

An excess of excesses examined via dark matter radio emissions from galaxies

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Abstract. Cosmic-ray and gamma-ray observations have yielded several notable excesses that often lend themselves to explanation by various dark matter annihilation/decay models. In particular, the AMS-02 anti-proton and positron excesses have continued to grow more robust with the collection of more data. This is supplemented by gamma-ray excesses in the Galactic Centre and a high-energy break in the spectrum of electron/positron cosmic rays seen by DAMPE. In this work we carefully model the magnetic field environments of M31 and M33 and use this to estimate expected synchrotron emissions from electrons produced via dark matter annihilation. By comparing this to available radio data we review simplifying assumptions used previously for dark matter hunting in these environments and produce novel constraints that are capable of fully ruling out dark matter models proposed to accommodate all the aforementioned excesses barring that of DAMPE. However, we do show that significant constraints can be placed upon the DAMPE parameter space with M31 data. In addition to this we project SKA non-observation constraints for the Reticulum II and Triangulum II dwarf galaxies and find these have potential to rule out cosmic-ray and gamma-ray excess-producing models of dark matter, even when the most conservative assumptions are employed.

Keywords: dark matter theory, dark matter experiments, semi-analytic modeling

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1 Introduction

Measurements of cosmic-ray fluxes have recently been of some importance in the hunt for the nature of Dark Matter (DM), with many proposed models to account for excess fluxes observed by AMS-02 [1, 2], PAMELA [3, 4], DAMPE [5], HESS¹ [6] and Fermi [7, 8]. For an admittedly non-exhaustive list of related works on the topic of DM and cosmic-rays see [9–16] and references therein.

A recent work [17] has re-analysed the significance of the observed AMS-02 anti-proton excess, finding much agreement with earlier work on AMS-02 electron-positron fluxes in [18]. This accompanied by other recent work in [19] which uses a 60 GeV neutralino to explain Large Hadron Collider missing energy as well as AMS-02 anti-proton and Galactic Centre GeV gamma-ray excesses. The authors of [17] demonstrate the robustness of the excess above expected backgrounds to between 5 and 7 σ confidence intervals, with models that include a DM component in addition to the background being strongly preferred to background-only models. The best-fit region of WIMP parameter space provided by [17] also evades most existing indirect detection constraints and was shown has some overlap with the favoured region for the GeV galactic-centre gamma-ray excess [20] and the models used in [18, 19]. These recent works display a resurgence of models with WIMP masses in the 40 to 100 GeV range initially favoured for the Galactic Centre GeV gamma-ray excess that had been somewhat superseded by more mundane astrophysical explanations [21–25]. In addition to this, the DArk Matter Particle Explorer (DAMPE) satellite announced an excess in the electron/positron flux around 1.4 TeV [5]. It has been proposed that this could be accounted for by leptophilic DM models via the annihilation of WIMPs and subsequent decay of a leptophilic mediator [26–28]. A common element of all these models is the required presence of an over-dense

¹<https://www.mpi-hd.mpg.de/hfm/HESS/>.

clump of DM within a radius < 1 kpc in order for the DM models to also satisfy DM cosmological abundance constraints on the annihilation cross-section. For example: [26] requires, for a WIMP mass range 1.4 to 1.7 TeV, an NFW local clump of DM with mass $2 \times 10^5 M_\odot$ to $10^8 M_\odot$, positioned at $d \sim 0.1$ kpc from the solar system.

This resurgence of old DM explanations as well as new ones from [26] can benefit from a multi-frequency and multi-target approach. Where we use the near-universality of DM in cosmic structures to subject DM explanations to scrutiny in situations outside those that suggested the excess in the first place. In particular it has been argued in the literature that radio frequency observations can play a strong role in probing the nature of DM, especially in light of the upcoming Square Kilometre Array (SKA) [25, 29–33]. This stance has received some experimental backing with recent works such as [34, 35] (for a summary of DM hunting via radio in dwarf galaxies see [36]). In particular the faintness of diffuse radio emissions around many cosmic structures means they can provide very strong constraints on DM emissions [25, 34, 35].

In this work we will focus in particular on nearby galaxies as target environments. This is due to recent work done in [32, 33], which used radio-frequency data on M33 and M31 to produce constraints on viable DM models through synchrotron emissions produced by secondary electrons from DM annihilation. In particular claiming to be able rule out models designed to account for AMS-02 positron [18] and Galactic Centre GeV gamma-ray [37] excesses. We reconsider both the data used in the aforementioned works and additional data from other sources such as [38–43]. Unlike [32, 33] our modelling of DM emissions includes all diffusive effects and, by numerical means, is somewhat less approximate in its treatment of DM synchrotron emissions. We demonstrate that the impact of the results in [32, 33] is somewhat predicated on their assumption of a constant magnetic field and no diffusion. A radially varying magnetic field produces results that differ noticeably from [33], while the same occurs for light lepton annihilation channels in M33. We also derive novel constraints under more realistic assumptions in both M31 and M33. These, with the aid of larger data sets, show that radio DM constraints can rule out the entire parameter space for both the GeV gamma-ray [20], AMS-02 positron [18], and anti-proton excesses [17] without the dependence on assumptions of substructure boosting, halo density profile, or magnetic field geometry. In addition to this the DM parameter space for the DAMPE excess from [26] can be partially explored by M31 data. Notably, the constraints found here allow many annihilation channels to be probed to below the thermal relic cross-section [44] for WIMP masses up to 100–1000 GeV. In this regard we agree with, and advance upon previous work done on radio-frequency DM hunting in M31 in [45].

This work is structured as follows: sections 2 and 3 detail the formalism for calculating DM radio emissions, section 4 displays the properties used for all the DM halos studied, section 5 lists the radio data used here, sections 6, 7, 8 show our results, and conclusions are drawn in section 9.

2 Electrons from dark matter annihilation

The source function for the production of electrons and positrons (hereafter just referred to as electrons) by WIMP annihilation in a DM halo is taken as

$$Q_e(r, E) = \frac{1}{2} \langle \sigma V \rangle \sum_f \frac{dN_e^f}{dE} B_f \left(\frac{\rho_\chi(r)}{m_\chi} \right)^2, \quad (2.1)$$

where f denotes the annihilation channel, $\langle\sigma V\rangle$ is the velocity-averaged annihilation cross-section, $\frac{dN_i^f}{dE}$ is the production spectrum of electrons per annihilation [46, 47], B_f is the branching fraction of channel f , $\rho_\chi(r)$ is the DM density at distance r from the halo centre, and m_χ is the WIMP mass.

3 Synchrotron emission from dark matter annihilation

Since electrons can be produced, either as primary or secondary products, in the annihilation of DM particles and magnetic fields are ubiquitous in cosmic structure we can expect synchrotron radiation to result. The average power of the synchrotron radiation at observed frequency ν emitted by an electron with energy E in a magnetic field with amplitude B is given by [48, 49]

$$P_{\text{synch}}(\nu, E, r, z) = \int_0^\pi d\theta \frac{\sin\theta}{2} 2\pi\sqrt{3}r_e m_e c \nu_g F_{\text{synch}}\left(\frac{\kappa}{\sin\theta}\right), \quad (3.1)$$

where m_e is the electron mass, $\nu_g = \frac{eB}{2\pi m_e c}$ is the non-relativistic gyro-frequency, $r_e = \frac{e^2}{m_e c^2}$ is the classical electron radius, and the quantities κ and F_{synch} are defined as

$$\kappa = \frac{2\nu(1+z)}{3\nu_g \gamma^2} \left[1 + \left(\frac{\gamma\nu_p}{\nu(1+z)} \right)^2 \right]^{\frac{3}{2}}, \quad (3.2)$$

with the plasma frequency $\nu_p \propto \sqrt{n_e}$, γ as the electron Lorentz factor, and

$$F_{\text{synch}}(x) = x \int_x^\infty dy K_{5/3}(y) \approx 1.25 x^{\frac{1}{3}} e^{-x} (648 + x^2)^{\frac{1}{12}}. \quad (3.3)$$

The emissivity from a population of electrons and positrons with energy spectra $\frac{dn_{e^-}}{dE}$ and $\frac{dn_{e^+}}{dE}$ is then found via

$$j_{\text{synch}}(\nu, r, z) = \int_{m_e}^{m_\chi} dE \left(\frac{dn_{e^-}}{dE} + \frac{dn_{e^+}}{dE} \right) P_{\text{synch}}(\nu, E, r, z), \quad (3.4)$$

note that in this work $\frac{dn_{e^-}}{dE}$ is the equilibrium electron distribution from DM annihilation (see below). The flux density spectrum within a radius r is then written as

$$S_{\text{synch}}(\nu, z) = \int_0^r d^3r' \frac{j_{\text{synch}}(\nu, r', z)}{4\pi(D_L^2 + (r')^2)}, \quad (3.5)$$

where D_L is the luminosity distance to the target DM halo.

