concluded that this line can be perfectly described in LTE (Asplund 2005).

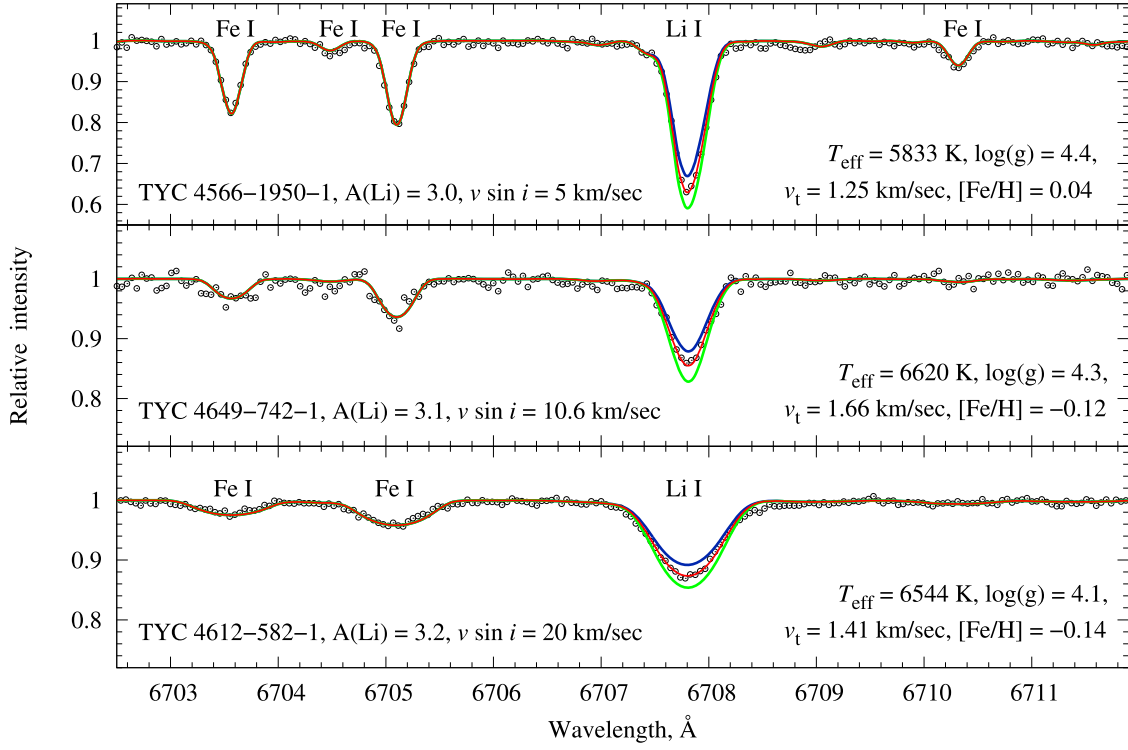


Figure 2. Examples of the observed and synthetic spectra around the lithium 6708 Å line with the lithium abundance $A(\text{Li}) \sim 3$ dex. The red solid line shows the best fit while the green and blue solid lines indicate ± 0.1 dex. The atmospheric parameters and values of the rotational velocities $v \sin i$ (km s^{-1}) are also indicated.

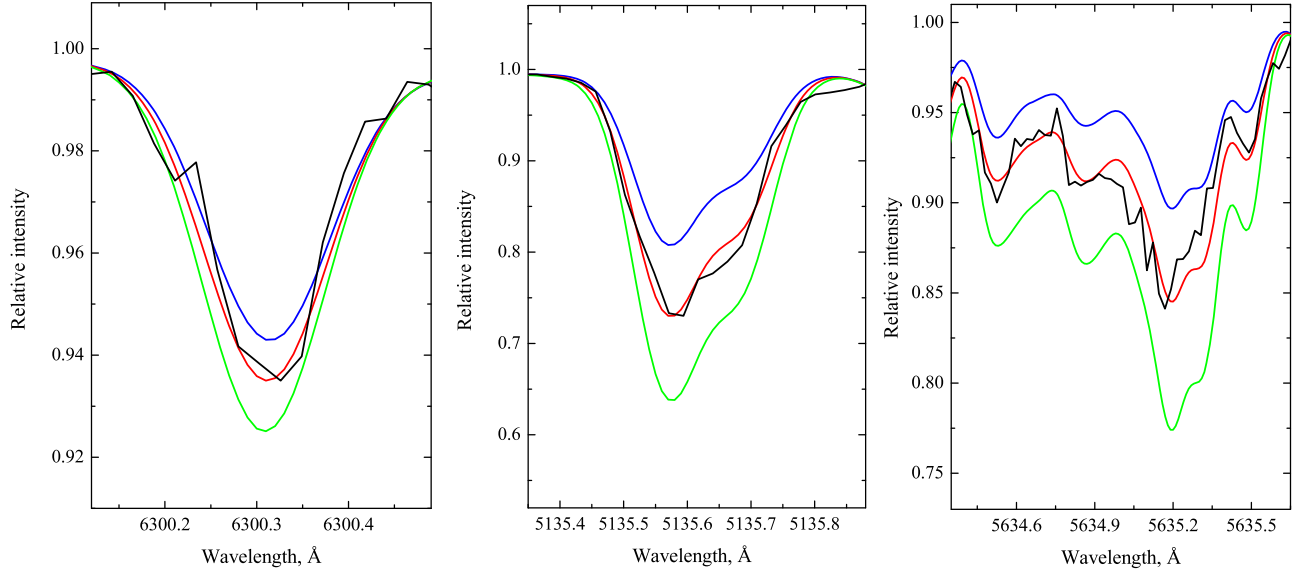


Figure 3. Examples of the synthetic spectrum fits to the forbidden [O I] line at 6300 Å and to the C_2 Swan (1, 0) band head at 5135 Å, and to the C_2 Swan (0, 1) band head at 5635 Å in the star TYC 4493-620-1. The observed spectra of the star TYC 4493-620-1 are shown as the black solid lines. The red solid lines represent the best fit with $[\text{C}/\text{Fe}] = 0.05$ dex and $[\text{O}/\text{Fe}] = 0.02$ dex, while the blue and green solid lines represent a change in abundance by ± 0.1 dex to the corresponding elements.

All the synthetic spectra have been calibrated to the solar spectrum by Kurucz (2005) with $\log A_{\odot}(\text{Li}) = 1.05$, $\log A_{\odot}(\text{C}) = 8.39$, and $\log A_{\odot}(\text{O}) = 8.66$ taken from Grevesse et al. (2007). Several examples of the synthetic spectra fits for C and O lines are presented in Figure 3.

2.2. Uncertainties

There are two sources of errors in abundances determined in this work: (1) uncertainties caused by analysis of individual

lines, including random errors of atomic data or continuum placement and signal-to-noise (S/N) ratio, that affect a single line, and (2) uncertainties in the stellar parameters, that affect all lines together. Uncertainties coming from the atomic data are mostly eliminated because of the differential analysis relative to the Sun. The calculated medians of atmospheric parameter determination errors from all the stars in our sample are presented in Table 1. In this table medians of Li, C, and O abundance determination errors are introduced as well.

Table 1
The Median Values of Stellar Atmospheric Parameters and Their Errors

$T_{\text{eff}} \mid \sigma$ (K)	$\log g \mid \sigma$	[Fe/H] $\mid \sigma$	$v_t \mid \sigma$ (km s ⁻¹)	A(Li) $\mid \sigma$	[C/H] $\mid \sigma$	[O/H] $\mid \sigma$
6064 $\mid \pm 46$	4.22 $\mid \pm 0.30$	-0.10 $\mid \pm 0.11$	1.04 $\mid \pm 0.27$	2.44 $\mid \pm 0.04$	-0.03 $\mid \pm 0.02$	0.07 $\mid \pm 0.14$

We calculated changes in abundances caused by the error of each individual atmospheric parameter, keeping other parameters fixed. Results are presented in Table 2. As can be seen, the abundances are not very sensitive to the changes of atmospheric parameters, only the oxygen abundance sensitivity to the surface gravity is more visible.

Carbon and oxygen are bound together by the molecular equilibrium, therefore we measured abundances of these two elements in unison, even though carbon and oxygen abundances in dwarf stars are not very sensitive to each other. Exempli gratia in TYC 4349-205-1, $\Delta[\text{C}/\text{H}] = \pm 0.10$ causes $\Delta[\text{O}/\text{H}] = 0.00$ and $\Delta\text{C}/\text{O} = \pm 0.10$; $\Delta[\text{O}/\text{H}] = \pm 0.10$ causes $\Delta[\text{C}/\text{H}] = \pm 0.03$ and $\Delta\text{C}/\text{O} = \pm 0.06$. Thus changes in one element influence abundances in the other very little.

