

Pseudosmooth tribrid inflation in SU(5)

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Abstract. A realistic tribrid model of sneutrino inflation is constructed in an R -symmetric SU(5) grand unified theory (GUT). To avoid the monopole problem, a pseudosmooth inflationary trajectory is generated with the help of an additional Z_5 symmetry which is broken during and after inflation. The predictions of inflationary parameters are made at the central value of the scalar spectral index, $n_s = 0.968$. The largest possible value of the tensor to scalar ratio, $r \lesssim 0.0027$, is obtained with sub-Planckian field values ($\lesssim m_P$). A successful realization of reheating and leptogenesis is achieved by avoiding the gravitino problem with a reheat temperature as low as 10^6 GeV. The predicted range of the gauge symmetry breaking scale, $5 \times 10^{16} \lesssim M/\text{GeV} \lesssim 5 \times 10^{17}$, turns out to be somewhat larger than the typical GUT scale. With additional vector-like families, a successful gauge coupling unification is achieved by avoiding the no-go theorem related to R -symmetric SU(5) GUT.

Keywords: inflation, particle physics - cosmology connection

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1 Introduction

An interesting extension of supersymmetric hybrid inflation [1–5] is tribrid inflation [6–10] where a matter field can be employed to realize inflation. One of the simplest candidates for tribrid (matter) inflation could be a sneutrino, the superpartner of the right handed neutrino. An early model of sneutrino inflation was proposed in [11] as a chaotic model of inflation. [Also see [12, 13] where the various predictions of this model were compared with the available experimental data.] This model, however, is plagued with the common problems of realizing chaotic inflation in a supergravity framework [14]. The first model of sneutrino tribrid inflation was introduced by in [6]. This framework, however, is not suited for realizing inflation in a grand unified theory (GUT) model since the gauge symmetry associated with the waterfall GUT Higgs field breaks down at the end of inflation, and so the monopole problem is not resolved. In addition, a domain wall problem arises from the spontaneous breaking of a Z_4 symmetry which is introduced to constrain the structure of the superpotential. In order to resolve this problem higher order Z_4 symmetry breaking terms are introduced. For a general discussion of tribrid inflation see [9, 15], where three types of tribrid inflation are identified, depending on terms of different origin dominating the scalar potential. From these scenarios, only pseudosmooth tribrid inflation [9] is well suited for GUTs with the potential monopole problem. In pseudosmooth tribrid inflation, a shifted smooth track is employed for inflation with the GUT symmetry broken during inflation, such that the monopoles produced during inflation are inflated away.

In this paper we study the possibility of realizing sneutrino tribrid inflation in $SU(5)$ GUT. A pseudosmooth tribrid inflation model employing a Z_5 symmetry is particularly suited for the $SU(5)$ case. A non-minimal Kähler potential is required for the realization of this model. Including a supergravity mass term for the waterfall GUT Higgs field a shifted smooth track, suitable for inflation, can be generated. On this track the $SU(5) \times Z_5$ symmetry is broken and, therefore, any defects produced during inflation are inflated away. Another common problem in an R -symmetric $SU(5)$ GUT is the presence of light triplet and octet fields [16, 17], so that a successful gauge coupling unification in minimal supersymmetric standard model (MSSM) is spoiled. According to a no-go theorem discussed in [18, 19],

this is a generic problem of R -symmetric GUTs based on a simple group. This problem is, however, circumvented in our model with the help of additional vector-like families and as we shall show, a successful gauge coupling unification is achieved. Moreover, assuming relatively large squarks/sleptons masses of order 10 TeV or so, the dimension five proton decay rate is suppressed in accordance with the experimental bound [20].

An attractive feature of sneutrino tribrid inflation is the realization of reheat temperature as low as 10^6 GeV. This feature naturally avoids the gravitino problem usually encountered in supergravity models of inflation. A model of non-thermal leptogenesis [21] is employed in order to explain the observed baryon asymmetry. The numerical predictions of the various inflationary observables are found to be in perfect agreement with the latest Planck 2018 results [22, 23]. In particular, a tensor to scalar ratio $r \approx 0.0027$ can be obtained, and this hopefully can be tested in future experiments [24, 25].

2 Superpotential for tribrid inflation in $SU(5) \times Z_5$ model

The minimal supersymmetric standard model (MSSM) matter content with right handed neutrinos are embedded into $\bar{5}_i$, 10_i and 1_i dimensional representations of supersymmetric $SU(5)$ as

$$\begin{aligned}\bar{5}_i &= D_i^c(\bar{3}, 1, 1/3) + L_i(1, 2, -1/2), \\ 10_i &= Q_i(3, 2, 1/6) + U_i^c(\bar{3}, 1, -2/3) + E_i^c(1, 1, 1), \\ 1_i &= N_i = \nu_i^c(1, 1, 0),\end{aligned}\tag{2.1}$$

where i is the generation index ($i = 1, 2, 3$) and $N_i = \nu_i^c$ represents the right handed neutrino superfield. The GUT Higgs superfield, 24_H , is responsible for the breaking of $SU(5)$ into MSSM whereas the electroweak Higgs doublets (H_u , H_d) contained in the 5_H and $\bar{5}_H$ Higgs superfields trigger the electroweak breaking. The decomposition of minimal Higgs sector in terms of MSSM superfields is given by

$$\begin{aligned}5_H &= H_T(3, 1, -1/3) + H_u(1, 2, 1/2), \\ \bar{5}_H &= \bar{H}_T(\bar{3}, 1, 1/3) + H_d(1, 2, -1/2), \\ 24_H &= H_{24}(1, 1, 0) + W_H(1, 3, 0) + G_H(8, 1, 0) + X_H(3, 2, -5/6) + \bar{X}_H(\bar{3}, 2, 5/6).\end{aligned}\tag{2.2}$$

The desired superpotential of an R -symmetric $SU(5) \times Z_5$ model, including a gauge-singlet superfield S , can be written as

$$\begin{aligned}W &= \kappa S \left(\mu^2 + \frac{\text{Tr}(24_H^5)}{m_P^3} + \alpha \frac{\text{Tr}(24_H^2) \text{Tr}(24_H^3)}{m_P^3} \right) - \beta_{ij} \frac{\text{Tr}(24_H^3)}{m_P^2} N_i N_j \\ &\quad + \lambda_1 \frac{\text{Tr}(24_H^2)}{m_P} \bar{5}_H 5_H + \frac{\lambda_2}{m_P} \bar{5}_H 24_H^2 5_H \\ &\quad + y_{ij}^{(u)} 10_i 10_j 5_H + y_{ij}^{(d,e)} 10_i \bar{5}_j \bar{5}_H + \frac{\lambda_{ij}^\nu}{m_P^2} \text{Tr}(24_H^2) N_i \bar{5}_j 5_H + \frac{\tilde{\lambda}_{ij}^\nu}{m_P^2} N_i \bar{5}_j 24_H^2 5_H,\end{aligned}\tag{2.3}$$

where μ is a superheavy mass, $m_P = 2.43 \times 10^{18}$ GeV is the reduced Planck mass and all other couplings ($\kappa, \alpha, \beta_{ij}, \lambda_1, \lambda_2, \dots$) are dimensionless. The charge assignments of the various

superfields under $U(1)_R$ and Z_5 symmetries are respectively given by

$$\begin{aligned} R(S, 24_H, 5_H, \bar{5}_H, 10_i, \bar{5}_i, N_i) &= \left(1, 0, \frac{2}{5}, \frac{3}{5}, \frac{3}{10}, \frac{1}{10}, \frac{1}{2}\right), \\ q_5(S, 24_H, N_i, 5_H, \bar{5}_H, 10_i, \bar{5}_i) &= (0, 1, 1, 3, 0, 1, 4), \end{aligned} \quad (2.4)$$

with $R(W) = 1$. The terms in the first line of the superpotential W are relevant for tribrid inflation which is discussed below in detail. Owing to $SU(5)$ gauge invariance of the superpotential it is required to align Higgs 24_H superfield along the standard model (SM) gauge singlet direction, H_{24} , as

$$24_H \longrightarrow H_{24} = \frac{h}{\sqrt{15}}(1, 1, 1, -3/2, -3/2). \quad (2.5)$$