The equilibrium electron distribution is found as a stationary solution to the equation

$$\frac{\partial}{\partial t} \frac{dn_e}{dE} = \nabla \cdot \left(D(E, \mathbf{r}) \nabla \frac{dn_e}{dE} \right) + \frac{\partial}{\partial E} \left(b(E, \mathbf{r}) \frac{dn_e}{dE} \right) + Q_e(E, \mathbf{r}), \quad (3.6)$$

where $D(E, \mathbf{r})$ is the diffusion coefficient, $b(E, \mathbf{r})$ is the energy loss function, and $Q_e(E, \mathbf{r})$ is the electron source function from DM annihilation. In this case, we will work under the simplifying assumption that D and b lack a spatial dependence and thus we will include only

average values for magnetic field and gas densities. The solution, when diffusion is negligible, has the form [50]

$$\frac{dn_e}{dE} = \frac{1}{b(E)} \int_E^{m_\chi} dE' Q_e(r, E'). \quad (3.7)$$

When diffusion is not negligible, as in dwarf galaxies, a spherically symmetric solution can be found [31, 50–52]

$$\frac{dn_e}{dE}(r, E) = \frac{1}{b(E)} \int_E^{M_\chi} dE' G(r, E, E') Q(r, E'), \quad (3.8)$$

by means of a Green's function $G(r, E, E')$. This function is expressed as

$$\begin{aligned} G(r, E, E') = & \frac{1}{\sqrt{4\pi\Delta v}} \sum_{n=-\infty}^{\infty} (-1)^n \int_0^{r_h} dr' \frac{r'}{r_n} \\ & \times \left(\exp\left(-\frac{(r' - r_n)^2}{4\Delta v}\right) - \exp\left(-\frac{(r' + r_n)^2}{4\Delta v}\right) \right) \frac{Q(r')}{Q(r)}, \end{aligned} \quad (3.9)$$

where the sum runs over the indices of a set of image charges at positions given by $r_n = (-1)^n r + 2nr_h$. The radius r_h is the maximum radius we consider diffusion of particles out to. In this work we will use $r_h = 2R_{\text{vir}}$ with R_{vir} being the virial radius of the halo in question. Then Δv is defined

$$\Delta v = v(u(E)) - v(u(E')), \quad (3.10)$$

with

$$\begin{aligned} v(u(E)) &= \int_{u_{\min}}^{u(E)} dx D(x), \\ u(E) &= \int_E^{E_{\max}} \frac{dx}{b(x)}. \end{aligned} \quad (3.11)$$

In order to complete this solution we then need to define the diffusion and energy-loss functions. For $D(E)$ we follow [53] so that

$$D(E) = \frac{1}{3} cr_L(E) \frac{\overline{B}^2}{\int_{k_L}^{\infty} dk P(k)}, \quad (3.12)$$

where \overline{B} is the average magnetic field, r_L is the Larmour radius of a relativistic particle with energy E and charge e and $k_L = \frac{1}{r_L}$. This combined with the requirement that

$$\int_{k_0}^{\infty} dk P(k) = \overline{B}^2, \quad (3.13)$$

where $k_0 = \frac{1}{d_0}$, with d_0 being the smallest scale on which the magnetic field is homogeneous, yields the final form

$$D(E) = D_0 \left(\frac{d_0}{1 \text{ kpc}} \right)^{\frac{2}{3}} \left(\frac{\overline{B}}{1 \mu\text{G}} \right)^{-\frac{1}{3}} \left(\frac{E}{1 \text{ GeV}} \right)^{\frac{1}{3}}, \quad (3.14)$$

where $D_0 = 3.1 \times 10^{28} \text{ cm}^2 \text{ s}^{-1}$.

Profile	ρ_s ($M_\odot \text{ pc}^{-3}$)	r_s (kpc)	M_{vir} ($10^{10} M_\odot$)	R_{vir} (kpc)
NFW	$1.10 \pm 0.18 \times 10^{-2}$	16.5 ± 1.5	104	207
Burkert	$3.68 \pm 0.40 \times 10^{-2}$	9.06 ± 0.53	79	189
Einasto ($\alpha = 0.17$)	$8.12 \pm 0.16 \times 10^{-6}$	17.44	113	213

Table 1. M31 Density Profile Properties from [54]. M_{vir} and R_{vir} are the virial mass and radius.

The energy loss function is defined by [31, 50]

$$b(E) = b_{\text{IC}} E^2 (1+z)^4 + b_{\text{sync}} E^2 \bar{B}^2 + b_{\text{Coul}} \bar{n} \left(1 + \frac{1}{75} \log \left(\frac{\gamma}{\bar{n}} \right) \right) + b_{\text{brem}} \bar{n} \left(\log \left(\frac{\gamma}{\bar{n}} \right) + 0.36 \right), \quad (3.15)$$

where \bar{n} is the average gas density in the halo and is given in cm^{-3} , \bar{B} is the average magnetic field in μG , E is the electron energy in GeV , while b_{IC} , b_{synch} , b_{col} , and b_{brem} are the inverse Compton, synchrotron, Coulomb and bremsstrahlung energy loss factors. These are given, in units of $10^{-16} \text{ GeV s}^{-1}$, by 6.08/0.25, 0.0254, 6.13, and 1.51 respectively. The two values of b_{IC} represent the cases of a galactic inter-stellar radiation field, and that of the CMB only respectively.

4 Dark matter halos

4.1 M31

For the M31 DM halo we will follow the studies made in [54] and make use of three different halo profiles: Navarro-Frenk-White (NFW) [55], Burkert [56], and Einasto [57] summarised in eq. (4.1).

$$\begin{aligned} \rho_{\text{nfw}}(r) &= \frac{\rho_s}{\left(\frac{r}{r_s} \left(1 + \frac{r}{r_s} \right) \right)^2}, \\ \rho_{\text{burk}}(r) &= \frac{\rho_s}{\left(1 + \frac{r}{r_s} \right) \left(1 + \left[\frac{r}{r_s} \right]^2 \right)}, \\ \rho_{\text{ein}}(r) &= \rho_s \exp \left[-\frac{2}{\alpha} \left(\left[\frac{r}{r_s} \right]^\alpha - 1 \right) \right]. \end{aligned} \quad (4.1)$$

Note that r_s and ρ_s have different physical roles in each of the profiles, however, we label them in this manner for convenience of presentation in table 1 which is sourced from the modelling of [54]. The distance to the halo centre is taken to be 770 kpc. For the magnetic field and the gas distributions we will follow the model of [58] and take it so that our magnetic field strength is $B = 4.6 \pm 1.2 \mu\text{G}$ at $r = 14 \text{ kpc}$, with a more general profile following

$$B(r) = \frac{4.6r_1 + 64}{r_1 + r} \mu\text{G}, \quad (4.2)$$

where $r_1 = 200 \text{ kpc}$ is taken to follow the more conservative value from fitting in [58]. We take the gas density to be given by an exponential profile

$$n_e(r) = n_0 \exp \left(-\frac{r}{r_d} \right), \quad (4.3)$$

where $n_0 = 0.06 \text{ cm}^{-3}$ is the central density [59], and $r_d \approx 5 \text{ kpc}$ is the disk scale radius fitted by [58]. As argued in [33] spatial diffusion of secondary electrons can be neglected in this magnetic field environment, as its time-scale is far longer than that for energy losses (this was verified by calculation for all presented cases).

4.2 M33

For M33 we will implement both Burkert and NFW density profiles. In the case of NFW we will use $M_{\text{vir}} = 5.4 \pm 0.6 \times 10^{11} M_{\odot}$, and $r_s = 15.3 \pm 0.56 \text{ kpc}$ to fully specify the density profile [60]. For a Burkert form we will use $M_{\text{vir}} = 3 \pm 0.8 \times 10^{11} M_{\odot}$ and $r_s = 9.6 \pm 0.5 \text{ kpc}$ [60]. The distance to the halo centre is taken to be 840 kpc . The magnetic field radial profile will be assumed to follow an exponential form (as this is common for spiral galaxies [61]) and we will take B_0 such that the average field within 7.5 kpc is $8.1 \pm 0.5 \mu\text{G}$ following [32] based on the results from [62]. The scale-length of the magnetic field will be taken to be $\approx 5 \text{ kpc}$ following arguments in [61] that is $\approx 3.8 r_d$ where r_d is the scale-length of the baryonic matter distribution. We will consider diffusion with a minimal coherence length of 50 pc for the turbulent magnetic field to match spiral galaxies of this size [61]. For the purposes of energy-loss and diffusion we will consider the average magnetic field to be given by $8.1 \mu\text{G}$ to produce more conservative results. The gas central density is taken as 0.03 cm^{-3} [59] and is taken to follow an exponential radial profile, with scale-radius $r_d = 1.2 \pm 0.2 \text{ kpc}$ following [63].

4.3 Dwarf spheroidals

For Reticulum II and Triangulum II dwarf galaxies we will assume a Burkert density profile (as this is favoured for dwarf galaxies [64, 65]) normalised to a kinematically determined J-factor $J = 10^{19.6} \text{ GeV}^2 \text{ cm}^{-5}$ for 0.5 degrees [66] for Reticulum II and for Triangulum II we take $J = 10^{21.03} \text{ GeV}^2 \text{ cm}^{-5}$ [67]. J is given by

$$J(\Delta\Omega, l) = \int_{\Delta\Omega} \int_l \rho^2(\mathbf{r}) dl' d\Omega', \quad (4.4)$$

with the integral being extended over the line of sight l , and $\Delta\Omega$ is the observed solid angle. The magnetic field and electron densities will both be assumed to follow exponential distributions with a scale radius $r_d = 15 \text{ pc}$ in Reticulum II and $r_d = 35.68 \text{ pc}$ in Triangulum II. These scales being chosen as they are the stellar half-light radii of these targets [68–70]. We assume the magnetic fields in both dwarf galaxies have Kolmogorov turbulence spectra with a minimum coherence length of 1 pc . The central values for magnetic field and electron density will be assumed, in both cases, to be $B_0 \approx 1 \mu\text{G}$ and $n_0 \approx 10^{-6} \text{ cm}^{-3}$. Note that these are chosen to be slightly conservative following arguments made in [34, 35].