2.3. Comparison with Other Studies

Figures 4 and 5 show comparisons between the Li, C, and O abundances determined in this work and from two other studies: Luck (2017) and Ramírez et al. (2012) have 44 and 22 stars in common with our sample for ⁷Li, respectively; and Luck (2017) have 21 stars in common with our sample for the carbon and oxygen abundances. The points are color-coded by the difference of the determined effective temperatures for the same stars. The agreements between the studies are good, on average our A(Li) abundances are about 0.04 dex higher than the Luck (2017) and Ramírez et al. (2012) values, with a 1 σ dispersion of 0.25 dex. The higher Li values agree better, as the stronger lines provide a more accurate determination, whereas smaller Li values are usually determined from lower S/N spectra and weaker lines which result in only an upper limit determination or larger uncertainties in the values. Those studies use a different technique for the parameter and abundance determinations from the method used in our work. The temperatures in the works by Luck (2017) and Ramírez et al. (2012) were derived using the Casagrande et al. (2010) effective temperature calibration. It is encouraging that the effective temperatures in our and other works are in good agreement: comparing with Luck (2017) we obtain on average 20 K higher temperatures with a mean standard deviation of 70 K, and comparing with Ramírez et al. (2012) we get on average 16 K higher temperatures with a 50 K mean standard deviation.

Our [C/H] and [O/H] abundance values are on average about 0.08 and 0.04 dex lower with dispersions of 0.09 and 0.18 dex, respectively, than the Luck (2017) values for a sample of 21 common stars. We would like to note that for the comparison with Luck's (2017) work we use [O I] and C₂ data per line found in their Table 5. Their Table 5 contains absolute values for carbon abundance from the C₂ Swan line primary indicator at 5135 Å and oxygen abundances from the forbidden line at 6300 Å, the same lines that we used for the carbon and oxygen abundance determinations. We decided to scale their absolute values for C₂, oxygen and iron (Fe I) using the solar values from Grevesse et al. (2007) as in our work to be on the same elemental abundance scale.

Table 2
Effects on the Derived Chemical Abundances Resulting from the Atmospheric Parameter Uncertainties for the Program Stars

Species	ΔT_{eff} ± 46 K	$\Delta \log g$ ± 0.30 dex	$\Delta[\text{Fe}/\text{H}]$ ± 0.11 dex	Δv_t ± 0.27 km s ⁻¹
A(Li)	± 0.05	± 0.00	∓ 0.00	± 0.01
C (C ₂)	∓ 0.02	± 0.01	∓ 0.00	± 0.00
O [O I]	± 0.02	± 0.14	± 0.01	± 0.01

There are several other studies on carbon and oxygen abundance determinations with which we have fewer stars in common. For example, in the work by Petigura & Marcy (2011) we find eight stars in common and our [C/H] abundance values are on average 0.05 ± 0.10 dex lower than theirs. However, we see no difference for [O/H] abundance determination between the two studies.

3. Results and Discussion

Table 3 shows an example of the online table containing the resulting lithium, carbon, and oxygen abundances, C/O number ratios as well as atmospheric parameters, whether the star belongs to the thin or thick disk, and other information on the stars in our analyzed sample. For column names and details see the Appendix.

Our stars have metallicities, effective temperatures, and ages between $(-0.7 \div 0.4)$ dex, $(5000 \div 6900)$ K, $(1 \div 12)$ Gyr, accordingly. Accurately measurable lithium lines were found in spectra of 149 stars, the upper limits were determined for the other 100 stars (60% and 40% of the sample, respectively). The largest fraction of stars has $A(\text{Li}) \geq 2.0$ dex and makes up $\approx 51\%$ of the sample. Carbon and oxygen abundances were determined for 78 stars.

3.1. Lithium Abundances versus Effective Temperature

In Figure 6 we show the distribution of lithium abundance A(Li) as a function of effective temperature T_{eff} . Stars from our sample are color-coded by their ages. The lithium abundance versus effective temperature plane could be divided into six regions of interest: (1) the region of lithium-rich stars with $A(\text{Li}) \geq 2.0$; (2) the area of Li-depletion at $5500 \text{ K} \leq T_{\text{eff}} \leq 5950 \text{ K}$; (3) the region of the lithium-poor stars at $T_{\text{eff}} > 6000 \text{ K}$; (4) the area of the lowest values $A(\text{Li}) \leq 1.0$ at $T_{\text{eff}} < 5500 \text{ K}$; (5) the region of Li-abundant chromospherically active stars with $5200 \text{ K} < T_{\text{eff}} < 5700 \text{ K}$; and (6) the lithium desert with $A(\text{Li}) \approx 1.8$ between $5950 \text{ K} \div 6100 \text{ K}$. For the comparison in Figure 6 we plotted results from the work by Ramírez et al. (2012) as gray circles and triangles. This A(Li) versus T_{eff} plane has similar trends found in previous studies (see, e.g., Ramírez et al. 2012; López-Valdivia et al. 2015; Luck 2017; Bensby & Lind 2018):

(1) Lithium-rich dwarfs area. Stars with $A(\text{Li}) \geq 2.0$ are located in a broad ($\Delta A(\text{Li}) \simeq 1.0$) and a rather vague band that stops at $T_{\text{eff}} \simeq 5600 \text{ K}$. The highest values of lithium

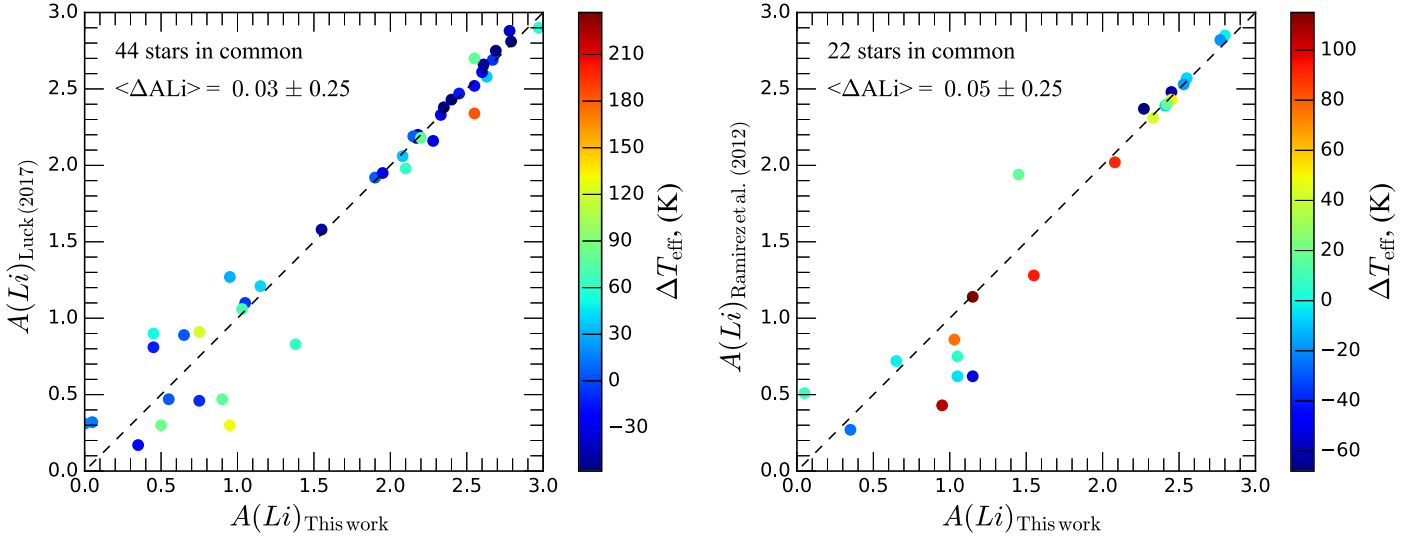


Figure 4. Comparison plots of $A(\text{Li})$ abundances from this work against those of Luck (2017) and Ramírez et al. (2012). The dashed line shows a 1:1 correlation. The average differences and standard deviations ($\Delta A(\text{Li})$ and ΔT_{eff}) are calculated as our values minus those of other authors. The stars have been color-coded according to a difference in the effective temperatures.