The global supersymmetric minimum, therefore, occurs at

$$\langle h^5 \rangle \equiv M^5 = \frac{8\sqrt{15}}{\left(\frac{13}{30} + \alpha\right)} \mu^2 m_P^3, \quad \langle S \rangle = 0, \quad \langle N_i \rangle = 0, \quad (2.6)$$

for the relevant superfields. The importance of the various terms in the superpotential can now be described conveniently in terms of h and its vacuum expectation value M .

The terms in the second line of eq. (2.3),

$$W \supset \frac{h^2}{m_P} \left(\left(\frac{\lambda_1}{2} + \frac{\lambda_2}{15} \right) \bar{H}_T H_T + \left(\frac{\lambda_1}{2} + \frac{3\lambda_2}{20} \right) H_u H_d \right) \supset \mu_2 H_u H_d + \mu_3 \bar{H}_T H_T, \quad (2.7)$$

are relevant for the doublet-triplet problem. Here, the mass parameter, μ_2 , is just the μ -parameter of MSSM which is usually taken to be of electroweak scale with $\lambda_1 \simeq -3\lambda_2/10$. On the other hand, the mass parameter, $\mu_3 \simeq -\lambda_1(M/m_P)M/12$, is taken to be order GUT scale in order to suppress dimension-5 proton decay amplitude mediated by the color triplet Higgs pair. This further requires the squark/slepton masses to be $\gtrsim 10$ TeV. Therefore, the doublet-triplet problem is solved, as usual, by fine tuning. Lastly, the couplings, $y_{ij}^{(u)}, y_{ij}^{(d,e)}, \lambda_{ij}^{(\nu)}, \tilde{\lambda}_{ij}^{(\nu)}$, in the third line of eq. (2.3) include the quark and lepton Yukawa couplings. In order to obtain the observed tiny neutrino masses, Majorana mass terms for the right handed neutrinos are required. Even though an explicit Majorana mass term is not allowed due to Z_5 symmetry, the spontaneous breaking of $SU(5)$ gauge symmetry generates an effective Majorana mass term, $(1/2)M_{ij}^R N_i N_j$, with

$$M_{ij}^R = \frac{\beta_{ij}}{2\sqrt{15}} \left(\frac{M}{m_P} \right)^2 M, \quad (2.8)$$

from the last term in the first line of eq. (2.3). Taking Majorana masses to be order 10^{13} GeV the light neutrino masses are naturally explained via type-I seesaw mechanism. As we discuss below, this term also plays an important role in realizing sneutrino inflation and subsequent reheating.

3 Inflationary scalar potential

To discuss inflation we consider the following superpotential terms from eq. (2.3),

$$\begin{aligned} W &\supset \kappa S \left(\mu^2 + \frac{\text{Tr}(24_h^5)}{m_P^3} + \alpha \frac{\text{Tr}(24_h^3) \text{Tr}(24_h^2)}{m_P^3} \right) - \beta_{ij} \frac{\text{Tr}(24_h^3)}{m_P^2} N_i N_j, \\ &\supset \mu^2 S \left(1 - \left(\frac{h}{M} \right)^5 \right) + \beta(\mu^2 m_P) \left(\frac{h}{M} \right)^3 \left(\frac{N}{M} \right)^2, \end{aligned} \quad (3.1)$$

where $\beta = \frac{2\beta_{11}}{\left(\frac{13}{30} + \alpha\right)}$, $N \equiv N_1$ and to achieve $M \ll m_P$ with a natural value of κ we set $\kappa = 1$. With smaller values of κ , the value of M becomes Planckian. As there is no contribution from the relevant fields in the D -term scalar potential, the global SUSY scalar potential obtained from the F -term is given by,

$$V_F = \mu^4 \left(\left| 1 - \left(\frac{h}{M} \right)^5 \right|^2 + \left| 3\beta \left(\frac{m_P}{M} \right) \left(\frac{N^2 h^2}{M^4} \right) - \frac{5Sh^4}{M^5} \right|^2 + \left| 2\beta \left(\frac{m_P}{M} \right) \left(\frac{Nh^3}{M^4} \right) \right|^2 \right), \quad (3.2)$$

where, $V_F = |\partial W / \partial z_i|^2$, with $z_i \in (S, h, N)$. To keep the discussion simple we assume that the phases of the fields have been stabilized before the start of observable inflation and, therefore, the above potential reduces to the following form,

$$V_F = \mu^4 \left((1 - z^5)^2 + \left(3\beta \left(\frac{m_P}{M} \right) y^2 z^2 - 5xz^4 \right)^2 + \left(2\beta \left(\frac{m_P}{M} \right) yz^3 \right)^2 \right), \quad (3.3)$$

where,

$$x = \frac{|S|}{M}, \quad y = \frac{|N|}{M}, \quad z = \frac{|h|}{M}. \quad (3.4)$$

Next we aim to find an effective single field form of the above potential, and to achieve this goal we need to include supergravity (SUGRA) corrections which are obtained from the following formula,

$$V_F = e^{K/m_P^2} \left(K_{ij}^{-1} D_{z_i} W D_{z_j}^* W^* - 3m_P^{-2} |W|^2 \right), \quad (3.5)$$

where

$$D_{z_i} W = \frac{\partial W}{\partial z_i} + \frac{1}{m_P^2} \frac{\partial K}{\partial z_i} W, \quad K_{ij} = \frac{\partial^2 K}{\partial z_i \partial z_j^*}, \quad D_{z_j}^* W^* = (D_{z_i} W)^*. \quad (3.6)$$

Here, we consider the following power-law expansion of the Kähler potential

$$\begin{aligned} K &= |S|^2 + |N|^2 + \text{Tr} |24_H|^2 \\ &+ \kappa_S \frac{|S|^4}{4m_P^2} + \kappa_N \frac{|N|^4}{4m_P^2} + \kappa_h \frac{(\text{Tr} |24_H|^2)^2}{4m_P^2} \\ &+ \kappa_{SN} \frac{|S|^2 |N|^2}{m_P^2} + \kappa_{Sh} \frac{|S|^2 \text{Tr} |24_H|^2}{m_P^2} + \kappa_{Nh} \frac{|N|^2 \text{Tr} |24_H|^2}{m_P^2} \dots \end{aligned} \quad (3.7)$$

Including only the relevant SUGRA correction terms, the scalar potential as a function of the three fields is given by

$$V_3(x, y, z) = \mu^4 \left((1 - z^5)^2 + \left(3\beta \left(\frac{m_P}{M} \right) y^2 z^2 - 5xz^4 \right)^2 + \left(2\beta \left(\frac{m_P}{M} \right) yz^3 \right)^2 - \kappa_S \left(\frac{M}{m_P} \right)^2 x^2 + \kappa_h \left(\frac{M}{m_P} \right)^2 z^2 + \gamma \left(\frac{M}{m_P} \right)^2 y^2 + \delta \left(\frac{M}{m_P} \right)^4 y^4 + \dots \right), \quad (3.8)$$

where $\gamma = 1 - \kappa_{SN}$ and $\delta = \frac{1}{2} + \kappa_{SN}^2 - \kappa_{SN} + \frac{1}{4}\kappa_N$.

