



A Variant Stellar-to-nebular Dust Attenuation Ratio on Subgalactic and Galactic Scales

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Abstract

The state-of-the-art geometry models of stars/dust suggest that dust attenuation toward nebular regions ($A_{V,\text{gas}}$) is always larger than that of stellar regions ($A_{V,\text{star}}$). Utilizing the newly released integral field spectroscopic data from the Mapping Nearby Galaxies at Apache Point Observatory survey, we investigate whether and how the $A_{V,\text{star}}/A_{V,\text{gas}}$ ratio varies from subgalactic to galactic scales. On a subgalactic scale, we report a stronger correlation between $A_{V,\text{star}}$ and $A_{V,\text{gas}}$ for more active H II regions. The local $A_{V,\text{star}}/A_{V,\text{gas}}$ is found to have moderate nonlinear correlations with three tracers of diffuse ionized gas (DIG), as well as indicators of gas-phase metallicity and ionization. The DIG regions tend to have larger $A_{V,\text{star}}/A_{V,\text{gas}}$ compared to classic H II regions excited by young OB stars. Metal-poor regions with a higher ionized level suffer much less nebular attenuation and thus have larger $A_{V,\text{star}}/A_{V,\text{gas}}$ ratios. A low- $A_{V,\text{gas}}$ and high- $A_{V,\text{star}}/A_{V,\text{gas}}$ sequence, which can be resolved into DIG-dominated and metal-poor regions, on the three BPT diagrams is found. Based on these observations, we suggest that besides the geometry of stars/dust, local physical conditions such as metallicity and ionized level also play an important role in determining the $A_{V,\text{star}}/A_{V,\text{gas}}$. On a galactic scale, the global $A_{V,\text{star}}/A_{V,\text{gas}}$ ratio has strong correlations with stellar mass (M_*), moderate correlations with star formation rate (SFR) and metallicity, and weak correlations with inclination and specific SFR. Galaxies with larger M_* and higher SFR that are more metal-rich tend to have smaller $A_{V,\text{star}}/A_{V,\text{gas}}$ ratios. Such correlations form a decreasing trend of $A_{V,\text{star}}/A_{V,\text{gas}}$ along the star-forming main sequence and mass–metallicity relation. The dust growth process accompanied by galaxy growth might be one plausible explanation for our observations.

Unified Astronomy Thesaurus concepts: [Interstellar dust extinction \(837\)](#); [Galaxy evolution \(594\)](#); [Interstellar medium \(847\)](#); [Warm ionized medium \(1788\)](#); [H II regions \(694\)](#)

1. Introduction

Almost half of optical emission is shielded by dust and reaches us in the form of infrared (IR) radiation (Dole et al. 2006), especially for stellar light from star-forming galaxies (SFGs) that contribute to the cosmic IR background up to about 95% (Viero et al. 2013). This so-called dust extinction/attenuation effect prevents us from accurately measuring some important galactic properties, such as star formation rate (SFR), stellar mass (M_*), and so on. Dust is produced from the late stage of stellar evolution—namely supernovae (SNe) and stellar winds of asymptotic giant branch stars (e.g., Dwek 1998; Aoyama et al. 2017; Gjergo et al. 2018)—grows in the interstellar medium (ISM), especially in dense molecular clouds (e.g., Hirashita & Voshchinnikov 2014); and is destroyed by SN shocks (e.g., Aoyama et al. 2017). Such dust processes substantially alter the dust abundance and the size distribution of dust grains (e.g., Asano et al. 2013; Aoyama et al. 2017; Hirashita & Aoyama 2019) and hence affect the absorption and scattering of light (see Galliano 2018 for a review). Therefore, to better understand the formation and evolution of galaxies, we should know how and how much stellar light is reprocessed by the dust extinction/attenuation effect.

Besides the dust properties, the geometry between dust and stars is also believed to be an important factor for dust attenuation in terms of variations in the shape of the dust attenuation law (Salmon et al. 2016; Narayanan et al. 2018a) and the IRX- β relation (Meurer et al. 1999; Kong et al. 2004; Ye et al. 2016; Popping et al. 2017; Narayanan et al. 2018b). Particularly, it is thought of as the main origin of the different levels of dust attenuation between stellar light and emission lines (e.g., Price et al. 2014; Reddy et al. 2015; Koyama et al. 2019).

Plenty of observations on both local (e.g., Calzetti 1997; Kreckel et al. 2013; Zahid et al. 2017; Koyama et al. 2019) and high-redshift (e.g., Garn et al. 2010; Wuyts et al. 2011, 2013; Price et al. 2014; Pannella et al. 2015; Theios et al. 2019) SFGs reveal that nebular emission lines tend to suffer more dust attenuation compared to stellar continuum. Such extra attenuation can be explained by a widely used two-component dust model (Charlot & Fall 2000; Wild et al. 2011; Chevallard et al. 2013) in which dust inside a galaxy consists of a diffuse, optically thin component and a denser, optically thick one (the birth cloud) that relates to star-forming regions. Thus, emission from ionized gas travels through the outer envelope of the birth clouds and then the ambient ISM, while the majority of stellar light propagates only the ambient ISM and suffers less dust reddening compared to emission lines (Charlot & Fall 2000).

By utilizing the observation of local starburst galaxies and bright star-forming regions, Calzetti (1997) presented the relation between the color excess $E(B-V)$ of stars and gas as $E(B-V)_{\text{star}} = 0.44E(B-V)_{\text{gas}}$, i.e., the reddening of stellar light is only nearly half of the one toward nebular regions. A similar relation based on V-band attenuation (A_V) was reported by the spatially resolved study of Kreckel et al. (2013) in which they found $A_{V,\text{star}} = 0.47 \pm 0.006A_{V,\text{gas}}$, assuming $R_V = 3.1$ for several relatively face-on local galaxies. More recent studies also suggest that $A_{V,\text{star}}/A_{V,\text{gas}}$ may be not a constant for all types of galaxies (Wild et al. 2011; Koyama et al. 2015; Zahid et al. 2017; Koyama et al. 2019; Qin et al. 2019). It is found to be correlated with M_* (Koyama et al. 2015; Zahid et al. 2017; Koyama et al. 2019), specific SFR (sSFR; Wild et al. 2011; Koyama et al. 2015, 2019; Qin et al. 2019), and inclination (Wild et al. 2011).

High-redshift studies also revealed a wide diversity of the stellar-to-nebular reddening ratio, varying from 0.44 (Wuyts et al. 2011) to ~ 1 (Pannella et al. 2015; Puglisi et al. 2016), as well as the dependences of $A_{V,\text{star}}/A_{V,\text{gas}}$ on the physical properties of galaxies (Price et al. 2014; Reddy et al. 2015; Puglisi et al. 2016).

However, most the aforementioned studies focused on dust attenuation on a global scale, deriving the $A_{V,\text{star}}/A_{V,\text{gas}}$ using the integrated fluxes of galaxies. With the development of the integral field unit (IFU) technique, spatially resolved spectra were observed for thousands of galaxies in several IFU surveys, such as SAMI (using the Sydney-Australian Astronomical Observatory Multi-object Integral Field Spectrograph; Croom et al. 2012), Calar Alto Large Integral Field Area (CALIFA; Sánchez et al. 2012), and Mapping Nearby Galaxies at Apache Point Observatory (MaNGA; Bundy et al. 2015). Taking advantage of these data sets, we are able to study the dust attenuation on subgalactic scales that would help us obtain a more comprehensive understanding of dust effects. However, only a few works paid attention to the subgalactic dust attenuation until now.

On the basis of spatially resolved spectra of several local galaxies, Kreckel et al. (2013) showed that the dust attenuation ratios within individual galaxies exhibit a dependence on the [S II] $\lambda 6717$ -to- $\text{H}\alpha$ line ratio, which is believed to be a tracer of the transition between diffuse ionized gas (DIG) and H II regions (e.g., Hoopes & Walterbos 2003). Adopting a critical [S II] $\lambda 6717/\text{H}\alpha$ ratio of 0.2, they found an $A_{V,\text{star}}/A_{V,\text{gas}}$ ratio of 0.7 and 0.5 for DIG- and H II-dominated regions, respectively. Recently, using the high-resolution IFU observations of NGC 5626, a lenticular galaxy with a redshift of 0.023, Viaene et al. (2017) found that the continuum attenuation is independent from the ionized gas attenuation in this early-type galaxy. This observation raises the question of whether the dust attenuations of the stellar continuum and ionized gas indeed correlate tightly with each other and whether the correlation could be broken under certain conditions.

On the other hand, adopting the ultraviolet (UV) slope β as an indicator of stellar attenuation, Calzetti et al. (1994) observed a linear correlation between β and the nebular attenuation, indicating a tight and monotonic relation between the stellar and nebular dust attenuation. Based on this property, the widely used dust attenuation law of Calzetti et al. (2000) was constructed. More recent works following the same methodology to derive the attenuation law either in the local universe (e.g., Battisti et al. 2016, 2017) or at high redshift (e.g., Reddy et al. 2015) also rely on this relation. Therefore, it is important to verify the applicability of this finding to a broader range of galaxies than those used in previous studies.

Making use of IFU data, this work aims at investigating whether the dust attenuation of the stellar continuum and ionized gas correlate with each other and, if so, whether their ratio correlates with local and global properties for H II regions. This paper is organized as follows. In Section 2, we give a brief overview of the data, as well as sample selection. We present results of correlations between the reddening ratio and local physical properties in Section 3. We further analyze how $A_{V,\text{star}}/A_{V,\text{gas}}$ varies with the global properties of galaxies in Section 4. Finally, we summarize in Section 5. Throughout this paper, we adopt a flat Λ CDM cosmology with $H_0 = 71 \text{ km s}^{-1} \text{ Mpc}^{-1}$, $\Omega_\Lambda = 0.73$, $\Omega_m = 0.27$, and a Salpeter (1955) initial mass function (IMF).

2. Data and Sample Selection

2.1. MaNGA Survey and Pipe3D VAC

MaNGA, as one of the major spectroscopic programs of the Sloan Digital Sky Survey IV (SDSS-IV; Blanton et al. 2017), plans to observe $\sim 10,000$ nearby galaxies with a redshift range of 0.01–0.15, spanning a wide range of M_* , SFR, and environment (Bundy et al. 2015; Yan et al. 2016). Spectral observations of the MaNGA survey cover a wavelength range of 3600–10300 Å with a spectral resolution of $R \sim 2000$ and a spatial resolution, i.e., the FWHM of the reconstructed point-spread function (PSF), of $2''.5$ (Law et al. 2016). In this work, we make use of MaNGA data from the SDSS DR15 (Aguado et al. 2019), which includes 4824 data cubes.

Our main data are from the MaNGA value-added catalog (VAC) of the Pipe3D pipeline (Sánchez et al. 2018), which is released as one part of the SDSS DR15.³ This pipeline is designed to implement full spectral fitting with stellar population models for IFU spectra. Details and algorithms can be found in Sánchez et al. (2016a, 2016b). Here we provide a brief summary of the procedure of stellar continuum fitting. Spatial binning is performed for each data cube to achieve a nominal continuum signal-to-noise ratio (S/N) of 50 before fitting. A set of 156 simple stellar population (SSP) templates from Cid Fernandes et al. (2013) accounting for 39 stellar ages (from 1 Myr to 13 Gyr) and four metallicities ($Z/Z_\odot = 0.2, 0.4, 1, \text{ and } 1.5$) is used to model the coadded spectra of each spatial bin. The fitting procedure includes two steps: first, a fitting is performed to extract the nonlinear parameters, such as velocity, velocity dispersion, and dust attenuation, assuming an extinction curve of Cardelli et al. (1989) and a selective extinction of $R_V = 3.1$; then, a robust modeling is applied to the corrected spectra after accounting for the nonlinear effects derived from the previous step, via a linear combination of the 156 SSPs with the Monte Carlo method to obtain the coefficients of the best-fit model and the corresponding uncertainties.

After the continuum modeling, the pipeline rescales the best-fit spectral energy distribution of one spatial bin by the V-band flux of each spaxel within that bin and takes it as the best-fit model of each spaxel. This resulting “fake” spectrum is then subtracted from the observed one to generate a pure emission line spectrum. Individual emission lines are fitted spaxel by spaxel with a single Gaussian profile. It is noteworthy that the continuum modeling is implemented for spatial bins, whereas the ionized gas analysis is performed for individual spaxels. Finally, the Pipe3D VAC provides data products for 4815 galaxies, including the results of the analyses of both stellar continua and emission lines (Aguado et al. 2019).

2.2. Reanalysis of Emission Lines

Given the aim of this study, we need dust attenuation of both stellar continuum and emission lines, where the former one is provided for spatial bins and the latter one is given for individual spaxels in the Pipe3D VAC. To reconcile the spatial scale of these two parameters, we determine to use spatial bins to study and reanalyze the emission lines on the same spatial scale. This choice can also prevent the oversampling of low-S/N spectra for which the stellar properties should be the same if they belong to the same spatial bin, as well as avoid any

³ <https://www.sdss.org/dr15/manga/manga-data/manga-pipe3d-value-added-catalog/>

potential uncertainties from the V-band flux-based assignment of the underlying stellar contribution. For simplicity, in the following, the term “spaxel” refers to a given spatial element regardless of the number of spaxels within it.

We reconstruct the best-fit spectra for each spaxel using the weights of 156 SSP templates from the Pipe3D VAC. The resulting spectra are in good agreement with the coadded observed spectra and subtracted to create pure emission line spectra for further emission line analysis. We fit each emission line with a single Gaussian function using the Python version of the MPFIT⁴ software (Markwardt 2009). The S/Ns of emission lines are calculated via the method introduced in Ly et al. (2014). We use the Balmer decrement to correct the galactic internal attenuation. Under the assumption of Case B recombination, we employ an intrinsic flux ratio of $H\alpha/H\beta = 2.86$ (Storey & Hummer 1995), an extinction curve of Cardelli et al. (1989), and $R_V = 3.1$ to do the dereddening. In short, we redo the analysis of ionized gas but on the spatial bin scale.

2.3. Sample Selection

To ensure a reliable estimate of the dust attenuation of gas ($A_{V,\text{gas}}$), we limit our analysis to spaxels with S/Ns of $H\alpha$ and $H\beta$ greater than 5 and S/Ns of $[O\text{ III}]\lambda 5007$, $[N\text{ II}]\lambda 6584$, and $[S\text{ II}]\lambda 6717$ greater than 3. Setting lower S/N cuts should not change our results (see the Appendix for a detailed discussion). Given that the spatial binning scheme performed by Pipe3D requires not only an S/N goal but also a continuity criterion, many of the spatial bins cannot reach the set S/N of 50 due to the balance between these two requirements (Sánchez et al. 2016b). Therefore, an additional continuum S/N cut, which we set to 10, is also needed to ensure an unbiased estimate of stellar attenuation ($A_{V,\text{star}}$; Sánchez et al. 2016a). The distribution of these selected spaxels on the standard Baldwin–Phillips–Terlevich (BPT) diagram (Baldwin et al. 1981), which is widely used to remove contamination of active galactic nuclei (AGNs) from star-forming regions, is shown in Figure 1. We further select star-forming spaxels based on the Kauffmann et al. (2003) demarcation. The quality control⁵ provided by the Pipe3D VAC is also considered; only galaxies flagged as zero (i.e., do not find any problem) are included in our sample. The above criteria result in more than 450,000 spaxels from 2975 galaxies, which is taken as our final sample. The median and 68% range of the physical size of these spaxels is $0.20^{+0.54}_{-0.12}$ kpc².

Recent IFU studies show that there is a subgalactic main-sequence (SGMS) relation between the stellar mass surface density (Σ_*) and the SFR surface density (Σ_{SFR}) held down to subkiloparsec scale (e.g., Wuyts et al. 2013; Cano-Díaz et al. 2016; Hsieh et al. 2017; Liu et al. 2018). The SFR is computed from the dust-corrected $H\alpha$ luminosity based on the conversion of Kennicutt (1998b). We find that our spaxels form an evident SGMS that shows remarkable agreement with another two SGMSs (Hsieh et al. 2017; Liu et al. 2018) derived from the MaNGA data of SDSS DR14, which suggests that our sample well represents the population of subgalactic star-forming regions.

The stellar dust attenuation $A_{V,\text{star}}$, which is a product of the continuum modeling process described in Section 2.1, is taken from the Pipe3D VAC. Its statistical uncertainty propagated from

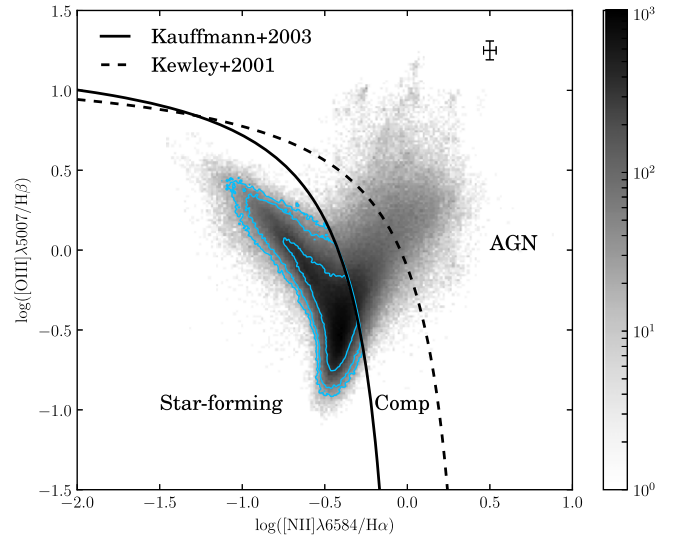


Figure 1. The BPT diagram of the selected spaxels. The dashed black curve represents the theoretical maximum starburst line of Kewley et al. (2001), and the solid black curve shows the demarcation between the pure star-forming region and the AGN region from Kauffmann et al. (2003). The gray region shows the density distribution of all of the selected spaxels, while the blue contours are drawn at 68%, 95%, and 99% of the star-forming spaxels only. The black error bar denotes the median uncertainties of all star-forming spaxels.

the uncertainties of the observed spectrum is derived by the Monte Carlo approach during the continuum modeling and is estimated to have a typical value of 0.007 mag. Therefore, the total uncertainty should be dominated by the systematic uncertainty. According to our S/N criterion, we expect a nearly unbiased $A_{V,\text{star}}$ estimation with a dispersion better than 0.26 mag (Sánchez et al. 2016a). The nebular dust attenuation $A_{V,\text{gas}}$ is a byproduct of the emission line dereddening in which the uncertainty propagation is carefully performed. Under the aforementioned assumptions, the median and 1σ dispersion of the statistical uncertainty of $A_{V,\text{gas}}$ are estimated be $0.08^{+0.07}_{-0.04}$ mag.