4.4 The effect of substructure

The expected presence of sub-halos within a DM halo can produce a substantial increase in fluxes resulting from DM annihilation [25, 71]. A common approach in the literature is to calculate a “boosting factor” by which the presence of sub-halos will increase the DM annihilation flux. In order to compare our results directly to [32, 33] we will employ the same boosting factors following the calculation method in [72]. These are, for M31 and M33 respectively, 5.28 and 4.86 . We will employ no boosting from substructure in the case of dwarf galaxies as this is not expected to be significant due to their low mass.

ν (MHz)	$S(\nu)$ (Jy)	$\Delta S(\nu)$ (Jy)	ROI	Ref.
408.0	0.22	0.02	5–50 arcmin	[38]
1.4e3	8.6	Not given	≈ 15 arcmin	[40]
4.8e3	1.863	Not given	40 kpc	[42]
3.0e6	2928.40	439.26	≈ 50 arcmin	[39]
5e6	536.18	80.427	≈ 50 arcmin	[39]
1.2e7	107.71	16.1565	≈ 50 arcmin	[39]
1.38e8	268.0	4.25	≈ 50 arcmin	[41]
1.82e8	3.14e2	4.96	≈ 50 arcmin	[41]
2.4e8	231.0	3.44	≈ 50 arcmin	[41]

Table 2. M31 flux data points used. See text for details.

5 Flux data sets

This work makes use of several sets of flux data. In the case of M31 we use the integrated flux found by [33], between 4.6 and 5 GHz with Region of Interest (ROI) covering a 40 kpc radius, as well as lower frequency fluxes integrated over a map of the target taken from the Nasa Extra-galactic Database (NED)² and sourced from [38–42] which will be compared to DM radio spectra integrated out to an ROI radius chosen to meet each case. We note that the low frequency 408 MHz point [38] has a significant impact on the results. However, it does not have a quoted ROI but it does have a maximum observable angular scale of at least 70 arcminutes according to the configuration quoted in [38]. However, due to the uncertainty we will use three cases for this data point: a pessimistic one with a 5 arcminute ROI, a median case of 20 arcminutes (4.5 kpc), and an optimistic one with a 50 arcminute (11 kpc) ROI. In the case of [39, 41] the observations were done with the MIPS/IRAC instruments on Spitzer with $\gtrsim 50$ [39]/51 [41] arcminute angular radii, so we use a 50 arcminute ROI in these cases. The observations for [40] come from the VLA in mode D, with a largest recoverable scale of 16 arcminutes, therefore we use 15 arcminutes as the ROI in this case. Additionally, [42] was conducted with the Effelsberg telescope so we need not consider the largest angular scale, so we use the 40 kpc ROI from [33] due to the similar frequency range and lack of additional sources at this frequency within the ROI [33]. We summarise the M31 data in table 2.

For M33 we employ two data sets. One using just upper limits from [32] integrated 7.5 kpc from the centre M33. The other uses the data from [32] as well as results from [43]. Note that, when comparing M33 DM spectra to [43] we only integrate over a 1.5 arcminute area around the halo centre (this small area will amplify the need to consider diffusive effects for this data in M33).

In the case of dwarf galaxies we will find non-observation constraints by using the sensitivity profile of SKA-1 [73] at the 2σ confidence interval with 50 hours of observation time.

6 Constraints from M31

In each case we will find the smallest value of $\langle\sigma V\rangle$ that is excluded by the data at a confidence level of 2σ . As a reference case we will also plot the constraints on the $b\bar{b}$ and $\tau^+\tau^-$ channels from Fermi-LAT studies of dwarf galaxies [74, 75].

²<http://ned.ipac.caltech.edu/>.

Note that we use the data from [33] with an ROI of 40 kpc. To compare to this we consider two cases, one where all of the data [38–42] is considered to have an ROI of 40 kpc, and the other where [38] has an ROI of 5 arcminutes while the others are unchanged.

In figures 1, 2, and 3 we display the constraints produced by comparing predicted fluxes to aforementioned radio data for NFW, Burkert, and Einasto halos respectively. These figures use a 50 arcminute (or ≈ 11 kpc) ROI at 408 MHz. The left panel of each figure displays a case with no halo boosting, comparing the quark channel in this case with the Fermi-LAT dwarf spheroidal limits we can see that there is almost an order of magnitude advantage over Fermi-LAT in the NFW case (solid lines showing the limits from section 5 including more data than just [33]) with Burkert and Einasto profiles respectively reducing and increasing the advantage over Fermi-LAT by a factor of 2 compared to NFW. This is significant as limits derived here for the differential case are similar to those reported for the integrated flux between 4.6 and 5 GHz in [33]. However, the integrated flux results found in this work (dot-dashed lines) are significantly weaker than those reported by [33] making use of the same data, being around an order of magnitude weaker than those shown for Fermi-LAT (comparing b-quark channels again). Significantly, the extension of our data set down to 408 MHz makes limits on low-mass WIMPs much stronger than when using the 4.6 to 5 GHz data only. In the right-hand column of these figures we use the same boosting factor as reported for M31 in [33] of 5.28, in this case the b-quark channel results for the radio data from [33] are only reaching parity with Fermi-LAT for WIMP masses > 100 GeV for an NFW density profile (contrary to the results in [33] where they exceed Fermi-LAT at all masses), the Burkert case remains roughly a factor of 2 weaker than Fermi-LAT for the same mass range. However, the Einasto case with boosting and [33] radio data surpasses Fermi-LAT for masses > 30 GeV.

These results suggest that the approximations used in [33] do not in fact provide an accurate conservative estimate for DM constraints, these being around an order of magnitude stronger than those found here for the same integrated flux data. The only major points of difference between this work and [33] is the use of differing magnetic field profiles and in our consideration of more energy-loss processes than just synchrotron emission (the latter likely causes around a factor of 3 reduction in DM radio flux [76]). In this work we allow the field to vary radially following [58] while [33] employs a constant magnetic field with $5 \mu\text{G}$ magnitude. This implies that the argument made in [32, 33], being that the likely higher central value of the magnetic field strength justifies the use of a flat profile, does not hold up in practice. As demonstrated by the fact that the central value for the magnetic field strength used here for M31 is $\gtrsim 10 \mu\text{G}$, fully compatible with the values argued for in [33]. Although [33] uses the halo profile from [77] we find that this produces a synchrotron flux smaller by a factor of ≈ 1.2 (with our magnetic field model), so it will not account for the differences found here.

While our above analysis has focussed on b-quark channels for the purpose of simple benchmarking against Fermi-LAT, we note that the other annihilation channels display very strong results as well. We see in figures 1, 2, and 3 that the differential constraints (solid lines) are extremely strong, allowing the leptonic annihilation channels to be constrained below the expected thermal relic cross-section all the way up to ~ 100 GeV WIMP mass. However, other channels produce more striking results, with constraints below the thermal relic level out to ~ 1000 GeV WIMP masses regardless of choice of halo density profile or the use of substructure boosting factors.

In the case of a 20 arcminute ROI (or ≈ 4.5 kpc) for the data from [38] (displayed in figures 4, 5, and 6) the constraints are very similar to those found for largest 50 arcminute ROI, differing by around a factor of ~ 1.5 at all WIMP masses. At this choice of ROI the

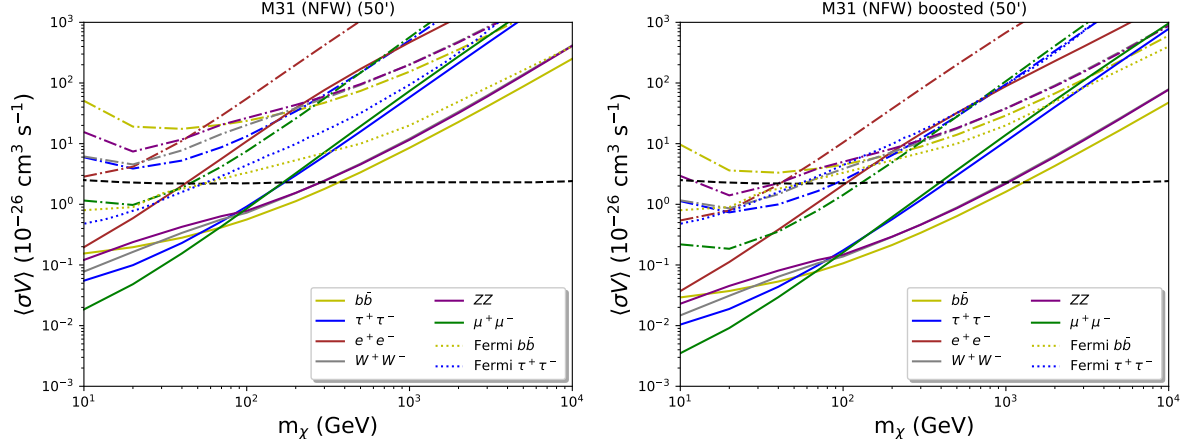


Figure 1. M31 cross-section upper limits at 2σ confidence level for NFW halos with a 50 arcminute ROI at 408 MHz. Left: unboosted. Right: boost factor 5.28. The black dashed line shows the thermal relic cross-section [44]. The dotted lines show Fermi-LAT dwarf galaxy limits from [75]. The solid lines show M31 results from this work for various annihilation channels using data points listed in section 5, while dash-dotted lines do the same for the integrated flux from [33].