Stabilization of S ($x = S/M$) field

In order to obtain an effective single-field potential we first minimize the three-field potential V_3 with respect to x . The potential minimum occurs at,

$$x_{\min} = \frac{15y^2 z^6 \beta m_P^3}{25m_P^2 z^8 - M^2 \kappa_S}, \quad (3.9)$$

with $\beta > 0$ and $\kappa_S < 0$. The mass squared of the S field, m_S^2 , in term of Hubble mass squared, $H^2 \simeq \frac{\mu^4}{3m_P^2}$, is given by

$$m_S^2/H^2 \simeq \left(75 \left(\frac{m_P}{M} \right)^2 z^8 - 3\kappa_S \right). \quad (3.10)$$

Therefore, the S field attains Hubble size mass for $\kappa_S \lesssim -\frac{1}{3}$ and quickly settles down to its minimum. This leads us to the following effective two-field potential,

$$\begin{aligned} V_2(y, z) &\equiv V_3(x_{\min}, y, z) \\ &= \mu^4 \left((1 - z^5)^2 + \left(3\beta \left(\frac{m_P}{M} \right) y^2 z^2 - 5x_{\min} z^4 \right)^2 + \left(2\beta \left(\frac{m_P}{M} \right) yz^3 \right)^2 - \kappa_S \left(\frac{M}{m_P} \right)^2 x_{\min}^2 + \kappa_h \left(\frac{M}{m_P} \right)^2 z^2 + \gamma \left(\frac{M}{m_P} \right)^2 y^2 + \delta \left(\frac{M}{m_P} \right)^4 y^4 + \dots \right), \end{aligned} \quad (3.11)$$

where x_{\min} is given by eq. (3.9). This two-field potential is displayed in figure 1 for values of the various parameters given in the caption. A smooth trajectory suitable for inflation is clearly visible in this figure. For greater clarity a closer look at this trajectory is displayed in figure 2. The smooth trajectory here actually ends at a waterfall point which is shown by a red dot in figure 2. This is the reason why inflation along this trajectory is termed as pseudosmooth inflation [9]. In this model the sneutrino N field actually plays the role of the inflaton whereas variation in the z field remains negligible during inflation.

Waterfall critical point

In pseudosmooth tribrid model inflation ends by a waterfall transition whereas in standard smooth hybrid model it ends by a slow-roll breaking with no waterfall along the complete

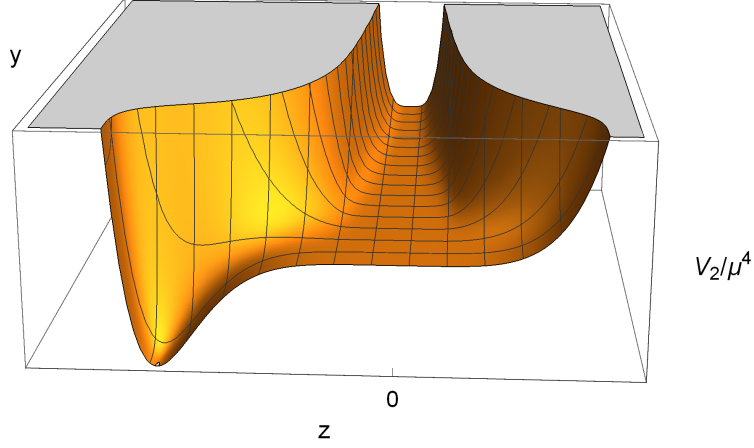


Figure 1. The normalized two-field potential V_2/μ^4 as a function of $z = \frac{h}{M}$ and $y = \frac{N}{M}$ with $x = \frac{S}{M} = x_{\min}$, $\kappa_h = -1$, $\kappa_S = -1/3$, $\gamma = \delta = 0$ and $M = 10^{17}$ GeV. The zoom-in plot of pseudosmooth inflationary valley with $z \neq 0$ is shown in figure 2.

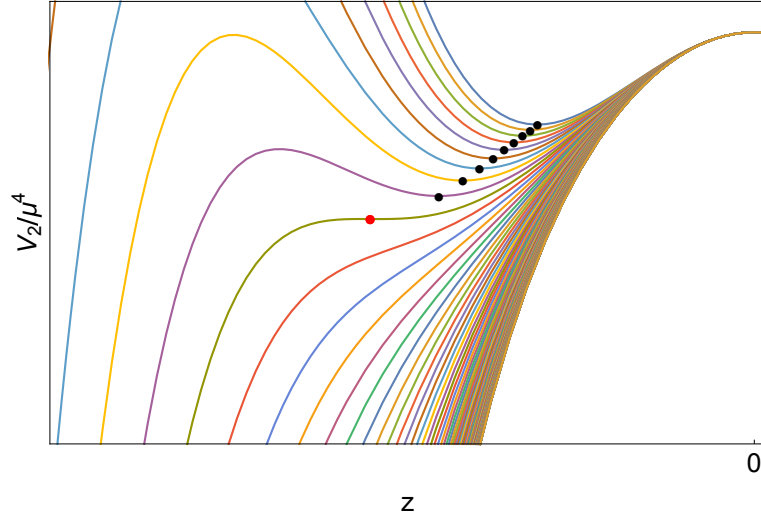


Figure 2. The normalized two-field potential V_2/μ^4 as a function of $z = \frac{h}{M}$ for various values of $y = \frac{N}{M}$. We fix $\gamma = \delta = 0$, $\kappa_S = -\frac{1}{3}$, $\kappa_h = -1$ and $M = 1 \times 10^{17}$ GeV.

smooth trajectory. The waterfall critical point (z_c, y_c) can be obtained from the following conditions,

$$\frac{\partial V_2(z_c, y_c)}{\partial z_c} = \frac{\partial^2 V_2(z_c, y_c)}{\partial z_c^2} = 0. \quad (3.12)$$

Applying these conditions we obtain the following critical point,

$$y_c = \left(\frac{5}{2^{\frac{5}{2}}}\right)^{\frac{1}{6}} \left(\frac{(-\kappa_h)^{\frac{1}{3}}}{3\beta^2 \left(\frac{M}{m_P}\right)^{4/3}}\right)^{\frac{1}{4}} \left(\frac{M}{m_P}\right), \quad z_c = \left(\frac{2}{5}\right)^{\frac{1}{3}} \left((- \kappa_h) \left(\frac{M}{m_P}\right)^2\right)^{\frac{1}{3}}, \quad (3.13)$$

which also defines the condition for the end of inflation.