The stellar and nebular attenuation are both derived on the basis of the Cardelli et al. (1989) extinction curve. The choice of extinction/attenuation curve might affect the resulting attenuation values. However, for $A_{V,\text{star}}$, Sánchez et al. (2016a) claimed that no major differences were found when the Cardelli et al. (1989) extinction curve was replaced by the Calzetti et al. (2000) attenuation law or a $\lambda^{-1.3}$ extinction curve due to the small differences between these extinction/attenuation laws in the optical wavelength range. The replacement of the extinction/attenuation curve in the calculation of $A_{V,\text{gas}}$ only introduces a constant scaled factor. For these reasons, the choice of extinction/attenuation curve could alter the absolute values of $A_{V,\text{star}}/A_{V,\text{gas}}$, but the relative values, as well as most of our following results, should remain almost unchanged.

3. Relation between $A_{V,\text{star}}/A_{V,\text{gas}}$ and Local Properties

3.1. Correlate or Not?

We first examine whether $A_{V,\text{gas}}$ correlates with $A_{V,\text{star}}$. For the whole spaxel population, the Pearson linear and Spearman rank correlation coefficients are 0.40 and 0.34, respectively. However, if the spaxels are divided into bins of $\Sigma_{H\alpha}$, we find that the correlation coefficients of the binned subsamples strongly change with $\Sigma_{H\alpha}$.

⁴ <https://code.google.com/archive/p/astrolibpy/>

⁵ https://data.sdss.org/datamodel/files/MANGA_PIPE3D/MANGADRP_VER/PIPE3D_VER/QC_MaNGA.html

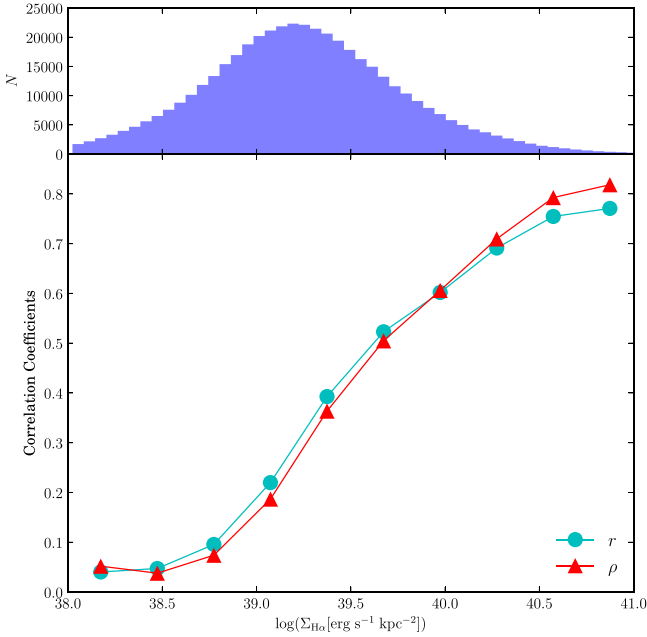


Figure 2. Pearson and Spearman rank correlation coefficients between $A_{V,\text{star}}$ and $A_{V,\text{gas}}$ as a function of $\Sigma_{\text{H}\alpha}$. The cyan circles denote the Pearson linear correlation coefficients, while the red triangles indicate the Spearman rank correlation coefficients. A histogram of $\Sigma_{\text{H}\alpha}$ is also given in the top panel.

We show how the Pearson (r) and Spearman rank (ρ) correlation coefficients between $A_{V,\text{star}}$ and $A_{V,\text{gas}}$ vary with $\Sigma_{\text{H}\alpha}$ in Figure 2. Obviously, both correlation coefficients display a substantial increase from $\lesssim 0.1$ at the low- $\Sigma_{\text{H}\alpha}$ end to ~ 0.8 at the high- $\Sigma_{\text{H}\alpha}$ end. Thus, the nebular attenuation indeed correlates with the stellar attenuation, although the strength of this correlation strongly correlates with the level of local star formation.

3.2. Correlation Analysis

In this subsection, we consider how the $A_{V,\text{star}}/A_{V,\text{gas}}$ ratio correlates with local physical properties, such as the $\text{H}\alpha$ luminosity (i.e., local SFR) and its surface brightness ($\Sigma_{\text{H}\alpha}$, i.e., Σ_{SFR}); stellar mass within each spaxel ($M_{*,\text{spaxel}}$) and its surface density (Σ_*); local sSFR and its observational indicator, equivalent width of $\text{H}\alpha$ ($\text{EW}_{\text{H}\alpha}$); stellar age tracer $D4000$ (Bruzual A., 1983); and several emission line indices.

The O3N2 and N2 indices, which are defined as

$$\text{O3N2} \equiv \log\left(\frac{[\text{O III}]\lambda 5007}{\text{H}\beta} \times \frac{\text{H}\alpha}{[\text{N II}]\lambda 6584}\right) \quad (1)$$

and $\text{N2} \equiv \log([\text{N II}]\lambda 6584/\text{H}\alpha)$, respectively, are widely used as diagnostics of gas-phase metallicity (e.g., Pettini & Pagel 2004; Kewley & Ellison 2008; Marino et al. 2013). Moreover, N2O2, defined as $\log([\text{N II}]\lambda 6574/[\text{O II}]\lambda 3727)$,⁶ is also considered as an excellent metallicity tracer (e.g., Dopita et al. 2000, 2013), especially when the DIG may have a nonnegligible contribution to emission lines (Zhang et al. 2017). However, the empirical calibrations between these indices and oxygen abundance are different among previous works, leading to discrepancies up to $\Delta[\log(\text{O}/\text{H})] = 0.7$ dex

⁶ Due to the limited spectral resolution of MaNGA, all $[\text{O II}]\lambda 3727$ mentioned in this work represents the sum of the $[\text{N II}]\lambda 3726, 3729$ doublet.

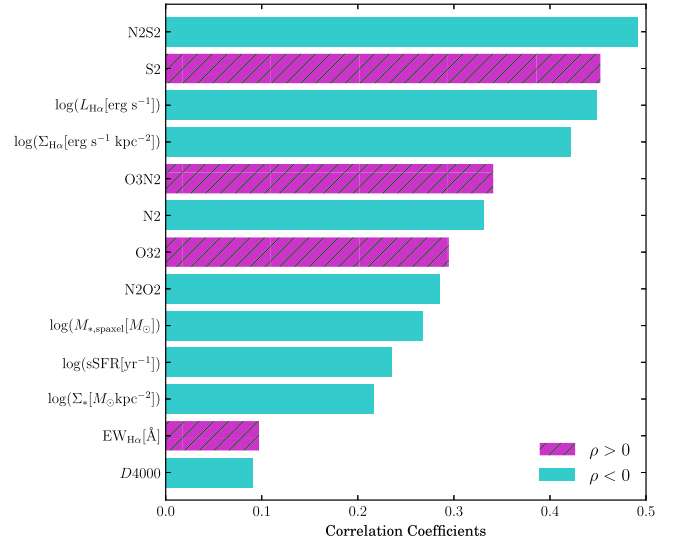


Figure 3. Spearman rank correlation coefficients between the $A_{V,\text{star}}/A_{V,\text{gas}}$ ratio and several local properties. The hatched magenta bars present positive correlation coefficients, while the cyan bars indicate the absolute value of negative correlation coefficients. The order of these parameters is sorted by the absolute value of their correlation coefficients.

(Kewley & Ellison 2008). Lin et al. (2017) also reported an inconsistency between metallicities derived from O3N2 and N2; even the calibrations were derived based on the same data set. For these reasons, here we directly use O3N2, N2, and N2N2 as indicators of metallicity but do not apply any empirical calibrations.

Given that the status of ionized gas may play an important role in determining the $A_{V,\text{star}}/A_{V,\text{gas}}$ ratio (Kreckel et al. 2013), we also include DIG indicators in the analysis. Besides the $[\text{S II}]\lambda 6717/\text{H}\alpha$ ratio, $\Sigma_{\text{H}\alpha}$ (Zhang et al. 2017) and $\text{EW}_{\text{H}\alpha}$ (Lacerda et al. 2018) are also used to separate H II- and DIG-dominated regions. Furthermore, as suggested by M. Lin et al. (2020, in preparation), the $[\text{N II}]/\text{H}\alpha$ versus $[\text{S II}]/\text{H}\alpha$ diagram can be a useful tool to infer the ionization source of DIGs. Thus, we define $\text{S2} \equiv \log([\text{S II}]\lambda 6717/\text{H}\alpha)$ and $\text{N2S2} \equiv \log([\text{N II}]\lambda 6584/[\text{S II}]\lambda 6717)$ and add these two indices in our correlation analysis, together with $\text{O32} \equiv \log([\text{O III}]\lambda 5007/[\text{O II}]\lambda 3727)$, to distinguish differences in ionization parameter (Baldwin et al. 1981; Dopita et al. 2013). Figure 3 shows the Spearman rank correlation coefficients (ρ) between the $A_{V,\text{star}}/A_{V,\text{gas}}$ ratio and all of the above parameters.

The DIG indicator of N2S2 gives the strongest correlation with the $A_{V,\text{star}}/A_{V,\text{gas}}$ ratio. However, its physical meaning is not clear at present; we defer discussion of this index to Section 3.5. The other two DIG tracers of S2 and $\Sigma_{\text{H}\alpha}$ also present moderate ($|\rho| > 0.4$) correlations with the attenuation ratio, which suggest that emission line regions with smaller $\Sigma_{\text{H}\alpha}$ (more diffuse) and higher $[\text{S II}]\lambda 6717$ -to- $\text{H}\alpha$ ratio (smaller ionization parameter; e.g., Haffner et al. 2009) would have larger $A_{V,\text{star}}/A_{V,\text{gas}}$ values. In other words, at fixed $A_{V,\text{star}}$, DIG regions tend to suffer less nebular attenuation with respect to classical H II regions. Such correlations are in good agreement with the result of Kreckel et al. (2013) and can be explained by the two-component dust model of Charlot & Fall (2000). In this model, the H II-dominated regions are embedded in dense and dustier clouds in which massive stars are born, surrounded by the diffuse ISM, where most of the DIG-dominated regions are located. Irrespective of what provides the

ionizing photons of the DIG, leaky H II regions (e.g., Haffner et al. 2009), or hot low-mass evolved stars (e.g., Flores-Fajardo et al. 2011; Zhang et al. 2017), the possible excitation mechanisms of the DIG imply that these regions should be more uniformly mixed with the diffuse ISM.

However, $EW_{H\alpha}$ (and sSFR), which is also considered an alternative DIG indicator by Lacerda et al. (2018), contrary to the other three DIG tracers, shows a much weaker correlation with the attenuation ratio. As elucidated in Zhang et al. (2017), the dependence of $EW_{H\alpha}$ on metallicity makes it a poor tracer of DIG in some cases, resulting in an almost flat distribution together with a very large scatter on the $EW_{H\alpha}$ (sSFR) versus $A_{V,star}/A_{V,gas}$ plane for our sample. By comparison, we present the contour of spaxel density on the $\Sigma_{H\alpha}$ – $A_{V,star}/A_{V,gas}$ plane in Figure 9. Obviously, the dispersion of spaxels significantly decreases with increasing $\Sigma_{H\alpha}$, suggesting the $A_{V,star}/A_{V,gas}$ ratio converges to a constant that close to the value of Calzetti (1997) at the high- $\Sigma_{H\alpha}$ end. The different behavior between $\Sigma_{H\alpha}$ (or $L_{H\alpha}$) and $EW_{H\alpha}$ implies that the dust attenuation of local star-forming regions is more sensitive to the current star formation activity rather than the past one, which is also implied by the weaker correlation between $A_{V,star}/A_{V,gas}$ and $M_{*,spaxel}$ (or Σ_*).

Moderate correlations are also found for emission line indices of both metallicity and ionization. Given the positive (negative) relation between oxygen abundance and the N2, N2O2 (O3N2) index (e.g., Kewley & Ellison 2008; Dopita et al. 2013; Marino et al. 2013), the correlations presented in Figure 3 imply that the ratio between $A_{V,star}$ and $A_{V,gas}$ tends to be smaller in more metal-rich regions. Moreover, regions with a higher ionized level (larger O32) tend to have a larger $A_{V,star}/A_{V,gas}$ ratio. We will return to these two parameters with a more detailed discussion in Section 3.3. The weakest correlation is shown by D4000, suggesting that $A_{V,star}/A_{V,gas}$ is stellar age-insensitive on the physical scale of our spaxels.

The Pearson correlation coefficients between the $A_{V,star}/A_{V,gas}$ ratio and these properties are also calculated. However, all of them are found to be smaller than the corresponding Spearman rank correlation coefficients, indicating weaker linear correlations between the $A_{V,star}/A_{V,gas}$ ratio and the explored parameters.

3.3. Correlation with Physical Condition of Ionized Gas

As mentioned in Section 3.2, the physical condition of ionized gas, such as metallicity and ionized level, may help us determine the dust attenuation ratio. However, it is hard to draw reliable conclusions from Figure 3 alone due to the small correlation coefficients. On the one hand, although we find moderate correlations between the $A_{V,star}/A_{V,gas}$ ratio and metallicity indicators, the relation with metallicity still has a large uncertainty because of the nonnegligible role of ionization in emission line ratios (Dopita et al. 2013). For example, Mao et al. (2018) suggested that the radial gradients of O3N2 and N2 indices across H II regions can be attributed to the variations of ionized level rather than the metallicity. On the other hand, there is a contradiction that a larger $A_{V,star}/A_{V,gas}$ ratio is preferred for highly ionized gas (i.e., large O32 index), as well as DIG, for which a smaller ionization parameter compared to classic H II regions is always suggested (e.g., Haffner et al. 2009). To better understand the role of metallicity and ionization parameter in determining the $A_{V,star}/A_{V,gas}$ ratio and reconciling the seemingly

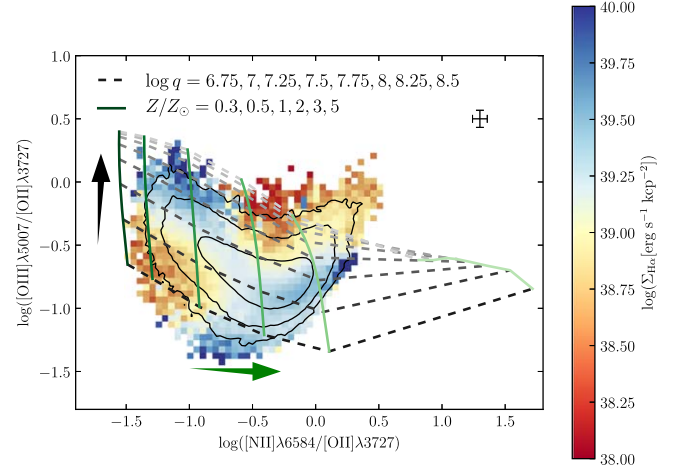


Figure 4. Median map of $\Sigma_{H\alpha}$ on the N2O2 vs. O32 diagram. The black solid lines indicate contours encompassing 68%, 95%, and 99% of all spaxels. The grid consisting of green solid lines and black dashed lines is an ionization model of the spherical H II region with $\kappa = \infty$ from Dopita et al. (2013), where the green solid lines represent constant metallicities and the black dashed lines represent constant ionization parameters. The colors of these lines become transparent as the values increase. The horizontal green arrow denotes the direction of increasing metallicity Z , and the vertical black arrow denotes the direction of increasing ionization parameter q . The values of Z and $\log q$ plotted in this diagram are also listed. The typical uncertainties of N2O2 and O32 are shown in the upper right corner.

conflicting observations, the connection between these two parameters should be disentangled.

3.3.1. N2O2 versus O32 Diagram

We first explore how the DIG tracer ($\Sigma_{H\alpha}$) distributes in the metallicity–ionization parameter space. The median map of $\Sigma_{H\alpha}$ on the N2O2 versus O32 diagram for all spaxels is shown in Figure 4, overlapping with density contour (black solid lines). In addition, we also plot the grid of the ionization model of the spherical H II region with $\kappa = \infty$ ⁷ from Dopita et al. (2013), in which emission line fluxes at different metallicities Z and different ionization parameters $\log(q)$ are computed. This diagram clearly separates the metallicity from the ionization parameter and will be suitable for examining the respective effect of abundance or ionization when the other is fixed. As we expected, most of the selected spaxels are covered by the theoretical grid of the H II region, except for a branch perpendicular to the main star-forming cloud at the upper right corner, which is identified as an AGN sequence (Dopita et al. 2013). At the metal-poor end, there is a trend that the median $\Sigma_{H\alpha}$ increases with increasing ionization. However, the trend apparently inverts at the metal-rich end, possibly due to the contamination from AGN-dominated spaxels.

We further separate all spaxels into H II- and DIG-dominated⁸ populations adopting the criterion of Zhang et al. (2017), i.e.,

⁷ Here κ describes how the energy distribution of electrons in H II regions departs from a standard Maxwell–Boltzmann equilibrium energy distribution, while $\kappa = \infty$ represents the Maxwell–Boltzmann distribution (Nicholls et al. 2012; Dopita et al. 2013).