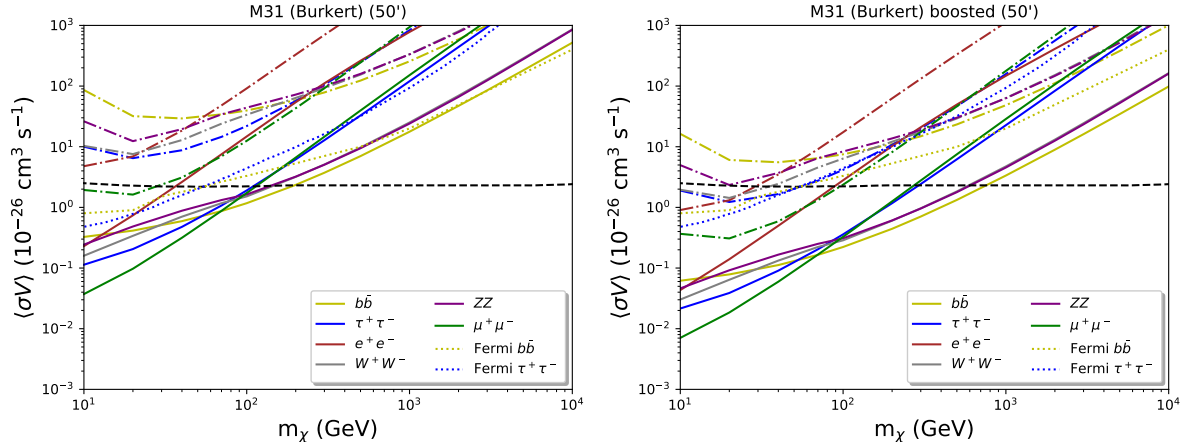


Figure 2. M31 cross-section upper limits at 2σ confidence level for Burkert halos with a 50 arcminute ROI at 408 MHz. Left: unboosted. Right: boost factor 5.28. The black dashed line shows the thermal relic cross-section [44]. The dotted lines show Fermi-LAT dwarf galaxy limits from [75]. The solid lines show M31 results from this work for various annihilation channels using data points listed in section 5, while dash-dotted lines do the same for the integrated flux from [33].

limits are still dominated by the data point from [38], again demonstrating the power of low frequency data to constrain DM radio emissions. This case displays NFW results that are very similar at all displayed WIMP masses to the joint constraints from an exponential magnetic field model and $B_0 = 50 \mu\text{G}$ used in [45].

These results should be contrasted with figures 7, 8, 9 which show the pessimistic ROI ($\sim 1 \text{ kpc}$) for the 408 MHz data point from [38]. In this case, we can only probe below the relic cross-section for NFW and Einasto halo profiles with $m_\chi \lesssim 30 \text{ GeV}$ or $\lesssim 100 \text{ GeV}$ with the use of the boosting factor. Both of these halo profiles, with boosts, produced results that are slightly stronger than Fermi-LAT in the compared channels. With or without the boosting factor, the NFW and Einasto halos are competitive with the bounds from [33]

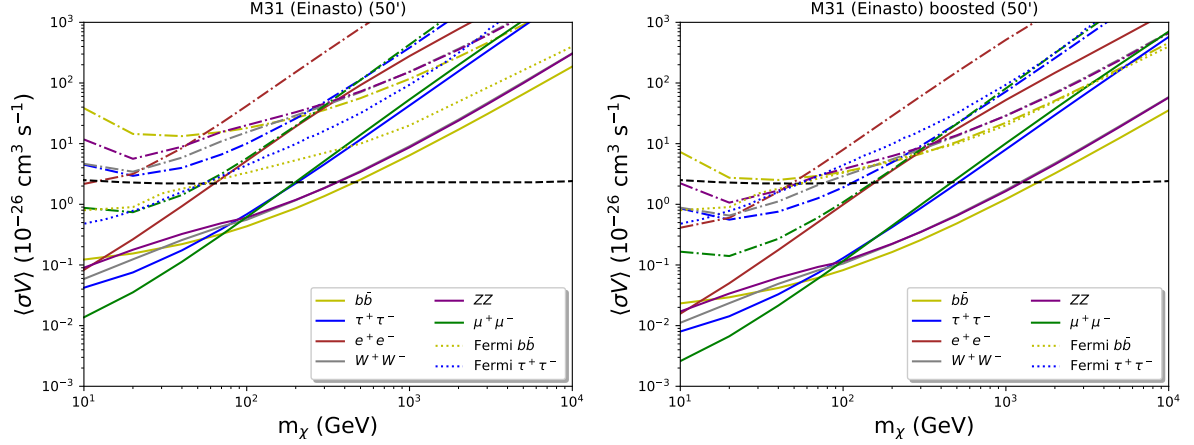


Figure 3. M31 cross-section upper limits at 2σ confidence level for Einasto halos with a 50 arcminute ROI at 408 MHz. Left: unboosted. Right: boost factor 5.28. The black dashed line shows the thermal relic cross-section [44]. The dotted lines show Fermi-LAT dwarf galaxy limits from [75]. The solid lines show M31 results from this work for various annihilation channels using data points listed in section 5, while dash-dotted lines do the same for the integrated flux from [33].

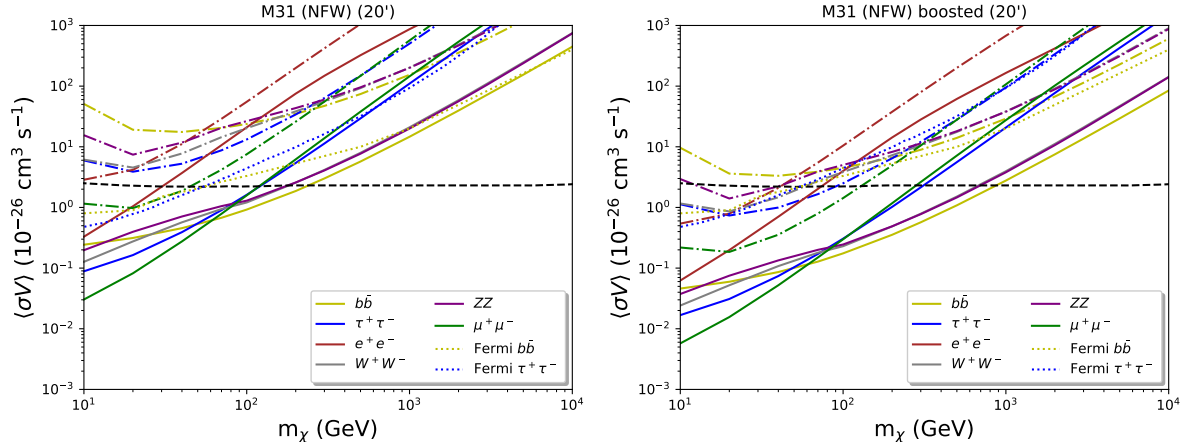


Figure 4. M31 cross-section upper limits at 2σ confidence level for NFW halos with a 20 arcminute ROI at 408 MHz. Left: unboosted. Right: boost factor 5.28. The black dashed line shows the thermal relic cross-section [44]. The dotted lines show Fermi-LAT dwarf galaxy limits from [75]. The solid lines show M31 results from this work for various annihilation channels using data points listed in section 5, while dash-dotted lines do the same for the integrated flux from [33].

data, and exceed them at low masses. In the Burkert case, the results are slightly weaker across the studied mass range than those obtained from [33] data. The NFW results for this pessimistic case agree quite closely with the joint constraints presented in [45] for the most pessimistic magnetic field estimate both the b quark and τ lepton channels (other channels are not displayed similarly in [45]). We note that now the limits are dominated by the data points from [39, 41, 42] at all masses and in all channels for the Burkert halo profile. In contrast, this is only true at masses above 70 GeV for NFW and Einasto halos in the muon and electron channels, as well as the same halo profiles in τ for $m_\chi > 200$ GeV. However, in the case of b quarks, Z , and W channels the data from [38] still provides the dominant limit across the entire studied mass range in the NFW and Einasto cases.