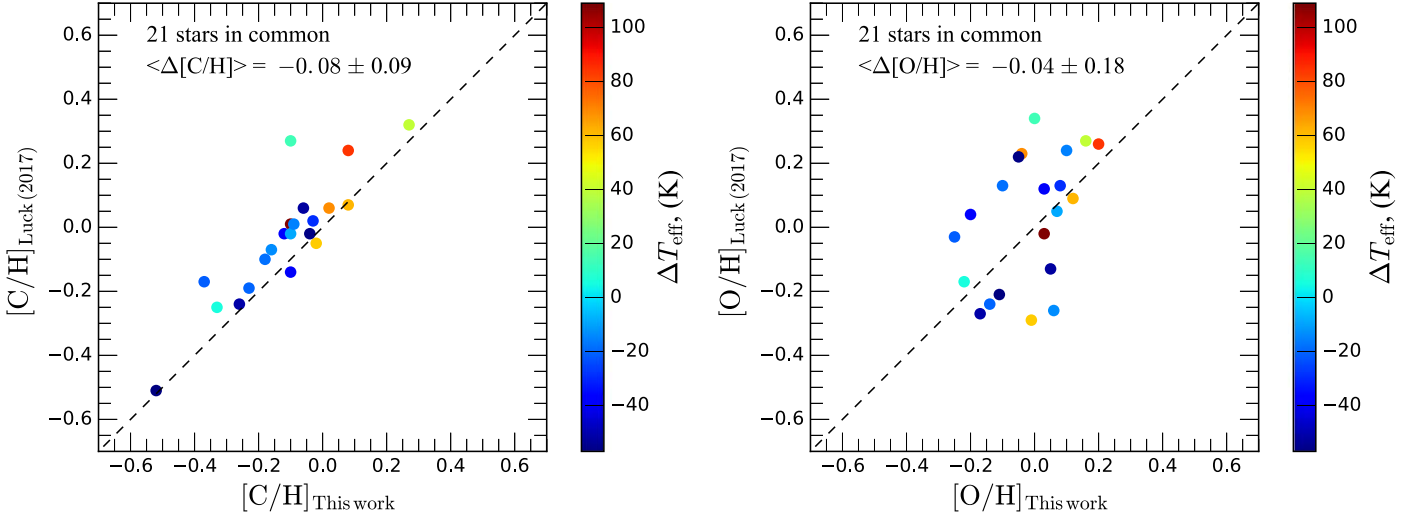


Figure 5. Comparison plots of $[\text{C}/\text{H}]$ and $[\text{O}/\text{H}]$ from this work against the work by Luck (2017). The dashed line shows a 1:1 correlation. The average differences and standard deviations ($\Delta [\text{Element}/\text{H}]$ and ΔT_{eff}) are calculated as our values minus the comparison values. The stars have been color-coded according to a difference in the effective temperatures.

Table 3
Abundances of Li, C, and O, Atmospheric and Other Parameters of the Sample Stars

Star ^a	$A(\text{Li})$ (dex)	$M U$	$v \sin i$ (km s^{-1})	$[\text{C}/\text{H}]$ (dex)	$e_{[\text{C}/\text{H}]}$ (dex)	n_{C}	$[\text{O}/\text{H}]$ (dex)	$e_{[\text{O}/\text{H}]}$ (dex)	n_{O}	C/O	Thin Thick	Planet	... ^a
TYC1547-367-1	0.86	<i>U</i>	6.0								0	0	...
TYC1563-3551-1	0.77	<i>U</i>	1.0								1	0	...
TYC1563-3552-1	2.87	<i>U</i>	11.0								0	0	...
TYC2057-709-1	1.98	<i>M</i>	1.2	-0.08	0.09	2	0.15	0.06	1	0.32	0	0	...
TYC2070-1061-1	2.49	<i>M</i>	2.0								0	0	...
...

Note.

^a For column names and details see the Appendix.

(This table is available in its entirety in machine-readable form.)

$A(\text{Li}) \geq 3$) were found for seven stars which are among the hottest and youngest objects in our sample with effective temperatures $T_{\text{eff}} \geq 6200$ K and ages ≤ 2.8 Gyr. A linear

regression fit to the area of $A(\text{Li}) \geq 2.0$ values shows that the slope per 1000 K equals to 0.4 ± 0.1 dex as indicated with the upper solid line in Figure 6. We find that this value is close to

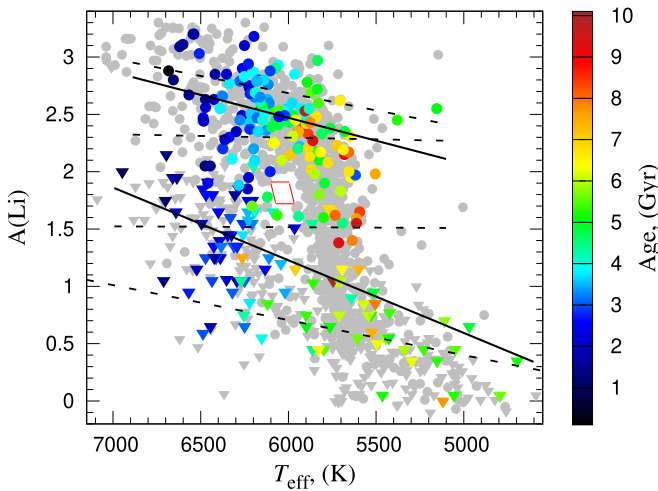


Figure 6. Abundance of lithium $A(\text{Li})$ as a function of effective temperature. The stars with measurable $A(\text{Li})$ and upper limits are shown as circles or triangles, respectively, and color-coded by their ages. For the comparison, we plotted results from Ramírez et al. (2012) as gray circles and triangles. The linear approximations of $A(\text{Li})$ as a function of T_{eff} are shown as the black solid lines for two groups of stars with $A(\text{Li}) \geq 2$ and $A(\text{Li}) < 2$. The dashed lines represent similar fits for the ranges $A(\text{Li}) \geq 2.5$; $2.0 \leq A(\text{Li}) < 2.5$; $1.0 \leq A(\text{Li}) < 2.0$, and $A(\text{Li}) < 1.0$. The quadrangle with the red perimeter shows an approximate location of the lithium desert near $T_{\text{eff}} \approx 6025$ K and $A(\text{Li}) \approx 1.8$.

the value of 0.45 ± 0.1 dex/1000 K found in the work of local metal-rich stars by Pavlenko et al. (2018).

Next, we divided the stars with $A(\text{Li}) \geq 2.0$ into two bins with abundances $2.0 \leq A(\text{Li}) < 2.5$ and $A(\text{Li}) \geq 2.5$ and determined gradients for both. We obtained that the slope per 1000 K for the $2.0 \leq A(\text{Li}) < 2.5$ bin is 0.01 ± 0.07 dex and for the $A(\text{Li}) \geq 2.5$ we got 0.28 ± 0.09 dex as indicated by upper dashed lines in Figure 6. Hence, the correlation between $A(\text{Li})$ and T_{eff} is stronger for the bin where $A(\text{Li}) \geq 2.5$ and is similar to the larger bin of $A(\text{Li}) \geq 2.0$ while it is almost flat or nonexistent for the $2.0 \leq A(\text{Li}) < 2.5$ bin.

(2) Lithium-depletion area. We note a region of strong lithium depletion toward lower T_{eff} in the relatively narrow range $5500 \text{ K} \leq T_{\text{eff}} \leq 5950 \text{ K}$ (or $M \sim (0.9 \div 1.1) M_{\odot}$). As is known, this depletion is due to a significant increase of convective depths of such stars since the deeper convection transports lithium to temperatures where it is destroyed. The mean age of these stars in our sample is 6.3 ± 1.5 Gyr.

(3) Lithium-poor stars with $T_{\text{eff}} > 6000$ K. Opposite to the mentioned area (2) we observe a fraction of dwarf stars with $A(\text{Li}) < 2$ with the effective temperature $T_{\text{eff}} > 6000$ K; however, the majority of dwarfs have only $A(\text{Li})$ upper limits and the mean age of these stars is equal to 2.9 ± 1.15 Gyr.

(4) Stars with $A(\text{Li}) \leq 1.0$ at $T_{\text{eff}} < 5500$ K. Moving toward the lower $T_{\text{eff}} < 5500$ K, we see a number of lithium-poor stars with $A(\text{Li}) \leq 1$. Some of stars have only upper limit values decreasing to $A(\text{Li}) \approx 0$. In this region ($A(\text{Li}) \leq 1.0$ at $T_{\text{eff}} < 5500$ K) lithium abundances reach a plateau that exists because the convection depth in such stars is not growing any more with decreasing temperatures (or masses; see, e.g., Luck 2017).

(5) Li-abundant chromospherically active stars with $5200 \text{ K} < T_{\text{eff}} < 5700$ K. This region was identified by Mishenina et al. (2012) while investigating a correlation between $A(\text{Li})$ and chromospheric activity of stars. We found two chromospherically active stars with $M \sim 1 M_{\odot}$, $[\text{Fe}/\text{H}] \sim -0.1$ and $A(\text{Li}) \sim 2.5$ in this temperature region, while

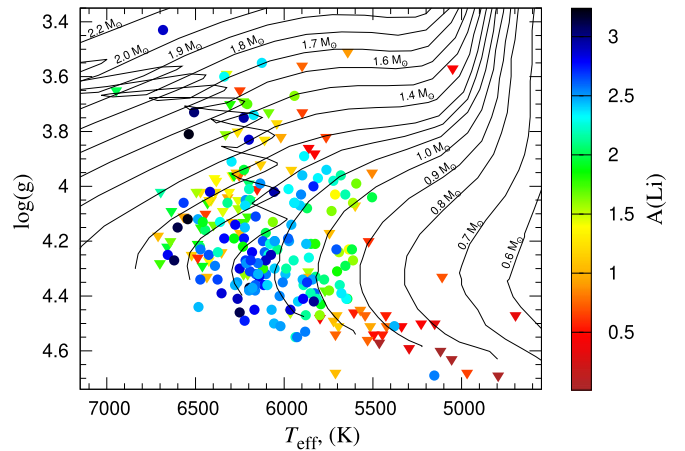


Figure 7. HR diagram in a surface gravity vs. effective temperature plane. Stars are color-coded according to their $A(\text{Li})$. The black lines show the evolutionary tracks, which were taken from the work by Girardi et al. 2000 with masses $(0.6 \div 2.2) M_{\odot}$ and the fixed $Z_{\text{ini}} = 0.019$.

other stars with similar atmospheric parameters had $A(\text{Li}) \sim 0.5$. Chromospheric activity of TYC-4566-120-1 (π^1 UMi B, HD 139813) was investigated by, e.g., Isaacson & Fischer (2010), Eisenbeiss et al. (2013, and references therein), and TYC-4609-535-1 (V368 Cep, HD 220140) was investigated by e.g., Zhang et al. (2015, and references therein). Thus, we confirm the conclusion that chromospherically active stars may have quite large lithium abundances in this effective temperature interval.