⁸ The distribution of all DIG-dominated spaxels on the BPT diagram might extend across the Kauffmann et al. (2003) boundary (Zhang et al. 2017) and covers the composite or even the AGN regime of Figure 1, where it is already excluded by our selection of star-forming spaxels. Thus, the DIG-dominated population in this work is incomplete. The following analysis of DIG-dominated spaxels is restricted to this subsample that resides in the star-forming regime.

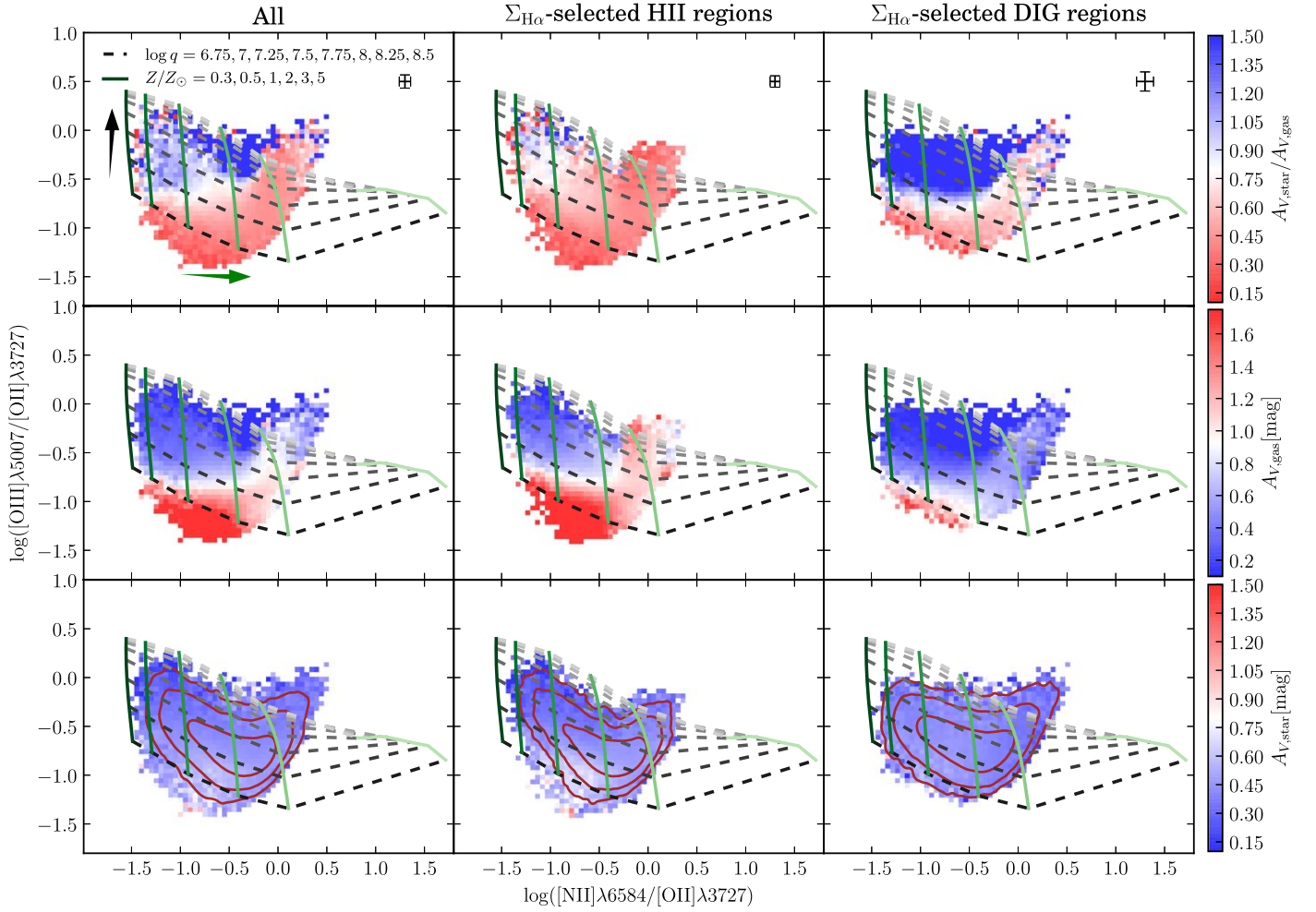


Figure 5. Median maps of $A_{V,\text{star}}/A_{V,\text{gas}}$ (top), $A_{V,\text{gas}}$ (middle), and $A_{V,\text{star}}$ (bottom) on the N2O2 vs. O32 diagram. In each case (row), we plot results for all spaxels (left), H II-dominated regions (middle), and DIG-dominated regions (right). The overlapping Z - q grid is the same as Figure 4. The contours overplotted in the bottom row enclose 68%, 95%, and 99% of the spaxels on each diagram. The typical uncertainties of N2O2 and O32 for each population are shown at the upper right corner of the top row.

H II-dominated spaxels with $\Sigma_{\text{H}\alpha} > 10^{39} \text{ erg s}^{-1} \text{ kpc}^{-2}$ and DIG-dominated spaxels with $\Sigma_{\text{H}\alpha} < 10^{39} \text{ erg s}^{-1} \text{ kpc}^{-2}$. Kreckel et al. (2013) adopted $[\text{S II}]\lambda 6717/\text{H}\alpha = 0.2$ to separate the spaxels into these two populations, which is based on observations in the Milky Way (Madsen et al. 2006). However, due to the beam-smearing effect resulting from the low spatial resolution of MaNGA (Zhang et al. 2017), as well as the metallicity dependence of the $[\text{S II}]\lambda 6717/\text{H}\alpha$ ratio, this simple boundary cannot be directly applied to the MaNGA data.

According to the $\Sigma_{\text{H}\alpha}$ criterion, the fractions of spaxels that are classified as H II- and DIG-dominated regions are 68.2% and 31.8%, respectively. Two DIG-dominated loci are found in Figure 4; one is at the low- Z , low- q regime, and the other is the AGN sequence. The distributions of deprojected galactocentric distances of spaxels in both loci span a wide range and peak at ~ 1.5 and ~ 2.5 kpc for the AGN sequence and the low- Z , low- q locus, respectively. In combination with the properties of the underlying low-S/N and low- $\Sigma_{\text{H}\alpha}$ population shown in the Appendix, we argue that these loci should be predominantly attributed to LI(N)ER-like regions or DIG excited by hot, low-mass, evolved stars. The coverage of the star-forming cloud is in good agreement with the distribution of SDSS galaxies from

Kewley et al. (2006) shown in Dopita et al. (2013) and extends from the low- Z , high- q end to the high- Z , low- q end.

3.3.2. Dust Attenuations on the N2O2 versus O32 Diagram

To see how dust attenuation (and ratio) vary with Z and q , we present the median map of $A_{V,\text{star}}/A_{V,\text{gas}}$, $A_{V,\text{gas}}$, and $A_{V,\text{star}}$ for all of the spaxels and the H II- and DIG-dominated spaxels in Figure 5, overplotting with the theoretical grid of star-forming regions from Dopita et al. (2013). The medians and 68% limits of $A_{V,\text{star}}/A_{V,\text{gas}}$ are $0.55^{+0.68}_{-0.26}$, $0.48^{+0.32}_{-0.22}$, and $0.93^{+7.89}_{-0.56}$ for all spaxels, H II-dominated, and DIG-dominated, respectively. As a comparison, Kreckel et al. (2013) found the $A_{V,\text{star}}/A_{V,\text{gas}}$ ratios of 0.5 and 0.7 for their selected H II- and DIG-dominated regions, respectively, which are consistent with ours.

Although the ionization level of DIG regions is found to be systematically lower compared to that of H II regions, the large overlap of the contours of both subsamples with the theoretical grid in Figure 5 indicates that both of them cover a similar and wide range in parameter space of either metallicity or ionization. Comparing the subplots in Figure 5 column by column, the overall features are very similar for all three populations except

for the systematic shift of absolute value between them. Comparing the subplots row by row, it is evident that the overall trend of the $A_{V,\text{star}}/A_{V,\text{gas}}$ ratio is dominated by that of $A_{V,\text{gas}}$, while the $A_{V,\text{star}}$ maps exhibit a relatively uniform distribution on the N2O2 versus O32 diagram (i.e., the metallicity–ionization parameter grid), regardless of the population.

The roles of gas-phase metallicity and ionization level in determining $A_{V,\text{star}}/A_{V,\text{gas}}$ are more obvious on this plot. For the $A_{V,\text{star}}/A_{V,\text{gas}}$ ratio, all three populations show a significantly higher value at the low- Z , high- q corner, leading to the moderate correlations between $A_{V,\text{star}}/A_{V,\text{gas}}$ and metallicity–ionization parameter found in Section 3.2. As for $A_{V,\text{gas}}$, the spaxels occupying the low- Z , high- q corner suffer smaller dust attenuation, indicating that $A_{V,\text{gas}}$ decreases with Z ($\log q$) being lower (higher). This trend might be a geometric result if the low- $A_{V,\text{gas}}$ spaxels are in the outer disks of their host galaxies. We adopt $\text{N2O2} < 0$ and $\text{O32} > -0.5$ to select the low- $A_{V,\text{gas}}$ regime and use the deprojected galactocentric distances normalized by the effective radii of host galaxies (i.e., r/R_e) to measure their location on the disks. We find that the r/R_e distribution of the low- $A_{V,\text{gas}}$ subsample is similar to that of the whole sample. However, spaxels in this low- $A_{V,\text{gas}}$ regime tend to reside in less massive galaxies compared to the whole sample. Therefore, the trend of the low- $A_{V,\text{gas}}$ subsample with lower metallicity and higher ionization parameter does not result from the geometric effect and is mainly due to the smaller M_* of their host galaxies.

Intriguingly, the low- Z , high- q corner of the DIG-dominated spaxels has $A_{V,\text{star}}/A_{V,\text{gas}}$ larger than 1, indicating that the stellar light is more attenuated than the emission lines. Spaxels in this regime only account for nearly 16% of the DIG-dominated population. Comparing with the H II-dominated spaxels in the same regime, these DIG-dominated spaxels have comparable $A_{V,\text{star}}$ but smaller $A_{V,\text{gas}}$, suggesting that $A_{V,\text{star}}/A_{V,\text{gas}} > 1$ mainly results from a smaller $A_{V,\text{gas}}$. The distributions of these spaxels on several emission line diagrams are consistent with the model predictions of leaky H II regions presented in Zhang et al. (2017; their Figure 18). Meanwhile, the $D4000$ and $\text{EW}_{\text{H}\alpha}$ of these DIG-dominated spaxels are also comparable to normal star-forming regions. Taking into account the relatively higher ionization level of these spaxels, we argue that this DIG might be attributed to the leaky H II region scenario in which ionized gas leaks from the dusty H II regions, while the stellar continuum is dominated by the background stellar populations that are physically unassociated with the leaky H II regions, leading to $A_{V,\text{star}} > A_{V,\text{gas}}$.

In summary, the smaller value of the ionization parameter of DIG regions with respect to H II regions is a systematic effect, within which they still have a large dispersion in this parameter, resulting in a large scatter of the local $A_{V,\text{star}}/A_{V,\text{gas}}$ ratio. Furthermore, the influences of metallicity and ionization level are independent of the classification of spaxels. Metal-poor regions with large $\log q$ tend to have smaller nebular attenuation and thus larger $A_{V,\text{star}}/A_{V,\text{gas}}$ values for either the H II- or DIG-dominated subsample. These two effects may originate from different aspects of dust attenuation, where the former reflects the different stars/dust geometry for H II- and DIG-dominated populations (Kreckel et al. 2013), and the latter implies the difference in dust properties (e.g., the size distribution of dust grains; Relaño et al. 2018) or the dust-to-gas ratio ($D/G \equiv m_{\text{dust}}/m_{\text{gas}}$; Engelbracht et al. 2008;

Galliano et al. 2011; Rémy-Ruyer et al. 2014) in different local environments.

3.3.3. Possible Reason for Correlations between $A_{V,\text{star}}/A_{V,\text{gas}}$ and Physical Conditions

Although most of the previous works attributed the integrated $A_{V,\text{star}}/A_{V,\text{gas}}$ to the geometry effect (e.g., Price et al. 2014; Reddy et al. 2015; Koyama et al. 2019), the conclusions drawn from Figure 5 imply that the local physical conditions of ionized gas may also have a significant effect on the $A_{V,\text{star}}/A_{V,\text{gas}}$ ratio. The dust attenuation at wavelength λ (A_λ) along the line of sight s can be written as

$$\begin{aligned} A_\lambda &= 1.086\tau_\lambda^{\text{att}} = 1.086 \times \int_s \kappa_\lambda^{\text{att}} \rho_{\text{dust}}(s) ds \\ &= 1.086 \times \int_s \kappa_\lambda^{\text{att}} \rho_{\text{gas}}(s) \cdot D/G ds, \end{aligned} \quad (2)$$

in which $\tau_\lambda^{\text{att}}$ is the attenuation optical depth at wavelength λ , $\kappa_\lambda^{\text{att}}$ is the corresponding total opacity (including the absorption and scattering effects), and $\rho_{\text{dust}}(s)$ and $\rho_{\text{gas}}(s)$ are the mass density of dust and gas along the line of sight, respectively. In fact, the term “extinction” should be distinguished from “attenuation.” The former always refers to dust effects for individual pointlike sources and includes absorption and scattering out of the line of sight, whereas the latter represents the net loss of light for extended emitters, including extinction and scattering into the line of sight in which the stars/dust geometry plays an important role (see Calzetti 2001 for a review). Therefore, the $\tau_\lambda^{\text{att}}$ we used here is something like an effective optical depth rather than the center-to-edge optical depth of a cloud that describes the amount of dust, and the $\kappa_\lambda^{\text{att}}$ depends on not only the dust properties but also the stars/dust geometry.

Here we further express A_λ as a function of ρ_{gas} and D/G . In the case of $A_{V,\text{star}}$ and $A_{V,\text{gas}}$, the differences between them that can be explained as geometry effects due to the two-component dust model (Charlot & Fall 2000; Wild et al. 2011; Chevillard et al. 2013) are in $\kappa_\lambda^{\text{att}}$, ρ_{gas} , and the integrated range s . However, variations in dust properties, including in $\kappa_\lambda^{\text{att}}$, and D/G can also change A_λ and thus the $A_{V,\text{star}}/A_{V,\text{gas}}$ ratio.

The metallicity is found to be the main driver of the variation of D/G (Engelbracht et al. 2008; Asano et al. 2013; Rémy-Ruyer et al. 2014), even within individual galaxies (e.g., Chiang et al. 2018). Given that the D/G increases with galaxies being more metal-rich, $A_{V,\text{gas}}$ is expected to be higher toward the metal-rich end at a fixed $\log q$, which is in good agreement with what we observe in Figure 5. However, we should note that the gas-phase metallicity measured here does not directly reflect the metal enrichment of the diffuse ISM in most of the line of sight for classical H II regions (except for the small portion surrounding H II regions). Besides, due to the small ρ_{gas} , the dust accretion is less efficient in the diffuse ISM (Draine 2009). As a result, the impact of the emission line–based metal enrichment on the diffuse ISM is much weaker. Hence, $A_{V,\text{star}}$ is not significantly altered by observed gas-phase metallicity, and only a very weak trend along constant $\log q$ is exhibited for the H II-dominated population in Figure 5.

Furthermore, Galliano et al. (2011) found that the D/G decreases with the interstellar radiation field (ISRF) being stronger in the Large Magellanic Cloud, mainly due to the

enhanced destruction of dust grains by shock. A similar but much weaker trend is also hinted at by Relaño et al. (2018) for M33. For this reason, the enhancement of the ISRF (elevation of $\log q$) would lead to a smaller $A_{V,\text{gas}}$, making the low- Z , high- q regime more transparent.

Therefore, the influences of the local physical conditions on either dust properties or D/G are one plausible explanation for the correlations between $A_{V,\text{star}}/A_{V,\text{gas}}$ and the metallicity–ionization parameter. The feature of Figure 5 is consistent with a picture in which less dust is surviving to shield the light in metal-poor ionized gas with a high ionized level; the resultant low nebular attenuation leads to a high $A_{V,\text{star}}/A_{V,\text{gas}}$ value.

3.4. Correlation with Location in BPT Diagrams

We have shown that the $A_{V,\text{star}}/A_{V,\text{gas}}$ value depends on the physical conditions of the ionized gas in a somewhat complex manner. Here we continue to study how this property correlates with the location of spaxels on the BPT diagrams that are always used to identify the excitation mechanism of the ionized gas (e.g., Kewley et al. 2001; Kauffmann et al. 2003; Kewley et al. 2006).

We present the median maps of $A_{V,\text{star}}/A_{V,\text{gas}}$, $A_{V,\text{gas}}$, and $A_{V,\text{star}}$ on the BPT diagrams for all selected spaxels in Figure 6. Predictions of the theoretical photoionization model of the spherical H II region with $\kappa = \infty$ from Dopita et al. (2013) are also plotted for all three diagnostic diagrams. Although the [N II]/H α –[O III]/H β BPT diagram (hereafter [N II]–[O III] diagram) combined with the Kauffmann et al. (2003) boundary is used to exclude spaxels with possible AGN contribution, we still find a small amount of spaxels lying beyond the Kewley et al. (2006) boundaries of star-forming regions on both the [S II]/H α versus [O III]/H β and [O I]/H α versus [O III]/H β diagrams (hereafter [S II]–[O III] and [O I]–[O III] diagram, respectively).