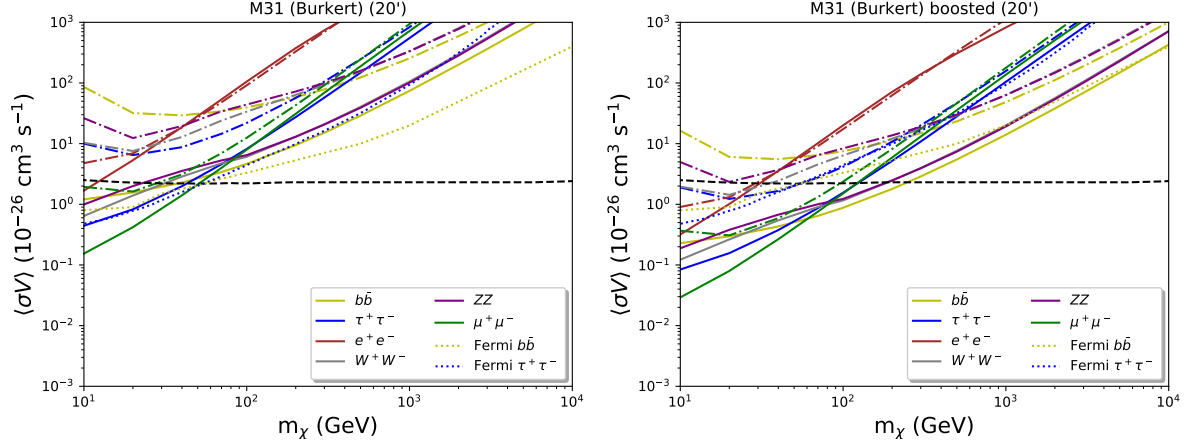


Figure 5. M31 cross-section upper limits at 2σ confidence level for Burkert halos with a 20 arcminute ROI at 408 MHz. Left: unboosted. Right: boost factor 5.28. The black dashed line shows the thermal relic cross-section [44]. The dotted lines show Fermi-LAT dwarf galaxy limits from [75]. The solid lines show M31 results from this work for various annihilation channels using data points listed in section 5, while dash-dotted lines do the same for the integrated flux from [33].

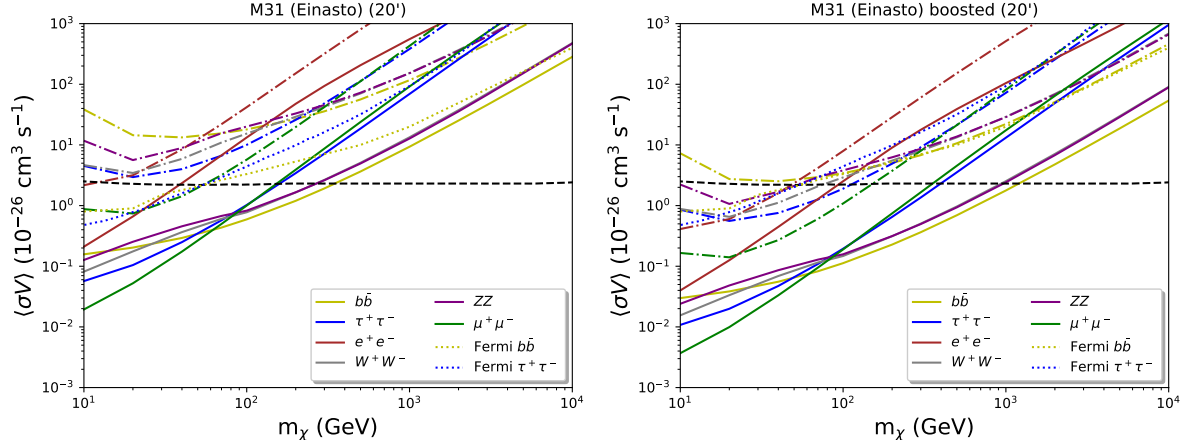


Figure 6. M31 cross-section upper limits at 2σ confidence level for Einasto halos with a 20 arcminute ROI at 408 MHz. Left: unboosted. Right: boost factor 5.28. The black dashed line shows the thermal relic cross-section [44]. The dotted lines show Fermi-LAT dwarf galaxy limits from [75]. The solid lines show M31 results from this work for various annihilation channels using data points listed in section 5, while dash-dotted lines do the same for the integrated flux from [33].

Importantly, we note that the ROI's used at 408 MHz are all substantially smaller than that used in the data of [33]. For the other data points from section 5 we used a 40 kpc ROI matching [33]. We justify this as they are taken at a similar frequency range and [33] notes a lack of other sources within this neighbourhood of the centre of M31 in their maps. In addition to this these fluxes are substantially larger, and so the impact of radial choice on our conclusions is very weak. We note that even had we used a 40 kpc ROI for the 408 MHz data point this would imply a -0.4 power-law index between the diffuse radio fluxes from [33] and [38]. The fact that we use substantially smaller radii is indicative then of the robustness of our results. We note that even in the case of a 40 kpc ROI our magnetic field is not necessarily unreliable, as the data used in the modelling of [58] extends out to 38 kpc from the centre of M31.

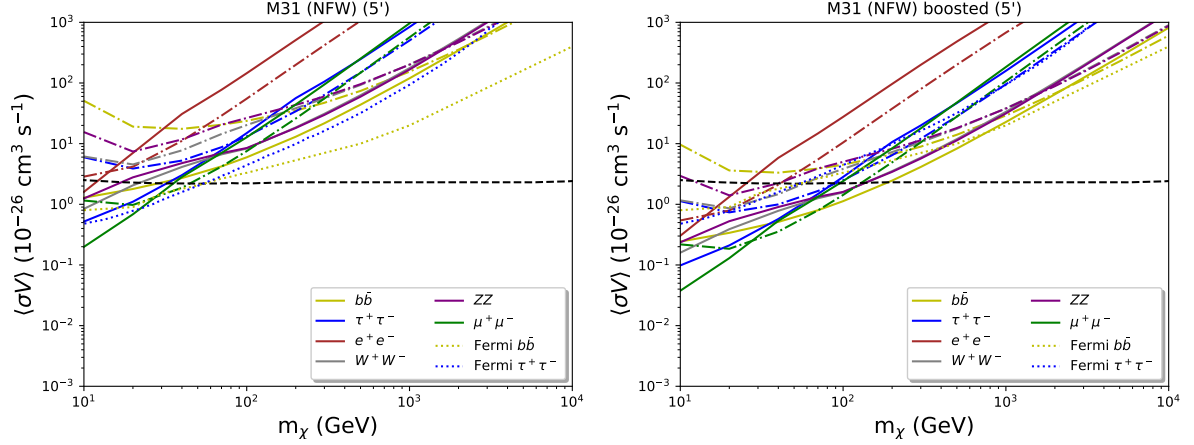


Figure 7. M31 cross-section upper limits at 2σ confidence level for NFW halos with a 5 arcminute ROI at 408 MHz. Left: unboosted. Right: boost factor 5.28. The black dashed line shows the thermal relic cross-section [44]. The dotted lines show Fermi-LAT dwarf galaxy limits from [75]. The solid lines show M31 results from this work for various annihilation channels using data points listed in section 5, while dash-dotted lines do the same for the integrated flux from [33].

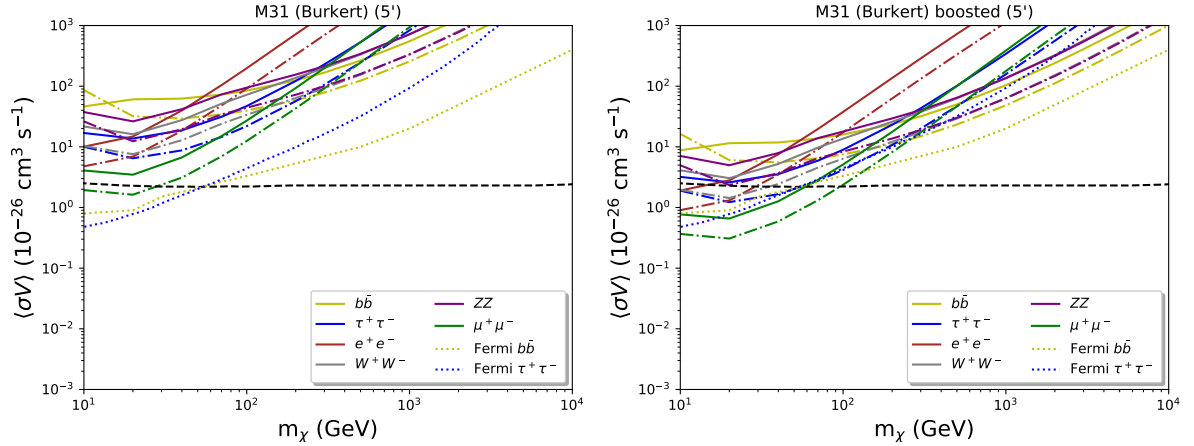


Figure 8. M31 cross-section upper limits at 2σ confidence level for Burkert halos with a 5 arcminute ROI at 408 MHz. Left: unboosted. Right: boost factor 5.28. The black dashed line shows the thermal relic cross-section [44]. The dotted lines show Fermi-LAT dwarf galaxy limits from [75]. The solid lines show M31 results from this work for various annihilation channels using data points listed in section 5, while dash-dotted lines do the same for the integrated flux from [33].

7 Constraints from M33

When studying M33 we use data points from [32] to compare to DM fluxes predicted within a radius of 7.5 kpc of the halo centre. The data point [43] is compared to an ROI of 1.5 arcmin only.