Effective single-field potential

Finally, minimizing V_2 with respect to field z gives us the following approximate form of the effective single-field potential,

$$V(y) \equiv V_2(y, z_{\min}) \simeq \mu^4 \left(1 - \frac{1}{18y^4} \left(\frac{M}{m_P} \right)^6 \left(\frac{\kappa_h}{\beta} \right)^2 + \gamma \left(\frac{M}{m_P} \right)^2 y^2 + \delta \left(\frac{M}{m_P} \right)^4 y^4 \right), \quad (3.14)$$

along the pseudosmooth trajectory approximately given by

$$z_{\min} \simeq \frac{1}{3} \sqrt{\frac{M^4(-\kappa_h)}{2\beta^2 m_P^4 y^4}}. \quad (3.15)$$

In the leading order slow-roll approximation defined below, the prediction for the various inflationary parameters can be calculated by employing the above form of the potential.

As natural values of the parameters are mostly involved in our analysis with field values of order GUT scale, a quick stabilization of all three phases is generally expected owing to their large masses. In the remaining 3-dimensional field space, any local minimum trajectory, if present, is expected to be isolated from the pseudosmooth trajectory and the global susy minimum. Any impact on inflationary predictions could be from the possible quantum tunneling transitions among these minima. A full analysis of estimating this impact on the model predictions can be quite cumbersome and lies beyond the scope of the current paper.

We have explicitly checked that the radiative corrections are negligibly small in our model. To see it with an order of magnitude estimate, we consider the following values of mass squared,

$$4 \left((-\kappa_h) \pm \sqrt{(-\kappa_h)^2 + \frac{2\kappa_h^3}{9\beta^2 y^6} \left(\frac{M}{m_P} \right)^4} \right) \frac{\mu^4}{m_P^2}, \quad 2(-\kappa_h) \frac{\mu^4}{m_P^2}, \quad 2(-\kappa_h) \frac{\mu^4}{m_P^2}, \quad (3.16)$$

for the inflaton-Higgs system in the limit $x = x_{\min} \ll 1$ and $z = z_{\min} \ll 1$. With an appropriate choice of renormalization scale, the radiative correction is proportional to $(-\kappa_h)^2(\mu/m_P)^4\mu^4$. As the quadratic mass term in the above potential plays equally important role in realizing inflation along with the other terms, the radiative correction can be ignored compared to this term for $\gamma \gg (-\kappa_h)^2(\mu/m_P)^2(\mu/M)^2$ with $N = M$. This constraint is naturally satisfied in our numerical estimates. [Also see [15] for a discussion of the smallness of radiative corrections in a typical model of tribrid inflation.]

The suppression of soft SUSY breaking terms with TeV scale soft masses is a common feature of tribrid inflation. In our model this can be seen with the following argument. As both x_{\min} and z_{\min} are very small during the bulk of the inflationary phase ($10^{-10} \lesssim x_{\min} \lesssim 10^{-5}$ and $0.005 \lesssim z_{\min} \lesssim 0.05$), both W and $z_i \partial W / \partial z_i \sim \mathcal{O}(W)$ turn out to be negligibly small. Hence, the soft SUSY breaking A -term is negligible. Furthermore, the soft mass term $m_{\text{soft}}^2 |N|^2$ can be ignored compared to the quadratic mass term in the above potential for $m_{\text{soft}} \ll \sqrt{\gamma} \mu (\mu/m_P)$. With $\mu \lesssim (10^{13} - 10^{16}) \text{ GeV}$, we obtain a soft mass $m_{\text{soft}} \ll \sqrt{\gamma} (10^5 - 10^{11}) \text{ TeV}$. Thus, the approximation of ignoring the soft SUSY breaking terms throughout our calculations is justified.

Inflationary slow-roll parameters

The slow-roll parameters are given below

$$\epsilon(y) = \frac{1}{4} \left(\frac{m_{\text{P}}}{M} \right)^2 \left(\frac{\partial_y V}{V} \right)^2, \quad \eta(y) = \frac{1}{2} \left(\frac{m_{\text{P}}}{M} \right)^2 \left(\frac{\partial_y^2 V}{V} \right), \quad \xi^2(y) = \frac{1}{4} \left(\frac{m_{\text{P}}}{M} \right)^4 \left(\frac{\partial_y V \partial_y^3 V}{V^2} \right), \quad (3.17)$$

where the subscript y on ∂ denotes the derivative with respect to y . In the leading order slow-roll approximation, with $(\epsilon, \eta, \xi^2) \ll 1$, the tensor-to-scalar ratio r , the scalar spectral index n_s and the running of the scalar spectral index $dn_s/d\ln k$ are given by

$$n_s \simeq 1 + 2\eta(y_0) - 6\epsilon(y_0), \quad r \simeq 16\epsilon(y_0), \quad (3.18)$$

$$\frac{dn_s}{d\ln k} \simeq 16\epsilon(y_0)\eta(y_0) - 24\epsilon^2(y_0) - 2\xi^2(y_0), \quad (3.19)$$

where y_0 is the field value at the pivot scale which is taken to be at $k_0 = 0.05 \text{ Mpc}^{-1}$. The amplitude of curvature perturbation is given by

$$A_s(k_0) = \frac{1}{24\pi^2} \left(\frac{V(y)/m_{\text{P}}^4}{\epsilon(y)} \right) \Big|_{y=y_0}, \quad (3.20)$$

where $A_s(k_0) = 2.142 \times 10^{-9}$ is the Planck normalization at $k_0 = 0.05 \text{ Mpc}^{-1}$ [22, 23]. This constraint can be used to express μ in terms of r ,

$$\mu \simeq \left(\frac{3A_s(k_0)\pi^2 r}{2} \right)^{\frac{1}{4}} m_{\text{P}}. \quad (3.21)$$

The number of efolds, ΔN , from the pivot scale to the end of inflation is given by

$$\Delta N = 2 \left(\frac{M}{m_{\text{P}}} \right)^2 \int_{y_e}^{y_0} \frac{V}{\partial_y V} dy, \quad (3.22)$$

where the field value at the end of inflation is $y_e = y_c$. Assuming standard thermal history we express the number of e-folds, ΔN , in terms of the reheat temperature, T_r , as

$$\Delta N \simeq 47 + \frac{1}{3} \ln \left(\frac{T_r}{10^6 \text{ GeV}} \right) + \frac{2}{3} \ln \left(\frac{\mu}{10^{13} \text{ GeV}} \right). \quad (3.23)$$

In estimating the numerical predictions of the various inflationary parameters we set $T_r = 10^6 \text{ GeV}$. The realization of such a low reheat temperature and related non-thermal leptogenesis is justified after the discussion of numerical results.