It is very interesting to find that the dust attenuation (and ratio) show some patterns on the BPT diagrams. Similar to the N2O2 versus O32 diagram, the maps of $A_{V,\text{star}}$ on all three diagrams are fairly smooth, so that the overall trends of the $A_{V,\text{star}}/A_{V,\text{gas}}$ ratio are dominated by those of the $A_{V,\text{gas}}$, regardless of the diagrams. Notably, there is an intriguing low- $A_{V,\text{gas}}$ regime and a consequent but more distorted high- $A_{V,\text{star}}/A_{V,\text{gas}}$ sequence on each diagnostic diagram. The low- $A_{V,\text{gas}}$ and high- $A_{V,\text{star}}/A_{V,\text{gas}}$ regime occupies the low-[N II]/H α region on the [N II]–[O III] diagram, the Kewley et al. (2006) boundary on the [S II]–[O III] diagram, and the upper right region, which also extends across the Kewley et al. (2006) boundary, on the [O I]–[O III] diagram. On the one hand, the closer locations of the low- $A_{V,\text{gas}}$ sequences to the Kewley et al. (2006) boundaries on both the [S II]–[O III] and [O I]–[O III] diagrams imply that the excitation sources of these spaxels may be something other than massive young stars, so that they would be more DIG-like and mixed well with the diffuse ISM. On the other hand, observations from either the Milky Way (Madsen et al. 2006) or nearby galaxies (e.g., Hoopes & Walterbos 2003; Kaplan et al. 2016; Zhang et al. 2017; Lacerda et al. 2018) revealed enhancement of [S II]/H α and [N II]/H α ratios in DIG regions compared to classical H II regions, indicating a lower ionization level of DIG. Although this low- $A_{V,\text{gas}}$ population indeed extends to the regime with a high [S II]/H α ratio, it seems to exhibit a smaller [N II]/H α ratio. The location of the sequence on the [N II]–[O III] BPT diagram differs from that of the DIG-dominated regions selected by either low- $\Sigma_{\text{H}\alpha}$ (Zhang et al. 2017) or low- $\text{EW}_{\text{H}\alpha}$ (Lacerda et al. 2018), as well as the theoretical prediction

of shocked gas from the photoionization model (Alatalo et al. 2016). We stress that such low $A_{V,\text{gas}}$ at the small-[N II]/H α end does not arise from the relatively larger uncertainties of DIG-dominated spaxels, since the S/Ns of both H α and H β for the small-[N II]/H α spaxels are slightly higher than those of the large-[N II]/H α spaxels if we simply separate the sample by $\log([\text{N II}]/\lambda 6584/\text{H}\alpha) = -0.75$. Therefore, the low- $A_{V,\text{gas}}$ and high- $A_{V,\text{star}}/A_{V,\text{gas}}$ sequences cannot be explained by DIG alone. We will show later that these sequences result from the combination of the DIG-dominated and metal-poor spaxels in which the latter ones lead to the low-[N II]/H α values on the [N II]–[O III] diagram.

Moreover, the theoretical grids of both the [N II]–[O III] and [S II]–[O III] diagrams are able to cover most of the spaxels in our sample, whereas the grid on the [O I]–[O III] diagram only occupies about half of the observed regime. As elucidated in Dopita et al. (2013), their calculations of the [O I] line are much less reliable than those of the [N II] and [S II] lines due to the fact that the [O I] emission comes out of a very narrow zone near the ionization front in which shocks may be an important heating source (Dopita 1997). Thus, the [O I]–[O III] grid should be treated with caution. As can be seen in Figure 6, the low- $A_{V,\text{gas}}$ and high- $A_{V,\text{star}}/A_{V,\text{gas}}$ sequences on these diagrams are all overlaid by the low- Z ends of the corresponding grids. In Section 3.3, we already showed that metal-poor spaxels tend to have smaller $A_{V,\text{gas}}$ and higher $A_{V,\text{star}}/A_{V,\text{gas}}$ ratios, irrespective of the populations we defined. Accordingly, metallicity may be another main factor to explain the location of the low- $A_{V,\text{gas}}$ and high- $A_{V,\text{star}}/A_{V,\text{gas}}$ sequences in the BPT diagrams.

To further separate the effects of DIG and metallicity, we plot similar median maps for the H II-dominated subsample in Figure 7. The most significant difference between Figures 6 and 7 is the truncation of the low- $A_{V,\text{gas}}$ and high- $A_{V,\text{star}}/A_{V,\text{gas}}$ sequences on the [S II]–[O III] and [O I]–[O III] diagrams. The removal of DIG contamination leads to the disappearance of the sequence at the low-[O III]/H β , high-[S II]/H α ([O I]/H α) regime on the [S II]–[O III] ([O I]–[O III]) diagram, indicating the low ionization state of DIG (e.g., Hoopes & Walterbos 2003; Haffner et al. 2009; Kaplan et al. 2016; Zhang et al. 2017). The remaining blue sequences in the H II-dominated population can be well described by the low- Z end of the theoretical grids, even on the “unreliable” [O I]–[O III] diagram. Interestingly, the extended region of the H II-dominated subsample is now well constrained by the Kewley et al. (2006) boundary on the [S II]–[O III] diagram.

In short, DIG contamination alone cannot explain the low- $A_{V,\text{gas}}$ and high- $A_{V,\text{star}}/A_{V,\text{gas}}$ sequences found on the [S II]–[O III] and [O I]–[O III] diagrams; metal-poor spaxels seem to be the origin of the remaining half of the sequences after removing DIG-dominated regions. This result reinforces the conclusions drawn from the analysis of the metallicity and ionization parameter in Section 3.3.

3.5. N2S2– $A_{V,\text{star}}/A_{V,\text{gas}}$ Correlation

In Section 3.2, we find that the N2S2 index has the strongest correlation with the $A_{V,\text{star}}/A_{V,\text{gas}}$ ratio among the local properties we explored. We use the predictions of the theoretical model of H II regions from Dopita et al. (2013) to examine how the N2S2 index varies with metallicity and ionization parameter and find that this index increases toward high- Z (high- q) end when q (Z) is fixed within the adopted parameter ranges (i.e., $Z = 0.1Z_{\odot}$ – $5Z_{\odot}$ and $\log q = 6.5$ – 8.5).

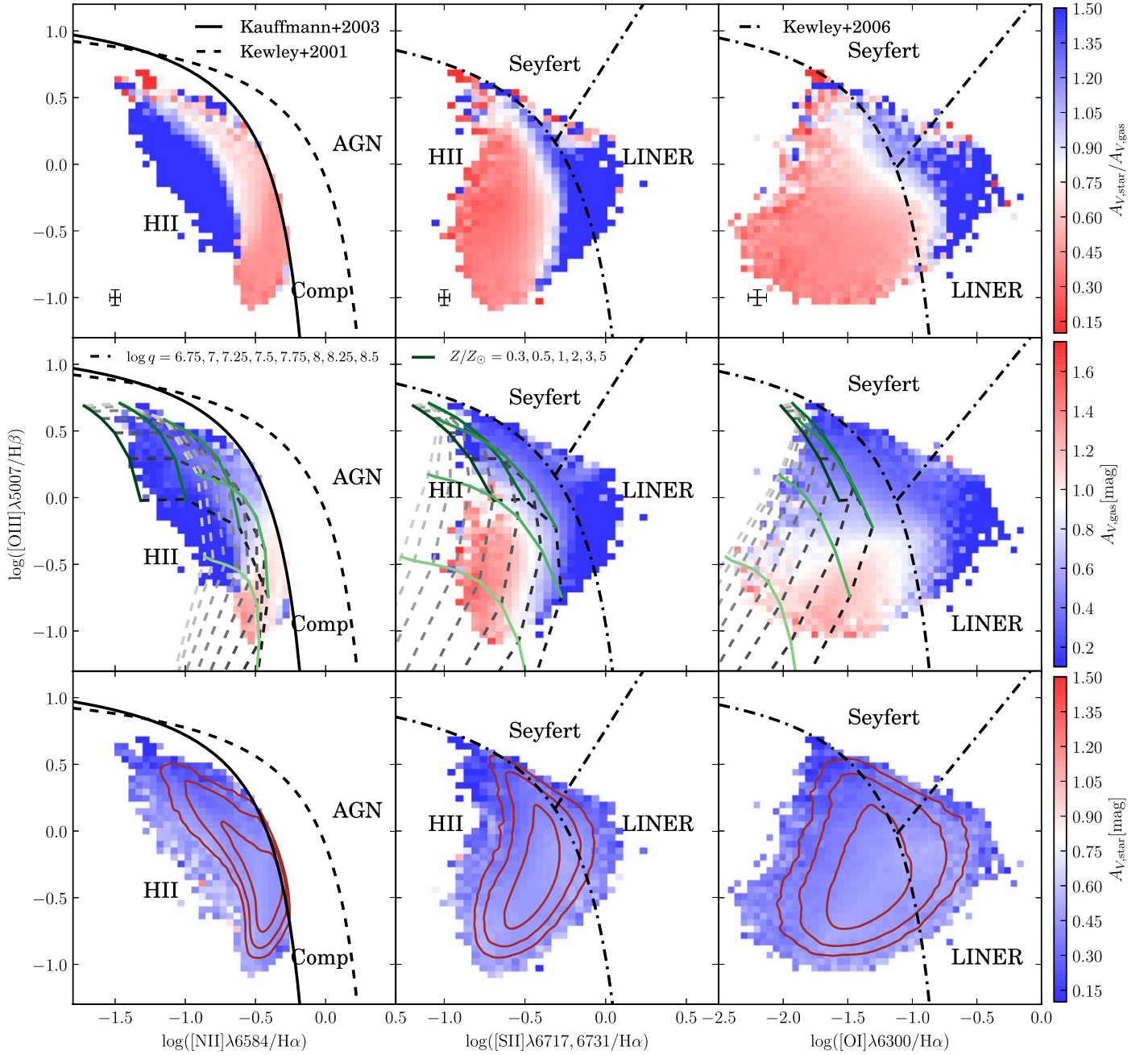


Figure 6. Median maps of $A_{V,star}/A_{V,gas}$ (top), $A_{V,gas}$ (middle), and $A_{V,star}$ (bottom) on the BPT diagrams for all selected spaxels. On the $[N II]/H\alpha$ vs. $[O III]/H\beta$ diagnostic diagrams, the dashed curves indicate the extreme starburst classification line of Kewley et al. (2001), while the solid curves are the division between pure star-forming and AGN-H II composite regions from Kauffmann et al. (2003). On the $[S II]/H\alpha$ vs. $[O III]/H\beta$ and $[O I]/H\alpha$ vs. $[O III]/H\beta$ diagnostic diagrams, the dashed-dotted curves are demarcations between star-forming regions, LINERs, and Seyfert galaxies from Kewley et al. (2006). The grid consisting of green solid lines (constant Z) and black dashed lines (constant $\log q$) denotes the predicted line ratios of the theoretical photoionization model of a spherical H II region with $\kappa = \infty$ from Dopita et al. (2013). The colors of these lines become faint as the value increases (see the directions of the arrows in Figure 4). The values of Z and $\log q$ plotted in this diagram are also given in the middle row. The contours overplotted in the bottom row encompass 68%, 95%, and 99% of the spaxels on each diagram. The typical uncertainties are shown in the lower left corner of the top row.

Observations of ionized gas also unveiled that DIG tends to have a smaller N2S2 ratio relative to the average value of traditional H II regions excited by massive OB stars due to the softer ionized field (Madsen et al. 2006; Haffner et al. 2009). Namely, the N2S2 ratio is sensitive to both metallicity and ionization parameter; we thus use the N2S2–O32 diagrams to separate the effect of the ionization parameter. In Figure 8, we present the median maps of $A_{V,star}/A_{V,gas}$ and $A_{V,gas}$ on the N2S2–O32 diagram for all spaxels and the H II-dominated

ones. Theoretical predictions of the ionization model of H II regions from Dopita et al. (2013) are also plotted. The median map of $A_{V,star}$ is not shown here due to the fact that, similar to the maps in other emission line ratio diagrams (e.g., Figures 5 and 6), it exhibits a uniform distribution across the diagram and does not reveal any more useful information.

The observed trends of $A_{V,star}/A_{V,gas}$ are dominated by those of $A_{V,gas}$. With the help of the O32 index, the role of the ionization parameter is well separated. We find that the N2S2

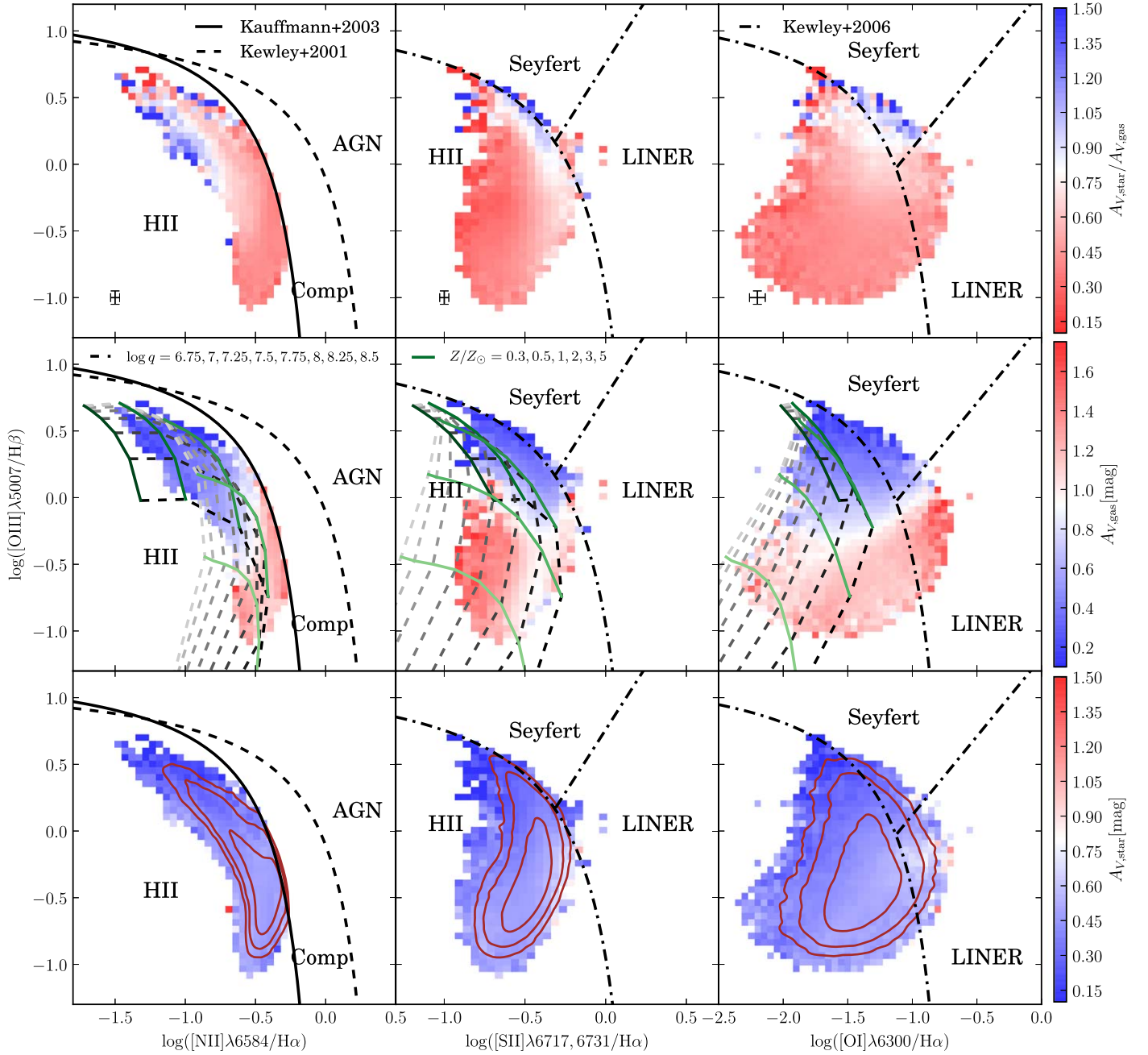


Figure 7. Median maps of $A_{V,star}/A_{V,gas}$ (top), $A_{V,gas}$ (middle), and $A_{V,star}$ (bottom) on the BPT diagrams for H II-dominated spaxels. The symbols are the same as in Figure 6.

index decreases toward the low-metallicity end at fixed $\log q$, leading to a larger $A_{V,star}/A_{V,gas}$ (smaller $A_{V,gas}$). Although DIG tends to have smaller N2S2 values, the removal of the DIG-dominated population cannot totally flatten the trend between the N2S2 index and $A_{V,star}/A_{V,gas}$, while the correlation between the N2S2 index and $A_{V,gas}$ is still significant ($\rho = 0.57$) for the H II-dominated population. Therefore, we ascribe the emergence of the N2S2– $A_{V,star}/A_{V,gas}$ correlation to the combination of DIG contamination (i.e., geometry effect) and the small N2S2 values of metal-poor spaxels.

3.6. Effect of Spatial Resolution

The two-component dust model, which we use to describe the role of geometry in determining the expected relation

between $A_{V,star}$ and $A_{V,gas}$, presumes that all line emission arises from H II regions (Charlot & Fall 2000). However, this presumption might not be true for all of our spaxels. Although the physical sizes of our spaxels have a median of 0.20 kpc^2 , the reconstructed PSF of the MaNGA data cubes has an FWHM of $\sim 2''.5$ (Law et al. 2016), corresponding to a physical size of $\sim 1.5 \text{ kpc}$ at the mean redshift of the MaNGA targets ($\langle z \rangle = 0.03$; Bundy et al. 2015). This scale is significantly larger than the typical size of the H II regions (Hunt & Hirashita 2009; Lawton et al. 2010), creating a problem regarding how the low spatial resolution of the MaNGA data cubes impacts the measured dust attenuations.