In figure 10 we display cross-section constraints found using an NFW density profile and the radio limits presented in [32] as well as [43]. Both boosted (left) and unboosted (right) cases are displayed, in the former case using a boost factor of 4.86 following [32]. Despite a strong central field value of $16.51 \mu\text{G}$, the [32]-only constraints for this case are only competitive with Fermi-LAT when a boosting factor is employed. However, when the

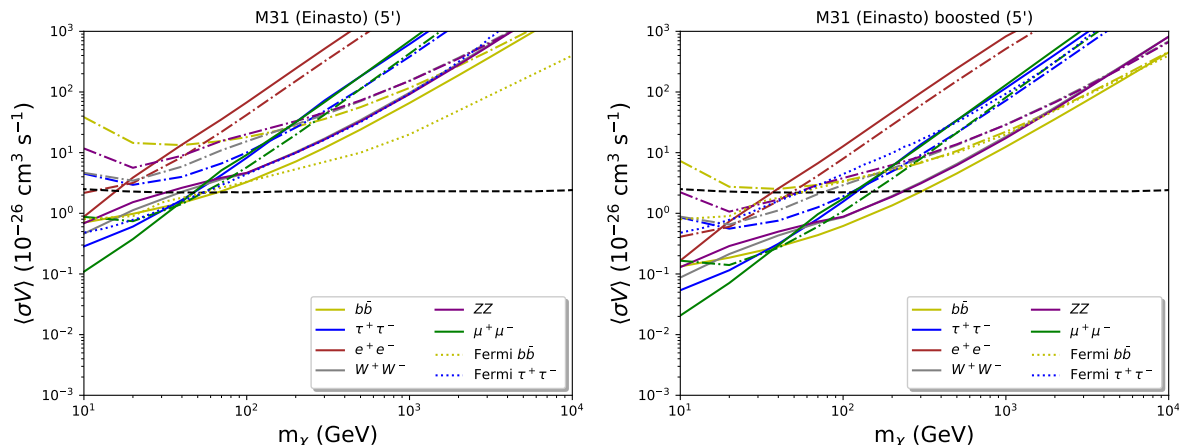


Figure 9. M31 cross-section upper limits at 2σ confidence level for Einasto halos with a 5 arcminute ROI at 408 MHz. Left: unboosted. Right: boost factor 5.28. The black dashed line shows the thermal relic cross-section [44]. The dotted lines show Fermi-LAT dwarf galaxy limits from [75]. The solid lines show M31 results from this work for various annihilation channels using data points listed in section 5, while dash-dotted lines do the same for the integrated flux from [33].

data from [43] is included the resulting limits are better than those from Fermi-LAT WIMP masses < 100 GeV by around an order of magnitude, becoming similar at larger masses. This also allows us to probe below the relic cross-section value up to around 100 GeV WIMP mass in all channels displayed. It is noteworthy that results using only data from [32] presented here agree quite closely those found in [32], being slightly stronger at low WIMP masses. However, the limits presented here are at least an order of magnitude weaker for the electron-positron annihilation channel. The difference between this work and [32] is our use of the radially varying field (with average values consistent with [32]), consideration of all energy-loss processes, and the inclusion of diffusion. This demonstrates that the seemingly reasonable assumptions made in [32], that the magnetic field is constant and diffusion is negligible, can result in an over-estimation of the predicted radio flux for some annihilation channels. However, the assumptions of [32] do seem to provide conservative limits for other channels, in agreement with ours for an exponential field and diffusion. It is also important to note that the limits depend very strongly on the assumed halo profile as the Burkert case in figure 11 is at least 1(2) orders of magnitude less constraining than for an NFW density profile with [32]([43]) data (the Burkert case being more sensitive to a steep field profile and diffusion). Interestingly, the Burkert limits are in agreement for both data sets used.

This has important consequences for the conclusions drawn in [32]. Notably that their claimed tension with AMS-02 positron excess models from [18] weakens due to the weakening of constraints for WIMP masses above 10 GeV with the electron annihilation channel. The muon channel can maintain a marginal tension only when a boost factor is assumed (tau is unchanged from [32]). Of course, these tensions disappear entirely when a Burkert density profile is used instead. However, it can be shown that the tension with models favoured by the Galactic Centre GeV excess [37] remains, although other models such as [20] are unaffected by the M33 data. The impactfulness of the M33 data can be recovered somewhat when using [43] data, so that the claimed tensions from [32] are once again in force for muon and quark channels but only for an NFW halo profile.

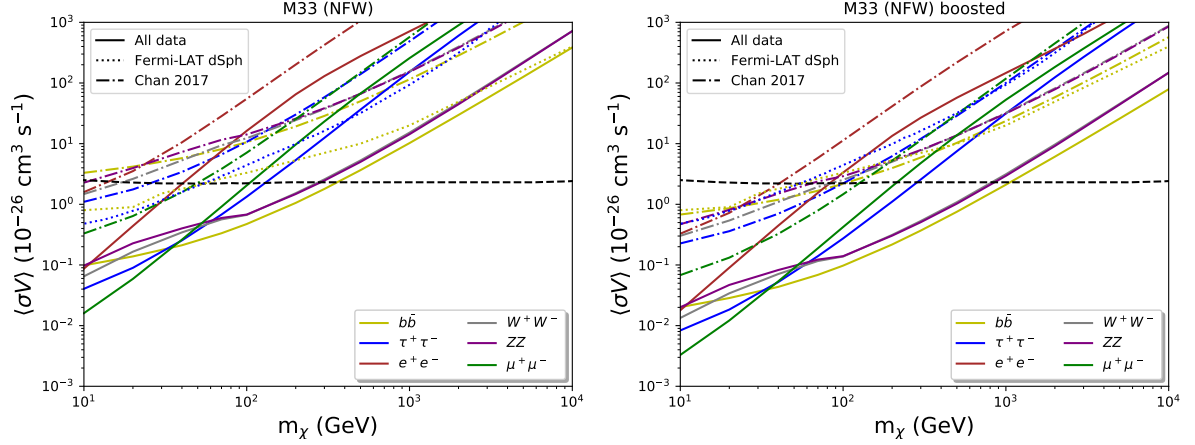


Figure 10. M33 cross-section upper limits at 2σ confidence level for NFW halos. Left: unboosted. Right: boost factor 4.86. The black dashed line shows the thermal relic cross-section [44]. The dotted lines show Fermi-LAT dwarf galaxy limits from [75]. The solid lines show M33 results from this work for various annihilation channels. The dot-dashed lines show our calculations using only the data points from [32].

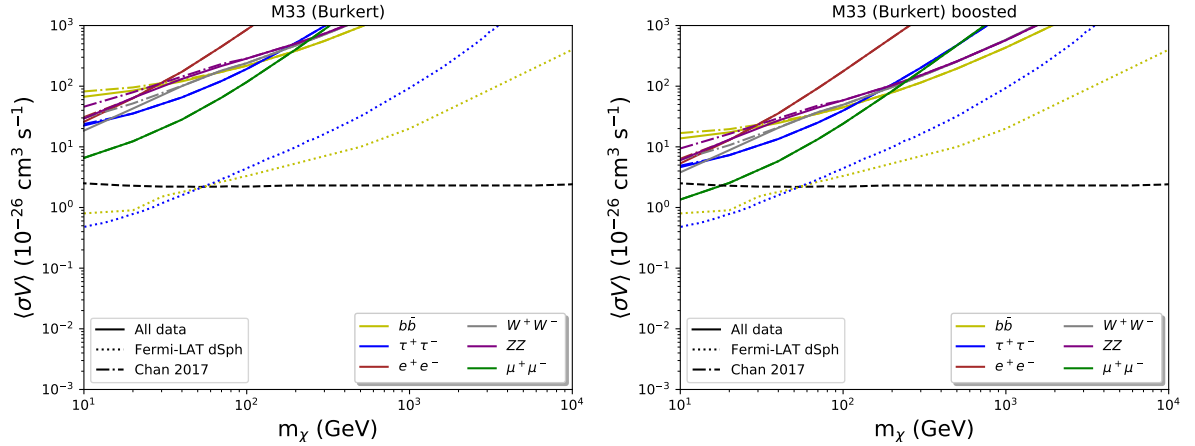


Figure 11. M33 cross-section upper limits at 2σ confidence level for Burkert halos. Left: unboosted. Right: boost factor 4.86. The black dashed line shows the thermal relic cross-section [44]. The dotted lines show Fermi-LAT dwarf galaxy limits from [75]. The solid lines show M33 results from this work for various annihilation channels. The dot-dashed lines show our calculations using only the data points from [32].

8 The excesses revisited

In this section we examine DM models proposed to explain the three excesses: AMS-02 [17], Galactic Centre GeV gamma-rays [20], and DAMPE [26].

In figure 12 we compare constraints from derived M31 to the favoured parameter space regions for both the GeV gamma-rays and the AMS-02 excesses. We use both the integrated flux for M31 from [33] (‘Chan 2019’ in the plots) and data points quoted in section 5 (‘All data’ in plots). Uncertainties for the cross-sections are also displayed using those from r_s , ρ_s , and B_0 . It is evident that for the integrated flux from [33] that only the Einasto case allows for potential constraint of the parameter space for AMS-02, however, this is very marginal.