4 Discussion of numerical results

The numerical predictions of inflationary parameters are estimated by fixing the scalar spectral index at its central value, $n_s = 0.968$, and by setting $y_e = y_c = 1$. Using $y_c = 1$ in eq. (3.13) the parameter β can be written in terms of the other parameters as,

$$\beta = \left(\frac{5}{2^{\frac{5}{2}}} \right)^{\frac{1}{3}} \left(\frac{(-\kappa_h)^{\frac{1}{3}}}{3 \left(\frac{M}{m_{\text{P}}} \right)^{4/3}} \right)^{\frac{1}{2}} \left(\frac{M}{m_{\text{P}}} \right)^2. \quad (4.1)$$

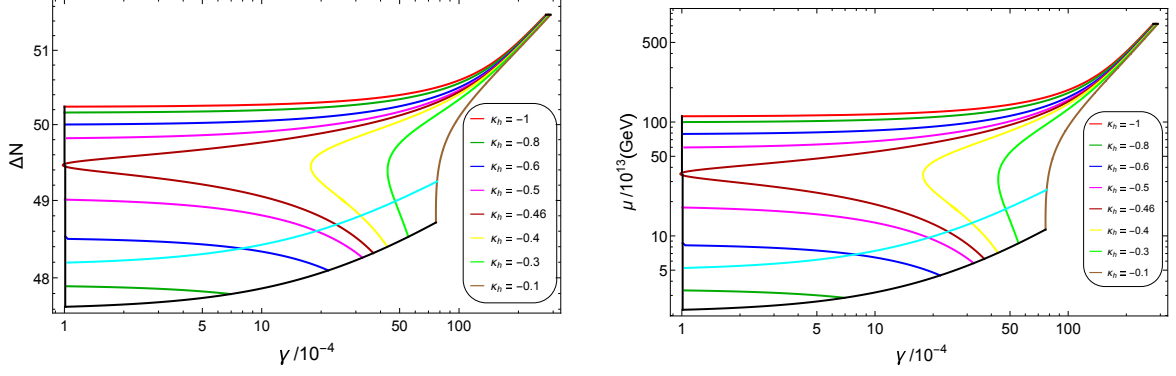


Figure 3. The number of e-folds ΔN (left panel) and the μ (right panel) versus the coupling γ . We set the scalar spectral index, $n_s = 0.968$ (central value of Planck's data), the end of inflation, $y_e = \frac{N_e}{M} = 1$, and the reheat temperature, $T_r = 10^6$ GeV.

We fix the reheat temperature to its lowest possible value, i.e. $T_r = 10^6$ GeV, allowed by the successful non-thermal leptogenesis, as discussed in section 5 below. This value avoids the gravitino problem for a relatively wider range of gravitino mass [26–28]. We further impose $N_0 \leq m_P$ as required by the reliability of supergravity corrections.

To identify a relatively natural region of parametric space we restrict $|\gamma| \gtrsim 10^{-4}$, $|\delta| \lesssim 1$ and $-1 \leq \kappa_h \leq -0.1$. We require $\kappa_h < 0$ in order to generate a smooth inflationary track while a successful realization of inflation further requires $\delta < 0$ and $\gamma > 0$. Assuming $\alpha \ll 13/30$, the parameter μ and tensor to scalar ratio, r , can be directly related to M , via eqs. (2.6) and (3.21), as

$$\mu \simeq \sqrt{\frac{13}{240\sqrt{15}}} \left(\frac{M}{m_P}\right)^{5/2} m_P, \quad r \simeq \left(\frac{1}{2 \times 10^{-4}}\right) \left(\frac{M}{m_P}\right)^{10}. \quad (4.2)$$

From eq. (3.23), we can also write down μ in terms of ΔN as $\mu/10^{13} \text{ GeV} = e^{3(\Delta N - 47)/2}$ with $T_r = 10^6$ GeV. After solving eq. (3.22) and eq. (3.23) numerically with the above mentioned constraints we obtain $\Delta N \simeq 47.6\text{--}51.5$ as shown in figure 3. This range of ΔN corresponds to the range $2 \lesssim \mu/10^{13} \text{ GeV} \lesssim 744$. This yields, via the above equations, $5.4 \times 10^{16} \lesssim M/\text{GeV} \lesssim 5.6 \times 10^{17}$ and $10^{-13} \lesssim r \lesssim 2.4 \times 10^{-3}$. Moreover, the approximate upper bound $M \lesssim 5.6 \times 10^{17}$ with $N_0 = m_P$ translates into the upper bound $y_0 \lesssim 4$. These approximate estimates are compatible with the exact numerical results displayed in figure 4.

The explicit dependence of γ and δ in terms of the remaining parameters can be obtained from eq. (3.17) as

$$\gamma \simeq \frac{1}{4} (1 - n_s) + \frac{3\sqrt{r}}{8y_0} \left(\frac{m_P}{M}\right) + \frac{8\kappa_h}{3y_0^6} \left(-\frac{2\kappa_h}{5} \left(\frac{M}{m_P}\right)^2\right)^{\frac{2}{3}}, \quad (4.3)$$

$$\delta \simeq \frac{(n_s - 1)}{8y_0^2} \left(\frac{m_P}{M}\right)^2 - \frac{\sqrt{r}}{16y_0^3} \left(\frac{m_P}{M}\right)^3 - \frac{\kappa_h}{y_0^8} \left(-\frac{2\kappa_h}{5} \left(\frac{m_P}{M}\right)\right)^{\frac{2}{3}}. \quad (4.4)$$

For relatively large values of r only the first two terms in the above expressions are important. This leads to the weak dependence of γ and δ on κ_h as depicted in figures 3–5. For instance,

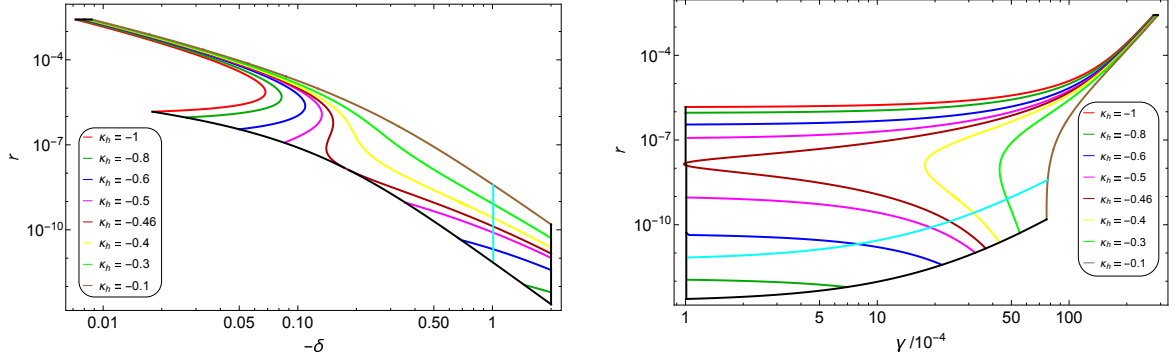


Figure 4. The tensor to scalar ratio r versus the coupling γ (left panel) and the coupling δ (right panel). We set the scalar spectral index $n_s = 0.968$ (central value of Planck’s data), the end of inflation $y_e = \frac{N_e}{M} = 1$ and the reheat temperature $T_r = 10^6$ GeV.

with $N_0 = m_P$ we obtain

$$\gamma \simeq \frac{1}{4}(1 - n_s) + \frac{3\sqrt{r}}{8} \simeq 0.026, \quad \delta \simeq -\frac{(1 - n_s)}{8} - \frac{\sqrt{r}}{16} \simeq -0.007. \quad (4.5)$$

This again is a very good approximation of the more precise numerical estimates shown in figures 3–5.