Based on a multiscale study of M33, Boquien et al. (2015) found that both nebular ($A_{H\alpha}$) and stellar (A_{FUV}) attenuation are scale-dependent. As the resolution downgrades (from 33 to

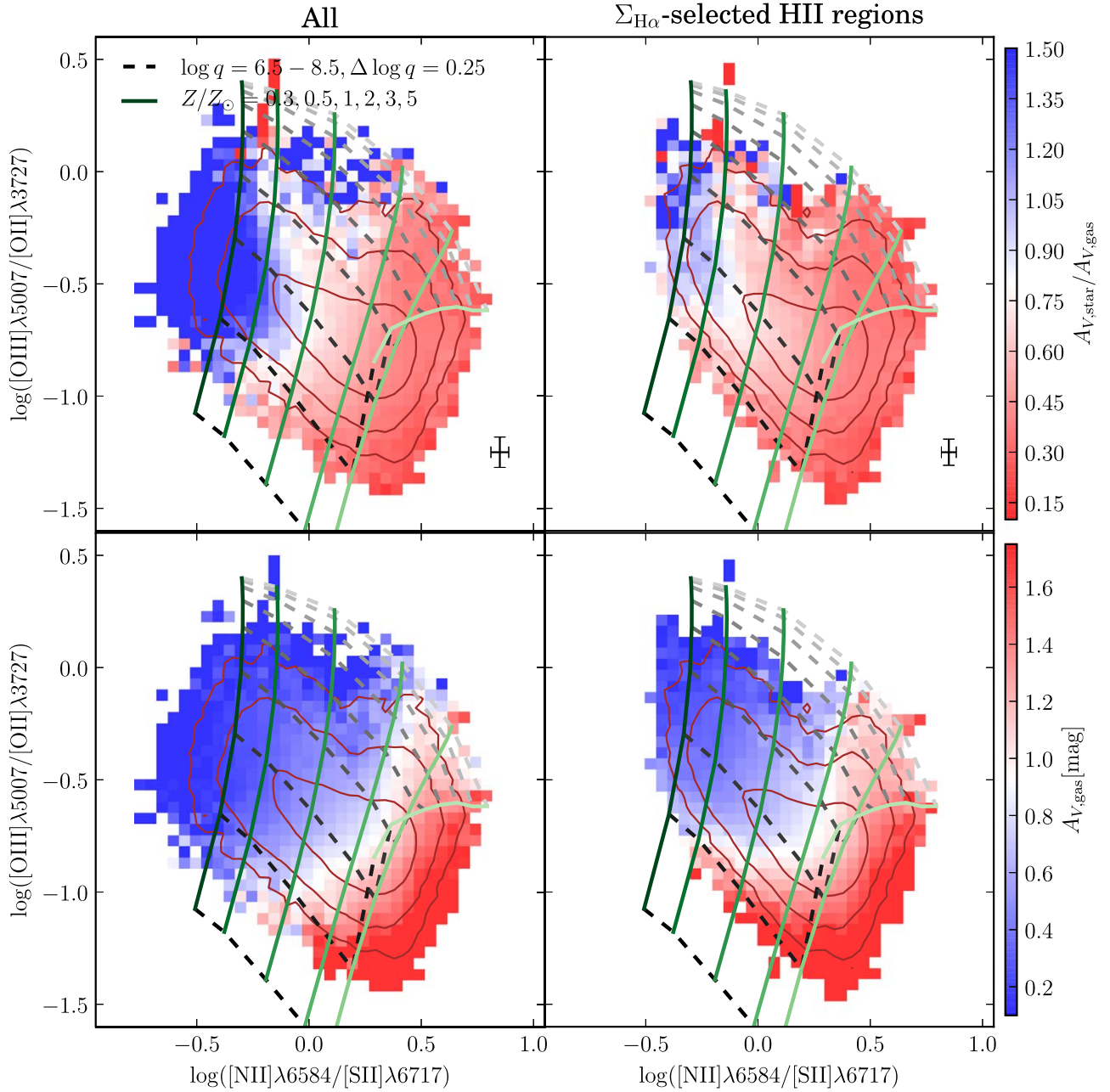


Figure 8. Median maps of $A_{V,\text{star}}/A_{V,\text{gas}}$ (top) and $A_{V,\text{gas}}$ (bottom) on the N2S2–O32 diagram for all (left) and the H II-dominated (right) spaxels. In each panel, the brown contours encompass 68%, 95%, and 99% of the spaxels, respectively, while the grid denotes the predicted line ratios of the theoretical photoionization model of a spherical H II region with $\kappa = \infty$ from Dopita et al. (2013). The typical uncertainties are shown in the lower right corner.

2048 pc), the dynamic ranges of both attenuations and the $A_{\text{FUV}}/A_{\text{H}\alpha}$ ratio become smaller, indicating a smooth effect at coarser resolutions where intense H II regions and quiescent regions merge.

For our H II-dominated spaxels, since $A_{V,\text{star}}$ is derived from the overall shape of the underlying stellar continuum and is not closely linked to the young stellar populations compared to A_{FUV} computed in Boquien et al. (2015), the smooth effect on our stellar attenuation should be weaker than the one reported by these authors. However, the surrounding more diffuse, less attenuated regions might have a nonnegligible contribution to the emission line measurements of the spaxels, possibly leading to a smaller $A_{V,\text{gas}}$, on average, and thus a slightly larger $A_{V,\text{star}}/A_{V,\text{gas}}$. Given that the systematic differences between

$A_{V,\text{star}}$ and $A_{V,\text{gas}}$ are still observed for the H II-dominated population, we conclude that the smooth effect stemming from the coarse spatial resolution is too weak to eliminate the geometry effect of the two-component dust model. Therefore, the large PSF size of the MaNGA data has little impact on the comparison with the dust model, as well as most of the conclusions drawn from this section.

3.7. The Smooth $A_{V,\text{star}}$ Maps

Since the derivation of $A_{V,\text{star}}$ is independent of emission lines, the relatively uniform distribution of $A_{V,\text{star}}$ on the (gas-phase) metallicity–ionization parameter space (Figure 5) and the BPT diagrams might be an effect of the methodology. However, it is noteworthy that the overall trends of $A_{V,\text{gas}}$

observed in Figure 5 still exist in the maps of $A_{V,\text{star}}$ but are much weaker. Particularly, the median of $A_{V,\text{star}}$ at the low- Z , low- q regime is slightly higher for the H II-dominated population, which is similar to the case of $A_{V,\text{gas}}$.

Another effect that might contribute to the uniformity of $A_{V,\text{star}}$ is the smooth effect arisen from the large PSF size of the MaNGA data. A multiscale study of M33 suggested that either stellar or nebular attenuation suffers a smooth effect at some level (Boquien et al. 2015). As discussed in Section 3.6, the ranges of both attenuations become narrower as the resolution becomes coarser. However, the median maps of $A_{V,\text{gas}}$ on either the N2O2 versus O32 diagram or the BPT diagrams (Figure 6) are not as smooth as those of $A_{V,\text{star}}$, implying that the coarse resolution should not be the main reason for the uniform $A_{V,\text{star}}$ distribution. Thus, we argue that the observed uniform distribution of $A_{V,\text{star}}$ indeed has a physical origin, indicating that the stellar attenuation is less relevant to the physical conditions of the ionized gas. In fact, as implied by the two-component dust model (Charlot & Fall 2000; Wild et al. 2011), $A_{V,\text{star}}$ might reflect the local environment of the diffuse ISM along the line of sight rather than the birth clouds. Studies of young star clusters (YSCs) confirmed that YSCs with ages of ~ 10 Myr are no longer embedded in their birth clouds (Whitmore et al. 2014; Hollyhead et al. 2015; Grasha et al. 2019). In other words, stellar populations older than this timescale might only suffer dust attenuation from the diffuse ISM.

Due to the lack of a second method to compute $A_{V,\text{star}}$ at present, we still cannot rule out the possibility that the derivation method has contributed to flattening the trends in the $A_{V,\text{star}}$ maps.

4. Relation between $A_{V,\text{star}}/A_{V,\text{gas}}$ and Global Properties

4.1. Determining the Global $A_{V,\text{star}}/A_{V,\text{gas}}$

As shown in Kreckel et al. (2013), the $A_{V,\text{star}}/A_{V,\text{gas}}$ ratio tends to converge to a constant at the high- $\Sigma_{\text{H}\alpha}$ flux end within individual galaxies. After visual inspection of the $A_{V,\text{star}}/A_{V,\text{gas}}$ versus $\Sigma_{\text{H}\alpha}$ plot of each galaxy, we find not only the similar trend presented in Kreckel et al. (2013) but also a galaxy-to-galaxy variation of the convergent value; i.e., the global $A_{V,\text{star}}/A_{V,\text{gas}}$ may vary between galaxies. To demonstrate this feature, we plot the density contours of all spaxels in the $A_{V,\text{star}}/A_{V,\text{gas}} - \Sigma_{\text{H}\alpha}$ plane in the upper panel of Figure 9. Since the convergent value may be a variable between galaxies, we calculate a representative value of $A_{V,\text{star}}/A_{V,\text{gas}}$ via σ -clipped statistics. The median of the 3σ clipping after three iterations is taken as the global value of $A_{V,\text{star}}/A_{V,\text{gas}}$ for each galaxy. Inspection of the histogram of $A_{V,\text{star}}/A_{V,\text{gas}}$ within individual galaxies confirms that this parameter can be a good estimate of the peak of the distribution. Then the global $A_{V,\text{star}}/A_{V,\text{gas}}$ is subtracted from the individual ones within the corresponding galaxy to obtain $\Delta(A_{V,\text{star}}/A_{V,\text{gas}})$. The density distribution of spaxels in the $\Delta(A_{V,\text{star}}/A_{V,\text{gas}}) - \Sigma_{\text{H}\alpha}$ plane is given in the bottom panel of Figure 9.

Both $A_{V,\text{star}}/A_{V,\text{gas}}$ and $\Delta(A_{V,\text{star}}/A_{V,\text{gas}})$ exhibit great reductions in scatters toward high $\Sigma_{\text{H}\alpha}$, indicating a common geometry of dust/star and local physical conditions for high- $\Sigma_{\text{H}\alpha}$ spaxels. The convergent $A_{V,\text{star}}/A_{V,\text{gas}}$ of 0.425, resulting from a peak detection utilizing the PeakUtils package⁹ based on the histogram of $A_{V,\text{star}}/A_{V,\text{gas}}$, is slightly smaller than the one of Calzetti (1997). Comparison between the two sets of

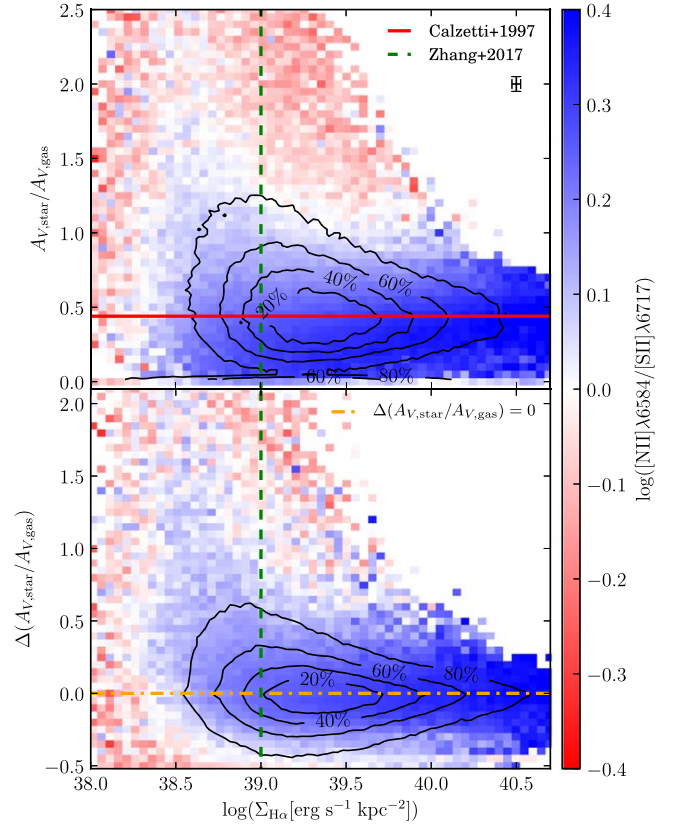


Figure 9. Density contours of all spaxels in the $A_{V,\text{star}}/A_{V,\text{gas}} - \Sigma_{\text{H}\alpha}$ (upper panel) and $\Delta(A_{V,\text{star}}/A_{V,\text{gas}}) - \Sigma_{\text{H}\alpha}$ (bottom panel) planes. The contours contain 20%, 40%, 60%, and 80% of the spaxels, respectively. The underlying maps are color-coded by the median value of $\log([\text{H II}]\lambda 6584/[\text{S II}]\lambda 6717)$. The red solid line is the canonical value of $A_{V,\text{star}}/A_{V,\text{gas}} = 0.44$ from Calzetti (1997), while the orange dashed-dotted line indicates $\Delta(A_{V,\text{star}}/A_{V,\text{gas}}) = 0$. The vertical green dashed lines show the criterion of H II-dominated spaxels from Zhang et al. (2017), i.e., $\Sigma_{\text{H}\alpha} = 10^{39} \text{ erg s}^{-1} \text{ kpc}^{-2}$. The typical statistical uncertainties are plotted.

contours in Figure 9 reveals that the scatter of $\Delta(A_{V,\text{star}}/A_{V,\text{gas}})$ is further reduced after normalization by the global $A_{V,\text{star}}/A_{V,\text{gas}}$ of each galaxy. This result implies that the global values of $A_{V,\text{star}}/A_{V,\text{gas}}$ indeed vary between galaxies, which was already suggested by many previous works on either local (e.g., Wild et al. 2011; Koyama et al. 2015; Zahid et al. 2017; Koyama et al. 2019) or high-redshift (e.g., Price et al. 2014; Puglisi et al. 2016) galaxies.

4.2. Observational Evidence of Correlations between $A_{V,\text{star}}/A_{V,\text{gas}}$ and Other Properties

Naturally, one would wonder whether the global $A_{V,\text{star}}/A_{V,\text{gas}}$ correlates with other physical properties of galaxies if it is thought of as a variable. Previous works demonstrated that this ratio may correlate with M_* (e.g., Koyama et al. 2015, 2019; Puglisi et al. 2016; Zahid et al. 2017), SFR (e.g., Reddy et al. 2015), sSFR (e.g., Wild et al. 2011; Price et al. 2014; Koyama et al. 2019; Qin et al. 2019), inclination (Wild et al. 2011), and so on. Making use of the Pipe3D VAC and the NASA-Sloan Atlas (NSA) catalog used by the MaNGA survey,¹⁰ we calculate the Pearson correlation

⁹ <https://peakutils.readthedocs.io/en/latest/index.html>

¹⁰ <https://www.sdss.org/dr15/manga/manga-target-selection/nsa/>

Table 1
Correlation Coefficients between Global Dust Attenuation (Ratio) and Some Physical Properties

Parameters ^a	$\log(M_*)^b$ (M_\odot)	$\log(\text{SFR})$ ($M_\odot \text{ yr}^{-1}$)	$\log(\text{sSFR})$ (yr^{-1})	$12+\log(\text{O}/\text{H})^c$ (O3N2)	$12+\log(\text{O}/\text{H})^c$ (N2)	b/a^d
$A_{V,\text{star}}/A_{V,\text{gas}}$	-0.66	-0.59	0.18	-0.59	-0.62	-0.27
	-0.72	-0.62	0.24	-0.60	-0.63	-0.32
$A_{V,\text{star}}$	-0.12	-0.08	0.06	0.11	0.14	-0.54
	-0.13	-0.09	0.07	0.05	0.05	-0.55
$A_{V,\text{gas}}$	0.49	0.46	-0.10	0.59	0.62	-0.23
	0.54	0.50	-0.15	0.62	0.65	-0.21

Notes.

^a For each parameter, the first row lists the Pearson correlation coefficients r , while the second row presents the Spearman rank correlation coefficients ρ .

^b Stellar mass from K -correction fit for Sérsic fluxes, taken from the NSA catalog.

^c Oxygen abundance $12+\log(\text{O}/\text{H})$ at the effective radius derived based on the O3N2 and N2 calibrations of Marino et al. (2013).

^d Axis ratio b/a from the two-dimensional, single-component Sérsic fit in the r band, taken from the NSA catalog.

coefficients r and the Spearman rank correlation coefficients ρ between the global $A_{V,\text{star}}/A_{V,\text{gas}}$ ratio and several physical properties and list them in Table 1. Taking the medians of $A_{V,\text{star}}$ and $A_{V,\text{gas}}$ of spaxels within individual galaxies as the corresponding global values, the correlation coefficients between these global attenuations and other physical properties are also given in Table 1. Note that these “global” dust attenuations (also the “global” $A_{V,\text{star}}/A_{V,\text{gas}}$ defined in Section 4.1) are the typical values of local H II regions within galaxies, which might be different from those derived from the integrated light. The global SFR is obtained by summing up the H α luminosities over the field of view (FOV) and applying the Kennicutt (1998b) conversion. The metallicities are taken from the Pipe3D VAC, while M_* and axis ratio b/a are from the NSA catalog after correcting for the differences in the adopted cosmology and IMF. Although the Pipe3D VAC provides a global stellar mass of galaxies, it is found to be slightly smaller at the high-mass end due to the limited FOV of MaNGA. The sSFR is calculated by SFR/M_* using the adopted parameters. Here $12+\log(\text{O}/\text{H})$ is measured at the effective radius R_e to represent the global oxygen abundance. Studies based on the CALIFA survey suggested that this parameter has a pretty good one-to-one relation with the average oxygen abundance over the entire FOV but with a smaller uncertainty, regardless of metallicity calibrations (Sánchez et al. 2013, 2017). To ensure a robust estimate of the global $A_{V,\text{star}}/A_{V,\text{gas}}$, we only include galaxies with more than 50 selected spaxels in the following analysis. In total, 1949 galaxies remain after this selection. We stress that changing this critical number of spaxels cannot significantly alter the following results.