(6) Lithium-desert area. Lastly, the astonishing region of the so-called lithium-desert is clearly visible in the effective temperature range of $5950 \text{ K} \div 6100 \text{ K}$ with $A(\text{Li}) \approx 1.8$ (see the red quadrangle in Figure 6). According to Aguilera-Gómez et al. (2018), stars below the lithium-desert have evolved from the lithium dip. A dip-like feature appears in $A(\text{Li})$ versus T_{eff} plane around 6700 K, first seen in old clusters and less evident in field stars. These stars are more massive, more evolved, and have lower $A(\text{Li})$ abundances compared to other stars of the same temperature, which are located above the desert and have higher Li abundances.

Looking at the Li abundance correlation with effective temperature for stars with $A(\text{Li}) < 2$, which is indicated in Figure 6 with a lower solid line, we see that the overall slope is quite similar to the one for stars with larger Li abundances ($A(\text{Li}) \geq 2$). The $A(\text{Li})$ slope per 1000 K equals 0.6 ± 0.1 . A corresponding slope found by Pavlenko et al. (2018) is a bit larger and equals 0.85 ± 0.18 dex but has a larger error. We divided the stellar sample with $A(\text{Li}) < 2$ in two bins as well. The slopes for $1 \leq A(\text{Li}) < 2$ and $A(\text{Li}) < 1$ bins per 1000 K are equal to 0.01 ± 0.11 dex and 0.31 ± 0.07 dex, respectively. Hence, the correlation between $A(\text{Li})$ and T_{eff} is stronger for the $A(\text{Li}) < 1$ bin and almost nonexistent for the $1 \leq A(\text{Li}) < 2$ bin (indicated by two lower dashed lines in Figure 6).

In Figure 7, we show investigated stars in the temperature-gravity plane. The stars with measurable $A(\text{Li})$ and upper limits are shown as circles and triangles, respectively, and color-coded by their $A(\text{Li})$. The black lines indicate evolutionary tracks, which are taken from the work by Girardi et al. (2000) with masses $(0.6 \div 2.2) M_{\odot}$ and fixed $Z_{\text{ini}} = 0.019$. We see that the lowest $A(\text{Li})$ dwarfs are on the lower right part of the MS in the area of lowest T_{eff} of the sample, whereas the highest Li abundances are spread in the region of the MS with

Table 4
The Stars with Similar Main Parameters but Different Lithium Abundances

Star	T_{eff} (K)	$\log g$	[Fe/H]	v_t (km s ⁻¹)	$v \sin i$ (km s ⁻¹)	Age (Gyr)	$A(\text{Li})$
TYC4612-582-1	6544	4.1	-0.14	1.41	20	1.7	3.20
TYC3890-1396-1	6445	4.2	0.13	1.59	25	1.6	2.05
TYC4426-975-1	6486	4.1	-0.20	1.39	9	2.4	2.47
TYC4488-687-1	6459	4.1	-0.23	1.46	8	2.5	1.90
TYC3132-2016-1	6228	3.8	0.00	1.60	11	2.7	2.94
TYC4414-1663-1	6227	3.9	-0.11	1.83	14	2.2	1.95
TYC2661-155-1	6297	3.9	0.26	1.54	9	1.9	2.35
TYC2592-1547-1	6294	4.1	0.28	1.57	12	1.7	1.05
TYC4566-1950-1	5833	4.4	0.04	1.25	5.0	4.7	2.97
TYC2639-2460-1	5817	4.4	-0.21	0.77	3.0	6.3	2.17
TYC2613-2218-1	5794	4.4	0.10	0.79	0.1	4.9	1.90
TYC4409-1023-1	5706	4.2	0.07	0.83	2	6.6	2.62
TYC4366-1351-1	5691	4.2	0.14	0.83	1	4.5	1.55

highest T_{eff} . $A(\text{Li})$ is predominantly higher among hotter and more massive stars. This main trend is in agreement with results of previous studies (see, e.g., Ramírez et al. 2012; Luck 2017; Bensby & Lind 2018; Fu et al. 2018; Pavlenko et al. 2018).

However, there are stars with the same atmospheric parameters and ages but different lithium abundances. We show some examples in Table 4. Within uncertainties, those stars have similar ages, atmospheric parameters, and rotational velocities, while their values of $A(\text{Li})$ differ significantly. It could mean that in addition to the abovementioned main parameters (T_{eff} , mass, age, etc.) lithium abundances in the stellar atmospheres can be affected by external factors such as presence of planets, planet migration, or inside-out disk formation and stellar radial migration (see Chen & Zhao 2006; Prantzos et al. 2017; Guiglion et al. 2019, and references therein).

3.2. Lithium Abundances versus Metallicity

To study the Li depletion and the Li abundance in the ISM one has to look at the most Li abundant stars which reflect the initial Li abundances in the ISM (e.g., Lambert & Reddy 2004; Delgado Mena et al. 2015; Guiglion et al. 2016, 2019). Following similar analyses, we binned our data every 0.15 dex in [Fe/H], except for the boundary regions as there are too few stars, and we made the metallicity bins of 0.3 dex for both the most metal-poor (-0.80 to -0.50 dex) and the metal-rich side (from +0.25 to +0.55 dex). In order to be consistent with previously mentioned studies, we took the six most lithium rich stars in every bin. The error bars represent the standard deviation between those six stars. The relation is shown in Figure 8. When we try to trace the lithium abundance in the environment from which stars were formed, one has to keep in mind that lithium is easily destroyed in stellar interiors and on the other hand, some lithium can be produced as well. Thus, only dwarf stars with the largest Li abundances should be considered, as such stars have only burned lithium in their photospheres and do not contain any that would be freshly synthesized. In Figure 8, we see that considering only the largest values of the Li abundances in dwarf stars, we confirm that there is a decrease of lithium at supersolar metallicities which cannot be explained by the most recent models (see Romano et al. 2001; Prantzos et al. 2017; Fu et al. 2018, and discussions therein). This work shows that the closest and

brightest stars in a small volume of space in the solar neighborhood also exhibit this trend. The largest lithium abundance of 3.0 dex in our study is reached at the solar metallicity and drops by 0.69 dex in the range of [Fe/H] from +0.10 to +0.55 dex. Furthermore, stars in Figure 8 are color-coded by $v \sin i$ and we find some evidence that $A(\text{Li})$ seems to increase with increasing stellar rotation speed.

3.3. Carbon and Oxygen Abundances in the Thin and Thick Galactic Disk

We determined carbon and oxygen abundances for 78 observed stars including the Sun. The forbidden oxygen [O I] line at 6300 Å was severely blended by telluric lines in many spectra of our sample of stars. This was the main reason why the total number of stars with determined carbon and oxygen abundances was lower than that for lithium. Stars in our sample were divided into the thin and thick disks according their chemical composition (see Mikolaitis et al. 2019 for more details).

Figures 9 and 10 show the carbon and oxygen abundances versus metallicity. The thin-disk stars are marked as circles, the squares show the thick-disk stars. All the stars are color-coded by the C/O number ratio (note that $\text{C/O} \neq [\text{C/O}]$). As expected, both [C/Fe] and [O/Fe] decrease with time due to the increasing production of iron by SNe Ia. The carbon clearly follows iron more closely than oxygen. The data obtained in our work are in agreement with previous studies (e.g., Bensby & Feltzing 2006; Brewer et al. 2017; Luck 2017). In theory, oxygen is synthesized mostly in massive stars, while carbon is produced in stars of all masses. Thus, the ejection of some carbon is delayed in time with respect to oxygen, so the [C/O] offers a relative age clock for stellar systems (Tinsley 1979).

Knowing that oxygen is mostly produced in massive stars on a relatively short timescale (see Cescutti et al. 2009), [C/O] as a function of [O/H] show results of the carbon evolution. The change in carbon-to-oxygen as a function of oxygen-to-hydrogen strongly depends on the yields and timescales of carbon production in various types of stars. In Figure 11 we see a systematic difference between the thin and thick disks in the [C/O] versus [O/H] plane. The thick-disk stars lie on a different [C/O] sequence compared to the thin-disk stars which is shifted by ~ 0.3 dex to lower [C/O] values. To explain this, carbon produced in both low-mass and massive stars has to be included.