In the small r limit the second terms in eqs. (4.3) and (4.4) become negligible. The last term in eq. (4.3) becomes comparable to the first term while making γ small compared to $(1 - n_s)/4$. This fact allows us to write y_0 and M in terms of δ and κ_h as

$$y_0 \simeq \left(\frac{2^7(-\kappa_h)^5}{3^3 \times 5^2(n_s - 1)\delta^2} \right)^{\frac{1}{22}}, \quad M \simeq \frac{1}{4\delta^{\frac{3}{2}}} \left(\frac{3^{\frac{3}{2}} \times 5((n_s - 1)\delta^2)^6}{2^9(-\kappa_h)^{\frac{5}{2}}} \right)^{\frac{1}{11}} m_P, \quad (4.6)$$

for $\gamma \ll (1 - n_s)/4$. Using these expressions with $y_0 \simeq 1$ and $-0.8 \lesssim \kappa_h \lesssim -0.5$ we obtain $-1.4 \lesssim \delta \lesssim -0.4$ and $6.5 \times 10^{16} \text{ GeV} \lesssim M \lesssim 1.2 \times 10^{17} \text{ GeV}$, which is in good agreement with our numerical estimates, as shown in figures 4–6. Note that in these figures the lower bounds on M and r are very sensitive to the upper bound on $|\delta|$. For example, increasing the value of $|\delta|$ above unity, represented by the cyan curve, can further reduce the lower bound on M and r as shown by the $\delta = -2$ black curve. Moreover, the bound on M with $-1 \leq \kappa_h \leq -0.1$ gives the range, $10^{-3} \lesssim \beta \lesssim 10^{-2}$, via eq. (4.1).

Finally, the running of spectral index $\frac{dn_s}{d \ln k}$ can be described in terms of r , n_s and other parameters as

$$\alpha_s \equiv \frac{dn_s}{d \ln k} \simeq r \frac{(n_s - 1)}{2} - \frac{3}{32} r^2 - 2\sqrt{r} \left(\frac{2^{2/3} 5^{1/3} (-\kappa_h)^{5/3}}{y_0^7} \left(\frac{M}{m_P} \right)^{1/3} + 3y_0 \delta \left(\frac{M}{m_P} \right) \right). \quad (4.7)$$

The largest possible value of $\frac{dn_s}{d \ln k} \lesssim 2 \times 10^{-3}$ appears in the large r limit (see figure 6). This shows that our results are perfectly consistent with the latest Planck data results.

5 Reheating and non-thermal leptogenesis

The reheating in the current model proceeds in analogy with the Z_4 sneutrino model [6], and the observed baryon asymmetry is explained by nonthermal leptogenesis [21]. The inflaton

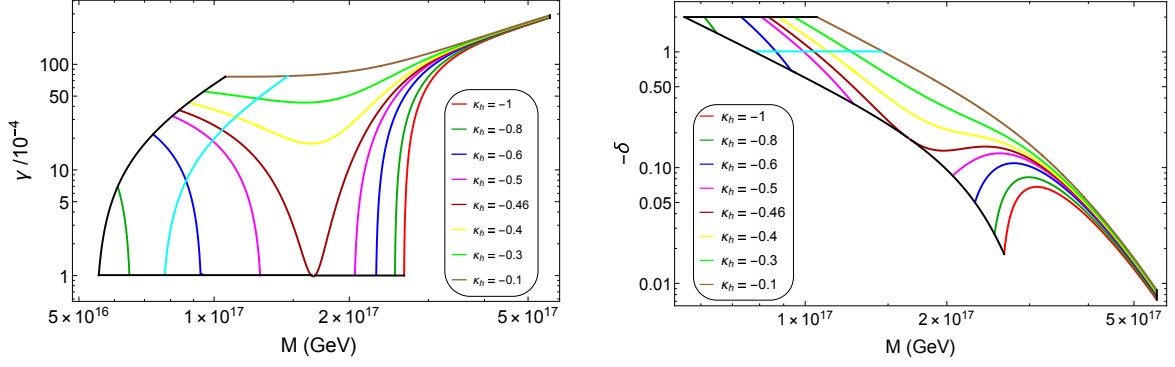


Figure 5. The coupling γ (left panel) and the coupling δ (right panel) versus the gauge symmetry breaking scale M . We set the scalar spectral index $n_s = 0.968$ (central value of Planck's data), the end of inflation $y_e = \frac{N_e}{M} = 1$ and the reheat temperature $T_r = 10^6$ GeV.

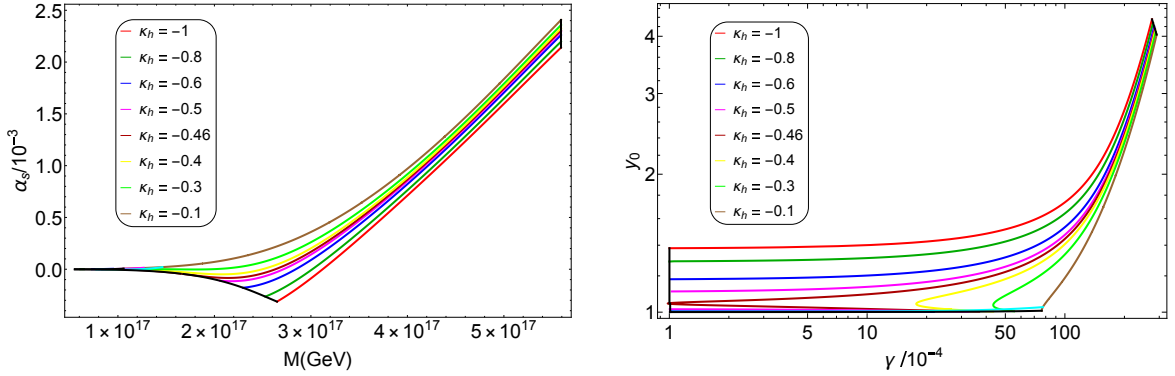


Figure 6. The running of spectral index $\alpha_s \equiv \frac{dn_s}{d \ln k}$ (left panel) and the normalized field value $y_0 = \frac{N_0}{M}$ (right panel) versus the coupling γ . We set the scalar spectral index $n_s = 0.968$ (central value of Planck's data), the end of inflation $y_e = \frac{N_e}{M} = 1$ and the reheat temperature $T_r = 10^6$ GeV.

N_1 , being the lightest sneutrino field, is assumed to play a dominant role in reheating after inflation and subsequent leptogenesis. This is possible if the Higgs field decays earlier than the inflaton as discussed below. From eq. (2.8), the mass of the singlet sneutrino inflaton N_1 is given by

$$M_R^I = \left(\frac{\beta_{11}}{2\sqrt{15}} \right) \frac{M^3}{m_P^2} = \left(\frac{13}{8y_c^2} \right) \left(\frac{(-\kappa_h)^{\frac{1}{3}}}{3^4 \times 2^{\frac{5}{3}} \times 5^{\frac{7}{3}}} \right)^{\frac{1}{2}} \left(\frac{M}{m_P} \right)^{10/3} M. \quad (5.1)$$