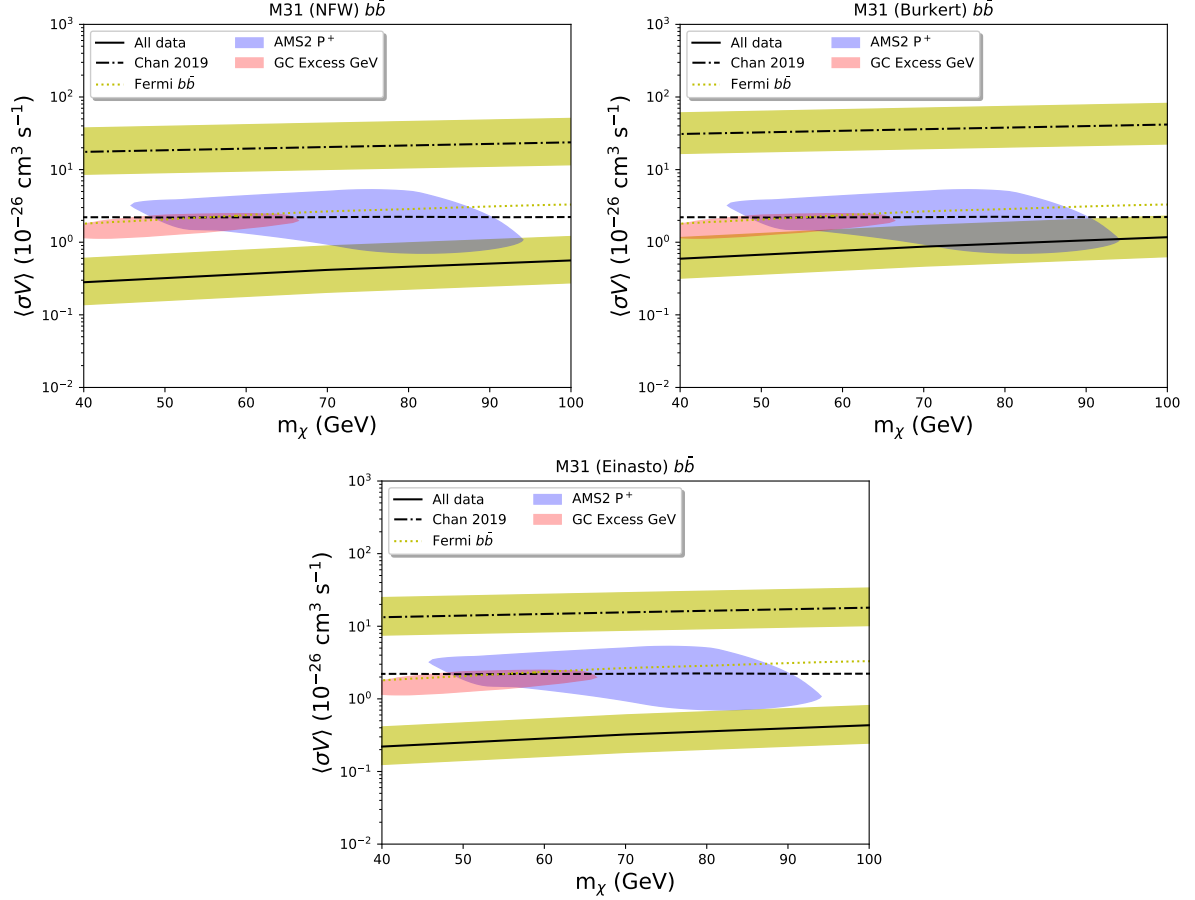


Figure 12. M31 cross-section upper limits at 2σ confidence level for NFW (upper left), Burkert (upper right), and Einasto (lower) halos. Solid lines show limits from the spectra listed in section 5 while dash-dotted lines are for the integrated flux from [33]. Shading around the lines displays uncertainties. The blue shaded region shows the parameter space for AMS-02 anti-proton excess DM models [17] and the red represents the Galactic Centre GeV gamma-ray excess models from [20]. The black dashed line shows the thermal relic cross-section [44].

In contrast, the data point from [38] at 408 MHz yields constraints that can rule out the entire parameter space of both excesses regardless of the choice of halo density profile (for the displayed WIMP mass range). This is significant as the observations made in [38] had an angular resolution at the level of a few arcminutes, and thus suited for probing large-scale diffuse emission in M31. In the pessimistic 5 arcminute ROI case for [38] we can rule out the parameter spaces only with the aid of a boosting factor in the NFW and Einasto halo profiles, the Burkert case provides no constraining power.

Figure 13 displays results analogous to figure 12 but for M33. For both NFW (left) and Burkert (right) density profiles, using only data from [32] (dash-dotted line) does not allow for the constraint of either of the parameter spaces of interest. However, the inclusion of data from [43] allows us to completely rule out all of the parameter space assigned by AMS-02 anti-proton and Galactic Centre gamma-ray excesses with an NFW density profile. However, the relevance of diffusion strongly affects the Burkert case.

In figure 14 we display SKA non-observation constraints at 2σ confidence level for two chosen dwarf spheroidal galaxies, these being Reticulum II and Triangulum II respectively.

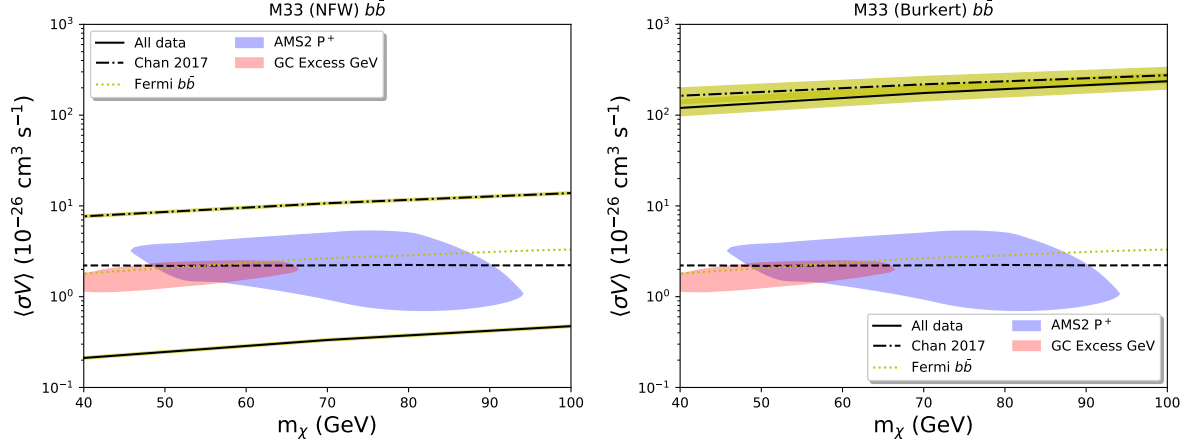


Figure 13. M33 cross-section upper limits at 2σ confidence level for NFW (left) and Burkert (right) halos. Solid lines show limits from the spectra listed in section 5 while dash-dotted lines are from [32] data only. Shading around the lines displays uncertainties. The blue shaded region shows the parameter space for AMS-02 anti-proton excess DM models [17] and the red represents the Galactic Centre GeV gamma-ray excess models from [20]. The black dashed line shows the thermal relic cross-section [44].

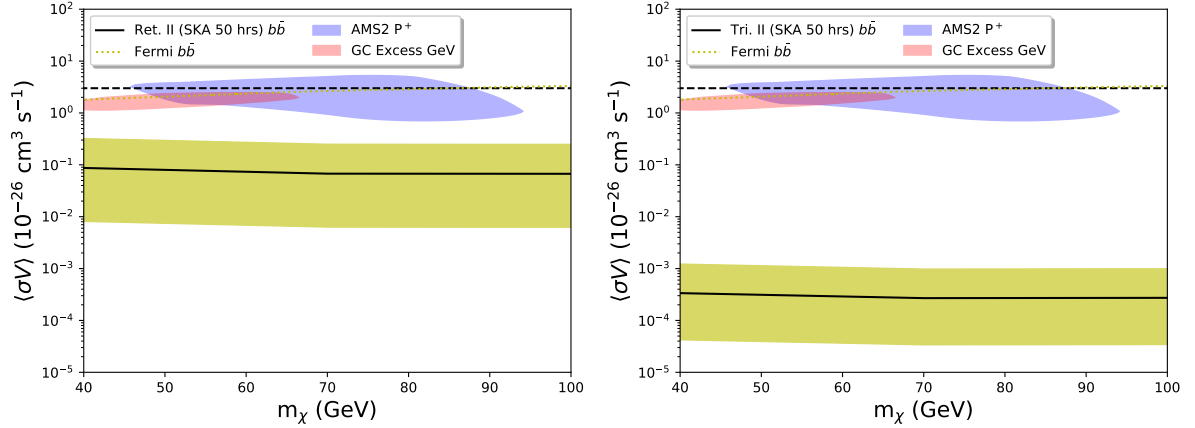


Figure 14. Dwarf galaxy non-observation cross-section upper limits at 2σ confidence level with 0.5 degree ROI. Left: Reticulum II. Right: Triangulum II. Shading around the upper-limit lines displays uncertainties. The blue shaded region shows the parameter space for AMS-02 anti-proton excess DM models [17] and the red represents the Galactic Centre GeV gamma-ray excess models from [20]. The black dashed line shows the thermal relic cross-section [44].

These are both capable of covering the entire parameter space of both excesses in the event of no radio signal observation by the SKA. Despite the significant J -factor and magnetic field uncertainties, these potential limits are at least 2σ away from the favoured parameter space regions for both excesses. This suggests that very robust constraints can be obtained in future from dwarf spheroidal targets, as is indeed indicated by early observational work in [34, 35].