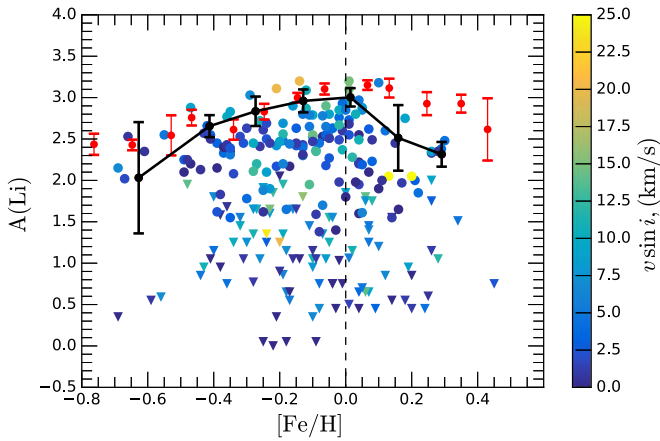


Figure 8. Abundances of lithium, $A(\text{Li})$ as a function of metallicity $[\text{Fe}/\text{H}]$. Our results (filled circles—measurable and triangles—upper limits) are color-coded according to their rotation $v \sin i$. The black dots indicate the mean lithium values for the six stars with the highest lithium abundance in their respective metallicity bins. The red circles show the results from the work by Guiglion et al. (2019).

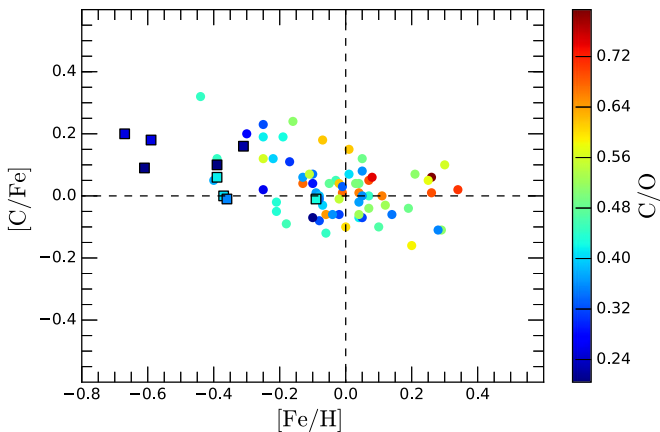


Figure 9. $[\text{C}/\text{Fe}]$ vs. $[\text{Fe}/\text{H}]$. Stars are color-coded by their C/O number ratio. The thin-disk stars are shown as circles, the thick-disk stars as squares.

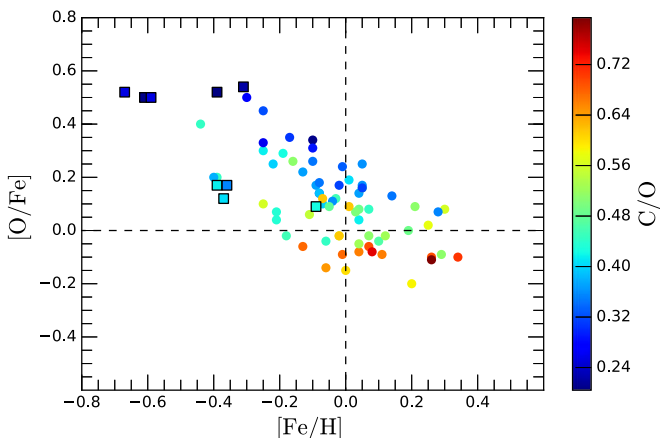


Figure 10. $[\text{O}/\text{Fe}]$ vs. $[\text{Fe}/\text{H}]$. The meaning of symbols are the same as in Figure 9.

The difference between thin- and thick-disk stars in the carbon-to-oxygen abundance ratio as a function of oxygen-to-hydrogen abundance ratio plane for dwarf stars has previously been demonstrated by Bensby & Feltzing (2006) using 35 thin- and 16 thick-disk stars and by Nissen et al. (2014) who

investigated 57 thin- and 25 thick-disk stars. Bensby & Feltzing (2006) used the forbidden carbon ($[\text{C I}]$ at 8727 Å) and oxygen ($[\text{O I}]$ at 6300 Å) lines. Nissen et al. (2014) investigated the high-excitation permitted C I lines at 5052 and 5380 Å, and O I at 7774 Å). The stars in the work by Bensby & Feltzing (2006) were assigned to populations based on statistical probability of their kinematics. Nissen et al. (2014) derived C and O from two data samples obtained using different spectrographs: HARPS, FEROS, UVES, and FIES, and stars were assigned to the separate disks by both the chemical and kinematical approach. Our results confirm the separation of thin- and thick-disk stars in the $[\text{C}/\text{O}]$ versus $[\text{O}/\text{H}]$ plane using the homogeneous analysis of C_2 bands at 5135 Å and 5635 Å and the $[\text{O I}]$ forbidden line at 6300 Å, and taking into account the molecular equilibrium of these elements.

In addition, we see that both the thin- and thick-disk dwarfs have no $[\text{C}/\text{O}]$ correlation with $A(\text{Li})$ in the $[\text{C}/\text{O}]$ versus $[\text{O}/\text{H}]$ plane (the right panel in Figure 11). To confirm this, we also checked the recent data by Luck (2017) and found the same.

3.4. Li, C, and O in Stars with Planets

Several previous studies have shown that stars with detected planets have different abundance patterns compared to those with so far no planets detected. The results, however, are still debatable (see, e.g., Israelian et al. 2009; da Silva et al. 2015; Aubé et al. 2014; Delgado Mena et al. 2014; Mishenina et al. 2016; Suárez-Andrés et al. 2017; Bensby & Lind 2018; Luck 2018b; Pavlenko et al. 2018; Carlos et al. 2019, and references therein).

We found that 10 stars in our sample have exoplanets detected. In Table 5 we collected data on the exoplanets identified around those stars (Akeson et al. 2013). For 5 out of 10 we have both carbon and oxygen abundance measurements and all 10 have lithium abundances determined. For convenience, in Table 5 we also present the determined Li abundances and C/O number ratios from our work as well as the Mg/Si number ratios from Mikolaitis et al. (2019), obtained from the same observational data and stellar atmospheric parameters as in our work.

In Figure 12, we show lithium abundances for the sample stars and mark with colored symbols the stars with confirmed exoplanets. Four out of 10 planet-hosting stars have confidently measured $A(\text{Li})$ values. Four stars with relatively high effective temperatures ($\langle T_{\text{eff}} \rangle = 6070 \pm 200$ K) with confirmed exoplanets have an average $\langle A(\text{Li}) \rangle \sim 2.4$. The remaining six stars have much lower $A(\text{Li})$ abundances ($\langle A(\text{Li}) \rangle = 0.50 \pm 0.3$) with only upper limits measured and on average have lower effective temperatures ($\langle T_{\text{eff}} \rangle = 5440 \pm 185$ K).

From the available data, we cannot see a discernible difference between the planet hosts and the stars with no planets detected regarding their lithium content (Figure 12). From a larger sample of F, G, and K dwarfs of the Southern Hemisphere containing 100 planet hosts, Luck (2018b) also came to the same conclusion. Moreover, Carlos et al. (2019) emphasize a strong connection between lithium depletion and stellar ages, and notice that there is no significant difference in lithium depletion between known planet-host stars and stars with no planets detected, when the lithium abundance and stellar age correlation is analyzed. However, there could be planet-hosting stars like our Sun with quite low lithium abundances.

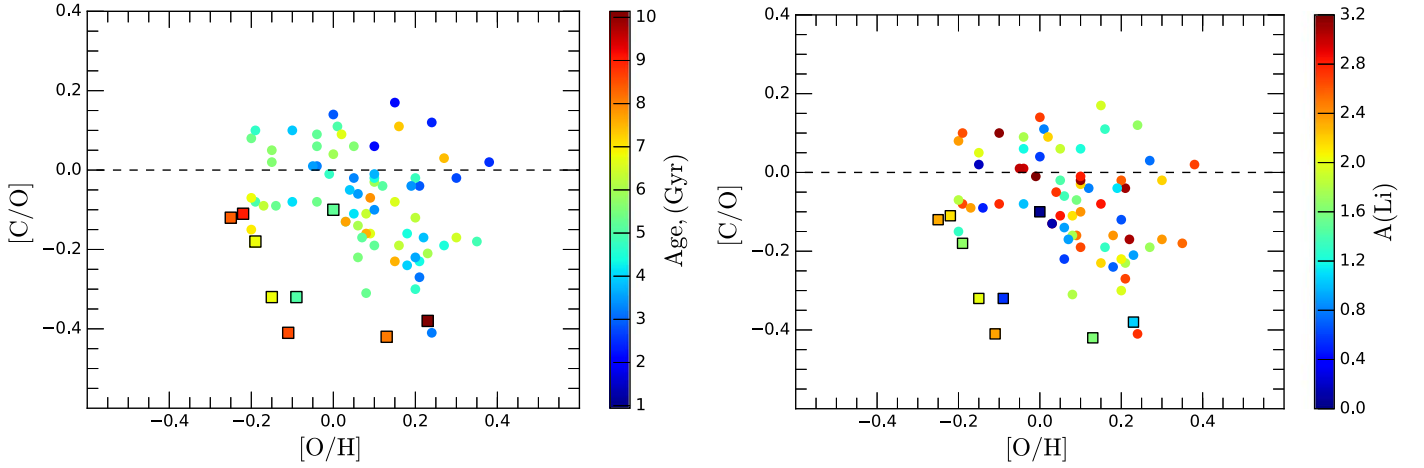


Figure 11. [C/O] vs. [O/H]. Left: our results for thin-disk stars are shown as circles and thick-disk stars as squares. The stars are color-coded by age. Right: the same plot with stars color-coded by their lithium abundances.