The inflaton decays through the effective Yukawa coupling,

$$\lambda_{1j}^\nu \frac{\text{Tr}(24_h^2)}{m_P^2} N_1 \bar{5}_j 5_h + \frac{\tilde{\lambda}_{1j}^\nu}{m_P^2} N_1 \bar{5}_j 24_h^2 5_h \supset Y_{1j}^\nu N_1 L_j H_u, \quad (5.2)$$

into sleptons and Higgs or into lepton and Higgsino with a decay width given by

$$\Gamma_{N_1} \simeq \frac{y_\nu^2}{4\pi} M_R^I = \frac{y_\nu^2}{4\pi} \left(\frac{13}{8y_c^2} \right) \left(\frac{(-\kappa_h)^{\frac{1}{3}}}{3^4 \times 2^{\frac{5}{3}} \times 5^{\frac{7}{3}}} \right)^{\frac{1}{2}} \left(\frac{M}{m_P} \right)^{10/3} M, \quad (5.3)$$

where,

$$y_\nu^2 \equiv (Y_\nu Y_\nu^\dagger)_{11}, \quad Y_{ij}^\nu = \left(\frac{\lambda_{ij}^\nu}{2} + \frac{\tilde{\lambda}_{ij}^\nu}{15} \right) \left(\frac{M}{m_P} \right)^2. \quad (5.4)$$

Compared to the Z_4 sneutrino tribrid model, we have an extra suppression factor $(M/m_P)^2$ which can make the fundamental Yukawa couplings $(\lambda_{ij}^\nu, \tilde{\lambda}_{ij}^\nu)$ relatively natural. Assuming the Higgs decay rate to be larger than the inflaton decay rate we obtain the following bound on y_ν ,

$$y_\nu^2 \ll \left(\frac{M_R^{(2,3)}}{M} \right) \left(\frac{M_R^{(2,3)}}{M_R^I} \right), \quad (5.5)$$

where $M_R^{(2,3)}$ are the masses of the heavier neutrinos $N_{(2,3)}$. This bound is easily satisfied in our model for the numerical data displayed in figure 7. After inflation, the universe reheats via inflaton decay to a temperature,

$$T_r \simeq \left(\frac{90}{g_* \pi^2} \right)^{\frac{1}{4}} \sqrt{\Gamma_{N_1} m_P}, \quad (5.6)$$

where $g_* = 228.75$.

The lepton asymmetry generated by the inflaton decay can be partially converted into the observed baryon asymmetry through sphaleron processes. We assume $M_R^I \gg T_r$ in order to suppress the washout factor of lepton asymmetry. The baryon asymmetry can be estimated in terms of the lepton asymmetry factor ε_L as

$$\frac{n_B}{n_\gamma} \simeq -1.84 \varepsilon_L \frac{T_r}{M_R^I}, \quad (5.7)$$

where ε_L satisfies the following bound,

$$(-\varepsilon_L) \lesssim \frac{3}{8\pi} \frac{\sqrt{\Delta m_{31}^2} M_R^I}{\langle H_u \rangle^2}, \quad (5.8)$$

assuming a hierarchical structure of neutrino masses. Here, the atmospheric neutrino mass squared difference is $\Delta m_{31}^2 \approx 2.6 \times 10^{-3} \text{ eV}^2$ and $\langle H_u \rangle = 174 \text{ GeV}$ in the large $\tan \beta$ limit. Finally, the bound on ε_L translates into the bound on reheat temperature $T_r \gtrsim 10^6 \text{ GeV}$ for the observed baryon-to-photon ratio $n_B/n_\gamma = (6.10 \pm 0.04) \times 10^{-10}$ [29]. Thus, the reheat temperature is small enough to avoid the gravitino problem. We set $T_r = 10^6 \text{ GeV}$ in all numerical work and obtain $2.8 \times 10^9 \lesssim M_R^I/\text{GeV} \lesssim 4 \times 10^{13}$ and $6.3 \times 10^{-10} \lesssim y_\nu \lesssim 10^{-7}$, as depicted in the figure 7.

6 Gauge coupling unification in R-symmetric SU(5)

According to the no-go theorem mentioned in [18, 19], we obtain ‘massless’ fields in any R -symmetric grand unified theory (GUT) based on a simple gauge group after spontaneous breaking of the GUT symmetry. These fields, however, can acquire TeV scale masses from the soft SUSY breaking terms. In our R -symmetric SU(5) model we obtain light ($\sim \text{TeV}$) octet and triplet components from the 24_H Higgs field [16, 17]. The presence of these light fields, in turn, ruins the successful gauge coupling unification feature of MSSM. To circumvent

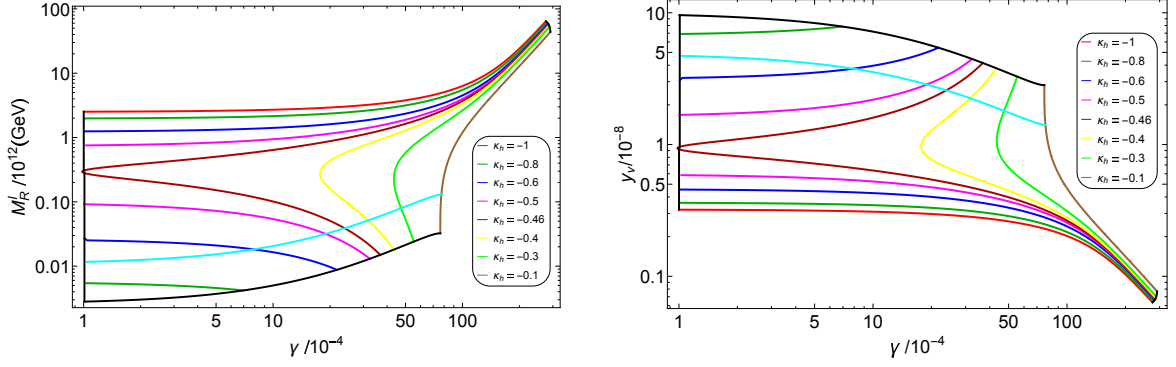


Figure 7. The mass of the (s)neutrino inflaton M_R^I (left panel) and the neutrino Yukawa coupling y_ν (right panel) versus the coupling γ . We set the scalar spectral index $n_s = 0.968$ (central value of Planck's data), the end of inflation $y_e = \frac{N_e}{M} = 1$ and the reheate temperature $T_r = 10^6$ GeV.

this problem we add copies of vectorlike families $5 + \bar{5} + 10 + \bar{10}$. This does not solve the problem of gauge coupling unification unless we allow mass splitting within their MSSM field components,

$$5 + \bar{5} = (D + \bar{D}, L + \bar{L}), \quad 10 + \bar{10} = (Q + \bar{Q}, U + \bar{U}, E + \bar{E}). \quad (6.1)$$