From Table 1, it is clear that M_* presents the strongest correlation with the global $A_{V,\text{star}}/A_{V,\text{gas}}$ in either r or ρ , implying smaller $A_{V,\text{star}}/A_{V,\text{gas}}$ values for more massive galaxies. We show the global $A_{V,\text{star}}/A_{V,\text{gas}}$ as a function of M_* for the selected galaxies in Figure 10, together with the best-fit result for all galaxies using a linear function in the form of

$$A_{V,\text{star}}/A_{V,\text{gas}} = a_0 + a_1 x, \quad (3)$$

in which $x = \log(M_*/M_\odot)$. The typical uncertainty of $A_{V,\text{star}}/A_{V,\text{gas}}$ shown in the figure is the median of the standard deviation of $A_{V,\text{star}}/A_{V,\text{gas}}$ within each galaxy. Since the NSA catalog does not contain a formal uncertainty for M_* , we give a rough estimation based on the comparison of M_* from

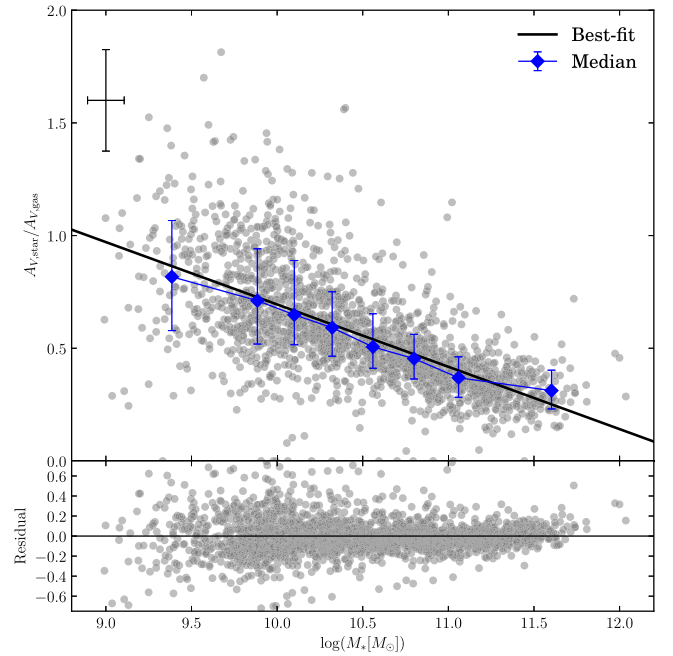


Figure 10. Global $A_{V,\text{star}}/A_{V,\text{gas}}$ ratio as a function of global stellar mass (M_*) of galaxies. The black solid line shows the best-fit result with a linear function for all galaxies. The upper panel shows the $A_{V,\text{star}}/A_{V,\text{gas}}$ vs. M_* distribution for individual galaxies, while the bottom panel presents the residuals between the observed and the predicted $A_{V,\text{star}}/A_{V,\text{gas}}$ from the best-fit line. The blue diamonds with error bars indicate the medians and corresponding 1σ scatters (i.e., 16%–84% range) of the binned distributions for all galaxies. These bins are adjusted to equalize the number of galaxies in each bin for each set of bins. The typical uncertainties are denoted in the upper left corner.

Kauffmann et al. (2003) and the kcorrect code from which the NSA M_* was calculated (Blanton & Roweis 2007).

The best-fit result (black solid line in Figure 10) is given by $a_0 = 3.460 \pm 0.075$, $a_1 = -0.277 \pm 0.007$, and is in good agreement with the binned median curve. The scatter of residuals is M_* -dependent, and more massive galaxies have smaller scatters. Comparing the observations with the predictions from the best-fit curve, the standard deviation of the residuals is 0.19.

Table 1 also suggests that the correlation between the attenuation ratio and oxygen abundance, regardless of indicators, is only slightly weaker than that of M_* . Taking the N2 index as the fiducial indicator, we show how the global $A_{V,\text{star}}/A_{V,\text{gas}}$ varies

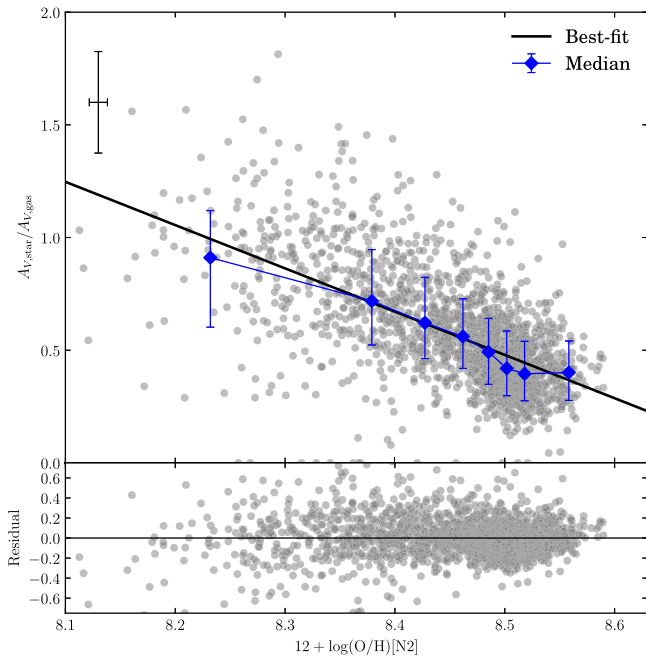


Figure 11. Global $A_{V,\text{star}}/A_{V,\text{gas}}$ ratio as a function of oxygen abundances derived from N2 indicators based on the Marino et al. (2013) calibrations. The symbols are the same as in Figure 10. The typical uncertainties are shown in the upper left corner.

with gas-phase metallicity in Figure 11. The oxygen abundance derived from the O3N2 diagnostic presents a very similar distribution as Figure 11. We fit the data points with a linear function of Equation (3) in which $x = 12 + \log(\text{O}/\text{H})$. The best-fit result is $a_0 = 16.804 \pm 0.486$, $a_1 = -1.921 \pm 0.058$, with substantial scatter for N2-based metallicities (black solid line in Figure 11), and is also consistent with the binned medians denoted by blue diamonds. The standard deviation of the residuals is 0.20.

The SFR shows a moderate correlation with the global $A_{V,\text{star}}/A_{V,\text{gas}}$, while sSFR and b/a (i.e., inclination) exhibit very weak correlations with the global $A_{V,\text{star}}/A_{V,\text{gas}}$. These correlations are shown in Figure 12, together with a median curve and the corresponding scatters around the medians of the binned distributions for each case. The $A_{V,\text{star}}/A_{V,\text{gas}}$ ratio decreases with increasing SFR, while the scatters become smaller toward the high-SFR end. Although the median curves of $A_{V,\text{star}}/A_{V,\text{gas}}$ in sSFR and inclination are fairly flat, slight elevations can be identified when $\log(\text{sSFR}[\text{yr}^{-1}]) > -10.5$ and toward smaller b/a (i.e., larger inclination). The global $A_{V,\text{star}}$ has a moderate anticorrelation with b/a , while the global $A_{V,\text{gas}}$ shows a very weak anticorrelation with b/a . A similar result was reported by Yip et al. (2010), in which the stellar attenuation is found to increase with inclination, while the $\text{H}\alpha/\text{H}\beta$ ratio remains nearly unchanged with b/a . Such behaviors are consistent with the picture of the two-component dust model (e.g., Charlot & Fall 2000; Wild et al. 2011). In this model, $A_{V,\text{star}}$ arises from the diffuse ISM that distributes more homogeneously in the stellar disk (e.g., Mosenkov et al. 2019), so that galaxies with larger inclinations would have larger $A_{V,\text{star}}$ due to the increasing dust column density. The $A_{V,\text{gas}}$ describes the combination of the dust effect from the dense birth clouds, which are more likely spherical, and the diffuse ISM but is dominated by the former (e.g., Wild et al. 2011). Therefore, the inclination only has a small effect on $A_{V,\text{gas}}$, and

the global $A_{V,\text{star}}/A_{V,\text{gas}}$ would show a weak correlation with the inclination.

Koyama et al. (2019) presented a detailed study of the relation between the global $A_{V,\text{star}}/A_{V,\text{gas}}$ ¹¹ and M_* , SFR, and sSFR using local SFG samples. They found that $A_{V,\text{star}}/A_{V,\text{gas}}$ decreases with increasing M_* , decreases at $-1 \lesssim \log(\text{SFR}[M_\odot \text{ yr}^{-1}]) \lesssim 0.5$, and then increases at $0.5 \lesssim \log(\text{SFR}[M_\odot \text{ yr}^{-1}]) \lesssim 1.5$ and toward the high-sSFR end for their SDSS/*Galaxy Evolution Explorer* (GALEX)/*Wide-field Infrared Survey Explorer* (WISE) sample. Within the parameter ranges explored by our sample, the overall trends of these three correlations are in good agreement with ours. By stacking the SDSS spectra of local SFGs, Zahid et al. (2017) reported that the trend between the global $A_{V,\text{star}}/A_{V,\text{gas}}$ and M_* increases first at $M_* \lesssim 10^{9.5} M_\odot$ and then decreases toward the high- M_* end with an increasing slope, whereas the same relation from both this work and Koyama et al. (2019) can be well described by a monotonic decreasing function with a gradual flattening at the high- M_* end. However, due to the different data treatment between our work (or Koyama et al. 2019) and Zahid et al. (2017; i.e., individual calculation versus stacking), it is difficult to find out the reasons for these conflicts at present. Such different observations need to be understood in future studies.

As for the global dust attenuations, Table 1 shows that $A_{V,\text{star}}$ only correlates with b/a , while $A_{V,\text{gas}}$ moderately correlates with M_* , SFR, and gas-phase metallicity; weakly correlates with b/a ; and does not correlate with sSFR. Similar to the local case, the correlations between $A_{V,\text{star}}/A_{V,\text{gas}}$ and the other physical properties discussed above are mainly driven by the correlations between $A_{V,\text{gas}}$ and these properties. However, we also note that, compared to $A_{V,\text{gas}}$, $A_{V,\text{star}}/A_{V,\text{gas}}$ has slightly stronger correlations (in terms of either r or ρ) with M_* , SFR, and b/a and a comparable correlation with gas-phase metallicity, indicating the nonnegligible role of $A_{V,\text{star}}$ in these correlations.

4.3. Correlations along Scaling Relations

All of the physical properties discussed above can be linked to M_* except the inclination (i.e., b/a) via two well-studied scaling relations of SFGs. On the one hand, the star-forming main sequence (SFMS; e.g., Speagle et al. 2014) connects the SFR and sSFR to M_* . On the other hand, the metal enrichment in the ISM is closely related to the growth of galaxies and can be characterized by the well-known mass–metallicity relation (MZR; e.g., Lequeux et al. 1979; Tremonti et al. 2004).

We present the SFMS and MZR of our sample in Figure 13, color-coding by their global $A_{V,\text{star}}/A_{V,\text{gas}}$ ratios and dust attenuations. Both plots exhibit a gradual trend of decreasing $A_{V,\text{star}}/A_{V,\text{gas}}$ along the sequence toward the high- M_* end. The trend of $A_{V,\text{star}}/A_{V,\text{gas}}$ in our SFMS resembles the one using the $\text{H}\alpha$ -based SFR ($\text{SFR}_{\text{H}\alpha}$) in Koyama et al. (2019); however, their fiducial total SFR is derived from a combination of UV and IR observations ($\text{SFR}_{\text{UV}+\text{FIR}}$) for which the global $A_{V,\text{star}}/A_{V,\text{gas}}$ presents an evident trend across the SFMS. Their SDSS/*GALEX*/*WISE* sample suggests that galaxies with higher SFR tend to have a larger $A_{V,\text{star}}/A_{V,\text{gas}}$ value at fixed M_* . By comparing the $\text{SFR}_{\text{H}\alpha}$

¹¹ Koyama et al. (2019) used $E(B-V)_{\text{star}}/E(B-V)_{\text{gas}}$ to indicate the ratio between stellar and nebular attenuation. In the following discussion, we use $A_{V,\text{star}}/A_{V,\text{gas}}$ instead of $E(B-V)_{\text{star}}/E(B-V)_{\text{gas}}$ when quoting their results due to the fact that $A_{V,\text{star}}/A_{V,\text{gas}} \propto E(B-V)_{\text{star}}/E(B-V)_{\text{gas}}$ and a constant R_V are assumed for all galaxies when either stellar or nebular attenuation is derived in Koyama et al. (2019).

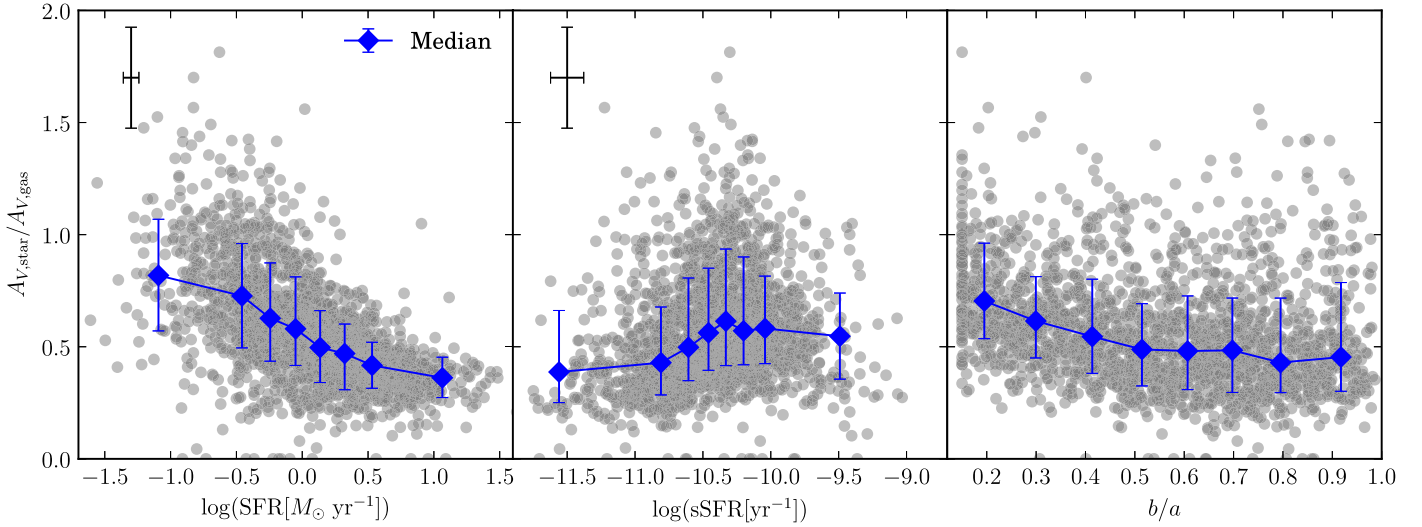


Figure 12. Global $A_{V,\text{star}}/A_{V,\text{gas}}$ ratio as a function of SFR (left), sSFR (middle), and b/a (right). The blue diamonds with error bars indicate the medians and corresponding 1σ scatters (i.e., 16%–84% range) of the binned distributions for all galaxies. These bins are adjusted to equalize the number of galaxies in each bin for each set of bins. The typical uncertainties are given in the upper left corner, except for the b/a subplot.

and $\text{SFR}_{\text{UV+IR}}$, they further found that galaxies with $\text{SFR}_{\text{H}\alpha} < \text{SFR}_{\text{UV+IR}}$ tend to have substantially larger $A_{V,\text{star}}/A_{V,\text{gas}}$. However, such a trend is inevitable because both the stellar and nebular attenuation are directly linked to the corresponding SFR in their calculation. The $E(B-V)_{\text{star}}$ in Koyama et al. (2019) is derived from the $\text{SFR}_{\text{UV+IR}}$ in the form of

$$\begin{aligned} E(B-V)_{\text{star}} &\propto \log(\text{SFR}_{\text{UV+IR}}/\text{SFR}_{\text{UV}}) \\ &= \log(1 + \text{SFR}_{\text{IR}}/\text{SFR}_{\text{UV}}), \end{aligned} \quad (4)$$

in which SFR_{IR} and SFR_{UV} are SFRs obtained from IR and UV data, respectively, while $E(B-V)_{\text{gas}}$ can be connected to $\text{SFR}_{\text{H}\alpha}$ via

$$\begin{aligned} A_{\text{H}\alpha} &\propto E(B-V)_{\text{gas}} \\ &\propto \log(1 + \text{SFR}_{\text{H}\alpha,\text{atten}}/\text{SFR}_{\text{H}\alpha,\text{obs}}), \end{aligned} \quad (5)$$

in which $\text{SFR}_{\text{H}\alpha,\text{atten}}$ and $\text{SFR}_{\text{H}\alpha,\text{obs}}$ are SFRs calculated from attenuated and observed $\text{H}\alpha$ fluxes, respectively. Then we can obtain

$$\frac{A_{V,\text{star}}}{A_{V,\text{gas}}} \propto \frac{\log(1 + \text{SFR}_{\text{IR}}/\text{SFR}_{\text{UV}})}{\log(1 + \text{SFR}_{\text{H}\alpha,\text{atten}}/\text{SFR}_{\text{H}\alpha,\text{obs}})}. \quad (6)$$

Although Koyama et al. (2019) claimed that the physical reason for the difference between $\text{SFR}_{\text{H}\alpha}$ and $\text{SFR}_{\text{UV+IR}}$ is unclear, we can expect that either the overestimation of SFR_{IR} , e.g., due to the large PSF of the *WISE* W4 band (Wright et al. 2010) and contamination from their neighbors, or the underestimation of $\text{SFR}_{\text{H}\alpha,\text{atten}}$, e.g., due to saturation in emission line attenuation (Calabrò et al. 2018), would lead to an elevation on $A_{V,\text{star}}/A_{V,\text{gas}}$ if the unattenuated parts are thought to be more reliable. Therefore, we suggest that an SFR-independent $E(B-V)_{\text{star}}$ (in calculation) and an IR-clear sample are required to solve this puzzle.

As suggested in Section 4.2, $A_{V,\text{star}}$ exhibits a relatively uniform distribution in either SFMS or MZR, while $A_{V,\text{gas}}$ dominates the observed trends along the two scaling relations for $A_{V,\text{star}}/A_{V,\text{gas}}$. These results imply that the properties of dust in the diffuse ISM do not significantly alter as galaxies grow,

while the dust in H II regions and/or the surrounding dense clouds is more sensitive to the galaxy evolution.