Table 5
Stars of Our Sample with Confirmed Exoplanets

TYC ID	Planet	Planet mass (M_{Jup})	Semimajor axis (au)	Orbital period (day)	$A(\text{Li})$	Host stars ^a C/O	Mg/Si
2099-2717-1	HD 164922 b	0.31 ± 0.05	2.10 ± 0.04	1155 ± 23	0.45^b	0.59	1.44
	HD 164922 c	0.04 ± 0.005	0.34 ± 0.0015	$75.77^{+0.058}_{-0.056}$			
2103-1620-1	HD 164595 b	0.05 ± 0.00856	0.23	40 ± 0.24	1.05^b	...	1.05
2595-1464-1	HD 155358 b	0.99 ± 0.08	0.63 ± 0.02	194.3 ± 0.3	2.02	0.26	1.55
	HD 155358 c	0.82 ± 0.07	1.02 ± 0.02	391.9 ± 1			
2648-2151-1	HD 178911 B b	8.03 ± 2.51	0.34 ± 0.01	71.48 ± 0.02	0.55^b	...	1.08
3067-576-1	14 Her b	4.66 ± 0.15	2.93 ± 0.08	1773.4 ± 2.5	0.75^b	...	1.60
3501-1373-1	HD 154345 b	0.82 ± 0.07	4.21 ± 0.10	3341.56 ± 93	0.05^b	...	0.95
4436-1424-1	ψ^1 Draconis B b	1.53 ± 0.10	4.43 ± 0.04	3117 ± 42	2.79	0.55	1.29
4494-1346-1	HD 7924 b	0.02 ± 0.00	0.06 ± 0.00	5.4 ± 0.00025	0.35^b	0.44	1.23
	HD 7924 c	$0.02^{+0.0023}_{-0.0022}$	$0.11^{+0.0013}_{-0.0014}$	$15.3^{+0.0032}_{-0.0033}$			
	HD 7924 d	$0.02^{+0.0025}_{-0.0025}$	$0.16^{+0.0018}_{-0.0019}$	$24.45^{+0.015}_{-0.017}$			
4532-2096-1	HD 33564 b	9.1	1.1	388 ± 3	2.28	...	1.11
4575-1336-1	HD 150706 b	$2.71^{+1.14}_{-0.66}$	$6.7^{+4.0}_{-1.4}$	5894^{+5584}_{-1498}	2.55	0.45	1.12

Notes. Data were taken from NASA Exoplanet Archive (Akeson et al. 2013) on 2019 July 31.

^a Our study results.

^b Upper limits.

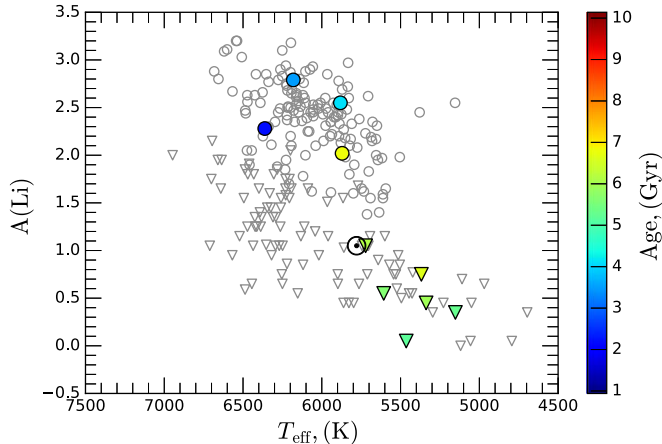


Figure 12. Lithium abundance as a function of the effective temperature. The stars with so far detected planets with measurable $A(\text{Li})$ and upper limits are shown as circles or triangles, respectively, and color-coded by their age values.

Theoretical studies suggest that the C/O number ratio as well as Mg/Si are important in determining the mineralogy of extraterrestrial planets. While the C/O ratio controls the amount of carbides and silicates formed in planets, the Mg/Si can tell us about the silicate mineralogy (e.g., Bond et al. 2010; Madhusudhan et al. 2014; Piso et al. 2016). For stars with $\text{Mg/Si} < 1$, terrestrial planets will have a magnesium depleted mineralogy different from that of the Earth. For stars with $\text{Mg/Si} \sim 1 \div 2$, Mg is distributed between olivine and pyroxene, leading to rocky planet compositions similar to the Earth. In case of $\text{Mg/Si} > 2$, all available Si is consumed to form olivine with excess Mg available to bond with other elements as MgO or MgS (Bond et al. 2010). Furthermore, the C/O number ratio as well as Mg/Si are important in the formation, atmospheric chemistry, interior structure for all types of exoplanets, and for the plate tectonics and habitability of terrestrial planets. For example, in order to constrain the interior structure of rocky exoplanets the stellar elemental

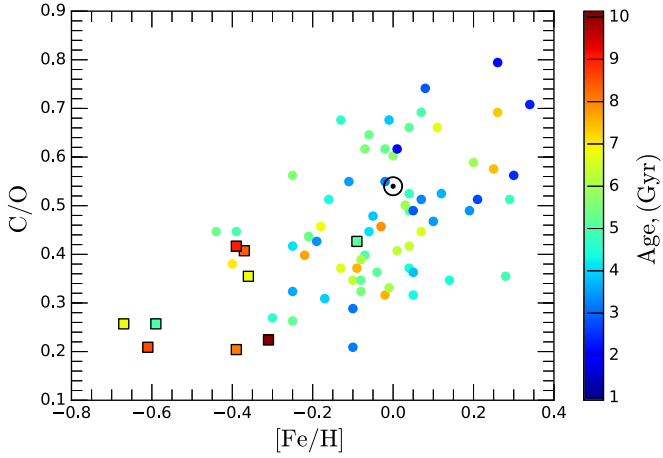


Figure 13. C/O number ratio as a function of metallicity. Stars are color-coded by age. The thin-disk stars are shown as circles, the thick-disk stars as squares.

abundances (such as Fe, Si, and Mg) are key constraints to reduce degeneracy in interior structure models and to constrain the mantle composition (Dorn 2015). Recent observational data showed that C/O varies among planetary atmospheres and host stars, and even among planets within the same planetary system (Brewer et al. 2017; Suárez-Andrés et al. 2018).

Figure 13 shows the C/O number ratio as a function of metallicity. Stars are color-coded by age. First of all, we notice the nine stars with predominantly supersolar metallicities that have the C/O number ratios larger than 0.65, thus according to, e.g., Moriarty et al. (2014) they may form carbon-rich rocky planets. This makes an unexpectedly large percentage, 12% of the sample stars, which is a hundred times larger than it was found by Brewer & Fischer (2016). The percentage of 0.13% for the carbon-enhanced dwarfs in the solar neighborhood predicted by Brewer & Fischer (2016) is indeed probably underestimated since, e.g., Suárez-Andrés et al. (2018) found about 1% of their sample stars having C/O ratio even larger than 0.8. Our stellar sample does not contain stars with $C/O > 0.8$, which necessarily should have carbon-rich rocky planets if detected. The mean C/O ratio for 53 sample stars within ± 0.2 dex of the solar metallicity is 0.47 ± 0.12 , this value falls exactly to the C/O peak found for the solar-neighborhood stars by Brewer & Fischer (2016) and Suárez-Andrés et al. (2018).

In Figure 14 we show C/O number ratios as a function of Mg/Si. Stars with detected planets are color-coded by age. The Sun with our determined values of $C/O_{\odot} = 0.54$ and $Mg/Si_{\odot} = 1.05$ is marked as well. About 89% of stars have Mg/Si values ranging from 1.0 to ~ 2.0 with the mean value of $Mg/Si = 1.26 \pm 0.17$, which would lead to rocky planet compositions close to that of the Earth. The remaining $\sim 10\%$ of the stars have $Mg/Si < 1.0$, where terrestrial planets would have a magnesium depleted mineralogy with the mean $C/O \simeq 0.48 \pm 0.10$. The thin-disk stars have C/O ratio in the range from 0.2 to 0.8 with the mean $Mg/Si = 1.16 \pm 0.16$. Whereas, the majority of the thick-disk stars (on average older stars) have C/O ratios between 0.2 and 0.4 with the mean $Mg/Si = 1.61 \pm 0.28$. The Mg/Si number ratios of all 249 stars are shown in Figure 15 as a function of the effective temperatures. Several thick-disk stars have $Mg/Si > 2.0$.