This splitting is achieved in a way similar to the doublet-triplet splitting but with far less fine tuning. With additional vectorlike families we obtain the following mass terms in the superpotential,

$$W \supset \frac{\lambda_{ij}^{(10, \bar{10})}}{m_P} \text{Tr}(24_h^2) \text{Tr}(10_i \bar{10}_j) + \frac{\tilde{\lambda}_{ij}^{(10, \bar{10})}}{m_P} \text{Tr}(10_i 24_h^2 \bar{10}_j) \quad (6.2)$$

$$+ \frac{\lambda_{ij}^{(5, \bar{5})}}{m_P} \text{Tr}(24_h^2) \text{Tr}(5_i \bar{5}_j) + \frac{\tilde{\lambda}_{ij}^{(5, \bar{5})}}{m_P} \text{Tr}(5_i 24_h^2 \bar{5}_j), \quad (6.3)$$

$$\supset M_Q Q \bar{Q} + M_U U \bar{U} + M_E E \bar{E} + M_D D \bar{D} + M_L L \bar{L}, \quad (6.4)$$

with Z_5 -charge, $q_5(5 \bar{5}, 10 \bar{10}) = (3, 3)$, and R -charge, $R(5 \bar{5}, 10 \bar{10}) = (1, 1)$. For simplicity, assuming $\lambda_{ij} = \delta_{ij} \lambda$ and $\tilde{\lambda}_{ij} = \delta_{ij} \tilde{\lambda}$, the masses of the MSSM field components of a vectorlike family are given by

$$M_E = \frac{30\lambda^{(10, \bar{10})} + 9\tilde{\lambda}^{(10, \bar{10})}}{30} \left(\frac{M^2}{m_P} \right), \quad (6.5)$$

$$M_Q = \frac{60\lambda^{(10, \bar{10})} + 13\tilde{\lambda}^{(10, \bar{10})}}{30} \left(\frac{M^2}{m_P} \right), \quad (6.6)$$

$$M_U = \frac{15\lambda^{(10, \bar{10})} + 2\tilde{\lambda}^{(10, \bar{10})}}{30} \left(\frac{M^2}{m_P} \right), \quad (6.7)$$

$$M_D = \frac{15\lambda^{(5, \bar{5})} + 2\tilde{\lambda}^{(5, \bar{5})}}{30} \left(\frac{M^2}{m_P} \right), \quad (6.8)$$

$$M_L = \frac{20\lambda^{(5, \bar{5})} + 6\tilde{\lambda}^{(5, \bar{5})}}{40} \left(\frac{M^2}{m_P} \right). \quad (6.9)$$

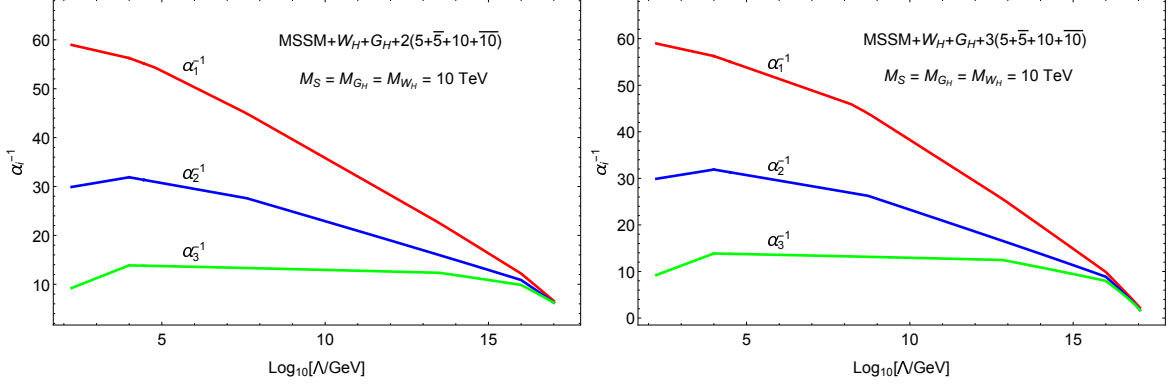


Figure 8. The evolution of the inverse gauge couplings versus the energy scale Λ in R -symmetric $SU(5)$ model, with two (left-panel) and three (right-panel) generations of vectorlike families ($5 + \bar{5} + 10 + \bar{10}$). The effective SUSY breaking scale is set at $M_S = 10$ TeV. The masses of vectorlike MSSM components are taken as $M_Q = M_U = 10^{16}$ GeV, $M_D = 10^{13.5}$ GeV, $M_L = 10^{7.6}$ GeV, $M_E = 10^{4.8}$ GeV (left-panel) and $M_Q = M_U = 10^{16}$ GeV, $M_D = 10^{12.845}$ GeV, $M_L = 10^{8.719}$ GeV, $M_E = 10^{8.2}$ GeV (right-panel). The GUT scale, $M_{\text{GUT}} = 10^{17}$ GeV, in both cases.

Now we can make a selected single field component to be light in both vectorlike multiplets of $SU(5)$. We choose light masses for $E + \bar{E}$, and $L + \bar{L}$ with $(30\lambda^{(10, \bar{10})} + 9\tilde{\lambda}^{(10, \bar{10})}) \sim 0$, and $(20\lambda^{(5, \bar{5})} + 6\tilde{\lambda}^{(5, \bar{5})}) \sim 0$. The other components can have masses as large as $M^2/m_P \sim 10^{16}$ GeV. A successful gauge coupling unification can be achieved with two or three generations of additional vectorlike families. This is shown in figure 8 with different mass splitting patterns described in its caption. As the triplet and the octet components of 24_H Higgs field attain masses of order $\langle S \rangle (M/m_P)^3$, we take their masses to be around the SUSY breaking scale M_S , which is fixed at $M_S = 10$ TeV in order to adequately suppress dimension five proton decay operator. The gauge coupling unification scale is set at $M_{\text{GUT}} \equiv (5/6)g_5 M = 10^{17}$ GeV, where g_5 is the unified gauge coupling of $SU(5)$.

7 Summary

We consider a pseudosmooth tribrid model of sneutrino inflation in an R -symmetric $SU(5) \times Z_5$ GUT model. With the help of an additional Z_5 symmetry and a non-minimal Kähler potential, a pseudosmooth trajectory is successfully generated to realize inflation while avoiding the monopole problem. The predicted values of the various inflationary parameters are calculated at the central value of the scalar spectral index, $n_s = 0.968$. The predictions for the tensor to scalar ratio, $2.7 \times 10^{-3} \lesssim r \lesssim 10^{-13}$, and for the running of the scalar spectral index, $-0.00031 \lesssim dn_s/d \ln k \lesssim 0.0024$, are in agreement with the latest Planck 2018 results. These ranges are obtained with $-1 \leq \kappa_h \leq -0.1$, $5.5 \times 10^{16} \lesssim M/\text{GeV} \lesssim 5.6 \times 10^{17}$, $5 \times 10^{16} \text{ GeV} \lesssim N_0 \lesssim m_P$, $|\gamma| \gtrsim 10^{-4}$, $|\delta| \lesssim 1$, $y_c = 1$ and $T_r = 10^6$ GeV. The gravitino problem is avoided with the realization of reheat temperature T_r as low as 10^6 GeV. A common problem of R -symmetric $SU(5)$ GUT is the appearance of light triplet and octet components from the GUT Higgs field, thus putting successful gauge coupling unification of MSSM in jeopardy. This problem is avoided with the help of additional vector-like families residing in complete multiplets, $(5 + \bar{5} + 10 + \bar{10})$, of $SU(5)$.

Before concluding we provide a few brief remarks related to proton decay and dark matter. Rapid proton decay from renormalizable superpotential couplings is not allowed in this SU(5) model thanks to the R -symmetry. Furthermore, with relatively large squark and slepton masses of order 10 TeV or so, and with $M_{\text{GUT}} \sim 10^{17}$ GeV, dimension five proton decay is adequately suppressed, and dimension six proton decay mediated by the superheavy gauge bosons is predicted to lie well beyond the scope of Hyper-Kamiokande [30, 31]. To realize the lightest supersymmetric particle (LSP) as a viable cold dark matter candidate, we need to separately invoke a Z_2 matter parity.

Acknowledgments

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