Since $A_{V,\text{gas}}$ can be connected to dust mass surface density and also gas mass surface density (e.g., Boquien et al. 2013; Brinchmann et al. 2013; Kreckel et al. 2013, 2016), one would speculate that the observed $A_{V,\text{star}}/A_{V,\text{gas}}$ –SFR relation might stem from the more basic Kennicutt–Schmidt star formation law (Schmidt 1959; Kennicutt 1998a) and leads to the $A_{V,\text{star}}/A_{V,\text{gas}}$ – M_* and $A_{V,\text{star}}/A_{V,\text{gas}}$ – $12+\log(\text{O}/\text{H})$ relations via the SFMS and MZR, respectively. We believe that the Kennicutt–Schmidt law can partly account for the $A_{V,\text{star}}/A_{V,\text{gas}}$ –SFR relation; however, it cannot fully explain the observed trends along the two scaling relations due to the following reasons. First, Table 1 shows that both M_* and gas-phase metallicity more strongly correlate with $A_{V,\text{gas}}$ compared to SFR. If the $A_{V,\text{gas}}$ –SFR relation is more fundamental and connects $A_{V,\text{gas}}$ with M_* and gas-phase metallicity via the SFMS and MZR, respectively, we would expect a stronger correlation with SFR compared to M_* and gas-phase metallicity. Second, comparing with $A_{V,\text{gas}}$, $A_{V,\text{star}}/A_{V,\text{gas}}$ displays stronger correlations with M_* and SFR and a comparable correlation with gas-phase metallicity, implying that although $A_{V,\text{star}}$ exhibits very weak correlations with these physical properties, its role is still nonnegligible. Third, our partial correlation analysis discussed below suggests that when SFR is controlled, M_* still moderately correlates with $A_{V,\text{star}}/A_{V,\text{gas}}$, and these three parameters seem to play comparable roles in the correlation.

Koyama et al. (2019) speculated that the decrease of $A_{V,\text{star}}/A_{V,\text{gas}}$ with increasing SFR at the low-SFR side originates from the relation between M_* and $A_{V,\text{star}}/A_{V,\text{gas}}$ because of the SFMS. Similarly, the correlation between $A_{V,\text{star}}/A_{V,\text{gas}}$ and gas-phase metallicity revealed by Figure 11 also might result from the M_* – $A_{V,\text{star}}/A_{V,\text{gas}}$ relation due to the tight MZR. To test these conjectures, a partial correlation analysis is performed utilizing the R ppcor package¹² (Kim 2015). We calculate the Pearson correlation coefficients r and Spearman rank correlation coefficients ρ to see how $A_{V,\text{star}}/A_{V,\text{gas}}$ correlates with M_* ,

¹² <https://cran.r-project.org/web/packages/ppcor/>

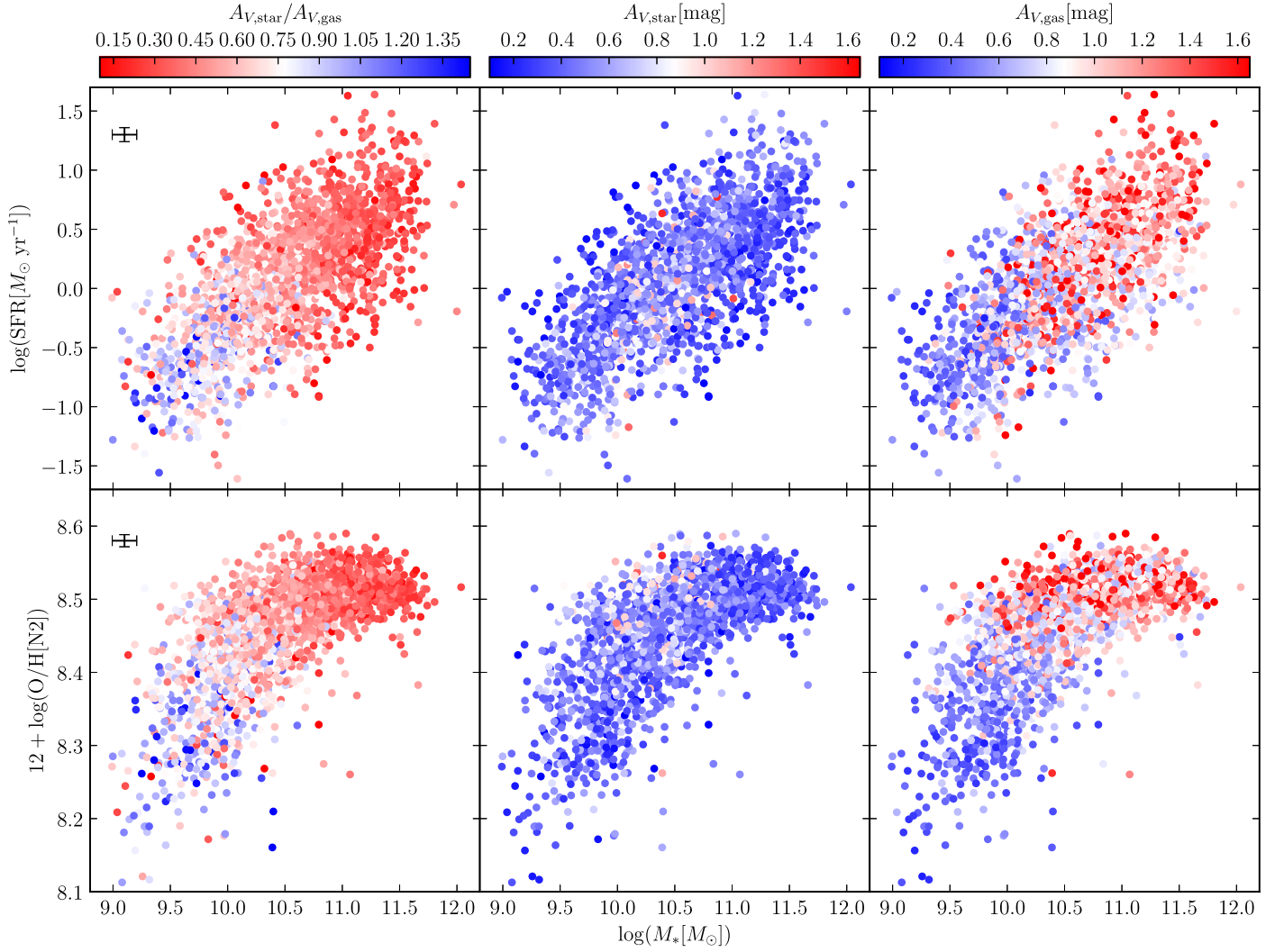


Figure 13. The SFMS (top row) and MZR (bottom row) of our selected galaxies, color-coded by their global $A_{V,\text{star}}/A_{V,\text{gas}}$ ratio (left), $A_{V,\text{star}}$ (middle), and $A_{V,\text{gas}}$ (right). The typical uncertainties are shown in the upper left corner of the left panels.

Table 2
Partial Correlations between $A_{V,\text{star}}/A_{V,\text{gas}}$ and Some Physical Properties

Variables ^a	M_*	M_*	SFR	$12+\log(\text{O}/\text{H})$	M_*	SFR	$12+\log(\text{O}/\text{H})$
Controlled ^b	SFR	$12+\log(\text{O}/\text{H})$	M_*	M_*	SFR and $12+\log(\text{O}/\text{H})$	M_* and $12+\log(\text{O}/\text{H})$	M_* and SFR
r	-0.42	-0.43	-0.25	-0.30	-0.16	-0.31	-0.35
ρ	-0.50	-0.52	-0.24	-0.23	-0.26	-0.29	-0.29

Notes.

^a For each column, partial correlation coefficients are calculated between the physical property in this row and $A_{V,\text{star}}/A_{V,\text{gas}}$, while the physical properties shown in the second row are controlled. Here M_* and SFR are in log scale, and $12+\log(\text{O}/\text{H})$ is the N2-based one. Pearson correlation coefficients r and Spearman rank correlation coefficients ρ are derived.

^b Physics properties that are controlled in partial correlation analysis.

SFR, and $12+\log(\text{O}/\text{H})[\text{N2}]$, respectively, when the other one or two physical properties are controlled.

The results are presented in Table 2, in which the first four columns show the cases of controlling only one parameter and the last three columns are the cases of controlling two parameters simultaneously. When only one parameter is controlled, M_* gives the strongest correlation with $A_{V,\text{star}}/A_{V,\text{gas}}$, regardless of the controlled parameters and calculation methods. Comparing with the coefficients listed in Table 1, it is evident that all three

parameters can partly explain the corresponding correlations, while M_* seems to have a dominant effect. In the cases of controlling two physical properties at the same time, the effect of M_* becomes much smaller compared to SFR and metallicity when considering linear correlation. However, if nonlinear correlation (i.e., ρ) is also taken into account, the three parameters have comparable roles in the correlations with $A_{V,\text{star}}/A_{V,\text{gas}}$. Given the bridge role of M_* in connecting SFR and metallicity, we argue that the apparently dominant effect of

M_* when controlling only one parameter has a significant contribution from metallicity (SFR) via the MZR (SFMS). Therefore, we conclude that M_* , SFR, and gas-phase metallicity play comparable roles in the correlations with these physical properties, and either of them can fully account for the trends unveiled by Figure 13.

4.4. A Possible Picture Based on Dust Growth

Both Table 1 and Figure 13 suggest that the correlations between $A_{V,\text{star}}/A_{V,\text{gas}}$ and other galactic properties (i.e., M_* , SFR, and gas-phase metallicity), as well as the trends along the scaling relations, are mainly driven by the global $A_{V,\text{gas}}$. Integrating along the line of sight of Equation (2) results in (e.g., Galliano 2018)

$$A_\lambda = 1.086 \times \Sigma_{\text{dust}} = 1.086 \times \Sigma_{\text{gas}} \cdot \text{D/G}, \quad (7)$$

where Σ_{dust} and Σ_{gas} are the mass surface densities of dust and gas, respectively. Obviously, an increasing $A_{V,\text{gas}}$ with M_* implies more dust in more massive galaxies, which is equivalent to more gas and/or a larger D/G ratio. The increasing dust mass and total gas mass with M_* in SFGs were reported by many previous works (e.g., De Vis et al. 2017a; Casasola et al. 2019). Considering the factor related to dust, here we only focus on the possible role of the D/G ratio in the observed correlations, attempting to provide an alternative origin of these correlations in the view of dust evolution.

In the analysis of the local $A_{V,\text{star}}/A_{V,\text{gas}}$, we suggest that in addition to the geometry of stars/dust, the local physical conditions of ionized gas also play an important role in determining the $A_{V,\text{star}}/A_{V,\text{gas}}$ value. The D/G ratio is used to connect the local gas-phase metallicity and ionization parameter to dust attenuation. On a galactic scale, the behaviors of the global D/G ratio are found to be controlled by some dust processes accompanied by galaxy growth (Asano et al. 2013; De Vis et al. 2017b, 2019; Aoyama et al. 2018; Galliano et al. 2018).

As we mentioned in Section 3.3.3, metallicity is the main physical property that drives the D/G ratio (Engelbracht et al. 2008; Asano et al. 2013; Rémy-Ruyer et al. 2014). This metallicity dependence can be described by the following picture (Galliano 2018): when galaxies are metal-poor, dust grains are mainly produced by stellar ejecta, and the D/G is proportional to metallicity with a small dust-to-metal ratio; at intermediate metallicities, metal accretion onto dust grains becomes efficient, and the D/G increases steeply; and when galaxies become more metal-rich, dust growth is limited by metallicity again, but with a larger dust-to-metal ratio, due to the depletion of metals in the ISM.¹³

Such a picture is well reproduced by hydrodynamic simulation. As displayed in Aoyama et al. (2018), for intermediate-mass galaxies ($10^{8.5} M_\odot \lesssim M_* \lesssim 10^{10} M_\odot$ in their simulation), dust growth by accretion becomes active and the D/G–Z relation becomes steeper. A large scatter of the relation in this region emerges, suggesting that the D/G is strongly

impacted by local internal conditions and star formation histories (SFHs) in this mass range (Asano et al. 2013; Rémy-Ruyer et al. 2014). For the most massive galaxies ($M_* \gtrsim 10^{10} M_\odot$), dust growth by accretion is limited by metallicity and saturated. Here the MZR is already reproduced in their simulation and naturally connects the D/G to M_* . In addition, Table 1 shows that the global $A_{V,\text{gas}}$ has slightly weaker (but opposite sign) correlations with M_* and SFR compared to those of $A_{V,\text{star}}/A_{V,\text{gas}}$ and comparable correlations with oxygen abundances, irrespective of metallicity diagnostics. However, the global $A_{V,\text{star}}$ shows almost no correlations with all of the properties listed in Table 1 except the b/a ratio. Combining the predictions from simulations with the observed evidence, we speculate that, similar to the local case, the elevated global D/G ratio in more massive SFGs leads to the increase of the overall $A_{V,\text{gas}}$ and thus a smaller $A_{V,\text{star}}/A_{V,\text{gas}}$.

Furthermore, the M_* -dependent scatter shown in Figure 10 can be explained as the scatter of D/G, which is found to be large for $10^{8.5} M_\odot \lesssim M_* \lesssim 10^{10} M_\odot$ and much smaller for $M_* \gtrsim 10^{10} M_\odot$ in the simulation of Aoyama et al. (2018). Recent IFU studies also suggest that the global stellar mass has a nonnegligible role in regulating the local metallicity (Gao et al. 2018), especially for low-mass galaxies, probably due to their shallower gravitational potential wells for which it is easier to lose their metal by galactic outflows (Chisholm et al. 2018). Because the global $A_{V,\text{star}}/A_{V,\text{gas}}$ used in this work is derived from the mode of the local ones, the M_* effect on local metallicity should also impact the global $A_{V,\text{star}}/A_{V,\text{gas}}$. Hence, besides the local internal processes and SFHs, M_* itself may be another factor that drives the M_* -dependent scatter in the M_* – $A_{V,\text{star}}/A_{V,\text{gas}}$ relation.

Therefore, the evolution of dust growth in birth clouds accompanied by galaxy growth might be one plausible origin for our observed behaviors of the global $A_{V,\text{star}}/A_{V,\text{gas}}$.

5. Summary

In this work, we construct a sample of H II regions using the IFU data of the MaNGA survey released from the SDSS DR15 and investigate how the $A_{V,\text{star}}/A_{V,\text{gas}}$ ratio varies with other physical properties on subgalactic and galactic scales.

On a subgalactic scale, we find that the $A_{V,\text{gas}}$ indeed correlates with the $A_{V,\text{star}}$, while the correlation becomes much stronger for spaxels with higher $\Sigma_{\text{H}\alpha}$. The local $A_{V,\text{star}}/A_{V,\text{gas}}$ is found to have moderate correlations with DIG indicators (i.e., the N2S2 and S2 indices, $\Sigma_{\text{H}\alpha}$) and tracers of gas-phase metallicity and ionization parameter. Among these parameters, the N2S2 index shows the strongest correlation with the local $A_{V,\text{star}}/A_{V,\text{gas}}$ ratio. The $A_{V,\text{star}}/A_{V,\text{gas}}$ ratio of DIG-dominated spaxels tends to be systematically larger than that of classic H II-dominated spaxels. Local physical conditions of ionized gas show significant effects on the dust attenuation. Metal-poor spaxels with high ionization parameters tend to suffer less nebular attenuation (i.e., smaller $A_{V,\text{gas}}$) and thus larger $A_{V,\text{star}}/A_{V,\text{gas}}$ values for either H II- or DIG-dominated sub-samples. We argue that the systematic difference in $A_{V,\text{star}}/A_{V,\text{gas}}$ between DIG- and H II-dominated spaxels can be explained by the stars/dust geometry based on the two-component dust model (Charlot & Fall 2000), while the dependences on metallicity and ionization parameters can be attributed to the change of dust properties in different local environments. We further suggest that the metallicity- and ionization parameter-dependent dust-to-gas mass ratio could

¹³ We note that the range of our N2-based $12+\log(\text{O}/\text{H})$ is already in the most metal-rich (ISM growth) regime of the Galliano (2018) picture. However, the absolute value of gas-phase metallicity derived from different calibrations varies up to 0.7 dex (Kewley & Ellison 2008). On the other hand, both our M_* and $12+\log(\text{O}/\text{H})$ ranges cover the predicted regimes in which dust growth by accretion goes from being efficient to being saturated in the hydrodynamic simulation of Aoyama et al. (2018). Therefore, here we focus on the picture evolved with the relative value of metallicity, rather than the absolute value.

be one possible parameter to connect the physical conditions with dust attenuation. The median maps of dust attenuation (and the ratio) on the BPT diagrams form a low- $A_{V,\text{gas}}$ and high- $A_{V,\text{star}}/A_{V,\text{gas}}$ sequence that can be resolved into DIG-dominated and metal-poor spaxels. The local analyses reveal that both geometry between stars and gas/dust and physical conditions are important in determining the local $A_{V,\text{star}}/A_{V,\text{gas}}$ values.

The local $A_{V,\text{star}}/A_{V,\text{gas}}$ within each galaxy converges to one constant at the high- $\Sigma_{\text{H}\alpha}$ end, which is also the mode of the distribution of $A_{V,\text{star}}/A_{V,\text{gas}}$ and is taken as the global $A_{V,\text{star}}/A_{V,\text{gas}}$.

On a galactic scale, this global $A_{V,\text{star}}/A_{V,\text{gas}}$ ratio indeed varies from galaxy to galaxy and correlates with physical properties such as stellar mass, SFR, and metallicity. It is found to show a strong correlation with stellar mass, moderate correlations with SFR and metallicity, and weak correlations with b/a and sSFR. The SFGs that have larger M_* and higher SFR and are more metal-rich tend to have smaller $A_{V,\text{star}}/A_{V,\text{gas}}$ ratios. A gradual trend of decreasing $A_{V,\text{star}}/A_{V,\text{gas}}$ toward the high- M_* end along the SFMS and MZR is found. Partial correlation analysis demonstrates that M_* , SFR, and gas-phase metallicity have comparable effects in the observed correlations, and no single property can fully account for these correlations. We suggest that the metallicity-dependent dust-to-gas ratio, together with the dust growth process accompanied by galaxy growth, might partly explain these observations, although we cannot rule out any contribution from the geometry effect.