The percentage of stars with exoplanets in our sample so far is very small, but with the new data from the NASA TESS space mission, there is no doubt, our homogeneous

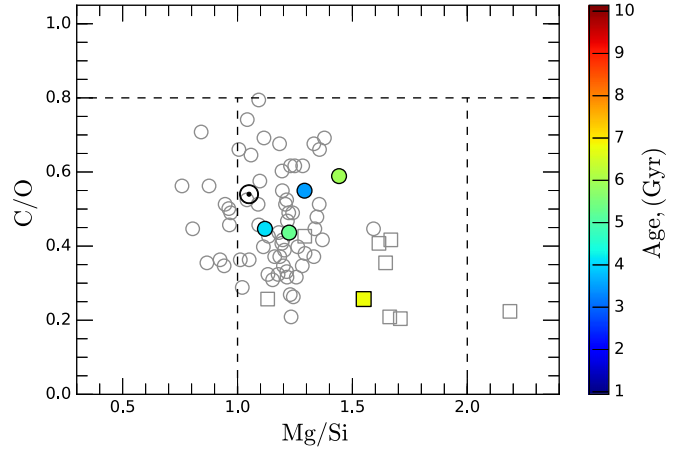


Figure 14. C/O number ratio as a function Mg/Si number ratio. Our results for the thin-disk stars are shown as circles, the thick-disk—squares. Mg/Si number ratios are taken from Mikolaitis et al. (2019). Stars with so far detected planets are color-coded by age. The dashed lines divide stellar parameter spaces where gaseous (if $Mg/Si < 1$) or rocky ($1 < Mg/Si < 2$) exoplanets may form (see Suárez-Andrés et al. 2018 and the text for more explanations).

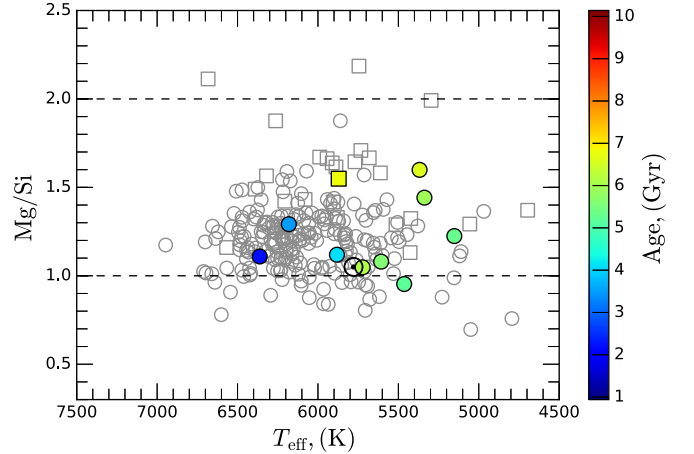


Figure 15. Mg/Si number ratio as a function of the effective temperature. All 10 stars with detected planets are color-coded by age.

high-resolution abundance analysis in planet hosts will present a useful case for future studies.

4. Summary and Conclusions

In the presented work, we homogeneously investigated abundances of lithium, carbon, and oxygen in a sample of 249 bright dwarf stars of the northern sky with a rather broad metallicity range. The investigated stars have metallicities, effective temperatures, and ages between $(-0.7 \div 0.4)$ dex, $(5000 \div 6900)$ K, and $(1 \div 12)$ Gyr, accordingly. We determined lithium abundances for all investigated stars and carbon and oxygen abundances for 78 stars of the sample. The stars belong mainly to the thin Galactic disk.

$A(Li)$ is predominantly higher in the hottest, youngest, and more massive stars of our sample as expected. The decrease of $A(Li)$ per 1000 K equals to 0.4 ± 0.1 dex for stars with $A(Li) \geq 2.0$. The lithium abundance correlation with the effective temperature is similar for both the lithium-rich stars with $A(Li) \geq 2$ and for those with $A(Li) < 2$.

We identified dwarf stars with similar ages, atmospheric parameters, and rotational velocities, but significantly different $A(\text{Li})$ values. We speculate that in addition to the above-mentioned main parameters, lowered lithium abundances in these solar-vicinity dwarfs could be caused by external factors such as the Galactic kinematic evolution and partially by presence of planets.

We confirm that chromospherically active dwarfs may have quite high lithium abundances. We found two such stars in our investigated sample.

Our work shows that even in a small volume of space in the solar neighborhood stars exhibit decrease of lithium at supersolar metallicities—the phenomenon that cannot be explained even by the most recent models. The highest lithium abundance of 3.00 dex in our study is reached at the solar metallicity and drops by 0.7 dex in the $[\text{Fe}/\text{H}]$ range from +0.10 to +0.55 dex.

In the investigated sample of stars, we cannot see a discernible difference between the planet hosts and the stars with no planets detected regarding their lithium content. Both the thin- and thick-disk dwarfs have no $[\text{C}/\text{O}]$ correlation with $A(\text{Li})$.

Nine stars with predominantly supersolar metallicities, i.e., about 12% of the sample of 78 stars with C and O abundances determined, have the C/O number ratios larger than 0.65, thus may form carbon-rich rocky planets. These stars could be interesting targets for space (e.g., *TESS* and *PLATO*) and ground-based exoplanet search surveys. About 89% of our sample stars have the mean value of $\text{Mg}/\text{Si} = 1.26 \pm 0.17$, i.e., Mg/Si ratios are in the range between one and two, which would lead to rocky planet compositions close to that of the Earth. The remaining $\sim 10\%$ of the stars have $\text{Mg}/\text{Si} < 1.0$, where terrestrial planets would have a magnesium depleted mineralogy with the mean $\text{C}/\text{O} \simeq 0.48 \pm 0.10$.

The percentage of stars with confirmed exoplanets in our sample is still small and we plan to contribute in the analysis of the key elements like lithium, carbon, and oxygen in planet hosts with the homogeneous high-resolution data at once the new data from the NASA *TESS* space mission will be available for stars in the Northern Hemisphere.

We gratefully acknowledge the grant from the European Social Fund via the Lithuanian Science Council (LMTLT) grant No. 09.3.3-LMT-K-712-01-0103. We thank the anonymous referee for the prompt and valuable report. This research has made use of the NASA Exoplanet Archive, which is operated by the California Institute of Technology, under contract with the National Aeronautics and Space Administration under the Exoplanet Exploration Program.

Appendix Information About the Online Table

Information about columns and meanings of acronyms used in presenting the full online table of results for the investigated 249 stars. The main atmospheric parameters (T_{eff} , $\log g$, $[\text{Fe}/\text{H}]$, v_t) of stars were taken from Mikolaitis et al. (2018, 2019).

Column	Label	Unit	Comment
(1)	Name		Star name TYC
(2)	$A(\text{Li})$	dex	$A(\text{Li})$ abundance
(3)	$M U$		M—measurable, U—upper limit of $A(\text{Li})$
(4)	$v \sin i$	km s^{-1}	Rotational velocity

(Continued)

Column	Label	Unit	Comment
(5)	$[\text{C}/\text{H}]$	dex	Carbon-to-hydrogen ratio
(6)	$e_{[\text{C}/\text{H}]}$	dex	Error on carbon-to-hydrogen ratio
(7)	n_c		Number of lines
(8)	$[\text{O}/\text{H}]$	dex	Oxygen-to-hydrogen ratio
(9)	$e_{[\text{O}/\text{H}]}$	dex	Error on oxygen-to-hydrogen ratio
(10)	n_o		Number of lines
(11)	C/O		Carbon-to-oxygen number ratio
(12)	Thin Thick		Thin-disk star “0” Thick-disk “1”
(13)	Planet		Have any planet “1” Have not “0”
(14)	T_{eff}	K	Effective temperature
(15)	$e_{T_{\text{eff}}}$	K	Error on effective temperature
(16)	$\log g$	dex	Surface gravity
(17)	$e_{\log g}$	dex	Error on surface gravity
(18)	$[\text{Fe}/\text{H}]$	dex	Metallicity
(19)	$e_{[\text{Fe}/\text{H}]}$	dex	Error on metallicity
(20)	V_t	km s^{-1}	Microturbulence velocity
(21)	e_{V_t}	km s^{-1}	Error on microturbulence velocity

ORCID iDs

E. Stonkutė  <https://orcid.org/0000-0002-8028-8133>
 G. Tautvaišienė  <https://orcid.org/0000-0001-7672-154X>
 A. Drazdauskas  <https://orcid.org/0000-0001-5145-254X>
 R. Minkevičiūtė  <https://orcid.org/0000-0001-6812-880X>
 Š. Mikolaitis  <https://orcid.org/0000-0002-1419-0517>
 E. Pakštienė  <https://orcid.org/0000-0002-3326-2918>