In figure 15 we compare the parameter space region favoured by the DAMPE excess [26] with the inclusion of a local DM over-density producing the excess cosmic-rays. This indicates that, due to the large mass of the WIMP, the DAMPE parameter space is challenging to probe. Indeed, as argued in [26], it is untouched by Fermi-LAT dwarf galaxy limits and

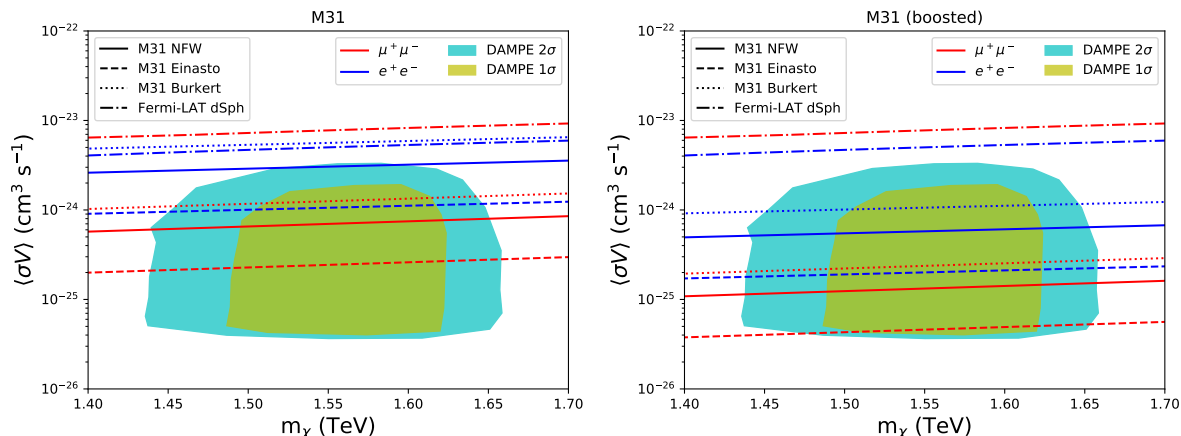


Figure 15. M31 cross-section upper limits at 2σ confidence level with 40 kpc ROI. The orange and cyan shaded regions show the 1σ and 2σ confidence interval best-fit models for the DAMPE excess from [26]. The dash-dotted line shows limits from [75]. The solid, dotted, and dashed lines show M31 results from this work for NFW, Burkert, and Einasto halos respectively. Note that the muon channel is shown in red and electrons in blue. Left: unboosted. Right: boost factor of 5.28 used.

is only strongly impacted by M31 data with NFW (muon channel only) and Einasto halo profile choices, with Burkert halos producing weaker constraints on the parameter space with the muon channel. This is largely due to the sensitivity of the frequency of the peak of DM synchrotron spectrum to WIMP mass [29]. In the right panel of figure 15 we can see that the principle uncertainty will be the boosting factor, as the use of the value from [33] results in strong constraints for NFW and Einasto halo profiles, with even the Burkert case making in-roads on the parameter space. However, in the pessimistic case where we assume [38] has a 5 arcminute ROI we can only minimally constrain the DAMPE parameter space. A similar exercise with the data points from [33] provide constraints around an order of magnitude weaker in each case, often meaning that no limits are put upon the parameter space without a boosting factor (which allows marginal limitation only).

Figure 16 displays results like those of figure 15 but for the case of M33. In the unboosted case even the data from [43] makes no impact on the parameter space and a boosting factor allows only marginal impact for an NFW halo and the muon channel.

9 Discussion and conclusions

This work has demonstrated, by studying M31 and M33, that the use of constant magnetic field profiles in calculating DM synchrotron emission should be approached with caution. Interestingly, despite a robust argument in [32] that diffusion should be insignificant in the M33 magnetic environment, we demonstrate that a steep radial profile for the magnetic field strength, and small ROI, make the inclusion of diffusive effects essential. Thus, diffusive effects can only be ignored when both the time-scale argument holds and when the magnetic field profile is sufficiently shallow. Interestingly, our results using data from [32], for all but electron annihilation channel agree well with those found in [32] despite the large difference in magnetic and diffusive environmental assumptions. The findings presented here reduce the impact of the limits on the DM annihilation cross-section from [32, 33] substantially in M31 as well as for electron channel in M33.

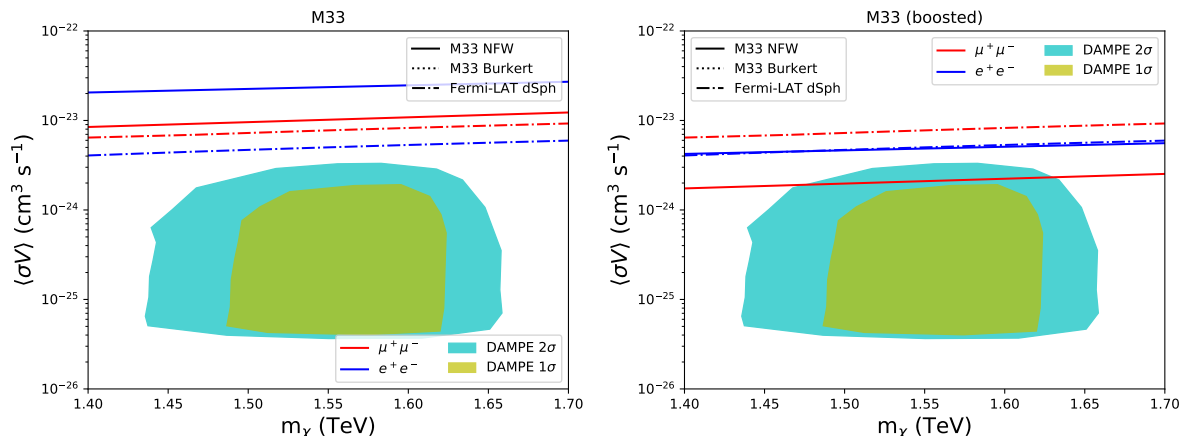


Figure 16. M33 cross-section upper limits at 2σ confidence level. The orange and cyan shaded regions show the 1σ and 2σ confidence interval best-fit models for the DAMPE excess from [26]. The dash-dotted line shows limits from [75]. The solid and dotted lines show M33 results from this work for NFW and Burkert halos respectively. Note that the muon channel is shown in red and electrons in blue. Left: unboosted. Right: boost factor of 4.86 used.

However, by expanding the radio data set for M31 and M33, the impact of resulting annihilation limits can be restored largely without recourse to uncertain substructure boosting factors. With the results presented here for M31 from the data enumerated in section 5 being very similar to those reported in [33], with the added advantage of extending down to lower WIMP masses due to the extent of the frequency range covered. In particular, the results presented here for M31 and M33 are shown to be capable of ruling out DM models proposed to account for the Galactic Centre gamma-ray GeV excess and the anti-proton and positron excesses observed by AMS-02 [17, 18] (in our pessimistic M31 scenario a boosting factor ≈ 5 is needed for this). This remains true regardless of the choice of halo density profile in M31, but only for an NFW halo in M33. We note that, for M31, this constitutes a conservative estimate as we use only the extrapolation of magnetic field strength from [58] without including an exponential component at the centre of M31. This would be conservative as the flux from DM-based emissions is strongly concentrated in the central regions of the halo by the ρ^2 dependence, making the central field value very influential. Additionally, as noted in [58], M31 is likely to have both a central exponential component and the outer flatter one modelled by [58]. Despite this, our results agree quite closely with the joint constraints presented for M31 in [45] (who use an exponential field profile). In particular, the optimistic M31 case matches closely (for an NFW halo profile) with more optimistic central magnetic field strength $B_0 = 50 \mu\text{G}$ results, while the pessimistic case agrees closely with the most pessimistic field estimates from [45]. We note particularly that our results display the power of low frequency data in constraining DM radio emissions, strongly in agreement with [45] where the 325 and 74 MHz data provide the strongest constraints over most of the studied WIMP mass range (particularly for the b quark channel).

Additionally, we show that the dwarf galaxies Reticulum II and Triangulum II both have the potential to rule out the favoured DM parameter spaces for the Galactic Centre GeV gamma-ray excess as well as for both the AMS-02 positron and anti-proton cases. This despite the inclusion of very significant uncertainties in halo J -factors, assuming only 50 hours of observation time, the use of a Burkert density profile, and conservative magnetic field assumptions.

Finally, we employed the M31 and M33 limits to examine the DM models proposed to account for the cosmic-ray excess seen by DAMPE. We find that the favoured parameter space from [26] cannot be strongly probed via M31 or M33 data without the assumption of a boosting factor. As we are required to reach the thermal relic annihilation cross-section for 1 TeV WIMPs in light lepton annihilation channels and DM radio spectra peak at a frequency sensitive to the WIMP mass. However, for NFW and Einasto density profiles in M31 some impact on the DAMPE DM parameter space is possible, particularly when the WIMP annihilates via muons. When a boosting factor is assumed a large majority of the parameter space can be explored for NFW as well as an Einasto cases for M31. The use of a Burkert density profile produces weaker limits even when a boosting factor is used. Notably, M33 does not strongly impact the parameter space for the DAMPE excess regardless of halo geometry or boosting. This DAMPE DM parameter space can also potentially be explored by the SKA hunting for radio emissions from the nearby DM clump, necessary to the DAMPE excess model from [26], as argued in [78].

Acknowledgments

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