Overall, our results highlight the importance of local physical conditions in dust processes, as well as the level of dust attenuation. High spatial resolution IR observations are required to give a quantitative conclusion. In addition, the most uncertain parameter in this study is the local $A_{V,\text{star}}$, for which more constraints from IR observation are required. Fortunately, the Mid-Infrared Instrument of the *James Webb Space Telescope* is able to provide mid-IR (MIR) images of nearby galaxies from 5 to 28.5 μm with unprecedented spatial resolution (Rieke et al. 2015). Such high-quality MIR observations, in combination with the optical IFU data and the well-calibrated monochromatic estimators of total IR luminosity (e.g., Boquien et al. 2010; Elbaz et al. 2010; Lin et al. 2016), will allow us to put more constraints on the derivation of $A_{V,\text{star}}$ and thus help us to better understand the relation between $A_{V,\text{star}}/A_{V,\text{gas}}$ and other parameters on a subkiloparsec scale.

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Software: Astropy (Astropy Collaboration et al. 2013, 2018), IPython (Pérez & Granger 2007), Matplotlib (Hunter 2007), Numpy (Oliphant 2006), MPFIT (Markwardt 2009).

Appendix Effects of S/N Criteria on Emission Lines

As mentioned in Section 2.3, we require our sample to have an S/N of the Balmer line ($\text{H}\alpha$ and $\text{H}\beta$; S/N_{Bal}) greater than 5 and an S/N of the other main emission lines ($[\text{O III}] \lambda 5007$, $[\text{N II}] \lambda 6584$, and $[\text{S II}] \lambda 6717$; S/N_{eml}) greater than 3 to ensure reliable estimates of $A_{V,\text{gas}}$ and emission line indices. Here we examine whether these S/N criteria affect our results.

In Figure 14, we compare the $\Sigma_{\text{H}\alpha}$ maps on the N2O2 versus O32 diagram for spaxels selected from three S/N cuts: $\text{S/N}_{\text{Bal}} > 5$ and $\text{S/N}_{\text{eml}} > 3$, $\text{S/N}_{\text{Bal}} > 3$ and $\text{S/N}_{\text{eml}} > 3$, and $\text{S/N}_{\text{Bal}} > 1$ and $\text{S/N}_{\text{eml}} > 1$, where the first one is the fiducial cut adopted by this work. Clearly, the selected sample is nearly unaffected by relaxing the S/N_{Bal} limit from 5 to 3 when the S/N_{eml} cut is fixed to 3. However, setting both the S/N_{Bal} and S/N_{eml} limits to 1 introduces a low- $\Sigma_{\text{H}\alpha}$ regime within which spaxels should be classified as DIG-dominated regions. This additional regime also cannot be covered by the model grid of H II regions from Dopita et al. (2013). To understand these low- $\Sigma_{\text{H}\alpha}$ spaxels, we further select spaxels included by $\text{S/N}_{\text{Bal}} > 1$ and $\text{S/N}_{\text{eml}} > 1$ but excluded by $\text{S/N}_{\text{Bal}} > 3$ and $\text{S/N}_{\text{eml}} > 3$ and plot their $\Sigma_{\text{H}\alpha}$ map in Figure 15.

Intriguingly, the selected subsample exhibits a significant trend of increasing $\Sigma_{\text{H}\alpha}$ from the high- $[\text{O III}] \lambda 5007$ and low- $[\text{N II}] \lambda 6584$ corner to the low- $[\text{O III}] \lambda 5007$ and high- $[\text{N II}] \lambda 6584$ corner. The bimodal 68% contour indicates that this subsample might be a combination of two populations with different average $\Sigma_{\text{H}\alpha}$ and even different excitation

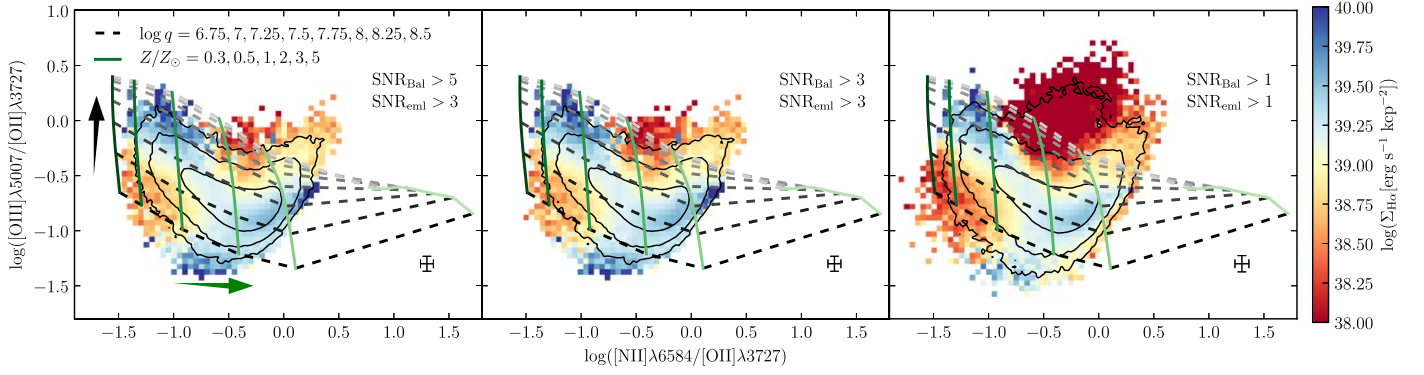


Figure 14. Median maps of $\Sigma_{H\alpha}$ on the N2O2 vs. O32 diagram for spaxels with $S/N_{\text{Bal}} > 5$ and $S/N_{\text{eml}} > 3$ (left), $S/N_{\text{Bal}} > 3$ and $S/N_{\text{eml}} > 3$ (middle), and $S/N_{\text{Bal}} > 1$ and $S/N_{\text{eml}} > 1$ (right). The symbols are the same as in Figure 4.

mechanisms. We visually define a demarcation expressed as

$$\log\left(\frac{[\text{O III}]\lambda 5007}{[\text{O II}]\lambda 3727}\right) = \log\left(\frac{[\text{N II}]\lambda 6584}{[\text{O II}]\lambda 3727}\right) - 0.1 \quad (8)$$

to cross two saddle points of the 68% contour and separate the subsample into low- $\Sigma_{H\alpha}$ (upper) and high- $\Sigma_{H\alpha}$ (lower) branches. The median $\Sigma_{H\alpha}$ and the corresponding 1σ ranges are $37.78^{+0.42}_{-0.32}$ and $38.71^{+0.41}_{-0.56}$ for the upper and lower branches, respectively.

To reveal the possible origins of the upper and lower branches, we present their histograms of $D4000$ and $\text{EW}_{H\alpha}$ in Figure 16 and the $\Sigma_{H\alpha}$ maps on the BPT diagrams in Figure 17. The two branches have distinct distributions for both $D4000$ and $\text{EW}_{H\alpha}$. The medians and 68% ranges of $D4000$ are $1.68^{+0.16}_{-0.30}$ and $1.44^{+0.14}_{-0.11}$ for the upper and lower branches, respectively, while the same parameters are $-1.05^{+0.66}_{-2.28}$ and $-7.16^{+3.33}_{-3.95}$ for $\text{EW}_{H\alpha}$. Obviously, the upper branch is dominated by an old stellar population for which star formation is almost quenched, whereas the lower one suggests a relatively young population with a moderate level of star formation. For the BPT diagrams, nearly all of the spaxels at the upper branch locate in the LINER regime on the $[\text{O I}]-[\text{O III}]$ diagram, and half of them lie beyond the H II regime on the $[\text{S II}]-[\text{O III}]$ diagram. For the lower branch, nearly all spaxels are within the H II regime in the $[\text{S II}]-[\text{O III}]$ plane, and only small fraction of spaxels with lower $\Sigma_{H\alpha}$ move toward the LINER regime of the $[\text{O I}]-[\text{O III}]$ plane.

The distinct distributions of the above properties of stellar population and ionized gas indicate the different physical origins of the two branches. The upper branch is comprised of an old stellar population with emission lines arising from LI(N) ER-like regions or the so-called hDIG in Lacerda et al. (2018), which is heated by hot, low-mass, evolved stars. Spaxels at the

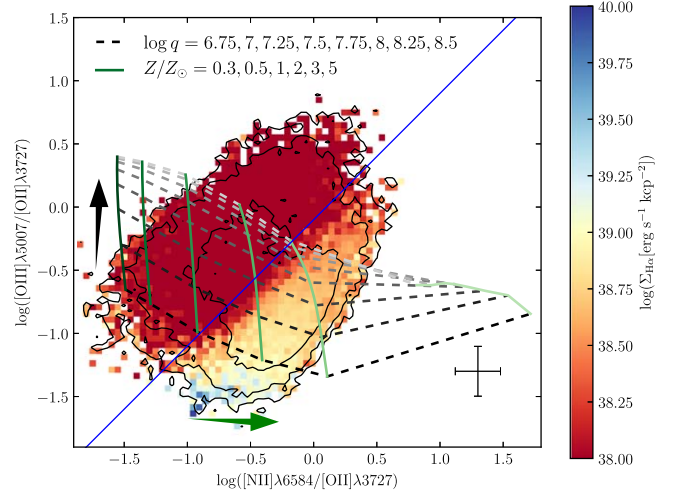


Figure 15. Median map of $\Sigma_{H\alpha}$ on the N2O2 vs. O32 diagram for spaxels included by $S/N_{\text{Bal}} > 1$ and $S/N_{\text{eml}} > 1$ but excluded by $S/N_{\text{Bal}} > 3$ and $S/N_{\text{eml}} > 3$. The blue line indicates a visually selected boundary of $\log([\text{O III}]\lambda 5007/[\text{O II}]\lambda 3727) = \log([\text{N II}]\lambda 6584/[\text{O II}]\lambda 3727) - 0.1$ that is used to divided the spaxels into low- $\Sigma_{H\alpha}$ (upper) and high- $\Sigma_{H\alpha}$ (lower) branches. Other symbols are the same as in Figure 4.

lower branch are more likely faint H II regions. Comparison between Figures 14 and 15 shows that our adopted S/N criteria indeed reduce the covered N2O2–O32 parameter space of non-star-forming origins, while the star-forming regime (i.e., the lower branch of Figure 15 or the main star-forming cloud revealed in Figure 4) is almost unaffected. Given the aim of studying the dust attenuation ratio for classical H II regions, we concluded that the fiducial S/N cuts (i.e., $S/N_{\text{Bal}} > 5$ and $S/N_{\text{eml}} > 3$) are suitable, and relaxing the criteria to smaller S/N values cannot change our results.

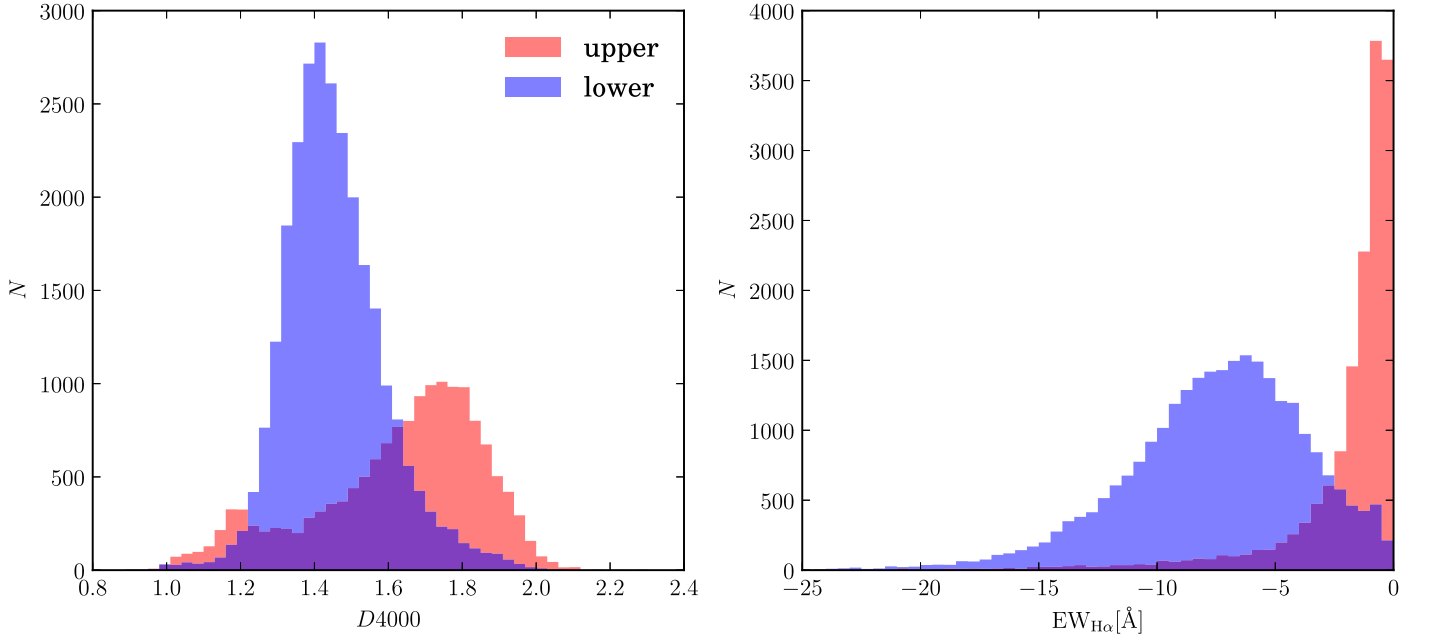


Figure 16. Histogram of $D4000$ (left) and $EW_{H\alpha}$ (right) for spaxels at the upper (red) and lower (blue) branches defined in Figure 15.

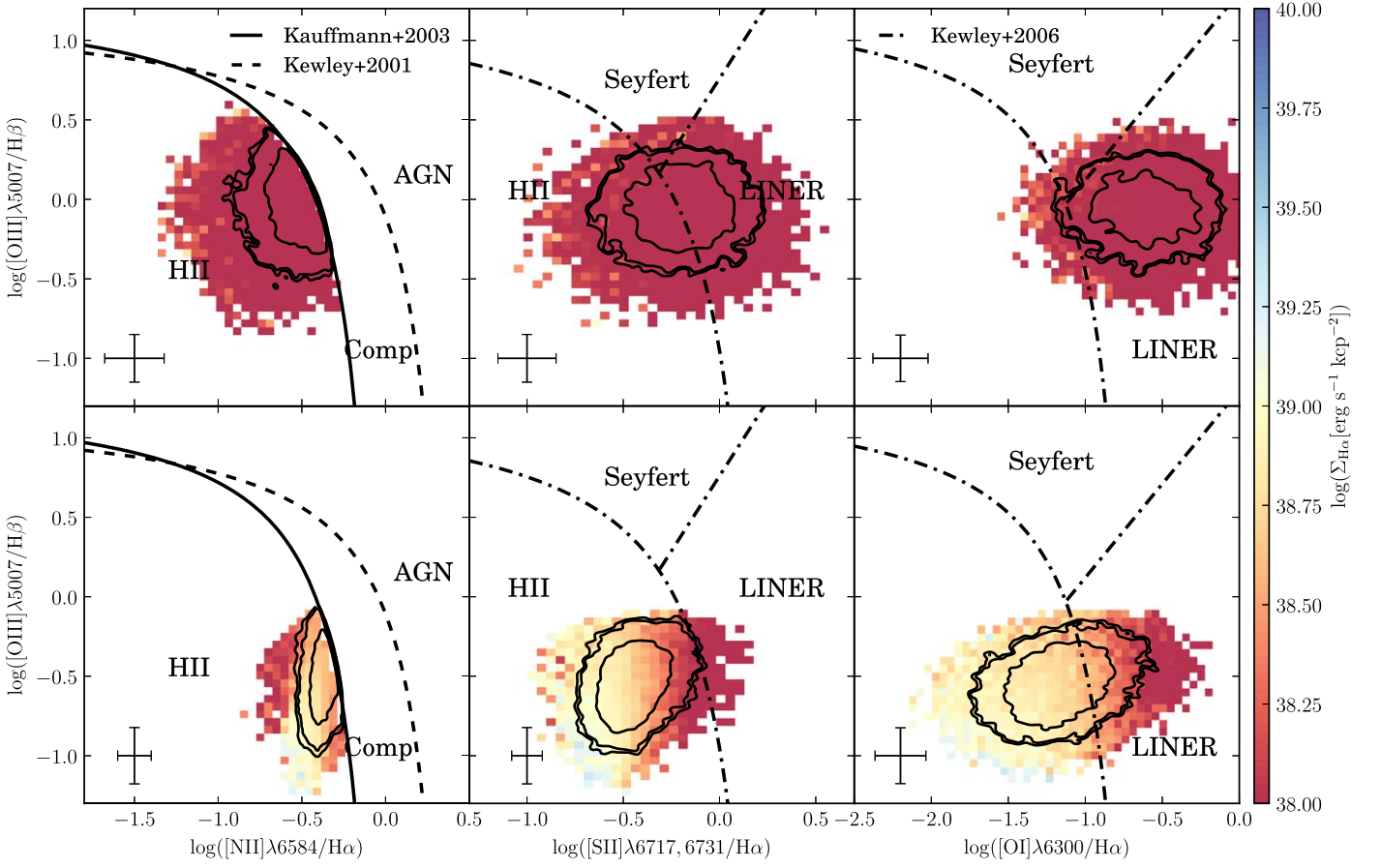


Figure 17. Median maps of $\Sigma_{H\alpha}$ on the BPT diagrams for spaxels at the upper (top) and lower (bottom) branches defined in Figure 15. Boundaries from Kauffmann et al. (2003), Kewley et al. (2001), and Kewley et al. (2006) are denoted by solid, dashed, and dashed-dotted black curves, respectively. The contours encompass 68%, 95%, and 99% of the spaxels on each diagram. The typical uncertainties are shown in the lower left corner.

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