References

- Adibekyan, V., de Laverny, P., Recio-Blanco, A., et al. 2018, *A&A*, **619**, A130
 Aguilera-Gómez, C., Ramírez, I., & Chanamé, J. 2018, *A&A*, **614**, A55
 Akesson, R. L., Chen, X., Ciardi, D., et al. 2013, *PASP*, **125**, 989
 Alvarez, R., & Plez, B. 1998, *A&A*, **330**, 1109
 Asplund, M. 2005, *ARA&A*, **43**, 481
 Aubé, M., Fortin, N., Turcotte, S., et al. 2014, *PASP*, **126**, 1068
 Bedell, M., Bean, J. L., Meléndez, J., et al. 2018, *ApJ*, **865**, 68
 Bensby, T., & Feltzing, S. 2006, *MNRAS*, **367**, 1181
 Bensby, T., & Lind, K. 2018, *A&A*, **615**, A151
 Bond, J. C., O’Brien, D. P., & Lauretta, D. S. 2010, *ApJ*, **715**, 1050
 Brewer, J. M., & Fischer, D. A. 2016, *ApJ*, **831**, 20
 Brewer, J. M., Fischer, D. A., & Madhusudhan, N. 2017, *AJ*, **153**, 83
 Brewer, J. M., Fischer, D. A., Valenti, J. A., & Piskunov, N. 2016, *ApJS*, **225**, 32
 Brugamyer, E., Dodson-Robinson, S. E., Cochran, W. D., & Sneden, C. 2011, *ApJ*, **738**, 97
 Caffau, E., Ludwig, H. G., Bonifacio, P., et al. 2010, *A&A*, **514**, A92
 Carlos, M., Meléndez, J., Spina, L., et al. 2019, *MNRAS*, **485**, 4052
 Casagrande, L., Ramírez, I., Meléndez, J., Bessell, M., & Asplund, M. 2010, *A&A*, **512**, A54
 Cescutti, G., Matteucci, F., McWilliam, A., & Chiappini, C. 2009, *A&A*, **505**, 605
 Cescutti, G., & Molaro, P. 2019, *MNRAS*, **482**, 4372
 Chen, Y. Q., & Zhao, G. 2006, *AJ*, **131**, 1816
 Cyburt, R. H., Fields, B. D., & Olive, K. A. 2008, *JCAP*, **11**, 012
 da Silva, R., Milone, A. d. C., Rocha-Pinto, H. J., et al. 2015, *A&A*, **580**, A24
 Delgado Mena, E., Bertrán de Lis, S., Adibekyan, V. Z., et al. 2015, *A&A*, **576**, A69
 Delgado Mena, E., Israelian, G., González Hernández, J. I., et al. 2010, *ApJ*, **725**, 2349
 Delgado Mena, E., Israelian, G., González Hernández, J. I., et al. 2014, *A&A*, **562**, A92
 Dorn, R. I. 2015, *Geomo*, **250**, 53
 Eisenbeiss, T., Ammler-von Eiff, M., Roell, T., et al. 2013, *A&A*, **556**, A53
 Figueira, P., Faria, J. P., Delgado-Mena, E., et al. 2014, *A&A*, **570**, A21
 Fu, X., Bressan, A., Molaro, P., & Marigo, P. 2015, *MNRAS*, **452**, 3256

- Fu, X., Romano, D., Bragaglia, A., et al. 2018, *A&A*, **610**, A38
- Girardi, L., Bressan, A., Bertelli, G., & Chiosi, C. 2000, *A&AS*, **141**, 371
- González Hernández, J. I., Israelian, G., Santos, N. C., et al. 2010, *ApJ*, **720**, 1592
- Grevesse, N., Asplund, M., & Sauval, A. J. 2007, *SSRv*, **130**, 105
- Guiglion, G., Chiappini, C., Romano, D., et al. 2019, *A&A*, **623**, A99
- Guiglion, G., de Laverny, P., Recio-Blanco, A., et al. 2016, *A&A*, **595**, A18
- Gustafsson, B., Edvardsson, B., Eriksson, K., et al. 2008, *A&A*, **486**, 951
- Gustafsson, B., Karlsson, T., Olsson, E., Edvardsson, B., & Ryde, N. 1999, *A&A*, **342**, 426
- Isaacson, H., & Fischer, D. 2010, *ApJ*, **725**, 875
- Israelian, G., Delgado Mena, E., Santos, N. C., et al. 2009, *Natur*, **462**, 189
- Israelian, G., Santos, N. C., Mayor, M., & Rebolo, R. 2004, *A&A*, **414**, 601
- Johansson, S., Litzén, U., Lundberg, H., & Zhang, Z. 2003, *ApJL*, **584**, L107
- Jurgenson, C., Fischer, D., McCracken, T., et al. 2016, *JAI*, **5**, 1650003
- King, J. R., Deliyannis, C. P., Hiltgen, D. D., et al. 1997, *AJ*, **113**, 1871
- Kurucz, R. L. 2005, *MSAIS*, **8**, 189
- Lambert, D. L., & Reddy, B. E. 2004, *MNRAS*, **349**, 757
- Larimer, J. W. 1975, *GeCoA*, **39**, 389
- Lind, K., Asplund, M., & Barklem, P. S. 2009, *A&A*, **503**, 541
- López-Valdivia, R., Hernández-Águila, J. B., Bertone, E., et al. 2015, *MNRAS*, **451**, 4368
- Luck, R. E. 2017, *AJ*, **153**, 21
- Luck, R. E. 2018a, *AJ*, **155**, 111
- Luck, R. E. 2018b, *AJ*, **156**, 171
- Madhusudhan, N., Agúndez, M., Moses, J. I., & Hu, Y. 2016, *SSRv*, **205**, 285
- Madhusudhan, N., Amin, M. A., & Kennedy, G. M. 2014, *ApJL*, **794**, L12
- Mikolaitis, Š., Drazdauskas, A., Minkevičiūtė, R., et al. 2019, *A&A*, **628**, A49
- Mikolaitis, Š., Tautvaišienė, G., Drazdauskas, A., et al. 2018, *PASP*, **130**, 074202
- Mishenina, T., Kovtyukh, V., Soubiran, C., & Adibekyan, V. Z. 2016, *MNRAS*, **462**, 1563
- Mishenina, T. V., Soubiran, C., Kovtyukh, V. V., Katsova, M. M., & Livshits, M. A. 2012, *A&A*, **547**, A106
- Moriarty, J., Madhusudhan, N., & Fischer, D. 2014, *ApJ*, **787**, 81
- Nissen, P. E., Chen, Y. Q., Carigi, L., Schuster, W. J., & Zhao, G. 2014, *A&A*, **568**, A25
- Nissen, P. E., & Gustafsson, B. 2018, *A&ARv*, **26**, 6
- Pavlenko, Y. V., Jenkins, J. S., Ivanyuk, O. M., et al. 2018, *A&A*, **611**, A27
- Pavlenko, Y. V., Kaminsky, B. M., Jenkins, J. S., et al. 2019, *A&A*, **621**, A112
- Petigura, E. A., & Marcy, G. W. 2011, *ApJ*, **735**, 41
- Piso, A.-M. A., Pegues, J., & Öberg, K. I. 2016, *ApJ*, **833**, 203
- Prantzos, N. 2012, *A&A*, **542**, A67
- Prantzos, N., de Laverny, P., Guiglion, G., Recio-Blanco, A., & Worley, C. C. 2017, *A&A*, **606**, A132
- Ramírez, I., Fish, J. R., Lambert, D. L., & Allende Prieto, C. 2012, *ApJ*, **756**, 46
- Rauer, H., Catala, C., Aerts, C., et al. 2014, *ExA*, **38**, 249
- Reddy, B. E., Lambert, D. L., & Allende Prieto, C. 2006, *MNRAS*, **367**, 1329
- Ricker, G. R., Winn, J. N., Vanderspek, R., et al. 2015, *JATIS*, **1**, 014003
- Romano, D., Matteucci, F., Ventura, P., & D'Antona, F. 2001, *A&A*, **374**, 646
- Schuler, S. C., Flateau, D., Cunha, K., et al. 2011, *ApJ*, **732**, 55
- Spina, L., Meléndez, J., & Ramírez, I. 2016, *A&A*, **585**, A152
- Spite, F., & Spite, M. 1982, *A&A*, **115**, 357
- Suárez-Andrés, L., Israelian, G., González Hernández, J. I., et al. 2017, *A&A*, **599**, A96
- Suárez-Andrés, L., Israelian, G., González Hernández, J. I., et al. 2018, *A&A*, **614**, A84
- Théado, S., & Vauclair, S. 2012, *ApJ*, **744**, 123
- Thiabaud, A., Marboeuf, U., Alibert, Y., Laya, I., & Mezger, K. 2015a, *A&A*, **574**, A138
- Thiabaud, A., Marboeuf, U., Alibert, Y., Laya, I., & Mezger, K. 2015b, *A&A*, **580**, A30
- Tinsley, B. M. 1979, *ApJ*, **229**, 1046
- Zhang, L.-Y., Pi, Q.-F., & Zhu, Z.-Z. 2015, *RAA*, **15**, 252